THE LIGO GRAVITATIONAL WAVE OBSERVATORIES: RECENT RESULTS AND FUTURE PLANS

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The LIGO interferometers are operating as gravitational wave observatories, with a noise level near an order of magnitude of the goal and the first scientific data recently taken. This data has been analyzed for four different categories of gravitational wave sources; millisecond bursts, inspiralling binary neutron stars, periodic waves from a known pulsar, and stochastic background. Research and development is also underway for the next generation LIGO detector, Advanced LIGO.
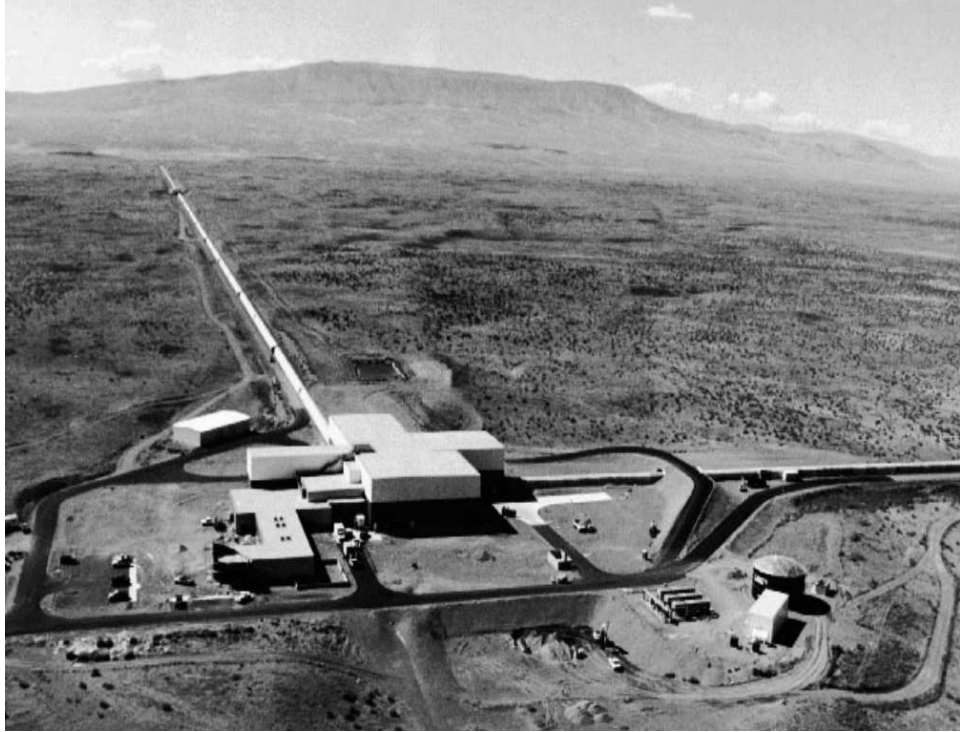
Figure 1. The LIGO Hanford Observatory

1 Introduction

The LIGO (Laser Interferometer Gravitational Wave Observatory) laboratory has constructed gravitational wave observatories at two sites in the continental United States; one near Richland, Washington on the Department of Energy Hanford Reservation (shown in Fig. 1), and the second near Baton Rouge, Louisiana in Livingston Parish (shown in Fig. 2). They are separated by 3000 km or about 10 ms of light travel time. Construction of the facilities began in 1994. These observatories consist of one (at Livingston) or two (at Hanford) laser interferometers designed to have high sensitivity to differential length changes in two perpendicular arms. The Livingston interferometer and one of the Hanford interferometers are 4 km in length, while the other Hanford interferometer is 2 km in length. The two observatories are as close to aligned as the curvature of the Earth will allow.

The tensor gravitational waves emitted by astronomical sources will cause a change in differential length between two perpendicular directions[1]. This strain, however, is expected to be extremely small, on the order of 10^{-21} in the acoustic frequency band that LIGO is sensitive to. The LIGO interferometers have been carefully designed so that they may be sensitive to small differential displacements. The interferometer measures the position of masses at the ends of each arm, which

Figure 2. The LIGO Livingston Observatory

are held fixed in the vertical direction, but are nearly free along the optical path of the laser. This is accomplished by hanging the masses as pendulums using wire slings. This freedom of motion allows the masses to respond to the changing space-time created by the wave without interference from terrestrial forces. The arms of each detector are kilometers long, so that the strain of the wave will create a relatively large change in displacement. This displacement is read out using a Michelson interferometer, so each mass is a mirror which allows the 1 micron laser light to reflect off of it. The gravitational wave signal increases with the interaction time with the light, so each arm includes a Fabry-Perot cavity made up of the end test mass and input test mass near the beamsplitter. The output port of the interferometer is held at a dark fringe. To further increase the signal, additional power buildup is accomplished by placing another mirror at the bright port. The entire optical path of the laser, including the kilometer long arms, must be held in a vacuum of below 10^{-8} mbar. The optical layout is shown in Figure 3.

The sensitivity is ultimately limited by fundamental noise sources inherent in the design. At the highest frequencies, above about 200 Hz, the sensitivity is limited by shot noise from the laser. This is reduced by increasing the laser power, as the signal rises faster with power than the noise. At intermediate frequencies, between about 40 Hz and 200 Hz, noise from thermally activated motion of the suspensions
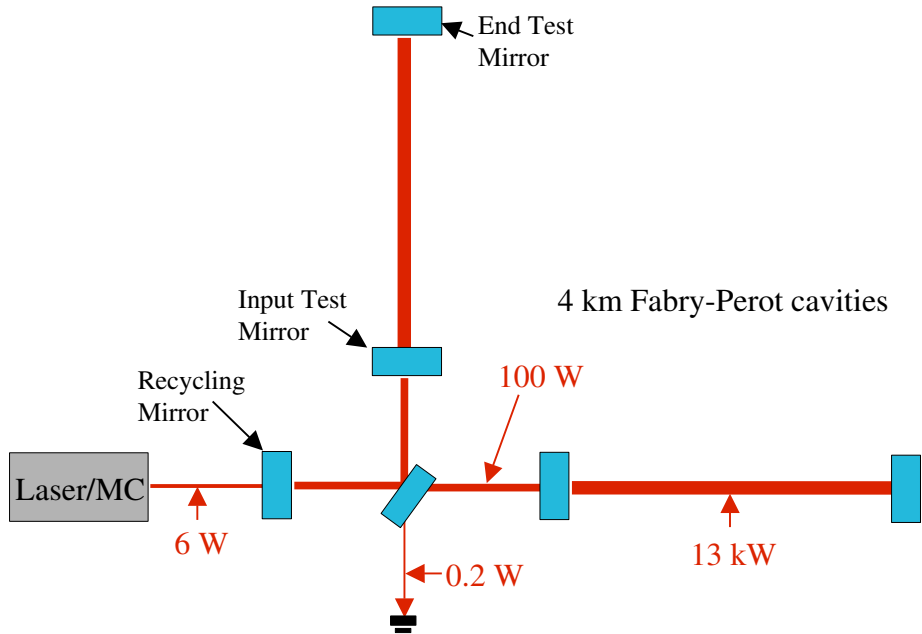
Figure 3. The optical layout of a LIGO interferometer showing the laser power at various points. A Michelson interferometer is supplemented with Fabry-Perot arm cavities 4 km long, and a power recycling mirror to reflect light back into the system as it leaves at the bright port.

dominate the noise budget. This noise is controlled by careful choice of materials and processes for the test masses, suspension wires, and nearby hardware. At low frequencies, below 40 Hz, seismic noise dominates. The coupling between the inherent motion of the ground and the mirror's motion is reduced by placing the test mass suspension on vibration isolation stacks. The sensitivity goal for the current 4 km long interferometers is shown in Figure 4. This goal reaches a noise of about 2×10^{-23} strain/$\sqrt{\mathrm{Hz}}$ at 150 Hz. More details about the LIGO interferometers is given below in Section 2. A more complete description of the LIGO detectors is available in Ref[2].

The first operation of a LIGO interferometer occurred in October of 2000. Since then, the interferometers have been in a commissioning phase, where work to reduce noise and improve performance has alternated with periods of running the interferometer to assess its status. A history of the noise in the Livingston interferometer is shown in Figure 5. These periods of operation, called engineering runs, were also used to collect data to test data handling, archiving, and the analysis software that was being developed in parallel with the instruments.

In the summer of 2002, the sensitivity had reached a level where data collected from a short period of operating could be expected to achieve limits on observable gravitational waves that would be comparable to or even better than previous experimental limits. This data collection run, referred to as the first LIGO Science
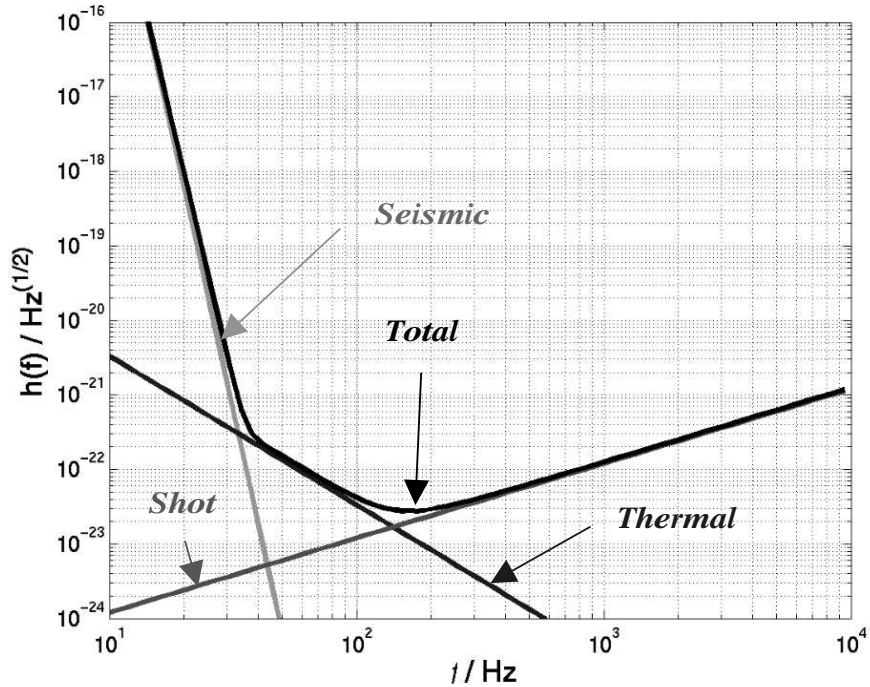
Figure 4. The sensitivity goal for the 4 km LIGO interferometers. Low frequency sensitivity is limited by seismic noise, the middle frequency band by thermal noise from both the suspension and test mass, and high frequencies by shot noise.

Run (S1), occurred between August 23 and September 9, 2002. GEO 600[3] also took data during this time period, to allow for coordinated data analysis. The data from all interferometers was analyzed for four different categories of gravitational waves; burst events, inspirals of binary neutron stars, continuous waves coming from a gravitational pulsar, and a stochastic background. Results of these analyses are given below in Section 3.

Simultaneous with commissioning and data taking, research and development has been going on throughout the LIGO Scientific Collaboration on technologies for the next generation interferometers. In 1998, a preliminary design was put forth for the Advanced LIGO interferometer, which incorporates technology advances made since the initial LIGO interferometers were constructed. Key among these technologies are a higher power laser, larger test masses made of single crystal sapphire, fused silica ribbon suspensions, and an additional mirror placed at the output port. Research for Advanced LIGO reached a mature enough state that a construction proposal was submitted to the US National Science Foundation in February of 2003. Details of Advanced LIGO are discussed below in Section 4.

Strain Sensitivity for the LLO 4km Interferometer
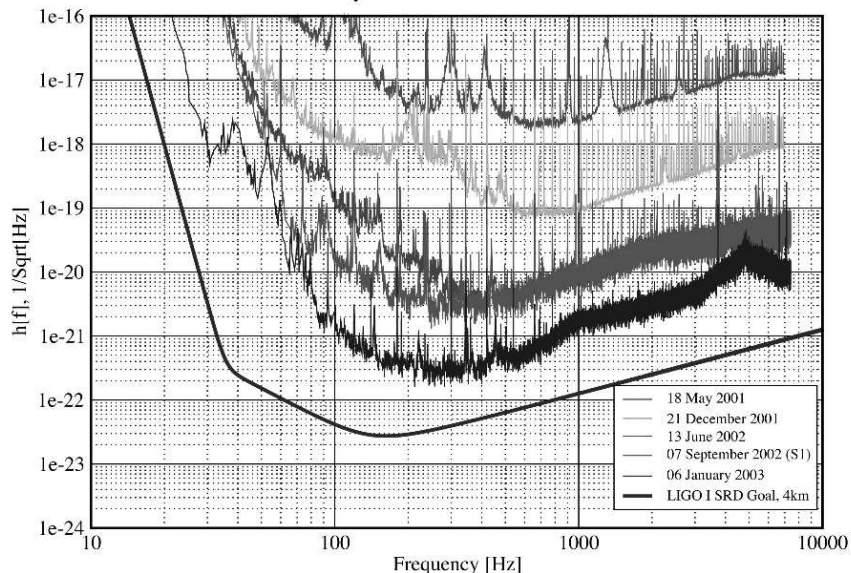
31 January 2003   LIGO-G030014-00-E

Figure 5. Noise in the LIGO Livingston interferometer through time.

2   The LIGO Interferometers

2.1   Laser and Input Optics

Each interferometer is fitted with a continuous-wave, Nd:YAG laser operating in the TEM_{00} Gaussian spatial mode at 1064 nm, and capable of delivering 10 W of power[4]. A small amount of the light coming out of the laser is sent to a reference cavity to stabilize the frequency to at least 0.1 $Hz/\sqrt{Hz}$ within the gravitational wave band. The remainder of the light goes through a triangular pre-mode cleaner to strip off light in modes other than TEM_{00}. This pre-mode cleaner also provides passive filtering of intensity noise. During the S1 run, intensity noise was around $10^{-7}/\sqrt{Hz}$. However, the intensity noise control system was only partially implemented, it is expected to be able to deliver intensity noise as low as $10^{-8}/\sqrt{Hz}$.

The light from the laser then passes into the vacuum and the input optics subsystem. During the S1 run, only 1 W of laser light entered the input optics in the Livingston and Hanford 4 km interferometers, and about 6 W for the Hanford 2 km interferometer. The light first passed through a mode cleaner, three suspended optics forming a triangular cavity 24 m in length with a finesse of about 1350. The mode cleaner stabilizes the beam size, position, and pointing of the laser light. It also acts as an additional frequency reference. Feedback to the laser from the mode

cleaner reduced the frequency noise to around 10^{-3} Hz/$\sqrt{\text{Hz}}$ during the S1 run. After the mode cleaner is a Faraday isolator, which diverts light returning from the interferometer away from the laser and into a photodetector. Finally, the light goes through a mode-matching telescope made of three suspended optics. This telescope increases the beam size to about 4 cm ($1/e^2$ radius in power) to match the mode of the arm cavities.

2.2   Interferometer Optics

The optics that form the main part of the interferometer[5,6] are made from fused silica glass. The end mirrors and input mirrors that define the Fabry-Perot cavities and the recycling mirror are 25 cm in diameter and 10 cm thick, while the beam-splitter is the same diameter by only 4 cm thick. The radii of curvature for the cavity optics are designed so the arm cavities have a stability factor

$$g = (1 - L/R_1)(1 - L/R_2) \tag{1}$$

of 0.33 for the 4 km long arms and 0.67 for the 2 km long arms. The surface figure of the polished optics is better than 1 nm.

The coatings on the mirrors are alternating layers of silica and tantala. The optical absorption is less than 1 part-per-million (ppm), and the scatter is less than 70 ppm. The transmission of the end mirrors is chosen to be as low as practical, about 12 ppm. The input mirrors have transmissions of 2.7%, which results in a pole frequency for the Fabry-Perot cavities of 85 Hz for the 4 km long arms and 170 Hz for the 2 km long arms. The beamsplitter transmission is 50 $\pm$ 0.5%. The recycling mirror has a 2.7% transmission, resulting in a recycling gain of about 50.

The absorption of optical power in the mirror substrates and coatings causes significant heating and thermal lensing of the optics. Thus, the radii of curvature specification assumed a level of thermal lensing. The low light level used during the S1 run caused reduced heating and thermal lensing in the 4 km interferometers. In addition, the Hanford 2 km interferometer had an anti-reflection coating for one of the input mirrors that did not meet the specification. These issues resulted in lower recycling gain for all interferometers during S1 than the design goal.

The silica substrates for the mirrors were chosen, in part, for their low mechanical loss. Through the fluctuation-dissipation theorem, this low loss results in low, off-resonance thermal noise. The mirror geometry assures that the lowest internal mode of the mirrors is above the gravitational wave band, at about 6700 Hz. The coating can also cause significant thermal noise[7], however for the particular materials[8] and beam radii used here, coating thermal noise is not expected to be a limiting noise source. Mechanical Q's of *in situ* mirrors have been measured to check this expectation. The addition of glued-on magnets and suspension attachments (see below) can degrade some modal Q's, but because of their distance from the reflected laser, will have a negligible effect on thermal noise. Modal Q's are found to be between 2×10^5 and 1.6×10^7, and finite element modeling predicts that the thermal noise will be near the design goal, and well below the limiting noise during the S1 run.

2.3 Seismic Isolation and Suspensions

To reduce the effect of ground motion on the optics, the optics are supported from a vibration isolation stack; an alternating series of masses and springs. This stack is composed of four layers of passive isolation[9]. The final stage is the optical table that holds the suspension. The internal modes of this isolation structure are between 1.5 and 12 Hz, with Q's typically between 10 and 30. The transfer functions between motion on the ground and motion at the optical table begins to roll off as f^{-8} above the highest frequency mode, reaching an isolation factor of about 10^6 at 50 Hz.

Each interferometer optic is suspended from a steel piano wire pendulum. A schematic drawing is shown in Fig. 6. The single loop of wire is loaded to near 40% of its yield strength. Small glass rods with a notch in them to hold the wire are glued onto the optic on each side a few millimeters above the center of mass. This helps define a clear break-off point for the wire, and reduces friction. This results in suspension modes of 0.74 Hz for the pendulum mode, 0.5 for yaw, 0.6 Hz for pitch, 12 Hz for the vertical bounce mode, 18 Hz for the roll mode, and 345 Hz (and multiples thereof) for the violin modes. All of these modes, except the first few violin modes, are outside of the gravitational wave band. Thermal noise from the violin modes is visible in the detection band in a very narrow region around the resonance. Off-resonance thermal noise is minimized by using the stressed steel wire and careful assembly of the suspension. Suspension thermal noise is expected to be a limiting noise source at the LIGO design goal (see Fig. 4). Measured Q's on the low frequency pendulum modes of typically 2 to 4 10^5 show that suspension thermal noise is below the sensitivity of the S1 run.

The necessary control forces for the interferometer optics are applied using the suspension system. Magnets made from Nd:Fe:B, 4 on the back and 1 one the side, are epoxied onto each optic. The back magnets have alternating polarity to reduce coupling to external magnetic fields. Attached to the suspension housing are assemblies next to each magnet containing a voice coil, LED, and a photodetector. These coils can be used to apply forces to the optics. The LED/photodector is used as a position sensor, allowing all the low frequency modes to be monitored. Feedback from the position sensors to the coils creates local damping at each mode. Signals from optical levers reflecting off the mirror faces and wavefront sensors are also fed back to these coils. All of these controls were used during the S1 run, but at the design sensitivity only wavefront sensing will be used.

2.4 Length Sensing and Control

The entire detector has four degrees of freedom that must be held in place so that all of the optical cavities remain on resonance; the length of each Fabry-Perot arm, the position of the beamsplitter, and the position of the recycling mirror. The beamsplitter is held to control the Michelson interferometer on resonance at a dark fringe on the output port and the recycling mirror controls the resonance in the recycling cavity. These lengths are sensed with an extended version of the Pound-Drever-Hall technique[10] using RF phase modulation sidebands on the carrier laser light. The modulation is applied using Pockels cells and the frequency is chosen so the sidebands are anti-resonant in the Fabry-Perot arms. The carrier light, however,
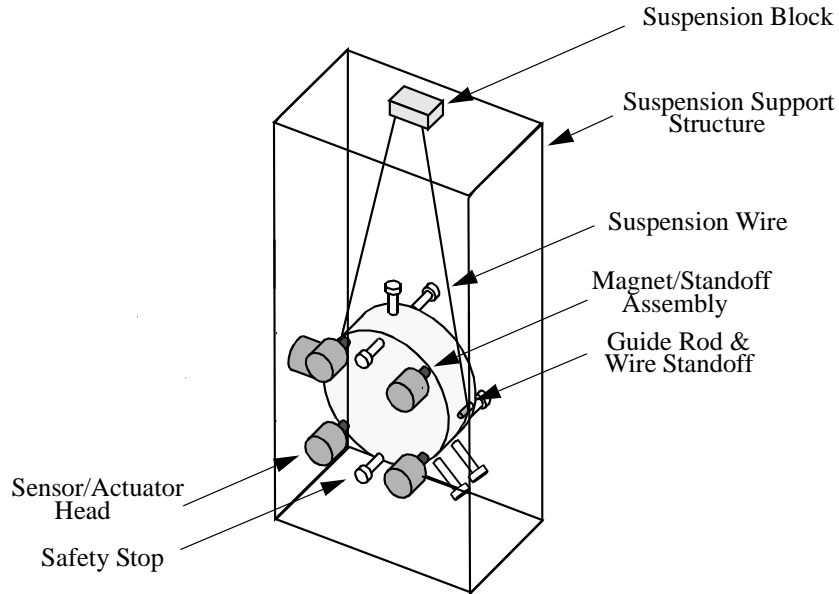
Figure 6. The suspension of a LIGO optic.

is strongly overcoupled in the arms. The recycling cavity is resonant for both the carrier and the sidebands. The sideband light is sent out through the dark port of the Michelson, where it is used as a local oscillator, by introducing a length asymmetry of about 30 cm between the two beamsplitter-input mirror distances.

Output beams from the interferometer are measured at three locations in the optical plant to determine these four length signals. At the antisymmetric port of the Michelson, the demodulated signal with quadrature phase gives the differential arm length. This is the signal that contains information about gravitational wave events. Light reflected from the recycling mirror demodulated in-phase gives the common mode motion of the arms. This signal is a measure of the laser frequency noise, and is fed back to the laser and mode cleaner to stabilize the laser frequency to the average arm length. Light from inside the recycling cavity, reflected off the anti-reflection coating of the beamsplitter, gives the remaining degrees of freedom. The in-phase part provides the recycling cavity length and the quadrature phase gives the differential length between the beamsplitter and the input masses. The positions of the photodetectors that measure these signals along with the interferometer layout and degrees of freedom is shown in Fig. 7.
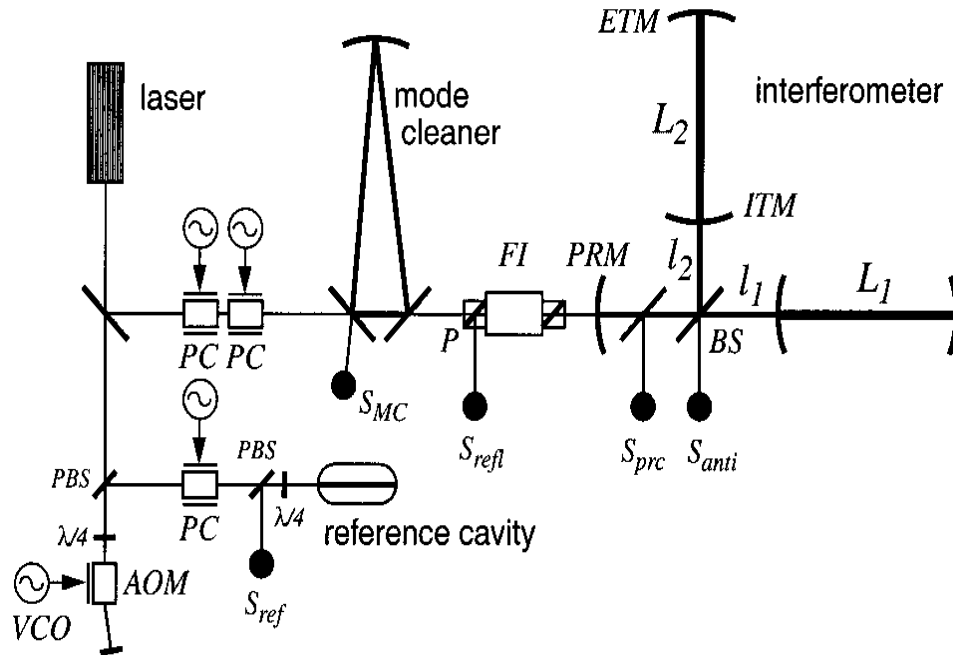
Figure 7. LIGO interferometer showing laser, mode cleaner, the Pockels cells, and photodetectors (S_{xxx}) used to sense the interferometer lengths. The arm cavities have lengths L_1 and L_2, the differential arm cavity length is $L_1 - L_2$ and the common mode is $L_1 + L_2$. The Michelson degrees of freedom are the separation of the beamsplitter from the input masses, l_1 and l_2. Abbreviations used; PC - Pockels Cell, VCO - Voltage Controlled Oscillator, S_{mc} - Mode Cleaner Signal, S_{ref} - Reference Cavity Signal, S_{refl} - Reflected Cavity Signal, FI - Faraday Isolator, PRM - Power Recycling Mirror, S_{prc} - Power Recycling Cavity Signal, S_{anti} - Antisymmetric Port Signal, ITM - Input Test Mass, ETM - End Test Mass.

3 Results from the First LIGO Science Run

3.1 Technical Results

The first science run was intended to collect an amount of real interferometer data to analyze for the presence of gravitational waves. The sensitivity of all interferometers during S1, shown in Fig. 8, was far from the design sensitivity and the run was planned for only 17 days. This made it highly unlikely that a confirmed detection of a gravitational wave would occur during S1, but this data does allow for the setting of upper limits on the event rate and amplitudes of various predicted waves. This upper limit analysis would also provide an opportunity to test the data analysis pipelines with actual data, including non-Gaussian and non-stationary noise. This science data run, similar to the previous engineering runs, also allowed for further evaluation of the interferometers in operation, and to train scientists and operators. The duty cycle of the interferometers was monitored, especially the combined duty cycle of two and three interferometer groupings. The duty cycle results are shown in Table 3.1.

Strain Sensitivities for the LIGO Interferometers for S1
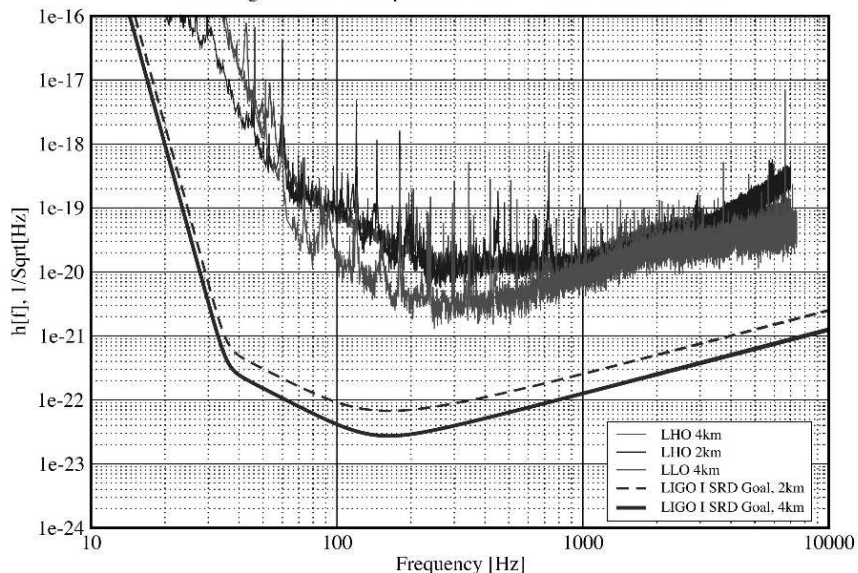### 23 August 2002 - 09 September 2002   LIGO-G020461-01-E

Figure 8. Strain noise in the three LIGO interferometers during S1.

Table 1. Total time locked and duty cycle for LIGO interferometers during S1. The GEO 600 interferometer had a duty cycle of 98% during S1, so its time in coincidence with LIGO interferometers is constrained by the LIGO duty cycle.

Detector/combination	Hours of Operation (Coincidence)	Duty Cycle
Hanford 4 km	235	57.6%
Hanford 2 km	298	73.1%
Livingston 4 km	170	41.7%
Hanford 4 km/Hanford 2 km	188	46.1%
Hanford 4 km/Livingston 4 km	116	28.4%
Hanford 2 km/Livingston 4 km	131	32.1%
All Three Interferometers	96	23.4%

3.2   Upper Limits on Gravitational Wave Events

Possible gravitational wave events were divided into four categories for purposes of analysis of S1 data; bursts, inspiralling compact objects, continuous waves, and stochastic background. The data was analyzed using different techniques in order to set upper limits on events in all of these categories. Constraints on time and personpower prevented an exhaustive look at all possible waveforms, so priorities were set based on likelihood of events and maturity of the analysis techniques

available.

All analyses (except the periodic sources search) used the LIGO Data Analysis System (LDAS). LDAS provides a distributed software environment with scalable hardware configurations to provide for LIGO's computational needs. The architectural design is based on multiple concurrent data analysis pipelines into which data is fed as it is collected. The data then proceeds down the pipeline where necessary signal analysis procedures are carried out[11].

3.3 Burst Events

Burst events, for purposes of the S1 analysis, are any gravitational wave event of any shape with duration 4 ms to 100 ms. No other assumptions are made about the waveform of a burst, so matched template analysis can not be used. This category may encompass a large number of astrophysically possible events, but gravitational waves from supernova and gamma ray bursts are of particular interest. A more complete description of the burst analysis can be found in Ref[12].

To optimize various parameters used in the burst search, a section of the data, called the "playground data", was set aside and not used for generating the final upper limit. This playground data was composed of about 10% of the total data set, and was chosen by hand to include as much data variety as possible. To eliminate especially noisy data, the total RMS noise in four frequency bands were monitored and compared to a threshold every six minutes. When band-limited RMS noise exceed the threshold in any band, the six minutes of data was removed from further analysis. To calibrate the data, three sinusoidal excitations were continuously applied to the mirrors during the run. To ensure the calibration of the data was valid, times when these calibration lines gave unreliable results were also rejected from the data analysis stream. Finally, only times when all three LIGO interferometers were providing useful data were included in the analysis. These cuts left 33.5 hours of data to be used in the burst analysis.

This data is then passed to two different event trigger generators; SLOPE[13,14] and TFCLUSTERS[15,16]. These are used to generate a list of candidate events for each interferometer. SLOPE differentiates the data in time, and compares the slope to a threshold. TFCLUSTERS identifies connected regions in time-frequency space where the power is not consistent with stationary Gaussian noise. After event trigger generation, the analysis pipeline allows for vetoes to be applied based on environmental and interferometer data channels. However, no vetoes were used in the S1 burst analysis. The final step in the pipeline is to generate a list of event candidates using the candidate lists from all three interferometers. Temporal coincidence is required to within the light travel time between the sites (10 ms) plus the uncertainty in the detection time. For the SLOPE events, a total coincidence window of 50 ms was used, for TFCLUSTERS 500 ms was used. For both trigger generators, all triggers that satisfy the triple coincidence that are within 500 ms of one another are considered the same event. Finally, for TFCLUSTERS events only, a coincidence of 80 Hz in frequency spacing or less is required. This creates the final list of event candidates.

The list of event candidates will include not only gravitational wave events, but

Table 2. Coincident events and confidence bands on the number of excess events in the S1 run from `TFCLUSTERS`.

Coincident Events	6
Background Events	10.1 ± 0.6
90% confidence band	0 - 2.3
95% confidence band	0 - 3.5
99% confidence band	0 - 5.9

some number of random noise events. To determine if the number of candidate events is consistent with only background noise, a background estimate is done. This is done by putting in an artificial time shift in one of the data streams, then performing the same analysis as is done originally. This ensures that gravitational waves will be unable to cause a coincidence, so any events will be solely due to noise. Time shifts between 8 s (to avoid correlated events) and 100 s (to minimize dependence on non-stationarity in the noise) were used. It is found that both event trigger generators give varying trigger rates. The results for `TFCLUSTERS` is shown in Table 3.3. In the presence of the non-stationary noise during S1, `SLOPE` gave such severe variability that it was not considered suitable for setting an upper limit on gravitational wave events. The problem was due to the use of a single fixed threshold, unlike our implementation of `TFCLUSTERS`. In future applications of the `SLOPE` algorithm, we may use an adaptive threshold instead.

Based on the background events in the time shift analysis, the number of co-incident events from `TFCLUSTERS` in Table 3.3 is low. None of these coincident events were seen by `SLOPE`. The probability of obtaining six or fewer events from `TFCLUSTERS` based on the measured background is 12%. There is no reason to suspect any systematic error. Alternate methods of estimating the background do not give significantly different background rates.

These results can be used to set an upper limit of the rate of gravitational wave events as a function of amplitude. This is shown in Figure 9. These results assume an isotropic ensemble of waves bathing the Earth with constant amplitude, with all linear polarizations, and a waveform of either Gaussian or sine-Gaussian shape. This model is not motivated by astrophysics, but is meant to be representative of generalized burst events. The amplitude in Figure 9 is shown as the root sum square amplitude spectral density,

$$h_{rss} \equiv \sqrt{\int |h(t)|^2 \mathrm{d}t}. \qquad (2)$$

Resonant detectors have set more stringent upper limits in the subset of these signals that intersect their frequency band[17,18]. LIGO results will improve both as the interferometer noise is lowered and as longer data sets are accumulated.

3.4 Inspiralling Binary Neutron Stars

Compact objects inspiralling together because of gravitational wave emission is one of the best understood astronomical sources of potentially detectable gravitational waves[1]. To create waves in the LIGO band, the inspiralling objects may be any
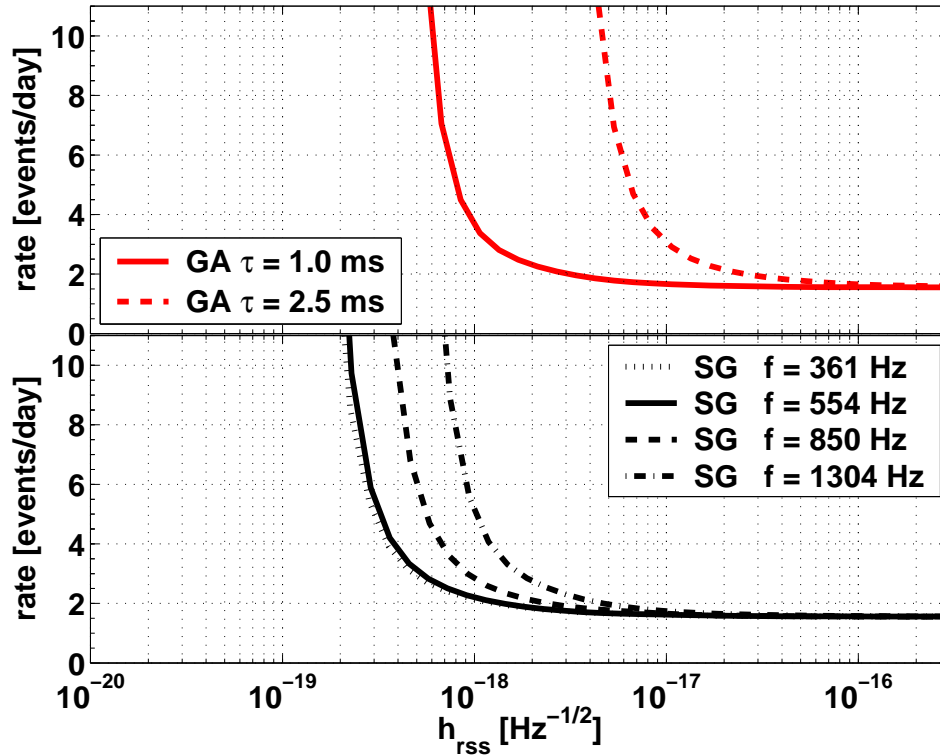
Figure 9. Rate vs. amplitude for detection of specific waveforms using TFCLUSTERS. The top graph shows rates for Gaussian waveforms, with τ of 1.0 ms or 2.5 ms. The bottom graph shows rates for sine-Gaussian with central frequencies 361 Hz, 554 Hz, 850 Hz, and 1304 Hz and Q of 9. In each case, the region above and to the right of the curves are excluded at the 90% confidence level.

combination of solar mass-scale black holes and neutron stars. In the S1 analysis, only binary neutron star waveforms were searched for. This is because of constraints on time and personpower, but also because such systems have been observed to exist[19,20,21]. From the known examples, estimates of the rate of these events can be made[22,23] which suggest that this source may be detectable by LIGO at the design sensitivity searching for only a few years. A more complete description of the inspiral analysis can be found in Ref[24].

The waveform from binary neutron star inspiral can be analytically predicted by using a post-Newtonian expansion[25], so a template based search can be used. This is the optimal way to search, providing the highest signal-to-noise ratio possible from a given data set. Only data from the LIGO Hanford and Livingston 4 km interferometers were used in the inspiral analysis.

The templates used were for non-spinning neutron stars, so depend on four parameters; the two neutron star masses m_1 and m_2, the arrival phase α, and an effective distance D_{eff}. This effective distance depends on the true distance to the

Table 3. Event candidates with signal-to-noise ratio (SNR) higher than 8. The first line shows candidates from when both interferometers were operating, while the last two lines area when only a single interferometer was operating. The first line is for candidates where the Livingston interferometer saw an event with a D_{eff} less than 51 kpc. Coincidence with Hanford was required in this case, and no events survived this coincidence.

Interferometers Operating	Number of candidates	Maximum SNR
Livingston and Hanford ($D_{\text{eff}} < 51$ kpc)	0	-
Livingston and Hanford ($D_{\text{eff}} > 51$ kpc)	418	15.6
Livingston	786	15.9
Hanford	274	12.0

source, the location of the source on the sky relative to LIGO's antenna pattern, and the orientation of the orbital plane of the stars. The masses m_1 and m_2 were allowed to vary between $1-3M_\odot$. The spacing of the templates in parameter space was chosen so that the most signal-to-noise ratio lost due to template mismatch would be less than 3%. This led to 2110 different templates being used. Each template, using a frequency dependent weighting due to the noise in the given interferometer, was integrated against the data for all possible arrival times.

Candidate inspirals from either the Hanford or Livingston 4 km interferometers are found by requiring the signal-to-noise ratio for some template to be higher than 6.5. They must be separated from each other by at least the length of a template. Then they must pass a time-frequency veto designed to reject high amplitude transients. Contributions from eight separate frequency bands were compared to the expected distribution from an inspiral signal[26]. The sum of the squared difference between the actual amount from each frequency and the expected must be less than a threshold. This is known as the χ^2 test. Candidates which pass all of these tests are stored as triggers.

Data was not analyzed from either interferometer if the calibration lines were not reliable or if the overall noise in the gravitational wave band was too high. This is similar to conditions used in the burst analysis in Section 3.3. However, unlike the burst search, a veto from an additional data channel was used here. The signal from the common, rather than differential, mode of the arms in the Hanford interferometer was compared to Hanford triggers. This signal is sensitive to frequency noise in the laser. Hanford triggers within one second of events in the common mode signal were vetoed. No instrument vetoes were used with Livingston triggers.

The sensitivity of the Livingston 4 km interferometer was significantly higher than the 4 km interferometer at Hanford (see Fig. 8). Because of this, coincidence between triggers from Hanford and Livingston was not required, unless the predicted D_{eff} from Livingston was less than 51 kpc and hence should be seen in the Hanford data. No coincident triggers were seen in the S1 data. Coincidence would be used to improve confidence in any detection event. The resulting triggers from all of these conditions are shown in Table 3.4.

An upper limit on inspiralling binary neutron star events during S1 was computed from these triggers. The largest signal-to-noise ratio found was used as a threshold. The probability of observing an inspiral signal with signal-to-noise ra-

tio ρ greater than this threshold ρ^* with some rate R for inspiral events and an observation time T is

$$P(\rho > \rho^*; R) = 1 - e^{-RT\epsilon(\rho^*)}, \qquad (3)$$

where $\epsilon(\rho)$ is the efficiency of the analysis for events with signal-to-noise ration ρ. The efficiency $\epsilon(\rho^*)$ for $\rho^* = 15.9$ was determined to be 0.53 from Monte Carlo simulations. A frequentist upper limit with 90% confidence on R can be found by solving $P(\rho > 15.9; R_{90\%}) = 0.9$ for $R_{90\%}$. With a time T of 236 hours from S1, this gives

$$R_{90\%} = 170/\text{y}/\text{MWEG}, \qquad (4)$$

where MWEG stands for Milky Way Equivalent Galaxy. Here, a value of the efficiency $\epsilon(\rho^*)$ at the lower end of the 1-σ error bars has been assumed to make a conservative estimate. This upper limit is better than previous limits from direct observation[27,26] by a factor of 26.

3.5 Periodic Sources of Gravitational Waves

Rapidly rotating neutron stars that have some deviation from sphericity will give off gravitational waves. The data from all three LIGO interferometers plus the GEO 600 interferometer were examined for gravitational waves from the fastest rotating pulsar known, J1939+2134. A more complete description of the periodic analysis can be found in Ref[28].

The expected signal from a rotating pulsar is an AM and FM modulated sinusoid, Doppler shifted according to the relative motion of the Earth with the pulsar, and amplitude modulated by its motion relative to the antenna pattern of the interferometer. This allows a template based search technique to be used, similar to that used for the inspiralling binary neutron stars. However, the continuous nature of the pulsar signal allows for integration of the signal-to-noise ratio over the complete time that data is available. There are known upper limits to all known pulsars based on their observed spin down rates, as gravitational wave emission cannot be removing energy and angular momentum faster than is seen. The sensitivities of the LIGO and GEO interferometers to periodic sources during S1 along with the spin down upper limits for all known pulsars is shown in Fig. 10.

Two separate data analysis methods were used to set an upper limit on periodic gravitational waves from J1939+2134; a frequency domain technique that yields a frequentist upper limit, and a time domain one that gives a Bayesian upper limit. The frequency domain approach is best for searches over a relatively large parameter space, while the time domain technique is best suited for searches where the phase evolution of the signal is known.

The frequency domain search uses the F statistic method[29]. The logarithmic likelihood function is defined as

$$\ln(\Lambda) = (s|h) - \frac{1}{2}(h|h), \qquad (5)$$

where

$$(s|y) = 4\Re \int_0^\infty \frac{\tilde{s}(f)\tilde{y}^\star(f)}{S_n(f)} \mathrm{d}f, \qquad (6)$$
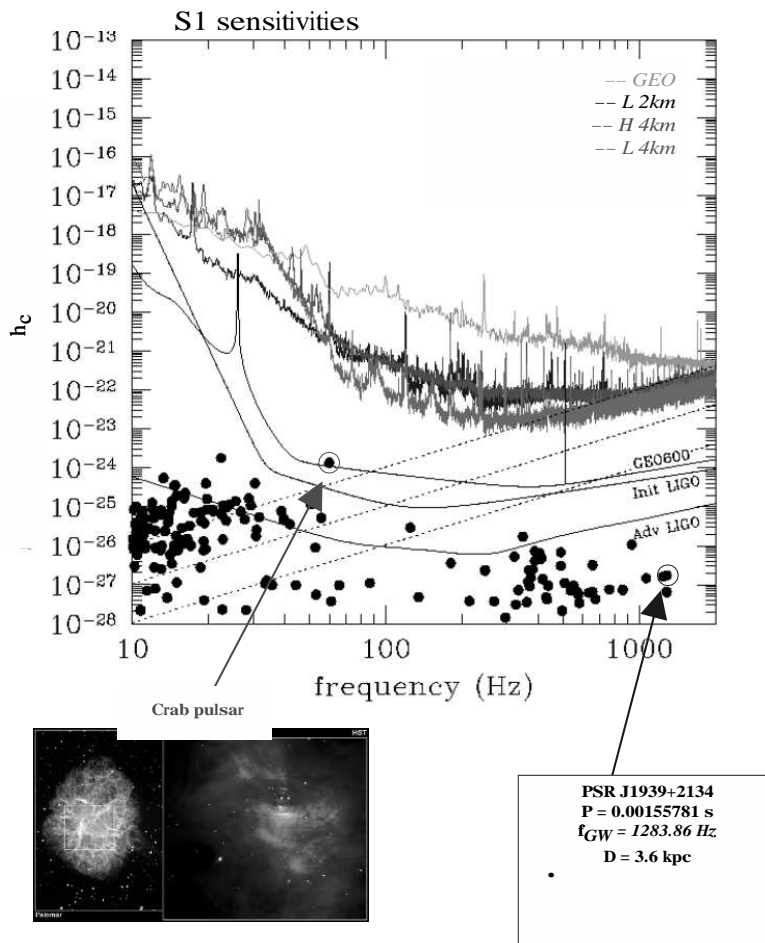
S1 sensitivities

Figure 10. The upper curves show the sensitivities of the four interferometers integrated over their respective up times during S1. The lower curves are sensitivities for a 4 km LIGO interferometer and GEO 600 with an integration time of one year. The filled circles are upper limits on gravitational wave emission from known radio pulsars based on their observed spin down rates. The circled points are the upper limits for the Crab pulsar and for J1939+2134.

where s is the calibrated detector output time series, h is the template, ˜ is the Fourier transform operator, and $S_n(f)$ is the one-sided power spectral density of the noise. The F statistic is the maximum value of $\ln(\Lambda)$ with respect to all the unknown signal parameters. For the pulsar J1939+2134, the known parameters are its position on the sky, the frequency f_s, and the spin down $\dot{f}_s$. The unknown parameters are the orientation angle $\imath$, the polarization state of the wave (angle ψ),

Table 4. Upper limits for the gravitational wave amplitude from J1939+2134 at the 95% confidence level. The time domain search gives a conservative frequentist limit, and the frequency domain search results in a Bayesian limit.

Interferometer	Time Domain Search	Frequency Domain Search
GEO 600	$(1.9 \pm 0.1) \times 10^{-21}$	$(2.2 \pm 0.1) \times 10^{-21}$
Livingston 4 km	$(2.7 \pm 0.3) \times 10^{-22}$	$(1.4 \pm 0.1) \times 10^{-22}$
Hanford 4 km	$(5.4 \pm 0.6) \times 10^{-22}$	$(3.3 \pm 0.3) \times 10^{-22}$
Hanford 2 km	$(4.0 \pm 0.5) \times 10^{-22}$	$(2.4 \pm 0.2) \times 10^{-22}$

the initial phase ϕ_0, and the amplitude h_0. The results shown in the right column of Table 3.5 are the frequentist upper limits that were derived from the F statistic for each interferometer.

The time domain search is better suited to situations where the phase evolution is known. The data to be analyzed for a pulsar signal is heterodyned with a unit amplitude complex function with phase evolution equal to but opposite in sign of the signal. The noise in the interferometer is estimated every minute, and a Bayesian upper limit is calculated. The results are shown in Table 3.5. For both searches, the upper limits are significantly above the limit imposed from the measured spin down rate; $h_0 \leq 10^{-27}$.

3.6 Stochastic Gravitational Waves

There may be a background of gravitational waves analogous to the cosmic microwave background for electromagnetic waves. They are unlikely to be thermally distributed, however. The origin could be cosmological, left over from the big bang or other early universe processes, or it could be astrophysical, an amalgam of more modern sources such as distant supernova or inspirals. See References[30,31] for more information about stochastic background sources. These sources would appear in LIGO as an otherwise unexplained noise source. Upper limits can be set using a single interferometer, but a much stronger limit comes from cross-correlating multiple detectors. Unlike true instrument noise, the effect of the stochastic background will be correlated in some frequency band. A more complete description of the stochastic background analysis can be found in Ref[32].

The spectrum of a stochastic background of gravitational waves is usually described by the dimensionless quantity $\Omega_{gw}(f)$, the gravitational wave energy density per logarithmic frequency divided by the critical density to close the universe ρ_c:

$$\Omega_{gw}(f) \equiv \frac{f}{\rho_c} \frac{d\rho_{gw}}{df}. \tag{7}$$

The one-sided power spectrum of the gravitational wave strain, $S_{gw}(f)$, is then

$$S_{gw}(f) = \frac{3H_0^2}{10\pi^2} f^{-3} \Omega_{gw}(f), \tag{8}$$

where H_0 is the Hubble expansion rate. It is customary to search for a stochastic background for which $\Omega_{gw}(f)$ is independent of frequency in the LIGO band, as such a spectral behavior is predicted by various cosmological models such as inflation or

Table 5. Upper limits for stochastic gravitational waves from pairs of LIGO interferometers. For the value of Ω_0 quotes, the Hubble constant was assumed to be 100 km/(s Mpc). It scales inversely with H_0.

Interferometer pair	90% confidence upper limit	Frequency range	Observation time
Hanford 4 km/Livingston 4 km	$\Omega_0 \leq 55 \pm 11$	40 -314 Hz	64 hr
Hanford 2 km/Livingston 4 km	$\Omega_0 \leq 23 \pm 4.6$	40 -314 Hz	51.25 hr

cosmic strings. In this case, the gravitatioanl wave strain power spectrum falls as $1/f^3$.

To search for stochastic signals in the S1 data, the background is assumed to be isotropic, unpolarized, stationary, and Gaussian. An incoherent superposition of local sources could cause an anisotropic and/or non-Gaussian background, and may require different data analysis techniques than are used in Ref[32]. Cross correlation of the output of different detectors provides the best search method, and allows for better limits than just analyzing data from a single interferometer. The cross correlation is done in the frequency domain, and narrow lines from correlated sources, e.g. power lines, are removed. This technique is limited by a frequency dependance set by c/d where d is the distance between the two detectors. This dependance is discussed in detail in Ref[32].

The result of this analysis is to set an upper limit on $\Omega_{\mathrm{gw}}(f)$. Arguments from nucleosynthesis already set an upper limit on this value of 10^{-5} in the LIGO band. The cross correlation between the two Hanford interferometers actually gives a negative value, which for this reason alone could not be due to a gravitational wave signal. This is due to correlations in the instrument noise, likely from coupling of acoustic noise to the output optics. Useful upper limits can be set using the correlations between Hanford and Livingston, these are shown in Table 3.6. This upper limit is about 10^4 better than the previous broadband limit using interferometric detectors and a factor of about 3 better than the narrowband results from resonant mass detectors.

4   Advanced LIGO

Advanced LIGO is the next generation detector currently being researched and designed. The goal is to reach a sensitivity of 200 Mpc for a single 4 km interferometer, or about 300 Mpc for the three interferometer detector system. At this sensitivity, it is expected that there should be multiple sources a year, if not a week. It will be a complete changeover from the initial LIGO interferometers, with all subsystems replaced, and only the vacuum envelope and other infrastructure retained. A conceptual design for Advanced LIGO was put forth by the LIGO Scientific Collaboration in 1998, and research and development has been ongoing to develop the needed technology. There will be significant improvements in the seismic isolation, suspensions, mirrors, laser, and interferometer configuration. These technologies are discussed below. The expected sensitivity is shown in Fig. 11. In February of 2003, a construction proposal was sent the the US National Science Foundation requesting funds to build Advanced LIGO.
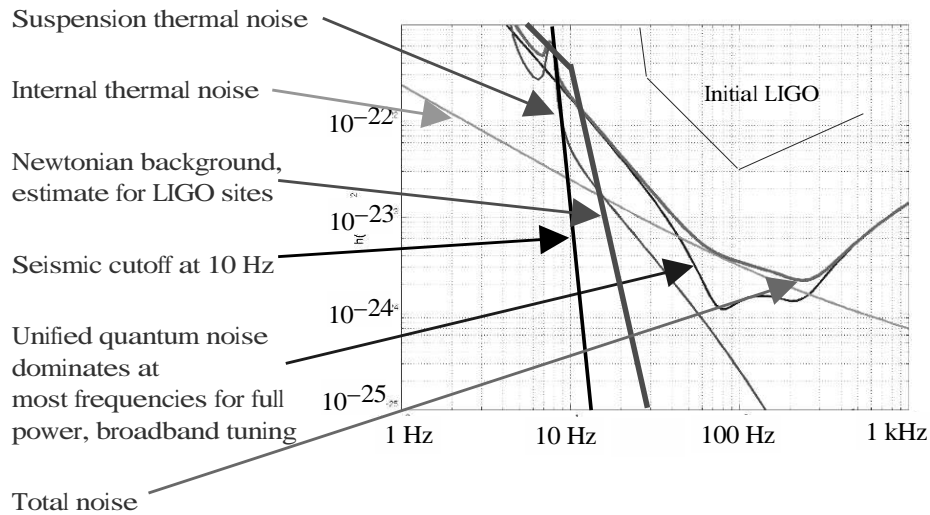
Figure 11. Sensitivity curve for a proposed Advanced LIGO interferometer. The addition of a signal recycling mirror makes radiation pressure and shot noise couple into a single quantum noise. The Newtonian background due to gravitational attraction between the mirrors and the moving ground is expected to dominate over seismic noise at low frequencies. In the middle frequencies, at the most sensitive region, thermal noise from the optical coatings and mirror substrates is the limiting noise source.

4.1 Seismic Isolation and Suspensions

The seismic isolation is designed to reduce mirror motion from seismic disturbances to below thermal noise for all frequencies above 10 Hz. This will be done using a two stage, in vacuum, seismic isolation platform with additional low frequency hydraulic pre-isolators external to the vacuum system. The hydraulic pre-isolators are being developed for use on the initial LIGO interferometers to reduce excess seismic noise at Livingston. The in-vacuum stages use a combination of position sensors, seismometers, and geophones to sense motion in the platform. These sensors are then used in feedback loops to apply forces which control all six degrees of freedom[34].

Below the seismic isolation, the interferometer mirrors hang from suspensions. These are quadruple pendulums, based on triple pendulums used in GEO 600 now[35]. The upper two pendulums use steel wires and blade springs to isolate the mirrors from motion in all directions. The penultimate mass will be made from sapphire or heavy glass and is suspended from above by steel wires. The mirror, which will be

sapphire or fused silica, then hangs below the penultimate mass from a fused silica suspension. The use of silica here rather than steel will reduce suspension thermal noise to a level below the quantum noise for most configurations (see Fig. 11). The suspension filaments will be shaped either as ribbons[36] or dumbbell fibers[37] to get the maximum dissipation dilution while still maintaining the strength to support the mirrors. The filaments will be welded to silica standoffs which are attached to the masses by hydroxy catalysis bonding[38] to create low thermal noise but high strength connections.

4.2 Lasers and Mirrors

To reduce shot noise, a more powerful laser is needed in Advanced LIGO. Plans are to use a 180 W injection-locked, end-pumped rod of Nd:YAG which is based on the current GEO 600 laser[39]. This will be able to meet the frequency and intensity noise requirements while greatly increasing the power available. This additional power will mean greater heating in the mirrors and beamsplitter, however. Plans are being developed to use active thermal compensation to reduce thermal lensing in the core optics. This will done using a ring heater, or a scanning CO_2 laser. Thermal compensation is also being added to the initial LIGO interferometers to solve thermal lensing problems there.

The interferometer mirrors in Advanced LIGO will be made of sapphire, for its greater ability to handle heat, or fused silica should problems arise with sapphire. The mirrors are planned to be 40 kg to reduce the effect of radiation pressure from the higher power laser. Single crystal sapphire has been grown this large, and been shown to have acceptable thermal noise characteristics. If the same tantala/silica coatings used in initial LIGO were used for the mirrors in Advanced LIGO, thermal noise from these coatings would place significant limits on the sensitivity. Research is ongoing to understand the cause of thermal noise in optical coatings, to measure it directly, and to develop a coating with lower thermal noise that still meets the optical and thermal requirements[8].

An additional mirror will be added to the initial LIGO configuration of a power-recycled Michelson interferometer with Fabry-Perot arm cavities for Advanced LIGO. This is at the output port, called the signal recycling mirror, and it allows the frequency response of the sensitivity curve to be changed and tuned. By changing the position of this mirror by a fraction of the laser wavelength, the maximum sensitivity can be made to occur at different frequencies. By changing the reflectivity, the bandwidth around the central frequency can be adjusted. The quantum noise shown in Fig. 11 shows a signal recycling choice optimized for binary neutron star inspirals, with a maximum sensitivity at about 150 Hz and a fairly wide bandwidth. A higher frequency, narrow bandwidth choice could be made for gravitational wave pulsar searches, where the particular frequency is known in advance. By lowering the laser power as well as adjusting the signal recycling mirror, Advanced LIGO could be optimized for low frequency performance to enhance searches for stochastic backgrounds, for example. The quantum noise can also be adjusted to best take advantage of the other noise sources, particularly mirror thermal noise, so sensitivity is enhanced in bands where other noise sources are not limiting. Figure 12 shows
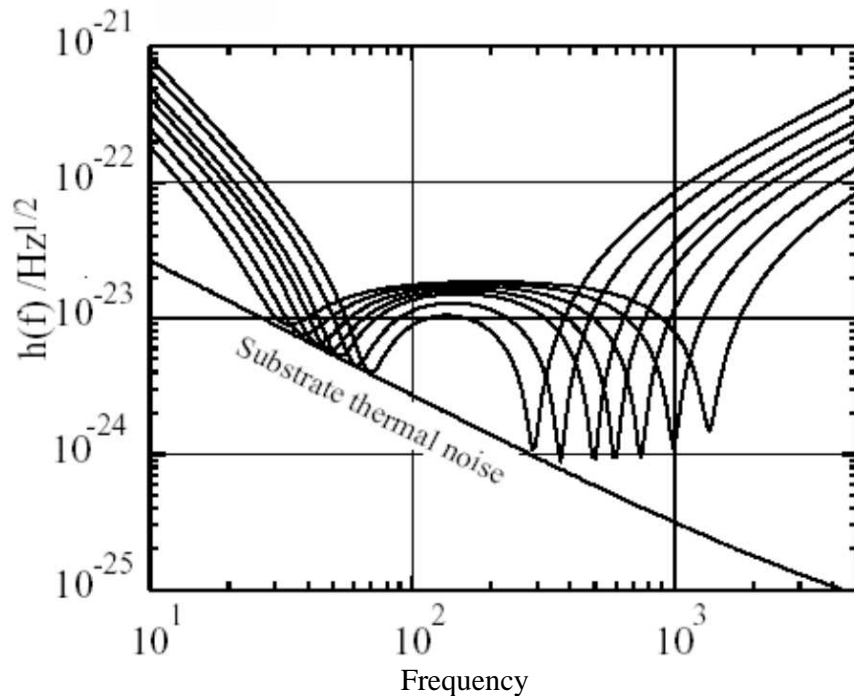
Figure 12. Sensitivity curves for a number of different positions and reflectivities of the signal recycling mirror in Advanced LIGO. These curves are chosen to have high sensitivity in narrow bandwidths, and to avoid the mirror thermal noise that could otherwise limit sensitivity.

a variety of possible narrowband options, as well as the role that mirror thermal noise can play.

5 Conclusion

The field of gravitational wave astronomy is rapidly becoming more of a reality, as the LIGO interferometers, in collaboration with GEO 600, are taking data at record sensitivities. The noise in these interferometers is getting close to the design goal, and achieving the target sensitivity is expected to happen soon. The data from the first science run has been analyzed for four categories of gravitational waves, setting upper limits on the rates and amplitudes of them all. Plans are well under way for Advanced LIGO, where every subsystem in the interferometer will be upgraded to take advantage of technological innovations that have occurred since initial LIGO was built. This project was proposed in February of 2003, and should allow for more than a magnitude improvement of sensitivity across a broad band.

Acknowledgments

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