

Inspiral-merger-ringdown waveforms of spinning, precessing black-hole binaries in the effective-one-body formalism

Yi Pan,¹ Alessandra Buonanno,¹ Andrea Taracchini,¹ Lawrence E. Kidder,²
Abdul H. Mroué,³ Harald P. Pfeiffer,^{3,4} Mark A. Scheel,⁵ and Béla Szilágyi⁵

¹*Maryland Center for Fundamental Physics & Joint Space-Science Institute,
Department of Physics, University of Maryland, College Park, MD 20742, USA*

²*Center for Radiophysics and Space Research, Cornell University, Ithaca, New York 14853, USA*

³*Canadian Institute for Theoretical Astrophysics, 60 St. George Street, University of Toronto, Toronto, ON M5S 3H8, Canada*

⁴*Canadian Institute for Advanced Research, 180 Dundas St. West, Toronto, ON M5G 1Z8, Canada*

⁵*Theoretical Astrophysics 350-17, California Institute of Technology, Pasadena, CA 91125, USA*

(Dated: May 22, 2014)

We describe a general procedure to generate spinning, precessing waveforms that include inspiral, merger and ringdown stages in the effective-one-body (EOB) approach. The procedure uses a precessing frame in which precession-induced amplitude and phase modulations are minimized, and an inertial frame, aligned with the spin of the final black hole, in which we carry out the matching of the inspiral-plunge to merger-ringdown waveforms. As a first application, we build spinning, precessing EOB waveforms for the gravitational modes $\ell = 2$ such that in the nonprecessing limit those waveforms agree with the EOB waveforms recently calibrated to numerical-relativity waveforms. Without recalibrating the EOB model, we then compare EOB and post-Newtonian precessing waveforms to two numerical-relativity waveforms produced by the Caltech-Cornell-CITA collaboration. The numerical waveforms are strongly precessing and have 35 and 65 gravitational-wave cycles. We find a remarkable agreement between EOB and numerical-relativity precessing waveforms and spins' evolutions. The phase difference is ~ 0.2 rad at merger, while the mismatches, computed using the advanced-LIGO noise spectral density, are below 2% when maximizing only on the time and phase at coalescence and on the polarization angle.

PACS numbers: 04.25.D-, 04.25.dg, 04.25.Nx, 04.30.-w

I. INTRODUCTION

An international network of gravitational-wave (GW) detectors operating in the frequency band $10\text{--}10^3$ Hz exists today. It is composed of the LIGO detectors in Hanford, WA and Livingston, LA, in the United States, the French-Italian Virgo detector [1], and the British-German GEO600 detector [2]. Those detectors have collected and analysed data for several years. Since 2010 they have been shut down to be upgraded to the advanced LIGO and Virgo configurations [3]. The design sensitivity for advanced detectors, which is planned to be achieved by 2019 [4], will be a factor of ten better than the one of the initial detectors. This improvement implies an increase in the event rates of coalescing binary systems of (roughly) one thousand, thus making very likely the first detection of gravitational waves [4] with the advanced detector network. Furthermore, efforts to build a gravitational-wave detector in space are underway [5].

Binary systems composed of compact objects, such as black holes and neutron stars (compact binaries for short) are the most promising sources for groundbased GW detectors. The signal detection and interpretation is based on a matched-filtering technique, where the noisy detector output is cross-correlated with a bank of theoretical templates.

Fueled by numerical relativity (NR) simulations, there has been substantial progress in building and validating accurate templates for the inspiral, merger and ringdown

stages of nonprecessing ¹ black-hole (BH) binaries [6–18] (see also Ref. [19] where several analytical templates have been compared to simulations produced by the NRAR collaboration). Despite this progress, template modeling for generic, spinning BH binaries is far from being complete. In this paper we focus on BH binary systems moving on quasi-spherical orbits, i.e., generic precessing orbits that are circularized and shrunk by radiation reaction.

During the last several years, the post-Newtonian (PN) formalism, which expands the Einstein equations in powers of v/c (v being the characteristic velocity of the binary and c the speed of light), has extended the knowledge of the dynamics and gravitational waveform for quasi-spherical orbits through 3.5PN [20, 21] and 2PN [22] order, respectively. Precession-induced modulations in the phase and amplitude of gravitational waveforms become stronger as the opening angle between the orbital angular momentum and the total angular momentum of the binary increases. Compact binaries with large mass ratios can have larger opening angles than those with comparable masses.

Pioneering studies aimed at understanding and modeling precession effects in inspiraling compact binaries

¹ Here, nonprecessing means that the BH spins are either zero or aligned/antialigned with the orbital angular momentum.

were carried out in the mid 90s [23, 24]. As GW detectors came online in early 2000, it became more urgent to develop template families for spinning, precessing binaries in which precession-induced modulations were incorporated in an efficient way, reducing also the dimensionality of the parameter space. In 2003, Buonanno, Chen and Vallisneri [25] introduced the *precessing convention* and proposed a template family for precessing binaries in which precessional effects are neatly disentangled from nonprecessing effects in both the amplitude and phase evolutions. The precessing convention was initially introduced to model phenomenological or detection template families [25], and then it was extended to physical templates of single-spin binary systems in Refs. [26, 27]. In the past few years, geometric methods have been developed to construct preferred precessing reference frames [28–31] for numerical or analytical waveforms, achieving a similar disentanglement of precessional effects. Waveforms decomposed in such frames exhibit relatively smooth amplitude and phase evolutions and are well approximated by nonprecessing waveforms [32, 33].

Here, we use the effective-one-body (EOB) formalism [34–37] to model precessing inspiral, merger and ringdown waveforms. The basic idea of the EOB approach is to map by a canonical transformation the conservative dynamics of two compact objects of masses m_1 and m_2 and spins \mathbf{S}_1 and \mathbf{S}_2 into the dynamics of an effective particle of mass $\mu = m_1 m_2 / (m_1 + m_2)$ and spin \mathbf{S}_* moving in a deformed Kerr metric with mass $M = m_1 + m_2$ and spin \mathbf{S}_{Kerr} , the deformation parameter being the symmetric mass ratio $\nu = \mu/M$. In the mid 2000s, Buonanno, Chen and Damour [38] modeled EOB inspiral, merger and ringdown waveforms including for the first time spinning, precessing effects.

In this paper we build on Refs. [25, 38], and also on the most recent analytical work [39–41] and the work at the interface between numerical and analytical relativity [11, 13], and develop a general procedure to generate spinning, precessing waveforms in the EOB approach. The procedure employs the precessing convention of Ref. [25] and an inertial frame aligned with the spin of the final BH. As a first application, we construct spinning, precessing waveforms that contain only the $\ell = 2$ gravitational mode and reduce to the nonprecessing waveforms calibrated to numerical-relativity (NR) waveforms [42–45] in Ref. [13]. We compare these EOB precessing waveforms to Taylor-expanded PN waveforms and to two NR waveforms recently produced by the Caltech-Cornell-CITA collaboration [46].

The paper is organized as follows. In Sec. II we discuss the main coordinate frames that are used to describe precessing waveforms and their physical characteristics. We also review different proposals that have been suggested in the analytical and numerical-relativity communities for the precessing source frame, in which precession-induced modulations in the waveform’s phase and amplitude are minimized. We also study how spin

components parallel to the orbital plane affect the energy flux and multipolar waveforms. In Sec. III we build EOB precessing waveforms using a precessing source frame aligned with the Newtonian orbital angular momentum and an inertial frame aligned with the direction of the final BH spin. In Sec. IV we compare EOB precessing waveforms computed in different precessing source frames and carry out comparisons between EOB, Taylor-expanded PN and NR precessing waveforms. Section V summarizes our main conclusions and future work.

II. MODELING PRECESSING WAVEFORMS

A. Conventions and inertial frames

Throughout the paper, we adopt geometric units $G = c = 1$ and the Einstein summation convention, unless otherwise specified. The masses of the BHs are m_1 and m_2 and we choose $m_1 \geq m_2$. The total mass, mass ratio and symmetric mass ratio are $M = m_1 + m_2$, $q = m_1/m_2$ and $\nu = m_1 m_2 / M^2$, respectively. The position, linear momentum and spin vectors of the BHs are $\mathbf{r}_i(t)$, $\mathbf{p}_i(t)$ and $\mathbf{S}_i(t) = \chi_i m_i^2 \hat{\mathbf{S}}_i$, where $i = 1, 2$ and χ_i is the dimensionless spin magnitude. In the EOB framework, we solve the time evolution of the relative (rescaled) position vector $\mathbf{r}(t) \equiv (\mathbf{r}_1(t) - \mathbf{r}_2(t)) / M$, the center-of-mass-frame (rescaled) momentum vector $\mathbf{p}(t) \equiv \mathbf{p}_1(t) / \mu = -\mathbf{p}_2(t) / \mu$, and the spins variables $\mathbf{S}_1(t)$ and $\mathbf{S}_2(t)$.

We start with an arbitrary orthonormal basis $\{\mathbf{e}_x, \mathbf{e}_y, \mathbf{e}_z\}$. Without loss of generality, we align the initial relative position vector \mathbf{r}_0 with \mathbf{e}_x and the initial orbital orientation $[\hat{\mathbf{L}}_N]_0 \equiv \hat{\mathbf{L}}_N(0) \equiv \mathbf{r}_0 \times \dot{\mathbf{r}}_0 / |\mathbf{r}_0 \times \dot{\mathbf{r}}_0|$ with \mathbf{e}_z , where we indicate with an over-dot the time derivative and $\dot{\mathbf{r}}_0 = \dot{\mathbf{r}}(0)$ is the initial relative velocity. The initial spin directions are specified by the spherical coordinates $\theta_1^S, \phi_1^S, \theta_2^S$ and ϕ_2^S associated with this basis.

In the nonprecessing case, $\hat{\mathbf{L}}_N$ is a constant and it is natural to choose a *source frame* described by the (unit) basis vectors $\{\mathbf{e}_1^S, \mathbf{e}_2^S, \mathbf{e}_3^S\}$ whose \mathbf{e}_3^S is aligned with $\hat{\mathbf{L}}_N$. In the precessing case, we choose to align the source frame $\{\mathbf{e}_1^S, \mathbf{e}_2^S, \mathbf{e}_3^S\}$ with $\{\mathbf{e}_x, \mathbf{e}_y, \mathbf{e}_z\}$. The basis vector \mathbf{e}_3^S is aligned with the initial orbital orientation $\hat{\mathbf{L}}_N(0)$ but not with $\hat{\mathbf{L}}_N(t)$ at later times.

The GW polarizations h_+ and h_\times can be obtained by projecting the strain tensor h_{ij} onto the *radiation frame* described by the (unit) basis vectors $\{\mathbf{e}_1^R, \mathbf{e}_2^R, \mathbf{e}_3^R \equiv \hat{\mathbf{N}}\}$, the basis vector $\hat{\mathbf{N}}$ being along the direction of propagation of the wave (see Fig. 1). That is

$$h_+ = \frac{1}{2} [\mathbf{e}_1^R \otimes \mathbf{e}_1^R - \mathbf{e}_2^R \otimes \mathbf{e}_2^R]^{ij} h_{ij}, \quad (1a)$$

$$h_\times = \frac{1}{2} [\mathbf{e}_1^R \otimes \mathbf{e}_2^R + \mathbf{e}_2^R \otimes \mathbf{e}_1^R]^{ij} h_{ij}, \quad (1b)$$

where the basis vectors \mathbf{e}_1^R and \mathbf{e}_2^R are defined by (see

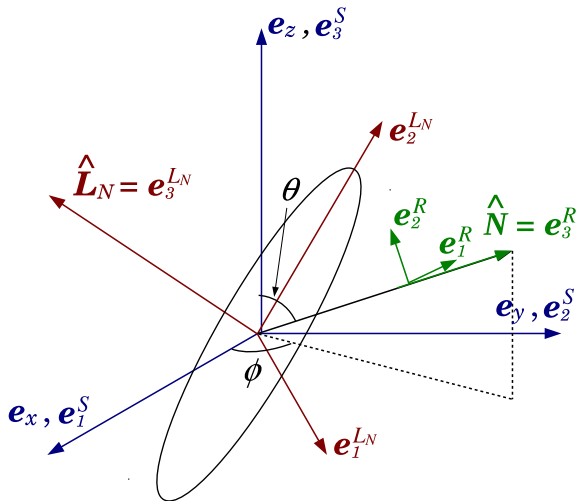


FIG. 1: We show the radiation frame $\{e_1^R, e_2^R, e_3^R\}$, the inertial source frame $\{e_1^S, e_2^S, e_3^S\}$ and the precessing source frame $\{e_1^{L_N}, e_2^{L_N}, e_3^{L_N}\}$ employed to describe a precessing BH binary and its GW radiation.

Fig. 1)

$$e_1^R \equiv \frac{e_z \times \hat{N}}{|e_z \times \hat{N}|}, \quad (2a)$$

$$e_2^R \equiv \hat{N} \times e_1^R. \quad (2b)$$

In the source frame $\{e_1^S, e_2^S, e_3^S\}$, we can decompose the polarizations h_+ and h_\times in -2 spin-weighted spherical harmonics ${}_{-2}Y_{\ell m}(\theta, \phi)$ as

$$h_+(\theta, \phi) - ih_\times(\theta, \phi) = \sum_{\ell=2}^{\infty} \sum_{m=-\ell}^{\ell} {}_{-2}Y_{\ell m}(\theta, \phi) h_{\ell m}. \quad (3)$$

The modes $h_{\ell m}$ can be calculated by applying the orthogonality condition valid for the ${}_{-2}Y_{\ell m}(\theta, \phi)$'s. Thus

$$h_{\ell m} = \int [h_+(\theta, \phi) - ih_\times(\theta, \phi)] {}_{-2}Y_{\ell m}^*(\theta, \phi) d\Omega, \quad (4)$$

where θ and ϕ are the inclination and azimuthal angles of the unit vector \hat{N} as measured in the source frame. Note that in the above expressions, we omit the dependence of the GW polarizations on time and binary parameters.

B. Precessing source frames in analytical-relativity waveforms

In the nonprecessing case, the orbital orientation is constant and coincides with the directions of the orbital

angular momentum $\mathbf{L} \equiv \mu M \mathbf{r} \times \mathbf{p}$, the Newtonian angular momentum $\mathbf{L}_N \equiv \mu M \mathbf{r} \times \dot{\mathbf{r}}$, and the total angular momentum $\mathbf{J} \equiv \mathbf{L} + \mathbf{S}_1 + \mathbf{S}_2$. If the source frame is aligned with the orbital orientation, the gravitational polarizations are quite simple and are described mainly by the (2, 2) mode and a few subdominant modes [12, 13]. In this case, the wave's amplitude and frequency increase monotonically during the inspiral and plunge stages, and the amplitudes of the GW modes display a clean hierarchy. By contrast, precessing waveforms decomposed in an inertial source frame show strong amplitude and phase modulations. In this case the amplitudes of the GW modes do not necessarily follow a clean hierarchy [47]. Ideally we would like to conduct comparisons and calibrations between analytical and numerical waveforms in a time dependent source frame that minimizes precession-induced modulations. Fortunately, this is possible if we choose a source frame that precesses with the binary orbital plane [25, 27–30].

Buonanno, Chen and Vallisneri [25] proposed the *precessing convention* that neatly disentangles precessional effects from both amplitude and phase evolutions in restricted (i.e., leading-order) PN waveforms. In the precessing convention, the precessing waveform is written as the product of a nonprecessing carrier waveform and a modulation term that collects all precessional effects. In Refs. [25, 27] the authors chose the precessing source frame aligned with the Newtonian orbital angular momentum \mathbf{L}_N . In this case, the basis vectors of the precessing source frame, $\{e_1^{L_N}(t), e_2^{L_N}(t), e_3^{L_N}(t)\}$ in Fig. 1, read [25]

$$e_3^{L_N}(t) = \hat{\mathbf{L}}_N(t), \quad (5a)$$

$$\dot{e}_1^{L_N}(t) = \boldsymbol{\Omega}_e(t) \times e_1^{L_N}(t), \quad (5b)$$

$$\dot{e}_2^{L_N}(t) = \boldsymbol{\Omega}_e(t) \times e_2^{L_N}(t), \quad (5c)$$

where

$$\boldsymbol{\Omega}_e(t) \equiv \boldsymbol{\Omega}_L(t) - [\boldsymbol{\Omega}_L(t) \cdot \hat{\mathbf{L}}_N(t)] \hat{\mathbf{L}}_N(t) = \dot{\hat{\mathbf{L}}}_N(t) \times \hat{\mathbf{L}}_N(t), \quad (6)$$

and $\boldsymbol{\Omega}_L(t)$ is the angular velocity of the precession of $\hat{\mathbf{L}}_N(t)$ and satisfies $\dot{\hat{\mathbf{L}}}_N(t) = \boldsymbol{\Omega}_L(t) \times \hat{\mathbf{L}}_N(t)$. Aligning the precessing source frame with $\hat{\mathbf{L}}_N(t)$ in Eq. (5a) removes the precession-induced amplitude modulations. Intuitively, Eqs. (5b)–(6) impose that $e_1^{L_N}(t)$ and $e_2^{L_N}(t)$ follow the precession of $e_3^{L_N}(t) = \hat{\mathbf{L}}_N(t)$, but do not precess around it. The key point of the precessing convention is the removal of all precession-induced modulations from the orbital phase $\Phi(t)$, so that $\Phi(t)$ is simply given by the integral of the (monotonic) orbital frequency Ω , i.e. $\Phi(t) = \int \Omega(t') dt'$ (see for details Sec. IVA in Ref. [25]). The freedom of choosing the constant of integration, or the initial phase, is degenerate with the only degree of freedom left in defining the precessing source frame, namely a constant rotation of $e_1^{L_N}$ and $e_2^{L_N}$ around $e_3^{L_N}$.

We want to test now the precessing convention on inspiraling PN waveforms computed beyond the restricted

approximation, i.e., beyond leading order. We employ the waveforms of Ref. [47] that have spin-amplitude corrections through 1.5PN order. We decompose the $h_{\ell m}$'s in two source frames: (i) the inertial source frame aligned with the initial total angular momentum \mathbf{J}_0 [47] and (ii) the precessing source frame $\{\mathbf{e}_1^{L_N}(t), \mathbf{e}_2^{L_N}(t), \mathbf{e}_3^{L_N}(t)\}$ defined by Eqs. (5a)–(6). The waveforms decomposed in the \mathbf{J}_0 -frame are given explicitly in Appendix B of Ref. [47]. We calculate the waveforms decomposed in the precessing $\mathbf{L}_N(t)$ -frame from the waveforms decomposed in the \mathbf{J}_0 -frame by properly rotating the $h_{\ell m}$ modes.

In general, given a set of spin-weighted spherical harmonics $h_{\ell m}^{(o)}$ decomposed in an original frame and the Euler angles (α, β, γ) that define the rotation from the original frame to a new frame, the modes $h_{\ell m}^{(n)}$ decomposed in the new frame are given by [47, 48]

$$h_{\ell m}^{(n)} = \sum_{m'=-\ell}^{\ell} D_{m'm}^{\ell*}(\alpha, \beta, \gamma) h_{\ell m'}^{(o)}, \quad (7)$$

where $D_{m'm}^{\ell*}(\alpha, \beta, \gamma)$ is the complex conjugate of the Wigner D -matrix

$$D_{m'm}^{\ell}(\alpha, \beta, \gamma) = (-1)^{m'} \sqrt{\frac{4\pi}{2\ell+1}} {}_{-m'}Y_{\ell m}(\beta, \alpha) e^{im'\gamma}, \quad (8)$$

where ${}_{-m'}Y_{\ell m}$ is the spherical harmonic of spin-weight $-m'$. The transformation is closed among modes with the same index ℓ . In this paper, we focus on the $\ell = 2$ modes both for simplicity and because even when precession is present the $\ell = 2$ modes still dominate. Nevertheless, the $\ell > 2$ modes are not negligible and we plan to extend the precessing EOB model to those modes in the future, following the same approach we propose and demonstrate here with the $\ell = 2$ modes.

In Fig. 2, we compare the \mathbf{J}_0 -frame and $\mathbf{L}_N(t)$ -frame h_{2m} inspiraling waveforms emitted by a BH binary with mass ratio $q = 6$ and spin magnitudes $\chi_1 = \chi_2 = 0.8$. We choose spin orientations that give strong precession-induced modulations. As we can see, there is a clear hierarchy among the h_{2m} 's amplitudes when decomposed in the $\mathbf{L}_N(t)$ -frame, but not when the decomposition is done in the \mathbf{J}_0 -frame. In fact, in the \mathbf{J}_0 -frame, the $(2, 1)$ and $(2, 0)$ modes have even larger amplitudes than the $(2, 2)$ mode. We notice that the strong amplitude and phase modulations of the modes in the \mathbf{J}_0 -frame almost completely disappear when the $\mathbf{L}_N(t)$ -frame is used.

Furthermore, we find it interesting to compare the modes of the precessing waveforms to the ones of a “nonprecessing” binary system having the same mass ratio and $\chi_i(t) \equiv \chi_i \hat{\mathbf{S}}_i(t) \cdot \hat{\mathbf{L}}_N(t)$, ($i = 1, 2$). That is to say, we keep only the components of the spin vectors along the direction of the Newtonian angular momentum $\mathbf{L}_N(t)$, and set all the other components to zero. For convenience, we refer to such “nonprecessing” waveforms as the *quasi-nonprecessing* waveforms. We use the adjective “quasi” because, differently from the nonprecessing waveforms,

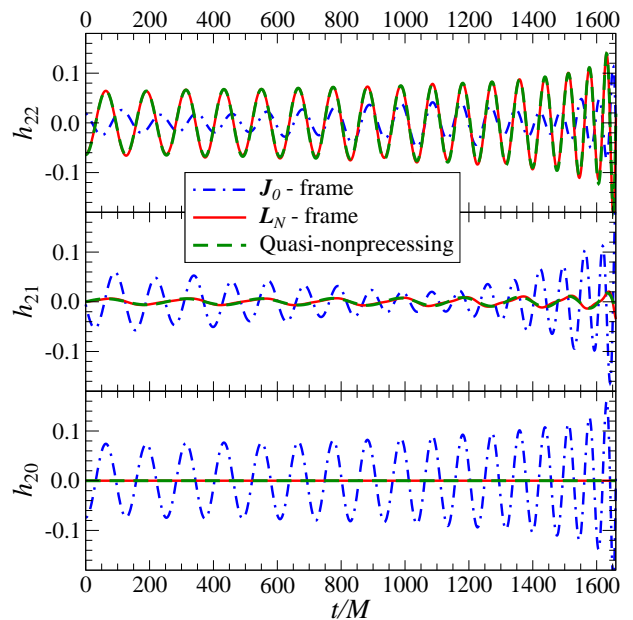


FIG. 2: We show inspiraling, precessing PN waveforms decomposed in the inertial source frame aligned with the initial total angular momentum \mathbf{J}_0 and in the precessing source frame aligned with the Newtonian orbital angular momentum $\mathbf{L}_N(t)$. For comparison, we show also the quasi-nonprecessing PN waveforms defined in Sec. II B. The three panels use the same scale on the y -axis so that the amplitudes of the modes h_{22} , h_{21} and h_{20} can be easily compared.

where the spins are aligned or antialigned with the orbital angular momentum and remain constant throughout the evolution, in the quasi-nonprecessing waveforms the spins evolve (according to Eqs. (11c)–(11d) given below) and over time change their projections onto \mathbf{L}_N . As can be seen in Fig. 2, the near-perfect agreement between $\mathbf{L}_N(t)$ -frame and quasi-nonprecessing waveforms indicates that the spin components along $\mathbf{L}_N(t)$ dominate the spin effects. This conjecture is reinforced by the observation that, because of parity symmetry, the spin-orbit couplings contribute to the GW energy flux to infinity (known through 3.5PN order) only through terms of the form $\mathbf{S}_i(t) \cdot \hat{\mathbf{L}}_N(t)$, ($i = 1, 2$) [20, 21, 49, 50]. The energy flux is a frame independent quantity. It is given, in the adiabatic assumption, by ²

$$\frac{dE}{dt} = \frac{(M\Omega)^2}{8\pi} \sum_{\ell=2}^{\infty} \sum_{m=-\ell}^{\ell} m^2 \left| \frac{\mathcal{R}}{M} h_{\ell m} \right|^2, \quad (9)$$

where \mathcal{R} is the distance to the source (and simply cancels the dependence on \mathcal{R} hidden in the $h_{\ell m}$'s). The fact

² Following the nonprecessing EOB model of Ref. [13], we include in dE/dt the spin-orbit terms through 1.5PN order [51], even if some of us have recently computed the spin-orbit and spin-spin terms in the factorized flux through 2PN order using results in Ref. [22].

that the spin-orbit effects in dE/dt depend on spins only through $\mathbf{S}_i(t) \cdot \hat{\mathbf{L}}_N(t)$, ($i = 1, 2$) suggests that the dependence of the modes on spin's components parallel to the orbital plane disappears when all the modes are summed up to make the total energy flux.

Therefore, beyond the leading-order results of Ref. [25], we find that PN precessing waveforms [47] reduce to nearly nonprecessing waveforms when decomposed to spin-weighted spherical harmonics in the source frame $\{e_1^{L_N}(t), e_2^{L_N}(t), e_3^{L_N}(t)\}$ of the precessing convention [25]. In addition, the PN quasi-nonprecessing waveforms are a good first approximations to the $h_{\ell m}$'s modes in the precessing frame.

C. Precessing source frames in numerical-relativity waveforms

The possibility of demodulating precessing waveforms using precessing source frames was recently verified with and generalized to NR waveforms in Refs. [28, 29, 32]. In particular, Schmidt *et al.* [28, 32] and O'Shaughnessy *et al.* [29] identified the preferred *radiation axis* at infinity and showed that if a precessing frame aligned with the radiation axis is chosen, then the amplitude and phase modulations of numerical waveforms are removed and a clean hierarchy among the modes is restored.

In particular, Schmidt *et al.* proposed the so-called quadrupole-preferred frame in which the power of the $(\ell, \pm m) = (2, 2)$ mode is maximized. O'Shaughnessy *et al.* proposed a more general and geometrical choice of the precessing frame in which the z component of the radiated angular momentum is maximized. The latter proposal reduces to the choice of Schmidt *et al.* when the radiated angular momentum is calculated using only the $(\ell, \pm m) = (2, 2)$ modes. Boyle *et al.* [30] then proposed the *minimal rotation* condition to remove the remaining arbitrariness in the azimuthal rotation of the precessing frame and in the phase modulations of the waveform. Given an inertial frame $\{e_x, e_y, e_z\}$ and the first two Euler angles $\alpha(t)$ and $\beta(t)$ that align e_z with the radiation axis, the minimal rotation condition on the third Euler angle $\gamma(t)$ is given by

$$\dot{\gamma}(t) = -\dot{\alpha}(t) \cos \beta(t). \quad (10)$$

This condition is equivalent to Eqs. (5b)–(6) above on the evolution of $e_1^{L_N}(t)$ and $e_2^{L_N}(t)$. If $\alpha(t)$ and $\beta(t)$ are the first two Euler angles that align e_z with $\hat{\mathbf{L}}_N$, then $\gamma(t) - \gamma(t_0)$ is the angle by which $e_1^{L_N}(t)$ and $e_2^{L_N}(t)$ shall rotate in the instantaneous orbital plane, relative to their positions at a reference time t_0 , to satisfy the precession convention. Recently, Boyle [31] proposed a geometric definition of the angular-velocity vector of a waveform, to determine a frame in which the modes' amplitudes become very simple and the phases are nearly constant.

Schmidt *et al.* [32] showed that precessing PN inspiral waveforms computed in the precessing source frame

aligned with the preferred radiation axis are well approximated by nonprecessing PN waveforms. Furthermore, they proposed that precessing waveforms can be generated with good accuracy by transforming nonprecessing waveforms from precessing source frames to inertial source frames. In a recent study, Pekowsky *et al.* [33] studied the mapping of precessing waveforms to nonprecessing ones using a large number of (short) numerical simulations and the analytical IMRPhenomB [17] waveforms. They found that precessional degrees of freedom that cannot be reproduced by nonprecessing models (such as spin's components perpendicular to $\hat{\mathbf{L}}_N$) give rise to corrections to the nonprecessing waveforms that are very small during inspiral, but they can become significant during merger and ringdown.

D. Strategy to build precessing effective-one-body waveforms

Motivated by the results discussed in Secs. II B and II C of a nearly complete separation of precession-induced modulations in precessing waveforms when using appropriate precessing source frames, we propose the following approach to generate generic EOB waveforms.

First, we evolve the EOB dynamics and solve Eqs. (5a)–(6) for the precessing source frame $\{e_1^{L_N}(t), e_2^{L_N}(t), e_3^{L_N}(t) = \hat{\mathbf{L}}_N(t)\}$. Since the difference between \mathbf{L}_N and \mathbf{L} starts at 1PN order, the leading-order conclusions achieved by the precessing \mathbf{L}_N -frame hold if we replace \mathbf{L}_N with \mathbf{L} in Eqs. (5a)–(6). We have verified that precessing waveforms decomposed in the \mathbf{L} -frame agree equally well with the quasi-nonprecessing waveforms generated by keeping only spin's components along \mathbf{L} . Furthermore, in Sec. IV we compare precessing EOB waveforms (generated either in the \mathbf{L}_N -frame or in the \mathbf{L} -frame) to NR waveforms, and find that their mutual difference is marginal. Without a more accurate calibration and comprehensive comparisons with NR waveforms, we do not know *a priori* whether the \mathbf{L}_N -frame or the \mathbf{L} -frame is more preferable, nor can we say which of them captures the precession effects more faithfully. Thus, at the current stage, we simply adopt the \mathbf{L}_N -frame as the default precessing source frame in the EOB model.

Second, because of the simple features of the inspiral-plunge modes in the precessing source frame — little modulation and clean hierarchy — we choose to model the precessing inspiral-plunge EOB modes in this frame, and generate modes in any arbitrary source frame through Eq. (7). Since factorized EOB modes for precessing spins are not available yet and EOB modes have been calibrated only to nonspinning and spinning, nonprecessing NR modes [12, 13], we choose to work in the precessing source frame and use quasi-nonprecessing modes as good approximations to precessing modes (as discussed in Secs. II B and II C). In particular, we employ the quasi-nonprecessing inspiral-plunge modes based on the latest

spinning, nonprecessing EOB model that was calibrated to NR modes in Ref. [13]. Note that we are not obliged to use in the future quasi-nonprecessing waveforms in the precessing source frame. As soon as factorized EOB modes for precessing spins become available, we shall relax the assumption of using quasi-nonprecessing inspiral-plunge modes³. The strategy that we present in this paper is generic and can easily be applied to future calibrations and analytical improvements of the EOB model.

Third, we rotate the quasi-nonprecessing modes from the precessing source frame to the inertial frame whose z -axis coincides with the direction of the total angular momentum \mathbf{J} at a time very close to merger when the direction of \mathbf{J} is a good approximation to the direction of the spin of the final BH. In this inertial frame we match the inspiral-plunge to merger-ringdown modes following the usual prescription in the EOB approach [13]. After generating inspiral-merger-ringdown modes in this frame, it is straightforward to calculate EOB modes $h_{\ell m}$ in any source frame or EOB polarizations $h_{+, \times}$ in any radiation frame.

III. PRECESSING EFFECTIVE-ONE-BODY MODEL

In this section we construct a generic, precessing EOB model following the general strategy outlined above — it employs the precessing source frame introduced in Ref. [25] and the quasi-nonprecessing waveforms based on the nonprecessing EOB model developed in Ref. [13].

A. Effective-one-body dynamics

Since we employ exactly the same EOB dynamics calibrated against NR simulations in Ref. [13], we review only the key ingredients of the dynamics and refer the readers to Ref. [13] for further details.

The EOB dynamics of spinning BH binary systems is obtained solving the following Hamilton equations

$$\frac{d\mathbf{r}}{d\hat{t}} = \{\mathbf{r}, \hat{H}_{\text{real}}\} = \frac{\partial \hat{H}_{\text{real}}}{\partial \mathbf{p}}, \quad (11a)$$

$$\frac{d\mathbf{p}}{d\hat{t}} = \{\mathbf{p}, \hat{H}_{\text{real}}\} + \hat{\mathcal{F}} = -\frac{\partial \hat{H}_{\text{real}}}{\partial \mathbf{r}} + \hat{\mathcal{F}}, \quad (11b)$$

$$\frac{d\mathbf{S}_1}{dt} = \{\mathbf{S}_1, \mu \hat{H}_{\text{real}}\} = \mu \frac{\partial \hat{H}_{\text{real}}}{\partial \mathbf{S}_1} \times \mathbf{S}_1, \quad (11c)$$

$$\frac{d\mathbf{S}_2}{dt} = \{\mathbf{S}_2, \mu \hat{H}_{\text{real}}\} = \mu \frac{\partial \hat{H}_{\text{real}}}{\partial \mathbf{S}_2} \times \mathbf{S}_2, \quad (11d)$$

where $\hat{t} \equiv t/M$ is the dimensionless time variable, \hat{H}_{real} is the reduced EOB Hamiltonian derived in Refs. [39–41] and reviewed in Sec. IIA of Ref. [13], and $\hat{\mathcal{F}}$ is the reduced radiation reaction force. Following Ref. [38], we use

$$\hat{\mathcal{F}} = \frac{1}{\nu \hat{\Omega} |\mathbf{r} \times \mathbf{p}|} \frac{dE}{dt} \mathbf{p}, \quad (12)$$

where $\hat{\Omega} \equiv M|\mathbf{r} \times \dot{\mathbf{r}}|/r^2$ is the dimensionless orbital frequency and dE/dt is the energy flux for quasi-spherical orbits. We use Eq. (9) with $\ell \leq 8$, namely

$$\frac{dE}{dt} = \frac{\hat{\Omega}^2}{8\pi} \sum_{\ell=2}^8 \sum_{m=-\ell}^{\ell} m^2 \left| \frac{\mathcal{R}}{M} h_{\ell m} \right|^2. \quad (13)$$

Because under a change of frame the $h_{\ell m}$ modes for a given ℓ transform into modes with the same ℓ , Eq. (13) is still frame-independent. We insert the quasi-nonprecessing modes $h_{\ell m}$, i.e., the modes decomposed in the precessing source frame aligned with $\mathbf{L}_N(t)$, into Eq. (13). The quasi-nonprecessing modes can be calculated directly in the inertial frame $\{\mathbf{e}_x, \mathbf{e}_y, \mathbf{e}_z\}$ in which we solve the EOB dynamics. The only difference from the procedure of Ref. [13] is to replace the constant χ_1 and χ_2 with $\chi_1 \hat{\mathbf{S}}_1(t) \cdot \hat{\mathbf{L}}_N(t)$ and $\chi_2 \hat{\mathbf{S}}_2(t) \cdot \hat{\mathbf{L}}_N(t)$.

B. Initial conditions

For applications in data analysis and comparisons with numerical or analytical waveforms, we need initial conditions that start the orbital evolution with sufficiently small eccentricity at a given orbital separation (or GW frequency) and spins orientation. The analytical quasi-spherical initial conditions proposed in Ref. [38] is a good first approximation.

In the nonprecessing case [11–13], the residual eccentricity can be further reduced by starting the evolution at a larger separation (smaller GW frequency) and waiting for the orbits to be better circularized by radiation reaction. In the precessing case, however, we can not easily reduce the eccentricity in this way because we need specific spin directions at the initial separation. In order to reduce eccentricity by starting the evolution at a larger separation, we need to figure out what are the spin directions at this larger separation to ensure the desired spin directions at a given (smaller) initial separation. To reach this goal and reduce the eccentricity for quasi-spherical initial conditions we employ the method developed in Ref. [52], which is based on [43, 53]⁴. Thus, we first evolve the binary for a few orbits and estimate the eccentricity through oscillations in orbital frequency

³ It remains to be investigated, though, whether it is necessary to include precessing effects in the EOB modes decomposed in the precessing source frame to meet more stringent accuracy requirements for advanced LIGO and Virgo searches.

⁴ This method has been employed to reduce eccentricity in NR simulations of BH binary systems [44, 45, 54–57].

Ω and separation r . We then apply corrections to the initial conditions following Eqs. (74) and (75) in Ref. [52]. We repeat these steps until the eccentricity is sufficiently small.

C. Nonprecessing effective-one-body waveforms

The EOB nonprecessing inspiral-plunge modes $h_{\ell m}^{\text{NP, insp-plunge}}$ developed in Ref. [13] are given by

$$h_{\ell m}^{\text{NP, insp-plunge}} = h_{\ell m}^{\text{F}} N_{\ell m}, \quad (14)$$

where $h_{\ell m}^{\text{F}}$ are the factorized modes derived in Refs. [51, 58, 59], and $N_{\ell m}$ are nonquasicircular (NQC) corrections that model deviations from the quasicircular motion, that is assumed when deriving $h_{\ell m}^{\text{F}}$. The factorized modes read

$$h_{\ell m}^{\text{F}} = h_{\ell m}^{(N, \epsilon)} \hat{S}_{\text{eff}}^{(\epsilon)} T_{\ell m} e^{i\delta_{\ell m}} (\rho_{\ell m})^\ell, \quad (15)$$

where ϵ is the parity of the mode. All the factors entering $h_{\ell m}^{\text{F}}$ can be explicitly found in the Appendix of Ref. [13]. As discussed above, when using these expressions to generate quasi-nonprecessing modes, the only minor modification we have to take into account is to replace the constant spin magnitudes χ_1 and χ_2 by their time dependent counterparts. More specifically, in the nonprecessing case, the leading order spin-orbit effects in $\rho_{\ell m}$ are parametrized by two linear combinations of the constant dimensionless spin parameters

$$\chi_S \equiv \frac{\chi_1 + \chi_2}{2}, \quad (16a)$$

$$\chi_A \equiv \frac{\chi_1 - \chi_2}{2}. \quad (16b)$$

In the precessing case, both χ_S and χ_A become linear combinations of the time varying spin vectors projected along $\hat{\mathbf{L}}_N(t)$,

$$\chi_S(t) \equiv \frac{1}{2} \left(\frac{\mathbf{S}_1(t)}{m_1^2} + \frac{\mathbf{S}_2(t)}{m_2^2} \right) \cdot \hat{\mathbf{L}}_N(t), \quad (17a)$$

$$\chi_A(t) \equiv \frac{1}{2} \left(\frac{\mathbf{S}_1(t)}{m_1^2} - \frac{\mathbf{S}_2(t)}{m_2^2} \right) \cdot \hat{\mathbf{L}}_N(t). \quad (17b)$$

In Ref. [13], the inspiral-merger-ringdown mode (2, 2) was calibrated against NR simulations. Studies in the test-particle nonspinning [59] and spinning, nonprecessing [51] cases suggest that the factorized modes $h_{\ell m}^{\text{F}}$ are good approximations of the inspiral-plunge modes even without any NQC correction or calibration. Thus, we model the inspiral-plunge (2, 1) mode with h_{21}^{F} . The (2, 0) mode has been computed in PN theory at leading order and its amplitude is $5/14\sqrt{6} \simeq 0.15$ times the amplitude of the leading order (2, 2) mode [60]. However, this prediction does not agree with NR results. In fact, we find [61] that for the nonspinning NR simulations of mass ratios $q = 1, 6$ [12], the amplitude of the (2, 0) mode

during the inspiral is smaller than the one of the (2, 2) mode by a factor $\sim 10^3$. Since we do not yet understand the origin of this discrepancy we have decided that in this first investigation we neglect the nonprecessing EOB (2, 0) mode.

The NQC correction to the (2, 2) mode, N_{22} , is given by

$$N_{22} = \left[1 + \left(\frac{p_{r^*}}{r \hat{\Omega}} \right)^2 \left(a_1^{h_{22}} + \frac{a_2^{h_{22}}}{r} + \frac{a_3^{h_{22}}}{r^{3/2}} + \frac{a_4^{h_{22}}}{r^2} + \frac{a_5^{h_{22}}}{r^{5/2}} \right) \right] \\ \times \exp \left[i \frac{p_{r^*}}{r \hat{\Omega}} \left(b_1^{h_{22}} + p_{r^*}^2 b_2^{h_{22}} + \frac{p_{r^*}^2}{r^{1/2}} b_3^{h_{22}} + \frac{p_{r^*}^2}{r} b_4^{h_{22}} \right) \right], \quad (18)$$

where the amplitude coefficients $a_i^{h_{22}}$ (with $i = 1..5$) and the phase coefficients $b_i^{h_{22}}$ (with $i = 1..4$) are obtained through the iterative procedure described in Sec. IIB of Ref. [13]. Since only equal-mass, equal-spin, nonprecessing NR simulations were used to calibrate the EOB model of Ref. [13], we have to map the N_{22} from generic spin configurations to equal-spin, nonprecessing configurations. Without further calibrations, we first adopt a mapping from precessing to nonprecessing configurations that equates the $\chi_S(0)$ and $\chi_A(0)$ of a precessing configuration (defined in Eqs. (17a) and (17b)) to the constant χ_S and χ_A of a nonprecessing configuration. Then, we apply the mapping from a generic nonprecessing configuration to an equal-spin, nonprecessing configurations as defined in Sec. IVA of Ref. [13].

D. Precessing source frame

Let $\mathbf{e}_3^{L_N}(t) = \hat{\mathbf{L}}_N(t)$ be the third (unit) basis vector of the precessing source frame. We solve the other two (unit) basis vectors $\mathbf{e}_1^{L_N}(t)$ and $\mathbf{e}_2^{L_N}(t)$ by applying the minimal rotation condition. We do it because the latter involves only one differential equation, namely Eq. (10), instead of Eqs. (5b) and (5c) for the precessing convention. Specifically, with the help of the inertial source frame $\{\mathbf{e}_x^S, \mathbf{e}_y^S, \mathbf{e}_z^S\}$, we define

$$\alpha(t) = \arctan \left[\frac{\mathbf{e}_3^{L_N}(t) \cdot \mathbf{e}_y^S}{\mathbf{e}_3^{L_N}(t) \cdot \mathbf{e}_x^S} \right], \quad (19a)$$

$$\beta(t) = \arccos \left[\mathbf{e}_3^{L_N}(t) \cdot \mathbf{e}_z^S \right], \quad (19b)$$

and solve ⁵

$$\dot{\gamma}(t) = -\dot{\alpha}(t) \cos \beta(t). \quad (20)$$

⁵ Following Boyle *et al.* [30], we integrate $\gamma(t)$ by parts and implement $\gamma(t) = -\alpha(t) \cos \beta(t) - \int \alpha(t) \dot{\beta}(t) \sin \beta(t) dt$ to avoid differentiating $\alpha(t)$, which can be noisy near the coordinate singularities at $\beta(t) = 0$ and $\beta(t) = \pi$. We note that Boyle [31] recently proposed a much more accurate and robust method to integrate $\gamma(t)$ using quaternions.

Those Euler angles $\alpha(t)$, $\beta(t)$ and $\gamma(t)$ describe the time-dependent rotation from the inertial source frame $\{\mathbf{e}_x^S, \mathbf{e}_y^S, \mathbf{e}_z^S\}$ to the precessing source frame $\{\mathbf{e}_1^{LN}(t), \mathbf{e}_2^{LN}(t), \mathbf{e}_3^{LN}(t)\}$ with the latter satisfying the minimal rotation condition. The only freedom in the definition of the precessing source frame is a constant shift in $\gamma(t)$ that is degenerate with the initial orbital phase.

E. Precessing effective-one-body waveforms

We build the complete inspiral-plunge-merger-ringdown waveforms in an inertial frame following the usual procedure in the EOB approach [6–14, 34, 35, 62]. More specifically, we join the inspiral-plunge waveform $h_{\ell m}^{\text{insp-plunge}}$ and the merger-ringdown waveform $h_{\ell m}^{\text{merger-RD}}$ at a matching time $t_{\text{match}}^{\ell m}$ as

$$h_{\ell m}^{\text{EOB}}(t) = h_{\ell m}^{\text{inspiral-plunge}}(t) \theta(t - t_{\text{match}}^{\ell m}) + h_{\ell m}^{\text{merger-RD}}(t) \theta(t - t_{\text{match}}^{\ell m}). \quad (21)$$

Given the quasi-nonprecessing inspiral-plunge modes $h_{\ell m}^{\text{NP, insp-plunge}}$ decomposed in the precessing source frame and the Euler angles (not necessarily those in Eqs. (19a)–(20), which are specific to $\{\mathbf{e}_x^S, \mathbf{e}_y^S, \mathbf{e}_z^S\}$) defining the rotation from the precessing source frame to any inertial frame, the inspiral-plunge modes in the inertial frame are given by Eqs. (7). To study the h_{2m} modes, we need all $\ell = 2$ modes in the precessing source frame. As discussed in Sec. III C, we employ the calibrated (2, 2) mode of Ref. [13], h_{21}^F for the (2, 1) mode, and zero for the (2, 0) mode. In the precessing source frame, since we use quasi-nonprecessing inspiral-plunge modes to approximate precessing modes, we further assume reflection symmetry, which, combined with parity invariance, gives modes with $m < 0$ through $h_{2, -m}^{\text{NP, insp-plunge}}(t) = h_{2m}^{\text{NP, insp-plunge}*}(t)$. Pekowsky *et al.* [33] discussed how this symmetry is broken by precessional effects, giving rise to a contribution to the (2, 2) mode which is odd under reflection. In the only example investigated in Ref. [33], the ratio between the component of the (2, 2) mode of the Weyl scalar $\Psi_{4,22}$ that is odd under reflection and the one that is even under reflection is ~ 0.01 , while the ratio between the former and the (2, 1) mode of the Weyl scalar $\Psi_{4,21}$ is ~ 1 . Since the Weyl scalar and the metric perturbation are related by $\Psi_{4,\ell m} = \ddot{h}_{\ell m} \simeq m^2 \hat{\Omega}^2 h_{\ell m}$ during the inspiral, the odd component of the h_{22} 's amplitude is about a fourth of the h_{21} 's. The odd component of the h_{22} 's amplitude becomes substantial only during the merger and ringdown. Thus, in this first study, we ignore the component of the (2, 2) mode that is odd under reflection when describing the inspiraling waveform in the precessing frame, but we include the odd component when building the merger-ringdown waveform.

It is convenient to choose an inertial frame in which the merger-ringdown waveforms take simple forms. A

natural choice is the frame aligned with the spin of the final BH, in which the merger-ringdown waveforms are expressed as linear combinations of the quasinormal modes (QNMs) [6–14, 34, 35, 62]. Barausse *et al.* [63] found strong evidence that the spin of the final BH is aligned with the initial total angular momentum of the binary. Using this assumption they derived accurate formulas for the final spin of a BH formed by merger. The success of their model verifies the PN-motivated assumption that the radiated angular momentum averaged over precessional cycles is almost aligned with the total angular momentum. Thus, the direction of the latter is preserved with high accuracy during the inspiral [23]. Here we employ the formulas in Ref. [63] to predict the magnitude of the spin of the final BH, and we align the final-spin direction with $\mathbf{J}(t_{\Omega_{\text{peak}}^{\text{EOB}}})$, which is the total angular momentum at the time the EOB orbital frequency reaches its peak ($t_{\Omega_{\text{peak}}^{\text{EOB}}}$). The time $t_{\Omega_{\text{peak}}^{\text{EOB}}}$ has been adopted in most previous EOB models as the reference time of merger [9–14]. Without further information from NR simulations of precessing, spinning BHs, we consider $\mathbf{J}(t_{\Omega_{\text{peak}}^{\text{EOB}}})$ our best prediction of the final-spin direction. We expect that not a lot of angular momentum is radiated during the swift transition from merger to ringdown [62] and the small amount being radiated is likely to be nearly aligned with $\mathbf{J}(t_{\Omega_{\text{peak}}^{\text{EOB}}})$.

The inspiral-plunge waveform $h_{\ell m}^{\text{insp-plunge}}$ in the inertial frame aligned with $\mathbf{J}(t_{\Omega_{\text{peak}}^{\text{EOB}}})$ contains NQC corrections from the nonprecessing (2, 2) mode $h_{\ell m}^{\text{NP, insp-plunge}}$. Those corrections are derived based on the assumption that the inspiral-plunge waveforms in the precessing frame are the calibrated nonprecessing waveforms generated with the specific mapping of spin parameters defined in Sec. III C. Although we expect that such assumption introduces systematic errors in $h_{\ell m}^{\text{insp-plunge}}$, we are not able to quantify them before comparing $h_{\ell m}^{\text{insp-plunge}}$ with precessing NR waveforms. Therefore, we do not apply any further correction to the inspiral-plunge waveform in this model. This choice also guarantees that $h_{\ell m}^{\text{insp-plunge}}$ modes reduce to the calibrated modes of Ref. [13] in the nonprecessing limit.

The merger-ringdown waveform is built following almost exactly the approach described in Ref. [13]. We first give a brief review of this approach and then describe the differences. The merger-ringdown waveform is modeled by a linear superposition of the QNMs of the final Kerr BH as

$$h_{\ell m}^{\text{merger-RD}}(t) = \sum_{n=0}^{N-1} A_{\ell mn} e^{-i\sigma_{\ell mn}(t - t_{\text{match}}^{\ell m})}, \quad (22)$$

where N is the number of overtones, $A_{\ell mn}$ is the complex amplitude of the n -th overtone of the (ℓ, m) mode, and $\sigma_{\ell mn}$ is the complex frequency of the n -th overtone. The complex frequencies are known function of the mass and spin of the final BH [64]. The mass of the final BH is given in Eq. (8) of Ref. [65]. The spin magnitude of the

final BH, as discussed earlier, is given in Eqs. (6), (8) and (10) of Ref. [63]. Following Ref. [13], we replace the highest physical overtone (the 7-th) of the (2, 2) mode with a pseudo QNM whose calibrated complex frequency is given in Eqs. (35a) and (35b) of Ref. [13]. Finally, we fix the complex amplitudes $A_{\ell mn}$ through a matching procedure [12] that imposes a C^1 -smooth connection over a time interval $\Delta t_{\text{match}}^{\ell m}$ between the merger-ringdown waveform and the inspiral-plunge waveform, in the inertial frame aligned with $\mathbf{J}(t_{\Omega_{\text{peak}}^{\text{EOB}}})$.

In Ref. [13], the matching time $t_{\text{match}}^{\ell m}$ and the time interval $\Delta t_{\text{match}}^{\ell m}$ were calibrated only for the (2, 2) mode. Here we need to specify those quantities also for the remaining $\ell = 2$ modes. We find that in order to keep the matching procedure stable when the binary is strongly precessing around merger, we have to introduce in $t_{\text{match}}^{\ell m}$ and $\Delta t_{\text{match}}^{\ell m}$ a dependence on how much the orbital and total angular momentum are misaligned at merger, i.e. on the quantity $\hat{\mathbf{L}}(t_{\Omega_{\text{peak}}^{\text{EOB}}}) \cdot \hat{\mathbf{J}}(t_{\Omega_{\text{peak}}^{\text{EOB}}})$. More specifically, in strongly precessing cases, the directions of $\mathbf{L}(t)$ and $\mathbf{J}(t)$ can be very different close to merger. As a consequence, the inspiral-plunge modes in the inertial $\mathbf{J}(t_{\Omega_{\text{peak}}^{\text{EOB}}})$ -frame can present strong amplitude and frequency oscillations around merger. [Technically those strong oscillations are generated by drastic time-dependent rotations from well-behaved quasi-nonprecessing inspiral-plunge modes as the merger is approached.] Thus, to keep the matching procedure stable in strongly precessing situations we set the matching time $t_{\text{match}}^{\ell m}$ earlier and make the matching interval $\Delta t_{\text{match}}^{\ell m}$ longer. We choose

$$t_{\text{match}}^{\ell m} = t_{\text{match}}^{22, \text{Cal}} - 10M (1 - |\kappa_{LJ}(t_{\Omega_{\text{peak}}^{\text{EOB}}})|), \quad (23)$$

$$\Delta t_{\text{match}}^{\ell m} = \Delta t_{\text{match}}^{22, \text{Cal}} (10 - 9|\kappa_{LJ}(t_{\Omega_{\text{peak}}^{\text{EOB}}})|), \quad (24)$$

where

$$t_{\text{match}}^{22, \text{Cal}} = t_{\Omega_{\text{peak}}^{\text{EOB}}} - \begin{cases} 2.5M & \chi \leq 0 \\ 2.5M + 1.77M \left(\frac{\chi}{0.437}\right)^4 & \chi > 0 \end{cases}, \quad (25)$$

$$\Delta t_{\text{match}}^{22, \text{Cal}} = 7.5M$$

are the calibrated values of the (2, 2) mode in Ref. [13],

$$\chi = \chi_S + \chi_A \frac{\sqrt{1 - 4\nu}}{1 - 2\nu} \quad (26)$$

is a linear combination of initial spin projections on \mathbf{L}_N , and

$$\kappa_{LJ}(t_{\Omega_{\text{peak}}^{\text{EOB}}}) = \hat{\mathbf{L}}(t_{\Omega_{\text{peak}}^{\text{EOB}}}) \cdot \hat{\mathbf{J}}(t_{\Omega_{\text{peak}}^{\text{EOB}}}) \quad (27)$$

is the cosine of the opening angle between the orbital and total angular momenta at the reference time of merger $t_{\Omega_{\text{peak}}^{\text{EOB}}}$. When $\kappa_{LJ}(t_{\Omega_{\text{peak}}^{\text{EOB}}}) = 0$, the matching time $t_{\text{match}}^{\ell m}$ is $10M$ earlier than that of the aligned case, and the time interval $\Delta t_{\text{match}}^{\ell m}$ is 10 times that of the aligned case. The choice of $10M$ and the factor of 10 made in this paper are rather arbitrary. They are based on the only requirement of producing qualitatively sound merger-ringdown waveforms.

Case	q	χ_1	χ_2	θ_1^S	θ_2^S	ϕ_1^S	ϕ_2^S	$M\Omega_0$
1	2	0.6	0.6	$\pi/3$	$\pi/3$	0	$\pi/2$	0.0112
2	6	0.8 ^a	0.6	$\pi/2$	$2\pi/3$	$\pi/2$	$\pi/2$	0.0112
3	3	0.500	0.499	0.499π	0.987π	0.767π	0.306π	0.0177
4	5	0.499	0	0.499π	0	-0.785π	0	0.0158

^aThe nonprecessing EOB model of Ref. [13] generates waveforms with any mass ratio and individual spin magnitudes $-1 \leq \chi_i \lesssim 0.7$. Although we consider here $\chi_1 = 0.8$ because $\chi_1(t) = \chi_1 \hat{\mathbf{S}}_1(t) \cdot \hat{\mathbf{L}}_N(t) < 0.7$ during the entire evolution, we do not find any problem in this case when generating the quasi-nonprecessing waveforms.

TABLE I: We list the binary parameters of the four precessing EOB waveforms that we consider in Sec. IV. Case 3 corresponds to the NR simulation SXS:BBH:0052 of Ref. [46], and case 4 corresponds to SXS:BBH:0058.

IV. COMPARISON BETWEEN PRECESSING WAVEFORMS

We generate four examples of EOB precessing waveforms using the model defined in Sec. III. The first two examples are a $q = 2$ BH binary system exhibiting moderate precession-induced modulations and a $q = 6$ binary system exhibiting strong modulations. The other two examples are chosen among the 171 NR simulations reported in Ref. [46] with the criterion of long and accurate waveforms exhibiting strong modulations. In these cases we compare NR, PN and EOB precessing waveforms. The physical parameters of the four binary configurations are listed in Table I.

A. Precessing and radiation-axis frames

In Sec. III D, we have proposed $\hat{\mathbf{L}}_N(t)$ and $\hat{\mathbf{L}}(t)$ as possible basis vectors for the precessing source frame. In this section, we compare their trajectories and the corresponding precessing waveforms generated through their respective precessing source frames. For convenience, we refer to waveforms generated in these precessing source frames as the $\hat{\mathbf{L}}_N(t)$ -frame and $\hat{\mathbf{L}}(t)$ -frame waveforms, respectively. Furthermore, we extract the quadrupole-preferred radiation axis [28] from the precessing waveforms and compare their trajectories with either $\hat{\mathbf{L}}_N(t)$ or $\hat{\mathbf{L}}(t)$.

Figure 3 shows for cases 1 and 2 of Table I the trajectories of the unit vectors $\hat{\mathbf{J}}$, $\hat{\mathbf{L}}$, $\hat{\mathbf{L}}_N$, $\hat{\mathbf{S}}_1$, and $\hat{\mathbf{S}}_2$ in the plane perpendicular to $\mathbf{J}(t_{\Omega_{\text{peak}}^{\text{EOB}}})$. In both cases, the BHs complete more than two cycles of precession and the directions of $\mathbf{J}(t)$ are well conserved during the entire inspiral phase. All other vectors precess around $\mathbf{J}(t)$. These are expected features of the well-known simple-precession picture of spinning binaries in PN theory [23]. Another common feature in both cases is the difference between the trajectories of $\mathbf{L}_N(t)$ and $\mathbf{L}(t)$, which has been pointed out in Ref. [66]. The trajectory of $\mathbf{L}_N(t)$

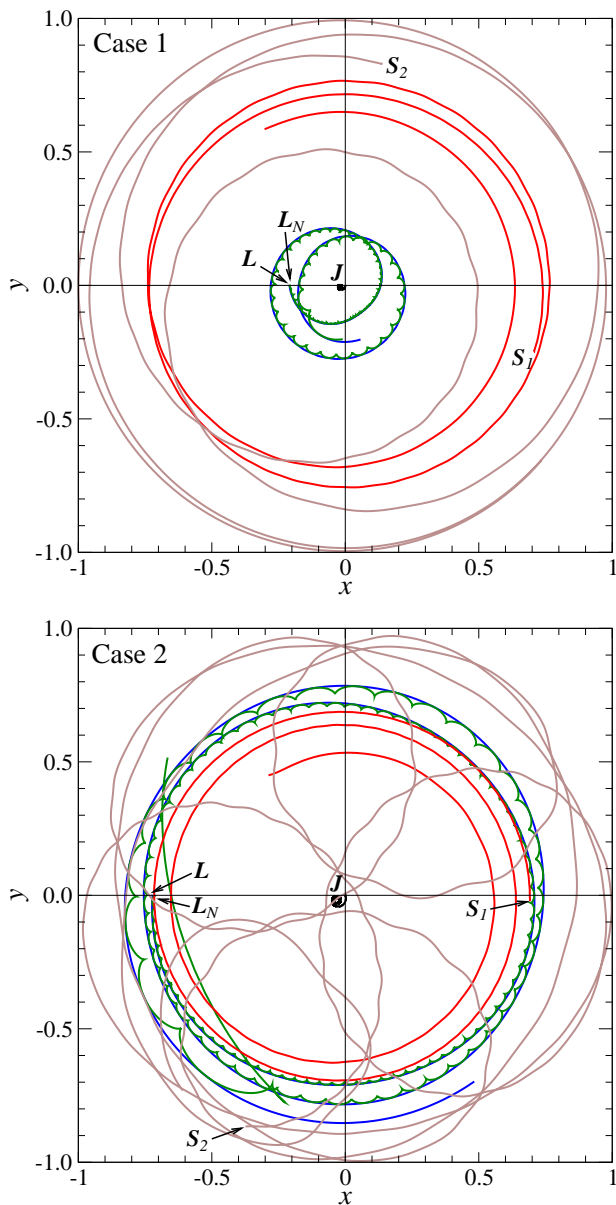


FIG. 3: We show the projections of $\hat{\mathbf{J}}(t)$, $\hat{\mathbf{L}}(t)$, $\hat{\mathbf{L}}_N(t)$, $\hat{\mathbf{S}}_1(t)$, and $\hat{\mathbf{S}}_2(t)$ on the x - y plane of the inertial frame whose z -axis is aligned with $\mathbf{J}(t_{\Omega^{\text{EOB}}})$. In the top and bottom panels we show trajectories of these unit vectors for cases 1 and 2 of Table I, respectively. The initial point of each trajectory is marked by its name. The trajectory of $\hat{\mathbf{J}}(t)$ ends at the origin, by definition. The trajectory of $\hat{\mathbf{L}}_N(t)$ follows that of $\hat{\mathbf{L}}(t)$ with oscillations due to nutation.

shows nutation at twice the orbital frequency and its average follows the smooth precession trajectory of $\hat{\mathbf{L}}(t)$. From PN theory [49]

$$\mathbf{L} = \mathbf{L}_N + \mathbf{L}_{\text{PN}} + \mathbf{L}_{\text{SO}} + \mathcal{O}(c^{-4}), \quad (28)$$

where

$$\mathbf{L}_{\text{PN}} \equiv \mathbf{L}_N \left[\frac{1}{2} v^2 (1 - 3\nu) + (3 + \nu) \frac{M}{r} \right], \quad (29)$$

$$\mathbf{L}_{\text{SO}} \equiv -\frac{2\mu}{r} \left[(\mathbf{S}_{\text{eff}} \cdot \hat{\mathbf{L}}_N) \hat{\mathbf{L}}_N + (\mathbf{S}_{\text{eff}} \cdot \hat{\boldsymbol{\lambda}}) \hat{\boldsymbol{\lambda}} \right], \quad (30)$$

with $v \equiv \hat{\Omega}^{1/3}$, $\hat{\boldsymbol{\lambda}} \equiv (\hat{\mathbf{L}}_N \times \mathbf{r})/r$ and

$$\mathbf{S}_{\text{eff}} \equiv \left(1 + \frac{3m_2}{4m_1} \right) \mathbf{S}_1 + \left(1 + \frac{3m_1}{4m_2} \right) \mathbf{S}_2. \quad (31)$$

Note that the unit vector $\hat{\boldsymbol{\lambda}}$ instantaneously rotates about $\hat{\mathbf{L}}_N$ at the orbital frequency Ω . In addition, \mathbf{L} obeys a simple precession equation about \mathbf{J} , i.e. $\dot{\mathbf{L}} \propto \mathbf{J} \times \mathbf{L}$ (see Eq. (2.13) of Ref. [49]). This, together with Eq. (28), implies that \mathbf{L}_N cannot simply precess about \mathbf{J} . When computing $\dot{\mathbf{L}}_N$, the spin-orbit term \mathbf{L}_{SO} generates contributions of the form

$$(\mathbf{S}_{\text{eff}} \cdot \dot{\hat{\boldsymbol{\lambda}}}) \hat{\boldsymbol{\lambda}} \quad \text{and} \quad (\mathbf{S}_{\text{eff}} \cdot \hat{\boldsymbol{\lambda}}) \dot{\hat{\boldsymbol{\lambda}}}, \quad (32)$$

which indeed oscillate at twice the orbital frequency, accounting for the nutations seen in Fig. 3.

The main difference between the two cases is the size of the opening angle between $\mathbf{J}(t)$ and $\mathbf{L}(t)$ and correspondingly the strength of the orbital precession. In the comparable-mass $q = 2$ case, $\mathbf{L}(t)$ always dominates over the BH spins during inspiral and the angle between $\mathbf{J}(t)$ and $\mathbf{L}(t)$ remains small. The orbital precession is therefore mild. In the $q = 6$ case, on the contrary, the contribution of $\mathbf{S}_1(t)$ to $\mathbf{J}(t)$ is comparable to that of $\mathbf{L}(t)$ initially and becomes more and more dominant. Because of the large opening angle between $\mathbf{J}(t)$ and $\mathbf{L}(t)$, the direction of $\mathbf{L}(t)$ changes more than $\pi/2$ during precession and an initially face-on binary becomes edge-on a few times during the inspiral.

In Fig. 4, we compare precessing waveforms generated in the $\hat{\mathbf{L}}_N(t)$ and $\hat{\mathbf{L}}(t)$ precessing source frames. Considering the oscillatory difference between the trajectories of $\hat{\mathbf{L}}_N(t)$ and $\hat{\mathbf{L}}(t)$ shown in Fig. 3, it is somewhat unexpected that the precessing waveforms agree quite well. In case 1, the waveforms are visually indistinguishable during inspiral — with relative amplitude difference below 1% and phase difference below 0.02 radians. Even in the $q = 6$ case 2, where precession is strong, the waveforms agree reasonably well. Although the relative amplitude and phase differences oscillate strongly when the amplitudes of the waveforms are small, their averages differ only by $< 5\%$ and < 0.15 radians over the $\sim 24000M$ long inspiral. The oscillations are due to the precession-induced modulation and are expected to be strong when the orbital plane goes through a nearly edge-on phase, corresponding to small waveform amplitudes.

In Fig. 5, we examine closely the waveforms as well as their phase evolutions over a time period of $500M$. The real and imaginary parts of h_{22} , i.e. its $+$ and \times polarizations in the radiation frame, show substantial amplitude difference, implying a deviation from circular polarization due to the orbital plane inclination. The phase evolves most rapidly when the dominant quadrature (the imaginary part in this example) goes through zero. Even a small difference in the times when this happens for the

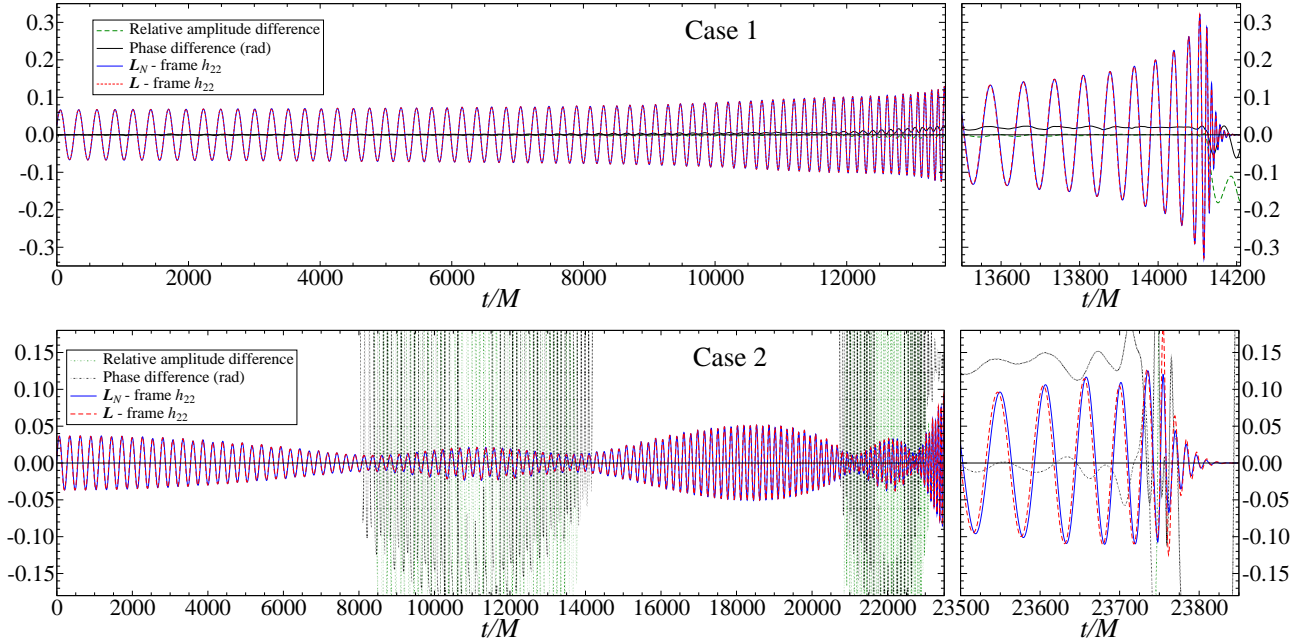


FIG. 4: We show the $\hat{L}_N(t)$ -frame and the $\hat{L}(t)$ -frame precessing waveforms, as well as their relative amplitude and phase differences. The top and bottom panels are waveforms for cases 1 and 2 of Table I. The left and right panels show the inspiral and the plunge-merger-ringdown stages of the waveforms, respectively.

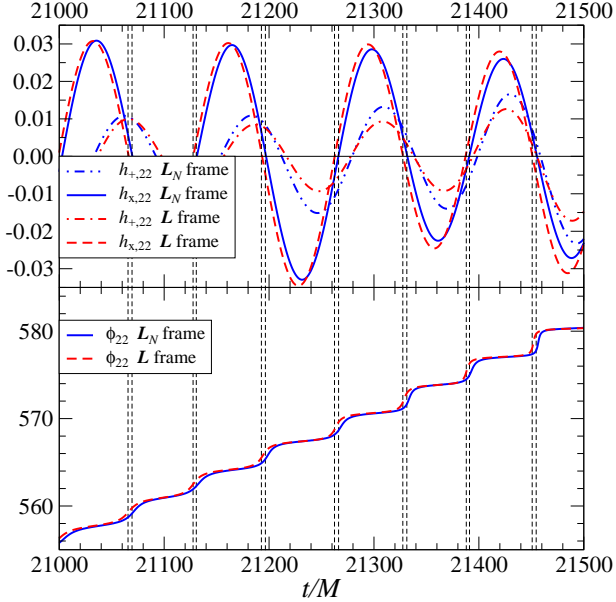


FIG. 5: For case 2 of Table I, we show the $\hat{L}_N(t)$ -frame and $\hat{L}(t)$ -frame waveforms in the top panel and their phase evolutions in the bottom panel, over a short time period from $t = 21\,000M$ to $21\,500M$. The vertical lines mark the time when the dominant quadrature (the imaginary part for this specific instance) of any waveform becomes zero. It coincides with the time when the corresponding phase evolution in the bottom panel experiences a rapid growth. The absolute phase values are not relevant.

two waveforms leads to a burst of phase difference. Such phase differences can be partly removed by time-shifting the two waveforms, but not through a phase shift. In spite of these bursts of amplitude and phase difference, the overall agreement of the waveforms is good. The overlaps between the $\hat{L}_N(t)$ -frame and $\hat{L}(t)$ -frame waveforms, optimized over time and phase of coalescence, are above 0.999 in case 1 and above 0.985 in case 2⁶. The lower overlaps in case 2 are due to the larger difference between $\hat{L}_N(t)$ -frame and $\hat{L}(t)$ -frame waveforms during merger and ringdown.

Finally, we examine the preferred radiation axis determined by the waveforms extracted at infinity. Since we developed only the $\ell = 2$ modes in the current EOB model, we calculate the quadrupole-preferred radiation axis [28] with a small modification. In Ref. [28], the quadrupole-preferred axis is determined by maximizing the power in the $(2, 2)$ and $(2, -2)$ modes of the Weyl scalar $\Psi_4(t)$. We determine the quadrupole-preferred axis by maximizing the power in the strain modes $h_{22}(t)$ and $h_{2,-2}(t)$. Specifically, given the $h_{2m}^{(o)}(t)$ modes in an arbitrary original frame, the $h_{2m}^{(n)}(t)$ modes in a new frame are given by Eq. (7); so we compute the Euler angles $\alpha(t)$, $\beta(t)$ and $\gamma(t)$ (defining the rotation from the original to the new frame) that maximize the quantity

⁶ The overlaps are calculated using the zero-detuned high-power Advanced LIGO noise curve [3] for the range of binary total masses from 20 to $200M_\odot$.

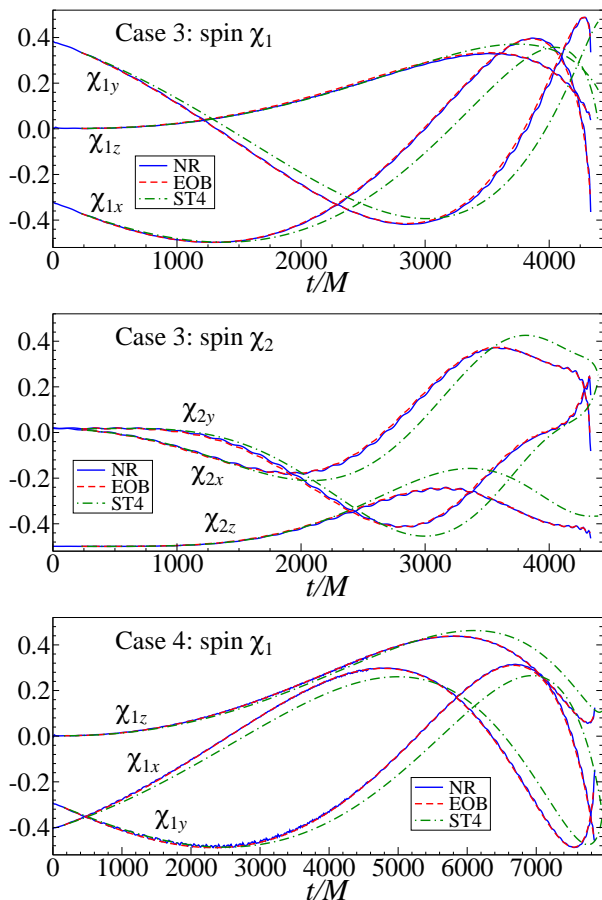


FIG. 6: We show evolutions of the dimensionless spin vectors $\chi_1 = \mathbf{S}_1(t)/m_1^2$ and $\chi_2 = \mathbf{S}_2(t)/m_2^2$ of the NR simulation and the EOB and ST4 models. Specifically, we show the projections of χ_1 and χ_2 on the basis vectors of the inertial source frame $\{\mathbf{e}_1^S, \mathbf{e}_2^S, \mathbf{e}_3^S\}$ that is aligned with the initial orbital orientation $[\hat{\mathbf{L}}_N]_0$ (see Fig. 1). The top two panels show χ_1 and χ_2 for case 3 of Table I. The bottom panel shows χ_1 ($\chi_2 = 0$) for case 4 of Table I. The EOB and ST4 data start at the after-junk-radiation time in the NR simulations, which are $t = 230M$ and $t = 160M$ for cases 3 and 4, respectively.

$|h_{22}^{(n)}(t)|^2 + |h_{2,-2}^{(n)}(t)|^2$. The quadrupole-preferred axis is then given by the z -axis of the new frame defined by these Euler angles. We find that the quadrupole-preferred axis computed from $\hat{\mathbf{L}}_N(t)$ -frame or $\hat{\mathbf{L}}(t)$ -frame waveforms agrees with $\hat{\mathbf{L}}_N(t)$ or $\hat{\mathbf{L}}(t)$ to within 0.3° during inspiral. That is to say, the preferred radiation axis determined by EOB waveforms coincides with the reference axis ($\hat{\mathbf{L}}_N(t)$ or $\hat{\mathbf{L}}(t)$) of the precessing frame determined by the EOB dynamics. Therefore, comparisons of preferred radiation axes determined by NR and EOB precessing waveforms will provide direct information for calibrating the precession dynamics, in particular the dynamics of $\hat{\mathbf{L}}_N(t)$ and $\hat{\mathbf{L}}(t)$, of the EOB model.

B. Comparison with numerical-relativity waveforms

The precessing EOB model defined in Sec. III is not calibrated to any precessing numerical simulations. The only nonperturbative information extracted from NR simulations and employed in this precessing EOB model is contained in the spinning, nonprecessing sector, which was calibrated to *only* two equal-mass, spinning, nonprecessing numerical simulations [44] and five nonspinning ones [42, 45] in Ref. [13]. It is therefore highly interesting to compare the EOB precessing waveforms to NR waveforms.

The Caltech-Cornell-CITA collaboration has recently produced a large number of long and accurate waveforms [46]. We choose among them two precessing waveforms that are sufficiently long (~ 35 and ~ 65 GW cycles) and display strong precessional modulations. The physical parameters of these two waveforms are given in the last two rows of Table I. We compare those numerical waveforms also with the PN SpinTaylorT4 (ST4) inspiraling waveforms [25], which are commonly used in the literature and in LIGO and Virgo software. We generate the ST4 waveforms at the highest PN order available today, namely spin-amplitude corrections through 1.5PN order [47]⁷ and phase corrections through 3.5PN order [21] using the LIGO Algorithm Library [67].

We extract the initial values of \mathbf{S}_1 , \mathbf{S}_2 and GW frequency from the NR data soon after the junk radiation, which typically carries away unphysical radiation present in the initial data. We then set EOB and ST4 initial conditions using these values and start their evolutions after the junk-radiation time, which is $t = 230M$ for case 3 of Table I and $t = 160M$ for case 4 of Table I. We align the orbital orientation $\hat{\mathbf{L}}_N$ at these after-junk-radiation times with the inertial source frame $\{\mathbf{e}_1^S, \mathbf{e}_2^S, \mathbf{e}_3^S\}$ (see Fig. 1) and use it as the default frame for our comparisons. Unlike the case of nonprecessing dynamics and waveforms, we must impose specific \mathbf{S}_1 and \mathbf{S}_2 directions relative to the initial binary separation vector \mathbf{r}_0 at a specific after-junk-radiation time. Thus, we *do not* apply any time or phase shifts when comparing numerical and analytical waveforms.

In Fig. 6 we compare the evolutions of the dimensionless spin vectors $\chi_1(t) = \mathbf{S}_1(t)/m_1^2$ and $\chi_2(t) = \mathbf{S}_2(t)/m_2^2$ ($\chi_2 = 0$ for case 4) for the NR, EOB and ST4 dynamics. Quite remarkably, the EOB spins follow the NR ones rather accurately all the way through the inspiral-plunge stage, while the ST4 spins, although capturing the qualitative precessional behavior of the NR ones, show quantitative differences in both the inspiral

⁷ The 2PN spin-amplitude corrections have been derived in Ref. [22]. Since they are not yet implemented in any ready-to-use software package and are not crucial for the purpose of our comparisons, we do not include them here.

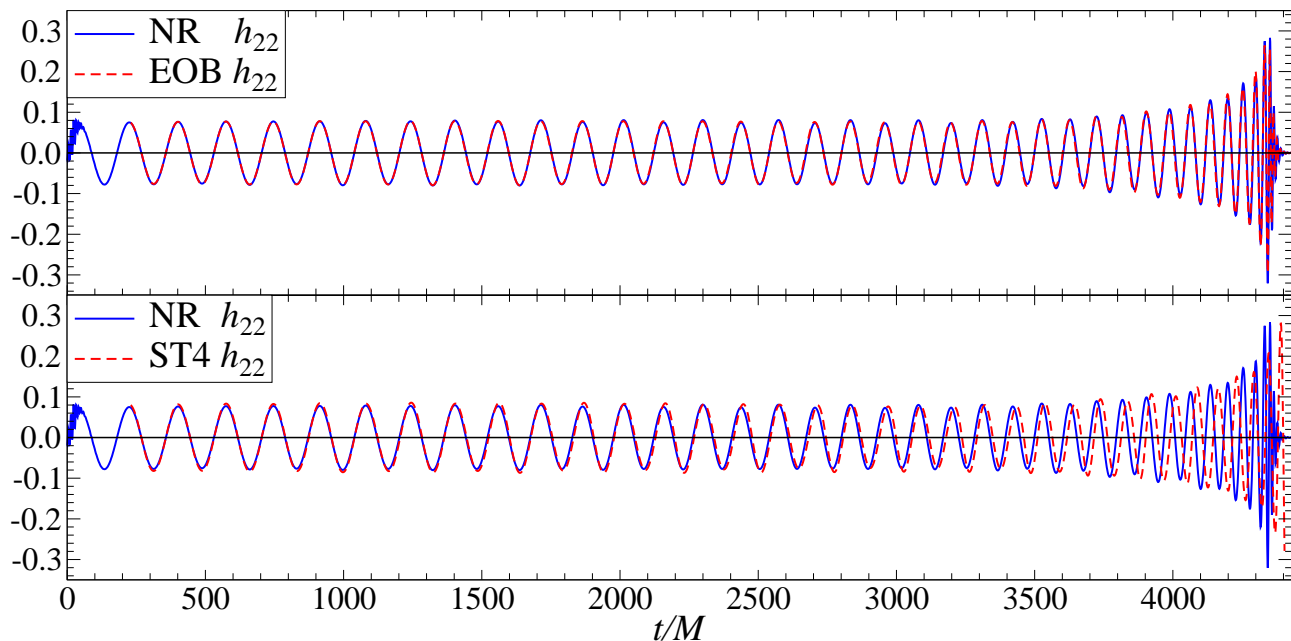


FIG. 7: We show for case 3 of Table I the h_{22} mode decomposed in the inertial source frame $\{e_1^S, e_2^S, e_3^S\}$ that is aligned with the initial orbital orientation $[\hat{\mathbf{L}}_N]_0$ (see Fig. 1). For clarity, we show the NR and EOB h_{22} in the top panel and the NR and ST4 h_{22} in the bottom panel. The EOB and ST4 data start at the after-junk-radiation time of $t = 230M$.

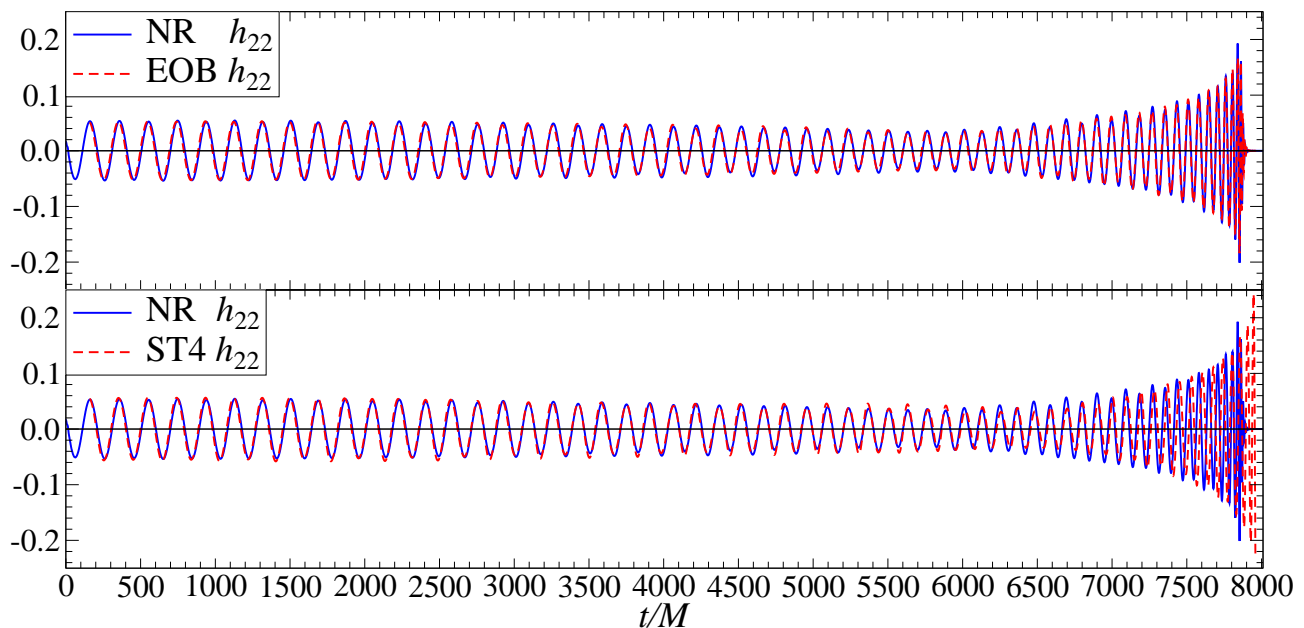


FIG. 8: We show for case 4 of Table I the h_{22} mode decomposed in the inertial source frame $\{e_1^S, e_2^S, e_3^S\}$ that is aligned with the initial orbital orientation $[\hat{\mathbf{L}}_N]_0$ (see Fig. 1). For clarity, we show the NR and EOB h_{22} in the top panel and the NR and ST4 h_{22} in the bottom panel. The EOB and ST4 data start at the after-junk-radiation time of $t = 160M$.

and precessional time scales.

In Figs. 7 and 8, we compare NR, EOB and ST4 h_{22} modes decomposed in the inertial source frame $\{e_1^S, e_2^S, e_3^S\}$. Since the source frame is aligned with the initial orbital orientation $[\hat{\mathbf{L}}_N]_0$ and the binary orbit precesses only moderately in case 3, there are only moder-

ate modulations on h_{22} in this case. The modulations in case 4 are strong, though. In both cases, the agreement between NR and EOB h_{22} modes is remarkable. Their amplitudes agree quite well and their phases, aligned at the initial time, differ by only ~ 0.2 rad at merger, i.e., at the peak of the NR (2,2) mode. The agreement be-

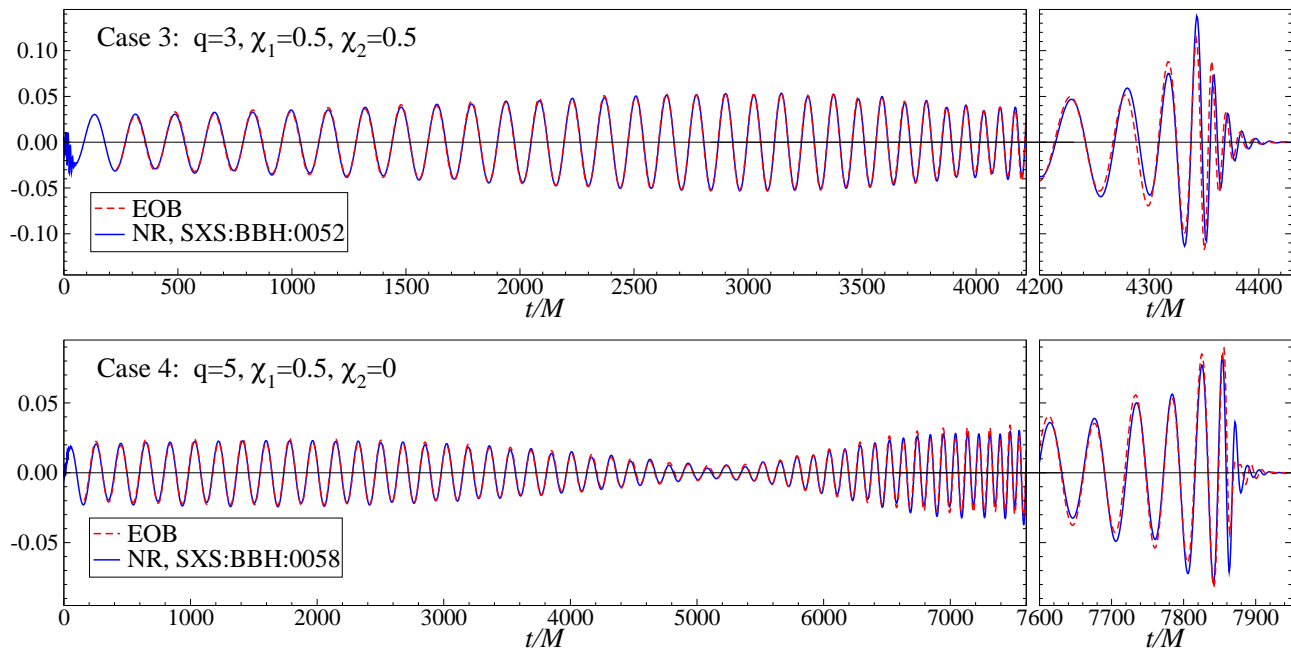


FIG. 9: We show for cases 3 and 4 of Table I the GW polarization h_+ , containing contributions from $\ell = 2$ modes, that propagates along a direction $\hat{\mathbf{N}}$ specified by spherical coordinates $\theta = \pi/3$ and $\phi = \pi/2$ associated with the inertial source frame $\{\mathbf{e}_1^S, \mathbf{e}_2^S, \mathbf{e}_3^S\}$. The EOB waveforms start at the after-junk-radiation times of $t = 230M$ and $t = 160M$, respectively.

tween NR and ST4 h_{22} modes, although not comparable with the agreement between NR and EOB, is also very good. Even though the amplitudes differ by $\sim 10\%$ during the inspiral, because amplitude corrections are known only through 1.5PN order in the spinning case [47], their phases agree quite well for tens of cycles but start departing from each other 10 GW cycles before merger. Quite interestingly, we have found that using the newly available 3.5PN spin-orbit effects [21] in the phasing of ST4 improves the agreement with the NR waveforms. If we were using the 2.5PN phasing, the end of the inspiral would occur $\sim 460M$ ($960M$) instead of $\sim 60M$ ($140M$) after the merger of the NR waveform, for case 3 (4). Moreover, we find for cases 3 and 4 that when we align the 3.5PN and NR phasing at the after-junk-radiation time, they accumulate a difference of 1 GW cycle only 1 cycle before merger. By contrast the 2.5PN phasing differs from the NR phasing by 1 GW after 28 (37) GW cycles [or 6 (16) GW cycles before merger] for case 3 (4).

The agreement between NR and EOB modes (2, 1) and (2, 0) modes are also very good. Rather than the modes, we show in Fig. 9 the NR and EOB polarizations $h_+(t)$ given by Eq. (3). Since only the $\ell = 2$ modes are available in the current precessing EOB model, we limit the summation over ℓ to only $\ell = 2$. To include substantial contributions from all $\ell = 2$ modes, we choose $\theta = \pi/3$ and $\phi = \pi/2$ for the direction of GW propagation \mathbf{N} (see Fig. 1). As expected from the very good agreement of the individual modes, the NR and EOB polarizations also agree remarkably.

Finally, we measure the difference between EOB and

NR polarizations with the *unfaithfulness* [68], defined as

$$\bar{\mathcal{F}} = 1 - \max_{t_c, \phi_c, \psi} \frac{\langle h_{\text{NR}} | h_{\text{EOB}} \rangle}{\sqrt{\langle h_{\text{NR}} | h_{\text{NR}} \rangle \langle h_{\text{EOB}} | h_{\text{EOB}} \rangle}}, \quad (33)$$

where the EOB waveform of the detector response is

$$h_{\text{EOB}}(t; t_c, \phi_c, \psi, \boldsymbol{\lambda}) \propto \cos \psi h_{\text{EOB},+}(t; t_c, \phi_c, \boldsymbol{\lambda}) + \sin \psi h_{\text{EOB},\times}(t; t_c, \phi_c, \boldsymbol{\lambda}), \quad (34)$$

and the maximization is over the time and phase of coalescence t_c and ϕ_c , as well as the polarization angle ψ that combines the $+$ and \times polarizations in the radiation frame. We do not optimize over the physical binary parameters $\boldsymbol{\lambda}$, i.e., we use the same $\boldsymbol{\lambda}$ in h_{NR} and h_{EOB} . Note that since we include modes with different m , ϕ_c and ψ are no longer degenerate and both of them have to be maximized over. We define the inner product between two waveforms through the following integral in the frequency domain

$$\langle h_1, h_2 \rangle \equiv 4\text{Re} \int_0^\infty \frac{\tilde{h}_1(f) \tilde{h}_2^*(f)}{S_h(f)} df, \quad (35)$$

where $\tilde{h}_1(f)$ and $\tilde{h}_2(f)$ are frequency domain waveforms and $S_h(f)$ is the noise power spectral density of the detector. We employ the zero-detuned high-power advanced LIGO noise curve ZERO_DET_HIGH_P given in [3]. The NR waveforms, although very long, cover the entire advanced LIGO frequency band only for $M \geq 100M_\odot$. Thus, to reduce artifacts when considering binaries with

$M < 100M_{\odot}$, we taper both ends of the NR and EOB waveforms using the Planck-taper window function [69] (see Ref. [12] for details). In Fig. 10, we show the EOB unfaithfulness when the total mass M varies between $20M_{\odot}$ and $200M_{\odot}$. We choose the same direction of GW propagation \hat{N} as is considered in Fig. 9, namely $\theta = \pi/3$ and $\phi = \pi/2$.

For each waveform we estimate the numerical error in the unfaithfulness results of Fig. 10 by calculating the unfaithfulness of the EOB waveform with two numerical waveforms: the extrapolated high-resolution waveform shown in Fig. 9 and the outermost finite-radius high-resolution waveform. We use the difference between these unfaithfulness results to estimate the extrapolation error. We might estimate the finite resolution errors in the same way by calculating the unfaithfulness of the EOB waveform with the extrapolated high- and medium-resolution numerical waveforms. However, medium resolution simulations for these two cases are not available, but we expect from previous studies that errors due to resolution are smaller than errors due to extrapolation [12].

Since the unfaithfulness of EOB waveforms is below $\sim 2\%$, we expect that the ineffectualness, which measures the difference between EOB and NR waveforms when minimizing also over the binary parameters λ , will be below 1% (with a loss of event rates less than 3%). Thus, for those two precessing binary configurations, the EOB waveforms are sufficiently accurate for detection with advanced LIGO detectors.

Although these very encouraging results refer only to two precessing binary configurations, they strongly suggest that the approach we have proposed for modeling precessing compact binaries within the EOB model is bound to succeed. A more comprehensive and careful comparison of the EOB model with a larger number of accurate NR simulations will be carried out in the near future using the entire catalog of simulations in Ref. [46].

V. CONCLUSIONS

So far, the EOB modeling of GWs emitted from compact binaries has focused primarily on nonprecessing binary configurations [6–14]. Nonspinning EOB waveforms have been employed in the first searches of GWs from high-mass binary BHs with LIGO and Virgo detectors [70–72]. Recently, studies carried out within the NRAR collaboration [19] have shown that nonprecessing EOB waveforms originally calibrated to seven NR waveforms [42, 44, 45] in Ref. [13] match very well also tens of new NR waveforms produced within the NRAR collaboration. The next, challenging task is to achieve a similar success also for generic, spinning binary configurations. In this paper we have started addressing this important problem.

Building on previous work [11, 13, 25, 38–41], we have proposed a strategy to generate EOB precessing wave-

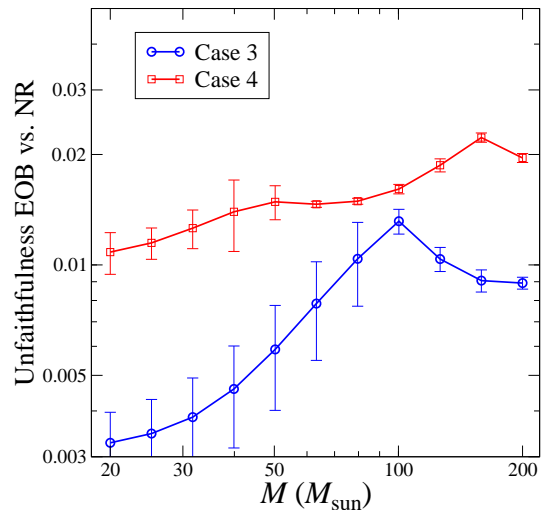


FIG. 10: Unfaithfulness of the EOB waveforms when compared to NR waveforms as a function of the total binary mass. Shown are cases 3 and 4 of Table I. The error bars are estimates of numerical errors. The direction of GW propagation \hat{N} is specified by the spherical coordinates $\theta = \pi/3$ and $\phi = \pi/2$.

forms. The procedure employs the precessing convention of Ref. [25] that minimizes the precession-induced modulations in the waveform’s phase and amplitude, and an inertial frame aligned with the spin of the final BH where the matching between the inspiral-plunge and merger-ringdown EOB waveforms is carried out.

When spins are aligned or antialigned with the orbital angular momentum, the EOB precessing waveforms that we have built reduce to the nonprecessing EOB waveforms calibrated to seven nonprecessing NR waveforms in Ref. [13]. Since the factorized energy flux is not yet available for precessing spins, we have included in the radiation-reaction force of the EOB dynamics only spin couplings whose projection along the orbital angular momentum is different from zero. This limitation will be relaxed in the future as soon as the radiation-reaction sector of the EOB model is improved. Furthermore, we have limited this first study to the EOB $\ell = 2$ modes.

Without recalibrating the EOB precessing waveforms, we have then compared them to two, long, strongly precessing NR waveforms that were recently produced in Ref. [46]. We have found a remarkable agreement both for the dynamics, that is the spins’ components, and the gravitational polarizations. In particular, when using the advanced-LIGO noise spectral density, the mismatches between the EOB and NR waveforms for binary masses $20\text{--}200M_{\odot}$ are below 2% when maximizing only on the time and phase at coalescence and on the polarization angle. Although those results only refer to two binary configurations, they are very encouraging and suggest that the EOB precessing model developed here is an excellent starting point for building a generic, spinning EOB model for advanced LIGO and Virgo searches. We have

also compared the two NR waveforms to PN ST4 waveforms that are largely used in the literature and in LIGO and Virgo software. We have found that the PN waveforms at 3.5PN order agree very well with NR waveforms for several GW cycles, and accumulate a phase difference of ~ 6 rad, starting about 10 GW cycles before merger.

Finally, several analyses were left out in this first study of precessing waveforms. They include (i) a more detailed comparison between spin variables in the numerical simulations and analytical models, (ii) the extension of precessing waveforms to modes higher than $\ell = 2$, (iii) a more systematic way of identifying the initial conditions in the numerical and analytical waveforms, and (iv) the inclusion of resolution errors when estimating numerical errors. We defer those important extensions to a future publication where many more NR waveforms will be also analysed.

Acknowledgments

We thank Anil Zenginoglu, Geoffrey Lovelace and Mike Boyle for their contributions to the productions of the

NR waveforms used in this paper. A.B., Y.P. and A.T. acknowledge partial support from NSF Grants No. PHY-0903631 and No. PHY-1208881. A.B. also acknowledges partial support from the NASA Grant NNX12AN10G. A.M. and H.P. gratefully acknowledge support from NSERC of Canada, the Canada Chairs Program, and the Canadian Institute for Advanced Research. L.K. gratefully acknowledges support from the Sherman Fairchild Foundation, and from NSF grants PHY-0969111 and PHY-1005426. Simulations used in this work were computed with the `SpEC` code [73]. Computations were performed on the Zwicky cluster at Caltech, which is supported by the Sherman Fairchild Foundation and by NSF award PHY-0960291; on the NSF XSEDE network under grant TG-PHY990007N; and on the GPC supercomputer at the SciNet HPC Consortium [74]. SciNet is funded by: the Canada Foundation for Innovation under the auspices of Compute Canada; the Government of Ontario; Ontario Research Fund—Research Excellence; and the University of Toronto.

-
- [1] F. Acernese, M. Alshourbagy, P. Amico, F. Antonucci, S. Aoudia, et al., *Class.Quant.Grav.* **25**, 114045 (2008).
- [2] H. Grote (LIGO Scientific Collaboration), *Class.Quant.Grav.* **25**, 114043 (2008).
- [3] D. Shoemaker (LIGO Collaboration), *Advanced LIGO anticipated sensitivity curves* (2010), LIGO Document T0900288-v3, URL <https://dcc.ligo.org/cgi-bin/DocDB/ShowDocument?docid=2974>.
- [4] J. Aasi et al. (LIGO Scientific Collaboration, Virgo Collaboration) (2013), 1304.0670.
- [5] *ESA LISA website*, <http://sci.esa.int/lisa>.
- [6] A. Buonanno, Y. Pan, J. G. Baker, J. Centrella, B. J. Kelly, et al., *Phys.Rev.* **D76**, 104049 (2007), 0706.3732.
- [7] T. Damour, A. Nagar, E. N. Dorband, D. Pollney, and L. Rezzolla, *Phys.Rev.* **D77**, 084017 (2008), 0712.3003.
- [8] T. Damour, A. Nagar, M. Hannam, S. Husa, and B. Bruegmann, *Phys.Rev.* **D78**, 044039 (2008), 0803.3162.
- [9] A. Buonanno, Y. Pan, H. P. Pfeiffer, M. A. Scheel, L. T. Buchman, et al., *Phys.Rev.* **D79**, 124028 (2009), 0902.0790.
- [10] T. Damour and A. Nagar, *Phys.Rev.* **D79**, 081503 (2009), 0902.0136.
- [11] Y. Pan, A. Buonanno, L. T. Buchman, T. Chu, L. E. Kidder, et al., *Phys.Rev.* **D81**, 084041 (2010), 0912.3466.
- [12] Y. Pan, A. Buonanno, M. Boyle, L. T. Buchman, L. E. Kidder, et al., *Phys.Rev.* **D84**, 124052 (2011), 1106.1021.
- [13] A. Taracchini, Y. Pan, A. Buonanno, E. Barausse, M. Boyle, et al., *Phys.Rev.* **D86**, 024011 (2012), 1202.0790.
- [14] T. Damour, A. Nagar, and S. Bernuzzi (2012), 1212.4357.
- [15] P. Ajith, S. Babak, Y. Chen, M. Hewitson, B. Krishnan, et al., *Class.Quant.Grav.* **24**, S689 (2007), 0704.3764.
- [16] P. Ajith, S. Babak, Y. Chen, M. Hewitson, B. Krishnan, A. M. Sintes, J. T. Whelan, B. Brüggmann, P. Diener, N. Dorband, et al., *Phys. Rev. D* **77**, 104017 (2008), 0710.2335.
- [17] P. Ajith, M. Hannam, S. Husa, Y. Chen, B. Bruegmann, et al., *Phys.Rev.Lett.* **106**, 241101 (2011), 0909.2867.
- [18] L. Santamaria, F. Ohme, P. Ajith, B. Bruegmann, N. Dorband, et al., *Phys.Rev.* **D82**, 064016 (2010), 1005.3306.
- [19] I. Hinder et al. (The NRAR Collaboration) (2013), 1307.5307.
- [20] L. Blanchet, A. Buonanno, and G. Faye, *Phys.Rev.* **D84**, 064041 (2011), 1104.5659.
- [21] A. Bohe, S. Marsat, and L. Blanchet (2013), 1303.7412.
- [22] A. Buonanno, G. Faye, and T. Hinderer, *Phys.Rev.* **D87**, 044009 (2013), 1209.6349.
- [23] T. A. Apostolatos, C. Cutler, G. J. Sussman, and K. S. Thorne, *Phys.Rev.* **D49**, 6274 (1994).
- [24] T. Apostolatos, *Phys.Rev.* **D52**, 605 (1995).
- [25] A. Buonanno, Y.-b. Chen, and M. Vallisneri, *Phys.Rev.* **D67**, 104025 (2003), gr-qc/0211087.
- [26] Y. Pan, A. Buonanno, Y.-b. Chen, and M. Vallisneri, *Phys.Rev.* **D69**, 104017 (2004), gr-qc/0310034.
- [27] A. Buonanno, Y. Chen, Y. Pan, H. Tagoshi, and M. Vallisneri, *Phys.Rev.* **D72**, 084027 (2005), gr-qc/0508064.
- [28] P. Schmidt, M. Hannam, S. Husa, and P. Ajith, *Phys.Rev.* **D84**, 024046 (2011), 1012.2879.
- [29] R. O’Shaughnessy, B. Vaishnav, J. Healy, Z. Meeks, and D. Shoemaker, *Phys.Rev.* **D84**, 124002 (2011), 1109.5224.
- [30] M. Boyle, R. Owen, and H. P. Pfeiffer, *Phys.Rev.* **D84**, 124011 (2011), 1110.2965.
- [31] M. Boyle, *Phys. Rev. D* **87**, 104006 (2013), 1302.2919.
- [32] P. Schmidt, M. Hannam, and S. Husa, *Phys.Rev.* **D86**, 104063 (2012), 1207.3088.

- [33] L. Pekowsky, R. O’Shaughnessy, J. Healy, and D. Shoemaker (2013), 1304.3176.
- [34] A. Buonanno and T. Damour, Phys.Rev. **D59**, 084006 (1999), gr-qc/9811091.
- [35] A. Buonanno and T. Damour, Phys.Rev. **D62**, 064015 (2000), gr-qc/0001013.
- [36] T. Damour, P. Jaranowski, and G. Schaefer, Phys.Rev. **D62**, 084011 (2000), gr-qc/0005034.
- [37] T. Damour, Phys.Rev. **D64**, 124013 (2001), gr-qc/0103018.
- [38] A. Buonanno, Y. Chen, and T. Damour, Phys.Rev. **D74**, 104005 (2006), gr-qc/0508067.
- [39] E. Barausse, E. Racine, and A. Buonanno, Phys.Rev. **D80**, 104025 (2009), 0907.4745.
- [40] E. Barausse and A. Buonanno, Phys.Rev. **D81**, 084024 (2010), 0912.3517.
- [41] E. Barausse and A. Buonanno, Phys.Rev. **D84**, 104027 (2011), 1107.2904.
- [42] M. A. Scheel, M. Boyle, T. Chu, L. E. Kidder, K. D. Matthews, et al., Phys.Rev. **D79**, 024003 (2009), 0810.1767.
- [43] M. Boyle, D. A. Brown, L. E. Kidder, A. H. Mroue, H. P. Pfeiffer, et al., Phys.Rev. **D76**, 124038 (2007), 0710.0158.
- [44] T. Chu, H. P. Pfeiffer, and M. A. Scheel, Phys.Rev. **D80**, 124051 (2009), 0909.1313.
- [45] L. T. Buchman, H. P. Pfeiffer, M. A. Scheel, and B. Szilagyi, Phys.Rev. **D86**, 084033 (2012), 1206.3015.
- [46] A. H. Mroue, M. A. Scheel, B. Szilagyi, H. P. Pfeiffer, M. Boyle, et al. (2013), 1304.6077.
- [47] K. Arun, A. Buonanno, G. Faye, and E. Ochsner, Phys.Rev. **D79**, 104023 (2009), 0810.5336.
- [48] J. Goldberg, A. MacFarlane, E. Newman, F. Rohrlich, and E. Sudarshan, J.Math.Phys. **8**, 2155 (1967).
- [49] L. E. Kidder, Phys.Rev. **D52**, 821 (1995), gr-qc/9506022.
- [50] L. Blanchet, A. Buonanno, and G. Faye, Phys.Rev. **D74**, 104034 (2006), gr-qc/0605140.
- [51] Y. Pan, A. Buonanno, R. Fujita, E. Racine, and H. Tagoshi, Phys.Rev. **D83**, 064003 (2011), 1006.0431.
- [52] A. Buonanno, L. E. Kidder, A. H. Mroue, H. P. Pfeiffer, and A. Taracchini, Phys.Rev. **D83**, 104034 (2011), 1012.1549.
- [53] H. P. Pfeiffer, D. A. Brown, L. E. Kidder, L. Lindblom, G. Lovelace, et al., Class.Quant.Grav. **24**, S59 (2007), gr-qc/0702106.
- [54] G. Lovelace, M. Scheel, and B. Szilagyi, Phys.Rev. **D83**, 024010 (2011), 1010.2777.
- [55] A. H. Mroue and H. P. Pfeiffer (2012), 1210.2958.
- [56] M. Purrer, S. Husa, and M. Hannam, Phys.Rev. **D85**, 124051 (2012), 1203.4258.
- [57] M. Purrer, M. Hannam, P. Ajith, and S. Husa (2013), 1306.2320.
- [58] T. Damour and A. Nagar, Phys.Rev. **D76**, 064028 (2007), 0705.2519.
- [59] T. Damour, B. R. Iyer, and A. Nagar, Phys.Rev. **D79**, 064004 (2009), 0811.2069.
- [60] L. E. Kidder, Phys.Rev. **D77**, 044016 (2008), 0710.0614.
- [61] A. Buonanno, L. Kidder, Y. Pan, and M. Scheel, Private Communication (2013).
- [62] A. Buonanno, G. B. Cook, and F. Pretorius, Phys.Rev. **D75**, 124018 (2007), gr-qc/0610122.
- [63] E. Barausse and L. Rezzolla, Astrophys.J. **704**, L40 (2009), 0904.2577.
- [64] E. Berti, V. Cardoso, and C. M. Will, Phys.Rev. **D73**, 064030 (2006), gr-qc/0512160.
- [65] W. Tichy and P. Marronetti, Phys.Rev. **D78**, 081501 (2008), 0807.2985.
- [66] E. Ochsner and R. O’Shaughnessy, Phys.Rev. **D86**, 104037 (2012), 1205.2287.
- [67] L. S. Collaboration, *LSC Algorithm Library software packages LAL, LALWRAPPER, and LALAPPS*, URL <http://www.lsc-group.phys.uwm.edu/lal>.
- [68] T. Damour, B. R. Iyer, and B. Sathyaprakash, Phys.Rev. **D57**, 885 (1998), gr-qc/9708034.
- [69] D. McKechnan, C. Robinson, and B. Sathyaprakash, Class.Quant.Grav. **27**, 084020 (2010), 1003.2939.
- [70] J. Abadie et al. (LIGO Scientific Collaboration and Virgo Collaboration), Phys. Rev. D **83**, 122005 (2011), 1102.3781.
- [71] J. Aasi et al. (LIGO Scientific Collaboration and Virgo Collaboration), Phys. Rev. D **87**, 022002 (2013), 1209.6533.
- [72] J. Abadie et al. (LIGO Scientific Collaboration and Virgo Collaboration), Phys. Rev. D **85**, 102004 (2012), 1201.5999.
- [73] <http://www.black-holes.org/SpEC.html>.
- [74] C. Loken, D. Gruner, L. Groer, R. Peltier, N. Bunn, M. Craig, T. Henriques, J. Dempsey, C.-H. Yu, J. Chen, et al., J. Phys.: Conf. Ser. **256**, 012026 (2010).