

Eccentric Binary Neutron Star Search Prospects for Cosmic Explorer

Amber K. Lenon and Duncan A. Brown

Department of Physics, Syracuse University, Syracuse NY 13244, USA

Alexander H. Nitz

*Max-Planck-Institut für Gravitationsphysik (Albert-Einstein-Institut), D-30167 Hannover, Germany and
Leibniz Universität Hannover, D-30167 Hannover, Germany*

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We determine the ability of Cosmic Explorer, a proposed third-generation gravitational-wave observatory, to detect eccentric binary neutron stars and to measure their eccentricity. We find that for a matched-filter search, template banks constructed using binaries in quasi-circular orbits are effectual for eccentric neutron star binaries with $e_7 \leq 0.004$ ($e_7 \leq 0.003$) for CE1 (CE2), where e_7 is the binary’s eccentricity at a gravitational-wave frequency of 7 Hz. We show that stochastic template placement can be used to construct a matched-filter search for binaries with larger eccentricities and construct an effectual template bank for binaries with $e_7 \leq 0.05$. We show that the computational cost of both the search for binaries in quasi-circular orbits and eccentric orbits is not significantly larger for Cosmic Explorer than for Advanced LIGO and is accessible with present-day computational resources. We investigate Cosmic Explorer’s ability to distinguish between circular and eccentric binaries. We estimate that for a binary with a signal-to-noise ratio of 8 (800), Cosmic Explorer can distinguish between a circular binary and a binary with eccentricity $e_7 \gtrsim 10^{-2}$ (10^{-3}) at 90% confidence.

I. INTRODUCTION

Cosmic Explorer is a proposed third-generation gravitational-wave observatory that will have an order of magnitude improved sensitivity beyond that of Advanced LIGO and will be able to explore gravitational-wave frequencies below 10 Hz [1]. Cosmic Explorer will be able to detect binary neutron stars with a signal-to-noise ratio of ≥ 8 out to a distance of ~ 2 Gpc [2]. Although most of the detected neutron-star binaries will be in circular orbits, measurement of eccentricity in neutron-star mergers allows us to explore their formation and to distinguish between field binaries, which are expected to be circular by the time they are observed [3], and binaries formed through other channels [4–17].

Dynamical interactions can form binary neutron stars with eccentricity that is measurable, although the predicted rate of these mergers detectable by current gravitational-wave observatories is small [18, 19]. The two binary neutron star mergers observed by Advanced LIGO and Virgo [20, 21] were both detected with searches that use circular waveform templates [22–30]. Constraints have been placed on the eccentricity of these binaries. At a gravitational-wave frequency of 10 Hz, the eccentricity of GW170817 is $e_{10} \leq 0.024$ [31] and GW190425 has an eccentricity $e_{10} \leq 0.048$ [31, 32] (90% confidence). Ref. [32] considered unstable mass transfer as a formation scenario for GW190425, but the measured eccentricity limit was insufficient to confirm this hypothesis. A search for eccentric binary neutron stars in the first and second Advanced LIGO and Virgo observing runs did not yield any candidates [33].

By extrapolating the upper limit on the rate of eccentric binary neutron stars from LIGO–Virgo observations, Ref. [33] estimates that the A+ upgrade [34] of Advanced

LIGO will require between half a year of observation and ~ 775 years of observation before the detectable rate is comparable with the optimistic [18] and pessimistic [19] rate predictions respectively, and an observation is plausible. However, with its increased sensitivity and bandwidth, Cosmic Explorer would need at most half a year of observations to achieve a detectable rate comparable to even the pessimistic models [33].

We investigate the ability of Cosmic Explorer to detect eccentric binary neutron stars and to measure their eccentricity. We find that at an eccentricity $e_7 = 0.004$, a matched-filter search using circular waveform templates begins to lose more than 3% of the signal-to-noise ratio due to mismatch between the circular and eccentric waveforms; this is an order of magnitude smaller than the equivalent limit for Advanced LIGO [35, 36]. We demonstrate that stochastic template placement [37, 38] can be used to construct a template bank that maintains a fitting factor greater than 97% to binaries with $e_7 \leq 0.05$. We will reference eccentricity at a reference frequency of 7 Hz in reference to eccentricity unless otherwise stated.

Using template banks constructed for Cosmic Explorer, we estimate the computational cost of matched-filter searches for binary neutron stars in circular and eccentric orbits and find that both are accessible with present-day computational resources. We then estimate the ability of Cosmic Explorer to measure and constrain the eccentricity of detected binary neutron star systems. For a binary neutron star with signal-to-noise ratio 8 (800), Cosmic Explorer will be able to measure eccentricities $\gtrsim 8 \times 10^{-3}$ (8×10^{-4}).

This paper is organized as follows: In Sec. II, we investigate the ability of a matched-filter search to detect eccentric binary neutron stars in Cosmic Explorer. We

calculate the lower-frequency cutoff required to obtain at least 99.9% of the available signal-to-noise ratio for binary neutron stars ($m_{1,2} \in [1, 3] M_{\odot}$). Using this frequency cutoff, we use geometric placement to construct a template bank using circular waveform for Cosmic Explorer that has a fitting factor of 97% and estimate the computational cost of performing a matched-filter search using this bank. In Sec. III we measure the loss in fitting factor when using a bank of circular waveform with neutron-star binaries with eccentricity $e \leq 0.05$. We use stochastic template placement to generate a bank containing circular and eccentric waveforms than has a fitting factor of 96.5% and estimate the computational cost of this eccentric binary search. In Sec. IV, we estimate the minimum eccentricity that can be measured by Cosmic Explorer as a function of the signal-to-noise ratio of the detected signal. We compare this to estimates of Advanced LIGO and the eccentricity constraints placed by the detection of GW170817. Finally, in Sec. V, we discuss the implications our results for measurement of eccentric binaries with Cosmic Explorer and extension of our work to higher eccentricities and binary black holes.

II. BINARY NEUTRON STAR SEARCHES IN COSMIC EXPLORER

Cosmic Explorer has a two-stage design [1, 39]. The first stage of Cosmic Explorer (CE1) assumes that the detector's core technologies will be similar to those of Advanced LIGO with the sensitivity gain from increasing the detector's arm length from 4 km to 40 km. The second stage (CE2) is a technology upgrade to the CE1 detector that further increases Cosmic Explorer's sensitivity. Estimates of the detector's noise power spectral density $S_n(f)$ are available for both CE1 and CE2 [40]; we consider both detector configurations in our analysis.

Compared to the low-frequency sensitivity limit of Advanced LIGO, which lies around 10 Hz, Cosmic Explorer pushes the low-frequency limit of the detector below this limit [1]. As for Advanced LIGO, the detector noise begins to rapidly increase as the gravitational-wave frequency reaches the seismic and Newtonian noise walls at low-frequency. The length of a binary neutron star waveform has a steep power-law dependence on its starting frequency f_{lower} , with the number of cycles between f_{lower} and the coalescence frequency scaling as $f_{\text{lower}}^{-8/3}$. A binary neutron star waveform that starts at $f_{\text{lower}} = 3$ Hz has a length of approximately 7 hours, presenting non-trivial data analysis challenges in searches and parameter estimation.

To determine the optimal starting frequency for binary neutron star searches in Cosmic Explorer, we consider the accumulation of the the signal-to-noise ratio in a matched filter search for a neutron star binary with $m_1 = m_2 = 1.4 M_{\odot}$; this accumulates linearly with frequency f as $f^{-7/3}/S_n(f)$ [23, 41]. Fig. 1 shows the fraction of the total signal-to-noise ratio accumulated at

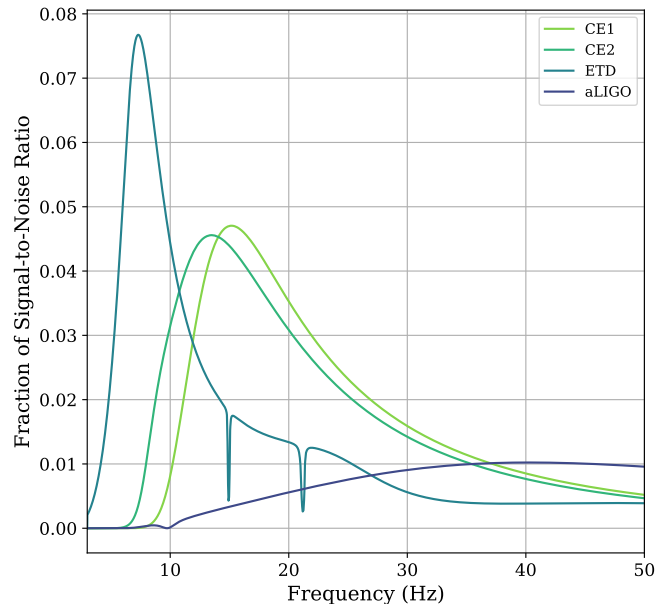


FIG. 1. The fraction of signal-to-noise ratio as a function of frequency for Cosmic Explorer (CE1/CE2), Einstein Telescope (ETD) and Advanced LIGO (aLIGO). This gives a visual representation of what the matched filter sees when it is integrating up the signal-to-noise ratio. A majority of the signal-to-noise ratio for Cosmic Explorer and Advanced LIGO is accumulated between 10 and 50 Hz, while the signal-to-noise ratio for Einstein Telescope is accumulated below 10 Hz.

a given frequency for Cosmic Explorer and Advanced LIGO. Advanced LIGO's most sensitive frequency lies around 40 Hz with almost no detectable signal below 10 Hz. For CE1 and CE2, the peak sensitivity of the detectors to binary neutron stars is shifted to lower frequencies, with a non-trivial amount of signal-to-noise available below 10 Hz. The fraction of signal-to-noise available drops rapidly as the frequency decreases due to the low-frequency noise wall of Cosmic Explorer.

To determine the optimal low-frequency cutoff, we consider the cumulative fraction of the total signal-to-noise ratio as a function of low-frequency cutoff, shown in Fig. 2. This is computed by comparing the ratio of the signal-to-noise obtained by integrating from 3 Hz to a fiducial low-frequency cutoff shown on the ordinate of Fig. 2. We find that for both the CE1 and CE2 detector sensitivity curves, the matched filter accumulates 99.97% (99.53%) of the signal-to-noise above 7 Hz for CE1 (CE2). We therefore use 7 Hz as an appropriate low-frequency cutoff for our analysis. At this starting frequency, the length of a binary neutron star waveform is reduced by two orders of magnitude to 4600 s (77 minutes). For a waveform of this length, the Doppler frequency modulation due to the diurnal and orbital mo-

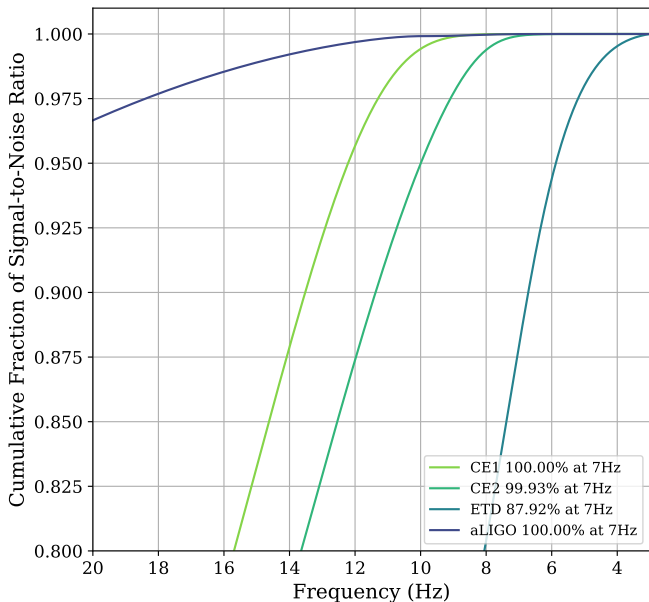


FIG. 2. The cumulative fraction of signal-to-noise ratio as a function of frequency. Cosmic Explorer (CE1/CE2) and Advanced LIGO (aLIGO) have accumulated more than 99.9% of their total signal-to-noise ratio from frequencies above 7 Hz. At 7 Hz, Einstein Telescope (ETD) accumulated more than 85% of their total signal-to-noise ratio. Since more than 99.9% of the total signal-to-noise ratio is accumulated, we use a low-frequency cutoff of 7 Hz to generate the waveforms in our template banks.

tion is $(\Delta f/f) \sim 10^{-8}$ and can be neglected in search algorithms. Similarly, the time dependence of the antenna response due to the Earth’s rotation can be neglected as the match between a waveform that neglected the time variation and a waveform that accounted for the variation is 98-99%. For comparison, we show the same result for the proposed E.U. third-generation detector Einstein Telescope [42]. Einstein Telescope has a significantly lower seismic-Newtonian-noise wall than Cosmic Explorer and so searches must be pushed to lower frequencies to accumulate all of the possible signal-to-noise ratio. We focus on Cosmic Explorer in this work as we have found that existing methodologies are sufficient to effectively address the challenges presented by the increased low-frequency sensitivity of the third generation observatory. Einstein Telescope has a significantly lower seismic-Newtonian-noise wall than Cosmic Explorer and so searches must be pushed to lower frequencies to accumulate all of the possible signal-to-noise ratio. For an optimal search of Einstein Telescope, this may require addressing how to best account for the time-dependent detector response.

Using a 7 Hz low-frequency cutoff, we generate a tem-

plate bank that can be used to search for binary neutron star merger with component masses $1.0 \leq m_1, m_2 \leq 3.0$. We first generate a template bank for binaries with zero eccentricity and component spin using the standard hexagonal lattice method of template placement [43–46]. The template bank is constructed so that it has a fitting factor of 97% [47]. We find that the number of templates required for the CE1 (CE2) sensitivity is 130,000 (209,000) to achieve a fitting factor of 97%. A template bank generated using the Advanced LIGO sensitivity and a 10 Hz low-frequency cutoff contains 77,000 points. Since the CE1 (CE2) template banks are only a factor of 1.7 (2.8) larger than the equivalent template bank for Advanced LIGO, we do not expect significant computational challenges executing these searches. We certainly expect no obstacles to implementing real-time searches a decade or more from now when Cosmic Explorer will be operational.

Before constructing a template bank for binaries with eccentricity, we determine how effective the non-eccentric template bank is at detecting signals from eccentric binary neutron star sources. We perform the standard test of measuring the match between a random set of eccentric gravitational-wave signals and maximizing this over the template bank to obtain the bank’s fitting factor to a population of eccentric signals. To model eccentric sources, we use the LIGO Algorithm Library implementation [48] of TaylorF2Ecc, a frequency-domain post-Newtonian model with eccentric corrections. This waveform is accurate to 3.5 pN order in orbital phase [49], 3.5 pN order in the spin-orbit interactions [50], 2.0 pN order in spin-spin, quadrupole-monopole, and self-interactions of individual spins [51, 52], and 3.0 pN order in eccentricity [53]. To model non-eccentric waveforms, we use the restricted TaylorF2 approximant, accurate to the same post-Newtonian order.

We test the template bank against 120,000 simulated signals that have detector-frame component masses $1.0 \leq m_1, m_2 \leq 3.0$ and eccentricity $0 \leq e \leq 0.05$. The results of the simulation are shown in Fig. 3 and Fig. 4 for CE1 and CE2, respectively. If the population of neutron star binaries has eccentricity less than 0.004 (0.003) in CE1 (CE2), then the non-eccentric template banks achieve a fitting factor of 97% and are effectual. However, for sources with larger eccentricities the effectualness of the template bank begins to rapidly decline; the effectualness of a non-eccentric binary neutron star bank fails at an eccentricity an order of magnitude lower than that of Advanced LIGO [35, 36]. To recover these signals, it is necessary to construct a template bank that captures eccentricity. We consider this in the next section.

III. EXTENSION TO ECCENTRIC TEMPLATE BANKS

The number of templates in an eccentric bank will depend on the bandwidth of the detector and the upper

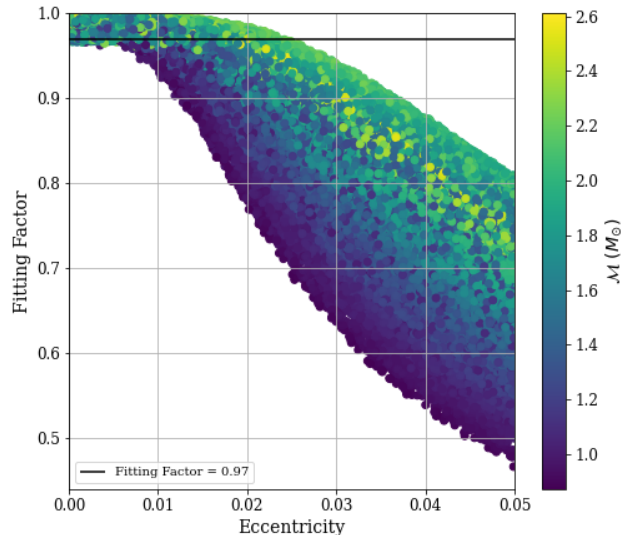


FIG. 3. The fitting factor as a function of eccentricity correlated with chirp mass for CE1. A non-eccentric template bank was used to calculate the fitting factor. For Cosmic Explorer the fitting factor decreases for increasing values of eccentricity. The non-eccentric template bank is effective in detecting eccentric systems with a fitting factor greater than 97% for $e \lesssim 0.004$.

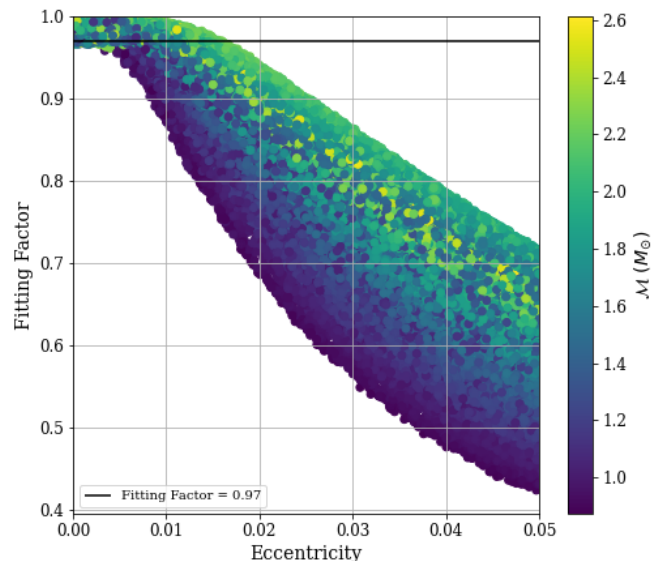


FIG. 4. As in Fig. 3, but we use the CE2 noise curve. A non-eccentric template bank was used to calculate the fitting factor. The non-eccentric template bank is effective in detecting eccentric systems with a fitting factor greater than 97% for $e \lesssim 0.003$.

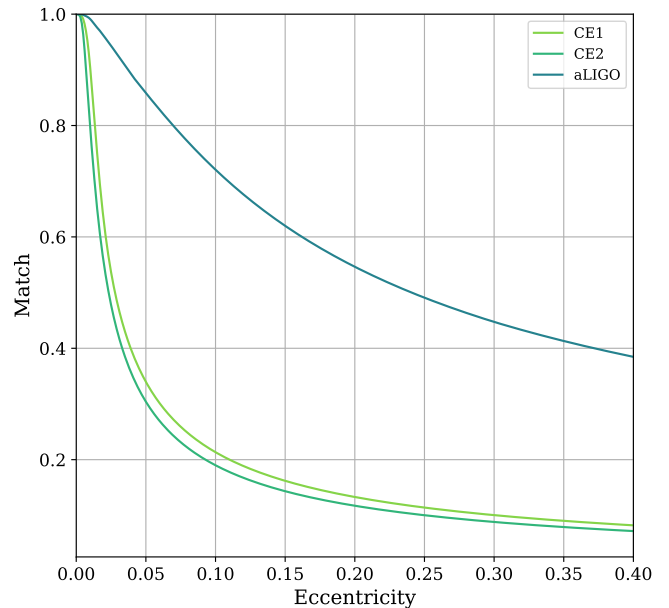


FIG. 5. The match as a function of eccentricity for Cosmic Explorer (CE1/CE2) and Advanced LIGO (aLIGO). This gives a representation of the match between an circular waveform and an eccentric waveform for various eccentricities. The match for Cosmic Explorer at an eccentricity of 0.05 is about a factor of 3 smaller than that of Advanced LIGO.

eccentricity boundary of the bank. To visualize the dependency on detector bandwidth, Fig. 5 shows the eccentricity ambiguity function for a $m_1 = m_2 = 1.4 M_\odot$ binary. This shows how quickly the loss in signal-to-noise ratio (match) changes as the eccentricity increases from 0 to 0.4 (referenced to 7 Hz). Without the use of eccentric templates, the match for CE1 (CE2) decreases to 34% (30%) at $e = 0.05$. In contrast, the Advanced LIGO match decreases much more slowly, reaching 38% at $e = 0.4$. Consequently, the density of an eccentric template bank will be significantly greater for Cosmic Explorer than Advanced LIGO.

Searches for eccentric binary neutron stars in Advanced LIGO used a template bank that covers the eccentricity range $0 \leq e \leq 0.4$ (referenced to 10 Hz) [33]; this bank contained 350,000 templates. To generate template banks of comparable density in eccentricity for Cosmic Explorer, we set the upper eccentricity of the template bank to $e = 0.05$ and keep the mass boundaries at $1.0 \leq m_1, m_2 \leq 3.0$ and the lower-frequency cutoff at 7 Hz. We then generated a template bank for eccentric gravitational-wave signals in this region using the stochastic placement method [37, 38] with a fitting factor of 96.5%.

We test the eccentric template bank against 25,000 simulated signals with detector-frame component masses

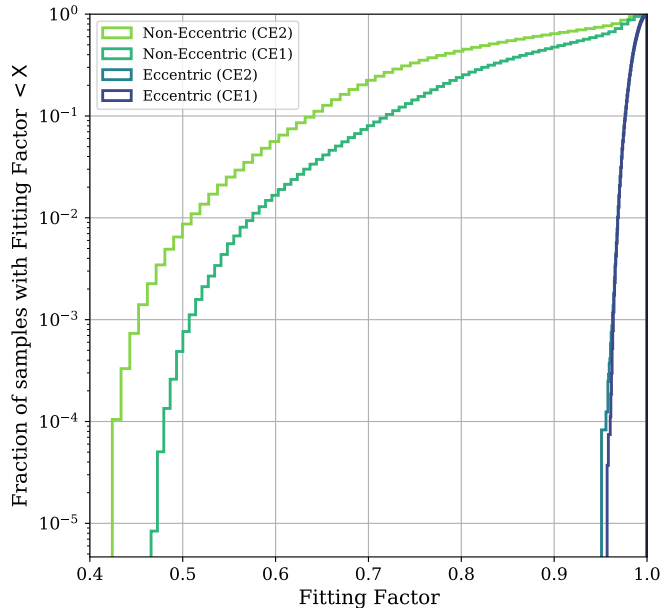


FIG. 6. A cumulative histogram that shows the fraction of points where the fitting factor is less than the value on the x-axis for each template bank. Using the eccentric template bank, a majority of the samples are at a fitting factor $\gtrsim 95\%$. For our eccentricity range, the eccentric template banks appear to do a better job at detecting eccentric systems than the non-eccentric template banks.

uniformly distributed between $1.0 \leq m_1, m_2 \leq 3.0$ and eccentricity uniformly distributed between $0 \leq e \leq 0.05$. The resulting fitting factor of these bank is shown in Fig. 6, with the fitting factor of the non-eccentric bank as reference. This result shows that the stochastic bank placement is effectual for signals with eccentricity in the target region, as all signals can be recovered with a fitting factor of $\gtrsim 95\%$ both the CE1 and CE2 banks. The number of eccentric templates generated using the CE1 (CE2) sensitivity is 1,900,000 (6,400,000), an order of magnitude larger than the non-eccentric template banks for CE and an order of magnitude larger than the the Advanced LIGO eccentric bank. We consider the size of a template bank with $e_{max} = 0.1$ to determine the increase in templates as eccentricity increases. A template bank with $0 \leq e \leq 0.1$ has 4,500,000 templates using the CE1 sensitivity, this is twice the size of the template bank we consider in this work. We expect that a bank of this size will present no computational challenges when Cosmic Explorer is operational in the 2030s; searches of similar magnitude are already regularly performed [54, 55].

IV. BINARY NEUTRON STAR PARAMETER ESTIMATION IN COSMIC EXPLORER

We can use our results to estimate the constraints that Cosmic Explorer will be able to place on the eccentricity of detected binary neutron stars with parameter estimation. We express this as the signal-to-noise ratio required to distinguish between an eccentric and circular binary at 90% confidence. This can be interpreted as the minimum detectable eccentricity at a given signal-to-noise ratio, or the upper limit that can be placed on the eccentricity of a circular a binary detected at a given signal-to-noise ratio.

To estimate this signal-to-noise required to distinguish between a circular binary and a binary with eccentricity e at 90% confidence, we use the method of Baird *et al.* [56]. This method relies on the fact that parameter estimation identifies regions of parameter space where a waveform is most consistent with the data. Ref. [56] uses the fact that high confidence regions in parameter estimation are associated with regions of high match between signal and template to obtain a relationship between the match and signal-to-noise ratio ρ , given by

$$M(h(\theta), h(\langle\theta\rangle)) \geq 1 - \frac{\chi_k^2(1-p)}{2\rho^2} \quad (1)$$

where k is the dimension of the parameter space of interest, $\chi_k^2(1-p)$ is the chi-square value for which there is $1-p$ probability of obtaining that value or larger. Here, we set $k = 4$ corresponding to intrinsic parameter space of an aligned spin binary neutron star merger with eccentricity $(m_1, m_2, \chi_{eff}, e)$, where χ_{eff} is the effective spin of the binary, and $p = 0.9$ for 90% confidence.

For Eq. (1) to provide a reasonable estimate of the signal-to-noise ratio, the match M must be maximized over the parameters of the signal. For eccentric binaries, there is a known degeneracy between the chirp mass $\mathcal{M} = (m_1 m_2)^{3/5} / (m_1 + m_2)^{1/5}$ of the binary and the eccentricity [31, 57]. Full parameter estimation naturally explores the likelihood and this degeneracy. Here, we use our method of eccentric template placement to place a fine grid of templates and brute-force maximize the match over this template bank to account for the chirp mass-eccentricity degeneracy.

Using this method, we estimate Cosmic Explorer's ability to constrain the eccentricity of a $m_1 = m_2 = 1.4 M_\odot$ binary as follows: Using a low frequency cutoff of 7 Hz, we generate a template bank with binary neutron star component masses $1.399 \leq m_1, m_2 \leq 1.401$, eccentricity $0 \leq e \leq 0.05$, an upper-frequency cutoff of 4096 Hz, and a minimal match of 99.9999%. We measure the match between a simulated eccentric gravitational-wave signal with component masses $m_1 = m_2 = 1.4$ and eccentricity $0 \leq e \leq 0.05$ and maximize over the chirp mass in the template bank to get the signal-to-noise ratio. From this we determine the signal-to-noise ratio needed to reach a 90% confidence interval [56] to measure the eccentricity.

V. CONCLUSION

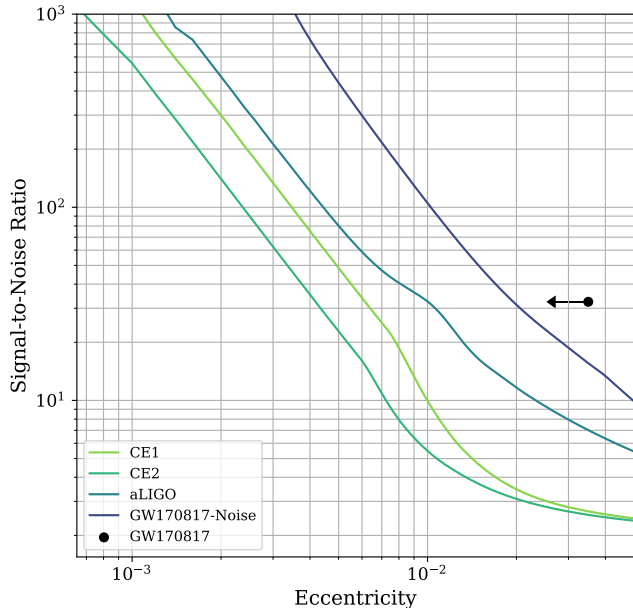


FIG. 7. The signal-to-noise ratio as a function of eccentricity. The black dot is at an signal-to-noise ratio of 32.4 [20] and eccentricity of 0.024 at 90% confidence [31]. For each detector, we show the signal-to-noise ratio needed to resolve the signal with eccentricity on the x-axis at 90% confidence. For $e \geq 10^{-3}$, a signal-to-noise ratio of 8 would be needed to resolve the signal at 90% confidence in Cosmic Explorer. As the eccentricity decreases, the signal-to-noise ratio needed to resolve the signal increases.

We apply the above method using the CE1, CE2, and Advanced LIGO design noise curves to obtain the signal-to-noise ratio as a function of eccentricity required to distinguish between a circular and eccentric binary. To check the accuracy of our estimation, we also compute this function using the detector noise around the time of GW170817 and compare the Baird *et al.* estimate to the 90% upper limit on the eccentricity of GW170817 computed using full parameter estimation [31]. These results are shown in Fig. 7. First, we note that our method provides a reasonable approximation when comparing to GW170817 and as the eccentricity increases the signal-to-noise ratio needed to resolve the signal decreases, as expected. Our results suggest that for $e \geq 8 \times 10^{-3}$ (8×10^{-4}), a minimum signal-to-noise ratio of 8 (800) would be needed to resolve the signal at 90% confidence in CE1 and CE2. This is an order of magnitude better than expected from Advanced LIGO operating at design sensitivity.

Our analysis used circular and eccentric template banks to determine the ability of Cosmic Explorer to detect eccentric binary neutron stars and to measure their eccentricity. The circular template banks are effective for detecting eccentric binaries with $e \leq 0.004$ ($e \leq 0.003$) in CE1 (CE2) at a reference frequency of 7 Hz. However, for larger eccentricities a template bank containing circular and eccentric waveform templates is required. This estimate is an order of magnitude smaller than estimates for Advanced LIGO [35, 36]. We determine the signal-to-noise ratio needed to constrain the eccentricity of a detected neutron star binary signal with 90% confidence. We also find that in Cosmic Explorer to measure binary neutron star with eccentricity $\gtrsim 8 \times 10^{-3}$ (8×10^{-4}) a signal-to-noise ratio of 8 (800) is needed to resolve the signal at a reference frequency of 7 Hz (90% confidence). Accurately constraining the eccentricity of the binary would provide valuable information on the formation of these mergers.

The computational cost of searches with template banks containing higher eccentricities will be challenging in Cosmic Explorer today as the density of the template bank increases with increasing eccentricity (see Fig. 2 of Ref. [33]). However, improvements in technology by the 2030s may make these searches a possibility. Along with the high computational cost, current waveform models for eccentricity break down for $e \geq 0.4$. To accurately search for higher eccentricity neutron-star binaries models that extend to high eccentricities will need to be developed or a burst search will need to be used. To extend this work to binary black holes, waveforms that better model the merger-ringdown will need to be accessible [58–62] and will be left for future work. Understanding the constraints that future observational limits place on eccentric binary formation channels will require computation of the rate as a function of eccentricity from population synthesis. As Cosmic Explorer will be able to aid in the understanding of the physics of binary neutron star mergers it is important to accurately constrain the eccentricity as the number of mergers increases.

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