

WhiskyMHD: a new numerical code for general relativistic magnetohydrodynamics

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Abstract. The accurate modelling of astrophysical scenarios involving compact objects and magnetic fields, such as the collapse of rotating magnetized stars to black holes or the phenomenology of γ -ray bursts, requires the solution of the Einstein equations together with those of general-relativistic magnetohydrodynamics. We present a new numerical code developed to solve the full set of general-relativistic magnetohydrodynamics equations in a dynamical and arbitrary spacetime with high-resolution shock-capturing techniques on domains with adaptive mesh refinements. After a discussion of the equations solved and of the techniques employed, we present a series of testbeds carried out to validate the code and assess its accuracy. Such tests range from the solution of relativistic Riemann problems in flat spacetime, over to the stationary accretion onto a Schwarzschild black hole and up to the evolution of oscillating magnetized stars in equilibrium and constructed as consistent solutions of the coupled Einstein-Maxwell equations.

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1. Introduction

Magnetic fields are ubiquitous in astrophysical objects and can play an important role, especially in those scenarios involving compact objects such as neutron stars and black holes. An accurate and consistent modelling of these scenarios, which are extreme both for the gravitational and the electromagnetic fields, cannot be done analytically and perturbative methods are also of limited validity. In the absence of symmetries, in fact, no dynamical and analytic solutions are known and it is only through the full solution of the equations of general-relativistic magnetohydrodynamics (GRMHD) that one can hope to improve our knowledge of these objects under realistic conditions.

As in general-relativistic hydrodynamics, the work in this area of research has started, more than 30 years ago with the pioneering work of Wilson [1] in lower spatial dimensions (see [2] for a review of the technical and scientific progress in relativistic hydrodynamics). Unlike general relativistic hydrodynamics, however, where both technical issues and scientific investigations have now reached an advanced stage of sophistication and accuracy, progress in GRMHD has yet to reach a comparable level of maturity. Indeed, it was only over the last few years that the slow but steady progress in GRMHD has seen a renewed burst of activity, with a number of groups developing a variety of numerical codes solving the equations of GRMHD under different approaches and approximations. This is partly due to the considerable added complexity of the set of equations to be solved in GRMHD and, partly, to the fact that only recently sufficient computational resources have become available to tackle this problem in two or three spatial dimensions and with sufficient resolution.

It does not come as a surprise, therefore, that most of the numerical codes developed in the last decade have been based on the same non-conservative formulation of the GRMHD equations introduced by Wilson, solving them on a fixed background to study accretion disks around black holes. In [3, 4], for instance, the effects of a Kerr black hole on magnetohydrodynamical accretion have been studied, with particular attention being paid to the transfer of energy and angular momentum. Koide *et al.* [5], on the other hand, have developed a numerical code based on the artificial-viscosity approach proposed by Davis [6] to perform the first simulations of jet formation in General Relativity [7] and to study the possibility of extracting the rotational energy from a Kerr black hole [8, 9]. Furthermore, a distinct numerical code has been constructed by De Villiers and Hawley [10] using the formulation proposed in [11, 12] to carry out a series of studies on accretion flows around Kerr black holes [13, 14, 15].

It is only rather recently that different groups have started to recast the system of GRMHD equations into a conservative form in order to benefit of the use of high-resolution shock-capturing schemes (HRSC) [16, 17, 18, 19, 20]. Such schemes, we recall, are essential for a correct representation of shocks, whose presence is expected in several astrophysical scenarios and in particular in those involving compact objects. Two mathematical results corroborate this view, with the first one stating that a stable scheme converges to a weak solution of the hydrodynamical equations [21], and with the second one showing that a non-conservative scheme will converge to the wrong weak solution in the presence of a shock [22].

All of these newly developed codes [16, 17, 18, 19] that make use of HRSC methods have been so far applied to the study of accretion problems onto black holes, for which the self-gravity of the accreting material introduces very small corrections to the spacetime and fixed spacetime backgrounds can be used satisfactorily.

Approaches alternative to that of constructing GRMHD codes have instead been based on the use of a modified Newtonian gravitational potential to mimic general relativistic effects without having to solve numerically Einstein equations (see [23] for an application to magnetorotational collapse of stellar cores) or on the use of different numerical methods, such as smoothed particle hydrodynamics and artificial viscosity, to study the merger of binary neutron star systems as a possible engine for short γ -ray bursts [24]. Although the use of these approximations has made it possible to investigate this astrophysical scenario for the first time including details about the microphysics, it is clear that equally important corrections coming from the dynamical evolution of the spacetime need to be introduced when trying to model the phenomenology that is thought to be behind γ -ray bursts. As a first step in this direction, two codes were recently developed to solve the full set of GRMHD equations on a dynamical background [25, 26]. These codes, in particular, were used to perform the first study, in two spatial dimensions, of the collapse of magnetized differentially rotating neutron stars [27, 28, 29] which are thought to be good candidates for short γ -ray bursts.

Here, we present WhiskyMHD, a new three-dimensional numerical code in Cartesian coordinates developed to solve the full set of GRMHD equation on a dynamical background. The code is based on the use of high-resolution shock-capturing techniques on domains with adaptive mesh refinements, following an approach already implemented with success in the general-relativistic hydrodynamics code Whisky [30], and which has been used in the study of several astrophysical scenario with particular attention to gravitational-wave emission from compact objects.

The paper is organized as follows: in Sect. 2 we recall the equations of GRMHD and the form they assume when recast in a conservative form, while in Sect. 3 we discuss in detail the numerical methods adopted for their solution. Section 4 is dedicated to the series of testbeds the code has passed both in special and in general relativistic conditions. Finally, Sect. 5 offers

a summary of the results and an overview on our future projects.

Throughout the paper we use a spacelike signature $(-, +, +, +)$ and a system of units in which $c = G = M_{\odot} = 1$. Greek indices are taken to run from 0 to 3, Latin indices from 1 to 3 and we adopt the standard convention for the summation over repeated indices. Finally we indicate 3-vectors with an arrow and use bold letters to denote 4-vectors and tensors.

2. Formulation of the equations

We adopt the usual 3-dimensional foliation of the spacetime so that the line element reads

$$ds^2 = -(\alpha^2 - \beta^i \beta_i) dt^2 + 2\beta_i dx^i dt + \gamma_{ij} dx^i dx^j, \quad (1)$$

where β^i is the *shift* vector, α is the *lapse* function and γ_{ij} are the spatial components of the four-metric $g_{\mu\nu}$.

As its predecessor Whisky [30], the WhiskyMHD code benefits of the Cactus computational toolkit [31] which provides an infrastructure for the parallelization and the I/O of the code, together with several methods for the solution of the Einstein equations. As a result, at each timestep our new code solves the MHD equations while Cactus provides the evolution of the metric quantities. The evolution of the field components is done using the NOK formulation [32, 33, 34] and details about its numerical implementation can be found in [35, 36, 37].

Here too we make use of the so-called ‘‘Valencia formulation’’ [38, 39] which was originally developed as a 3 + 1 conservative Eulerian formulation of the general relativistic hydrodynamic equations, but which has been recently extended to the case of GRMHD [18]. Following [18] we define the Eulerian observer as the one moving with four velocity \mathbf{n} perpendicular to the hypersurfaces of constant t at each event in the spacetime. This observer measures the following three-velocity of the fluid

$$v^i = \frac{h_{\mu}^i u^{\mu}}{-u^{\mu} n_{\mu}} = \frac{u^i}{W} + \frac{\beta^i}{\alpha}, \quad (2)$$

where $h_{\mu\nu} \equiv g_{\mu\nu} + n_{\mu} n_{\nu}$ is the projector orthogonal to \mathbf{n} , \mathbf{u} is the four-velocity of the fluid and $-u^{\mu} n_{\mu} = \alpha u^0 = W$ is the Lorentz factor which satisfies the usual relation $W = 1/\sqrt{1 - v^2}$, where $v^2 \equiv \gamma_{ij} v^i v^j$. The covariant components of the three-velocity are simply given by $v_i = u_i/W$.

2.1. Maxwell equations

The electromagnetic field is completely described by the Faraday electromagnetic tensor field $F^{\mu\nu}$ obeying Maxwell equations (cfr [40])

$$\nabla_{\nu} {}^*F^{\mu\nu} = 0, \quad (3)$$

$$\nabla_{\nu} F^{\mu\nu} = 4\pi \mathcal{J}^{\mu}, \quad (4)$$

where ∇ is the covariant derivative with respect to the four-metric g , \mathcal{J} is the charge current four-vector and *F is the dual of the electromagnetic tensor defined as

$${}^*F^{\mu\nu} = \frac{1}{2} \eta^{\mu\nu\lambda\delta} F_{\lambda\delta}, \quad (5)$$

$\eta^{\mu\nu\lambda\delta}$ being the Levi-Civita pseudo-tensor. A generic observer with four-velocity U will measure a magnetic field \mathbf{B} and an electric field \mathbf{E} given by

$$E^{\alpha} \equiv F^{\alpha\beta} U_{\beta}, \quad (6)$$

$$B^{\alpha} \equiv {}^*F^{\alpha\beta} U_{\beta}, \quad (7)$$

and the charge current four-vector \mathcal{J} can be in general expressed as

$$\mathcal{J}^\mu = qu^\mu + \sigma F^{\mu\nu} u_\nu, \quad (8)$$

where q is the proper charge density and σ is the electric conductivity.

Hereafter we will assume that our fluid is a perfect conductor (ideal MHD limit) and so that $\sigma \rightarrow \infty$ and $F^{\mu\nu} u_\nu = 0$ (i.e. the electric field measured by the comoving observer is zero) in order to keep the current finite. In this limit, the electromagnetic tensor and its dual can be written exclusively in terms of the magnetic field \mathbf{b} measured in the comoving frame

$$F^{\nu\sigma} = \eta^{\alpha\mu\nu\sigma} b_\alpha u_\mu, \quad {}^*F^{\mu\nu} = b^\mu u^\nu - b^\nu u^\mu. \quad (9)$$

with the Maxwell equations taking the simple form

$$\nabla_\nu {}^*F^{\mu\nu} = \frac{1}{\sqrt{-g}} \partial_\nu (\sqrt{-g} (b^\mu u^\nu - b^\nu u^\mu)) = 0, \quad (10)$$

In order to express these equations in terms of quantities measured by an Eulerian observer, we need to compute the relation between the magnetic field measured by the comoving and by the Eulerian observers, respectively \mathbf{b} and \mathbf{B} . To do that we introduce the projection operator $P_{\mu\nu} \equiv g_{\mu\nu} + u_\mu u_\nu$ orthogonal to \mathbf{u} . If we apply this operator to the definition of the magnetic field \mathbf{B} measured by an Eulerian observer, we can easily derive the following relations

$$b^0 = \frac{WB^i v_i}{\alpha}, \quad b^i = \frac{B^i + \alpha b^0 u^i}{W}, \quad b^2 \equiv b^\mu b_\mu = \frac{B^2 + \alpha^2 (b^0)^2}{W^2}, \quad (11)$$

where $B^2 \equiv B^i B_i$. The time component of equations (10) provides the divergence-free condition

$$\partial_i \tilde{B}^i = 0, \quad (12)$$

where $\tilde{B}^i \equiv \sqrt{\gamma} B^i$ and γ is the determinant of γ_{ij} . The spatial components of equations (10), on the other hand, yield the induction equations for the evolution of the magnetic field

$$\partial_t (\tilde{B}^i) = \partial_j (\tilde{v}^i \tilde{B}^j - \tilde{v}^j \tilde{B}^i), \quad (13)$$

where $\tilde{v}^i \equiv \alpha v^i - \beta^i$.

2.2. Conservation equations

The evolution equations for the rest-mass density ρ , the specific internal energy ϵ and for the three-velocity v^i can be computed, as done in relativistic hydrodynamics, from the conservation of the baryon number

$$\nabla_\nu (\rho u^\nu) = 0, \quad (14)$$

from the conservation of the energy-momentum

$$\nabla_\nu T^{\mu\nu} = 0, \quad (15)$$

and from an equation of state (EOS) relating the gas pressure p to the rest-mass density ρ and to the specific internal energy ϵ . We also assume that the fluid is perfect so that the total energy momentum tensor, including the contribution from the magnetic field, is given by

$$T^{\mu\nu} = (\rho h + b^2) u^\mu u^\nu + \left(p + \frac{b^2}{2} \right) g^{\mu\nu} - b^\mu b^\nu, \quad (16)$$

where $h = 1 + \epsilon + p/\rho$ is the specific relativistic enthalpy.

Following [18] and in order to make use of HRSC methods, we rewrite equations (14), (15) and (13) in the following conservative form

$$\frac{1}{\sqrt{-g}} [\partial_t(\sqrt{\gamma}\mathbf{F}^0) + \partial_i(\sqrt{-g}\mathbf{F}^i)] = \mathbf{S}, \quad (17)$$

where $\mathbf{F}^0 \equiv (D, S_j, \tau, B^k)^T$ is the vector of the conserved variables measured by the Eulerian observer, \mathbf{F}^i are the fluxes

$$\mathbf{F}^i = \begin{pmatrix} D\tilde{v}^i/\alpha \\ S_j\tilde{v}^i/\alpha + (p + b^2/2)\delta_j^i - b_j B^i/W \\ \tau\tilde{v}^i/\alpha + (p + b^2/2)v^i - \alpha b^0 B^i/W \\ B^k\tilde{v}^i/\alpha - B^i\tilde{v}^k/\alpha \end{pmatrix}, \quad (18)$$

and \mathbf{S} are the source terms

$$\mathbf{S} = \begin{pmatrix} 0 \\ T^{\mu\nu}(\partial_\mu g_{\nu j} - \Gamma_{\nu\mu}^\delta g_{\delta j}) \\ \alpha(T^{\mu 0}\partial_\mu \ln \alpha - T^{\mu\nu}\Gamma_{\nu\mu}^0) \\ 0^k \end{pmatrix}, \quad (19)$$

where

$$D \equiv \rho W, \quad (20)$$

$$S_j \equiv (\rho h + b^2)W^2 v_j - \alpha b^0 b_j, \quad (21)$$

$$\tau \equiv (\rho h + b^2)W^2 - (p + \frac{b^2}{2}) - \alpha^2 (b^0)^2 - D, \quad (22)$$

and $0^k = (0, 0, 0)^T$. While ready to make use of arbitrary EOS, the ones presently implemented in the code have a rather simple form and are given by either the polytropic EOS

$$p = K\rho^\Gamma, \quad (23)$$

$$\epsilon = \frac{p}{\rho(\Gamma - 1)}, \quad (24)$$

or by the ideal-fluid EOS

$$p = (\Gamma - 1)\rho\epsilon, \quad (25)$$

where K is the polytropic constant and Γ is the adiabatic exponent. In the case of the polytropic EOS (23), $\Gamma = 1 + 1/N$, where N is the polytropic index.

3. Numerical methods

As in the Whisky code, the evolution equations are integrated in time using the method of lines [41], which reduces the partial differential equations (17) to a set of ordinary differential equations that can be evolved using standard numerical methods, such as Runge-Kutta or the iterative Cranck-Nicholson schemes [42, 43].

3.1. Approximate Riemann solver

As mentioned in the Introduction, WhiskyMHD makes use of HRSC schemes based on the use of Riemann solvers to compute the fluxes between the numerical cells. More specifically, we have implemented the Harten-Lax-van Leer-Einfeldt (HLLC) *approximate* Riemann solver [44] which is simply based on the calculation of the eigenvalues of eqs. (17) and it does not require a basis of eigenvectors. In the HLLC formulation the flux at the interface between two numerical cells is therefore computed as

$$\mathbf{F}^i = \frac{c_{min}\mathbf{F}_r^i + c_{max}\mathbf{F}_l^i - c_{max}c_{min}(\mathbf{F}_r^0 - \mathbf{F}_l^0)}{c_{max} + c_{min}}, \quad (26)$$

where \mathbf{F}_r^μ and \mathbf{F}_l^μ are computed from the values of the primitive variables reconstructed at the right and left side of the interface \mathbf{P}_r and \mathbf{P}_l , respectively. The coefficients $c_{max} \equiv \max(0, c_{+,r}, c_{+,l})$, $c_{min} \equiv -\min(0, c_{-,r}, c_{-,l})$ and $c_{\pm,r}$, $c_{\pm,l}$ are instead the maximum left- and right-going wave speeds computed from \mathbf{P}_r and \mathbf{P}_l . In our implementation \mathbf{P}_r and \mathbf{P}_l are computed using a second order TVD slope limited method which can be used with different limiters such as minmod, Van Leer and MC [41].

An alternative to the use of an *approximate* Riemann solver could have been the use of the *exact* Riemann solver recently developed in GRMHD. We recall that in relativistic hydrodynamics the exact solution is found after expressing all of the quantities behind each wave as functions of the value of the unknown gas pressure p at the contact discontinuity. In this way, the problem is reduced to the search for the value of the pressure that satisfies the jump conditions at the contact discontinuity (see [45, 46, 47, 48] for the details). The procedure for the exact solution of the Riemann problem in relativistic MHD is based instead on the use of an hybrid approach that makes use of different set of unknowns depending on the wave. In the case of fast-magnetosonic waves, both shocks or rarefactions, all the variables behind the waves are rewritten as functions of the total pressure, *i.e.* $p + b^2/2$, while behind slow magnetosonic shocks or rarefactions the components of the magnetic field tangential to the discontinuity are used to compute all the other variables. The use of this strategy was essential in order to reduce the problem to the solution of a closed system of equations that can be solved with standard numerical routines such as Newton-Raphson schemes (see [49] for the details). The numerical code computing the exact solution is freely available from the authors upon request and it is now becoming a standard tool in the testing of both special and general relativistic MHD codes.

While the use of an exact Riemann solver could provide the solution of the discontinuous flow at each cell interface with arbitrary accuracy, the exact solver presented in [49] is still computationally too expensive to be used in ordinary multidimensional codes and we have found the HLLC algorithm to be sufficiently accurate for the resolution used in our tests.

3.1.1. Calculation of the eigenvalues An important difference with respect to relativistic-hydrodynamic codes is that in GRMHD the calculation of the eigenvalues required by the HLLC solver is made more complicated by the solution of a quartic equation. The characteristic structure of GRMHD equations is analyzed in detail in [40] and we simply report here the expressions for the calculation of the seven wave speeds associated with the entropic, the Alfvén, the fast and slow magnetosonic waves.

More specifically, the characteristic speed λ of the entropic waves is simply given by $\tilde{v} = \alpha v^i - \beta^i$, while the values for the left- and right-going Alfvén waves are

$$\lambda_{1,2}^i = \frac{b^i \pm \sqrt{(\rho h + b^2)u^i}}{b^0 \pm \sqrt{(\rho h + b^2)u^0}}. \quad (27)$$

Similarly, the four speeds that are associated with the fast and slow magnetosonic waves, and that are required in the calculation of the fluxes, can be obtained by the solution of the following quartic equation in each direction i for the unknown λ

$$\rho h \left(\frac{1}{c_s^2} - 1 \right) a^4 - \left(\rho h + \frac{b^2}{c_s^2} \right) a^2 G + \mathcal{B}^2 G = 0, \quad (28)$$

where

$$a \equiv \frac{W}{\alpha} (-\lambda + \alpha v^i - \beta^i), \quad (29)$$

$$\mathcal{B} \equiv b^i - b^0 \lambda, \quad (30)$$

$$G \equiv \frac{1}{\alpha^2} [-(\lambda + \beta^i)^2 + \alpha^2 \gamma^{ii}], \quad (31)$$

and c_s is the sound speed (Note that the convention on repeated indices should not be used for the last term in the expression for G , i.e. γ^{ii}). In the degenerate case in which $B^i = 0$, eq. (28) can be reduced to a simple quadratic equation that is solved analytically. In the more general case, however, eq. (28) cannot be reduced to the product of two quadratic equations as in Newtonian MHD and different methods are implemented in the code in order to solve it. The first one simply makes use of the analytic expression for a quartic equation [50], while the other two search the solution numerically either through an eigenvalue method or through a Newton-Raphson iteration [51]. The latter has shown to be the most accurate and robust and it is the one used by default.

We have also implemented an approximate method for the calculation of the eigenvalues associated with the fast magnetosonic waves (which are the only two roots needed by HLLE) which was introduced in [52] and which reduces the original quartic to a quadratic equation, that can be solved analytically, by imposing $B^i = 0$ and $B^j v_j = 0$ in equation (28). The values computed in this way differ by less than 1% with respect to the exact values and we have used them in those situations in which the solution of the quartic can be complicated by the presence of degeneracies or when two of the roots are very close to each other.

3.2. Constrained-Transport Scheme

Although an exact solution of eqs. (13) would guarantee that the constraint condition (12) is also satisfied identically and all times, any numerical solution of the induction equations (13) will inevitably produce a violation of the divergence-free condition which, in turn, may lead to unphysical results or even to the development of instabilities [53]. To avoid this problem several numerical methods were developed in the past starting from the so-called “staggered mesh magnetic field transport algorithm” first proposed by Yee [54] and then implemented in an artificial-viscosity scheme with the name of “constrained-transport” scheme (CT) by Evans & Hawley [55].

A modified version of the CT scheme [55] which is based on the use of the fluxes computed with HRSC methods has been proposed by Balsara & Spicer [56] (“flux-CT”) and is the one implemented in our code because of its simplicity and computational efficiency. An interested reader is referred to [57] for other possible methods to enforce the divergence-free condition with HRSC schemes.

We recall that the flux-CT method is based on the relation that exists in ideal MHD between the fluxes of the magnetic field \vec{B} and the value of the electric field $\vec{E} \equiv -\vec{v} \times \vec{B}$. In particular, if we define $\tilde{\mathbf{F}}^i \equiv \alpha \sqrt{\gamma} \mathbf{F}^i$ then the following relations hold

$$E^x = \tilde{F}^z (\tilde{B}^y) = -\tilde{F}^y (\tilde{B}^z), \quad (32)$$

$$E^y = -\tilde{F}^z (\tilde{B}^x) = \tilde{F}^x (\tilde{B}^z) , \quad (33)$$

$$E^z = \tilde{F}^y (\tilde{B}^x) = -\tilde{F}^x (\tilde{B}^y) , \quad (34)$$

where $\tilde{F}^i (\tilde{B}^j) \equiv \tilde{v}^i \tilde{B}^j - \tilde{v}^j \tilde{B}^i$. The induction equation (13) can then be written as

$$\partial_t \vec{\tilde{B}} + \vec{\nabla} \times \vec{E} = 0 . \quad (35)$$

Taking the surface integral of (35) across a surface Σ between two numerical cells, Stokes' theorem yields

$$\partial_t \int_{\Sigma} \vec{\tilde{B}} \cdot d\vec{\Sigma} + \int_{\partial\Sigma} \vec{E} \cdot \vec{l} = 0 , \quad (36)$$

where \vec{l} is the unit vector parallel to the surface boundary $\partial\Sigma$. Considering for simplicity the x -direction, with the surface Σ as located at $(i + \frac{1}{2}, j, k)$ and the integers (i, j, k) denoting the cell centers on our discrete grid (see figure 1), we can define

$$\tilde{B}_{i+\frac{1}{2},j,k}^x \equiv \frac{1}{\Delta y \Delta z} \int_{\Sigma} \vec{\tilde{B}} \cdot d\vec{\Sigma} , \quad (37)$$

and use the finite-difference expression of eq. (36) to obtain

$$\partial_t \tilde{B}_{i+\frac{1}{2},j,k}^x = - \left(E_{i+\frac{1}{2},j,k-\frac{1}{2}}^y - E_{i+\frac{1}{2},j,k+\frac{1}{2}}^y \right) / \Delta z - \left(E_{i+\frac{1}{2},j+\frac{1}{2},k}^z - E_{i+\frac{1}{2},j-\frac{1}{2},k}^z \right) / \Delta y , \quad (38)$$

where the values of the electric field on the edges of the surface are simply computed taking the arithmetic mean of the fluxes across the surfaces that have that edge in common [cf (32)–(34)], e.g.

$$E_{i+\frac{1}{2},j,k+\frac{1}{2}}^y = \frac{1}{4} \left(\tilde{F}_{i+\frac{1}{2},j,k}^x + \tilde{F}_{i+\frac{1}{2},j,k+1}^x - \tilde{F}_{i,j,k+\frac{1}{2}}^z - \tilde{F}_{i+1,j,k+\frac{1}{2}}^z \right) , \quad (39)$$

where the fluxes \tilde{F}^i are those computed with the approximate Riemann solver.

Since we are using HRSC methods, all the quantities are located at cells centers but in equation (38) we are effectively evolving the magnetic field at the surfaces between the cells. The relation between these two different values of the magnetic field is given by a simple average

$$\tilde{B}_{i,j,k}^x = \frac{1}{2} \left(\tilde{B}_{i+\frac{1}{2},j,k}^x + \tilde{B}_{i-\frac{1}{2},j,k}^x \right) , \quad (40)$$

$$\tilde{B}_{i,j,k}^y = \frac{1}{2} \left(\tilde{B}_{i,j+\frac{1}{2},k}^y + \tilde{B}_{i,j-\frac{1}{2},k}^y \right) , \quad (41)$$

$$\tilde{B}_{i,j,k}^z = \frac{1}{2} \left(\tilde{B}_{i,j,k+\frac{1}{2}}^z + \tilde{B}_{i,j,k-\frac{1}{2}}^z \right) . \quad (42)$$

To demonstrate that this method guarantees that $\nabla \cdot \vec{\tilde{B}} = 0$ and will not grow in time, we can integrate over the volume of a numerical cell this constraint and then using the Gauss theorem we obtain

$$\int_{\Delta V} \nabla \cdot \vec{\tilde{B}} dV = \sum_{i=1}^6 \int_{\Sigma_i} \vec{\tilde{B}} \cdot d\vec{\Sigma} , \quad (43)$$

where the sum is taken over all the six faces Σ_i that surround the cell. Taking now the time derivative of this expression and using eq. (36) we obtain

$$\partial_t \int_{\Delta V} \nabla \cdot \vec{\tilde{B}} dV = - \sum_{i=1}^6 \int_{\partial\Sigma_i} \vec{E} \cdot \vec{l} , \quad (44)$$

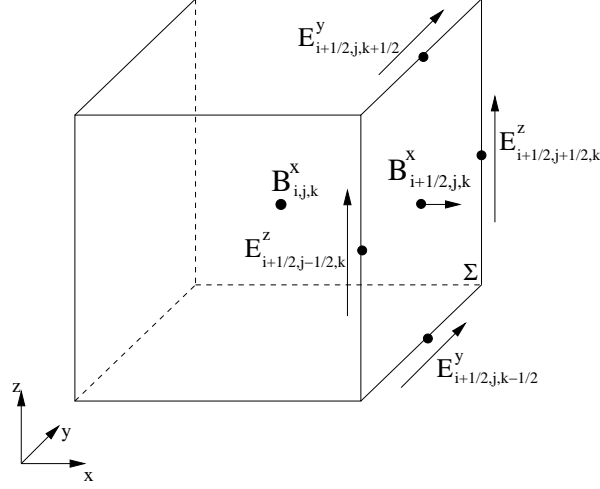


Figure 1. Schematic diagram of the data needed for the CT scheme. The evolution of $\tilde{B}^x_{i+1/2,j,k}$ is determined by the values of the electric field \vec{E} at the edges of the surface Σ located at $(i + 1/2, j, k)$.

and the sum on the right-hand side gives exactly zero since the value of $\vec{E} \cdot \vec{l}$ for the common edge of two adjacent faces has a different sign.

3.3. Primitive-variables recovery

Because the fluxes \mathbf{F}^i depend on the primitive variables \mathbf{P} and not on the evolved conservative variables \mathbf{F}^0 , the values for the primitive variables need to be recovered after each timestep and at each gridpoint. With the exception of the magnetic field variables B^i , the complexity of the system of equations to be solved prevents from an analytic solution relating the primitive to the conservative variables through simple algebraic relations and thus the system of equations (20)–(22) needs to be solved numerically. Several methods are available for this, the most obvious (and expensive) one consisting of solving the full set of 5 equations given by the expressions for (D, S_j, τ) in the 5 unknowns (ρ, v^i, ϵ) ; we refer to this as to the 5D method. Alternatively, and under certain conditions, it is possible to reduce the set of equations to be solved to a couple of nonlinear equations (2D method) or even to a single one (1D method). We review them briefly in the following Sections but a more detailed discussion can be found in [58].

3.3.1. 2D method The following procedure is the same used in [18] and it is an extension to full General Relativity of the method developed in [59] in special relativity. The idea is to take the modulus $S^2 = S^j S_j$ of the momentum instead of the expression for its three components reducing the total number of equations that one has to solve. Using the relations (11) it is possible to write S^2 as

$$S^2 = (Z + B^2)^2 \frac{W^2 - 1}{W^2} - (2Z + B^2) \frac{(B^i S_i)^2}{Z^2}, \quad (45)$$

where $Z \equiv \rho h W^2$. It is also possible to rewrite the equation for the total energy in a similar way

$$\tau = Z + B^2 - p - \frac{B^2}{2W^2} - \frac{(B^i S_i)^2}{2Z^2} - D . \quad (46)$$

Using then the definition of $D = \rho W$, eqs. (45) and (46) form a closed system for the two unknowns p and W , assuming the function $h = h(\rho, p)$ is provided. When using a polytropic EOS [i.e. eq. (23)], the integration of the total energy equation is not necessary (the energy density can be computed algebraically from other quantities) and the system reduces to the numerical solution of the equation (45). Once the roots for W , p and $\rho = D/W$ are found, it is possible to compute the values of v_i using the definition of the momentum S_i

$$v_i = \frac{B_i(B^j S_j) + S_i Z}{Z(B^2 + Z)} . \quad (47)$$

3.3.2. 1D method The basic idea of this method is to consider also the gas pressure p as a function of W reducing the total number of equations that must be solved numerically. When using an ideal-fluid [i.e. eq. (25)], Z can in fact be rewritten as

$$Z = DW + \frac{\Gamma}{\Gamma - 1} p(W) W^2 . \quad (48)$$

Using equation (48) it is possible to rewrite (46) as a cubic equation for $p(W)$ which admits only one physical solution. So at the end we need only to solve equation (45) for the only unknown W . Having obtained W , we can then compute $p = p(W)$ and the other quantities in the same way as done in the 2D method.

3.4. Atmosphere treatment

As already done in the Whisky code and in other full GRMHD codes [26, 25] we avoid the presence of vacuum regions in our domain by imposing a floor value to the rest-mass density. This is necessary because the routines that recover the primitive variables from the conservative ones may fail to find a physical solution if the rest-mass density is too small. The floor value used for the tests reported here is $\rho_{\text{atm}} = 10^{-7} \times \max(\rho_0)$, with ρ_0 being the value of the rest-mass density at $t = 0$, but a floor which is two orders of magnitude smaller works equally well. In practice, for all of the numerical cells at which $\rho \leq \rho_{\text{atm}}$, we simply set $\rho = \rho_{\text{atm}}$, $v^j = 0$, and do not modify the magnetic field. This is different from what done in other codes (e.g., [26, 25]), which set to zero the initial value of the magnetic field in the low density regions.

3.5. Excision

Many interesting astrophysical scenarios involve the presence of black holes and so of regions of spacetime where singularities are present. These regions are causally disconnected from the rest of the physical domain and the values of the fields inside should not affect the zone outside the event horizon. This is not true in numerical codes where it can happen that some information from inside the event horizon is used to compute the values of the variables outside. In order to avoid this, excision algorithms were developed in general relativistic hydrodynamics and they are based on the use of some kind of boundary condition applied to the boundary between the excised zone, where the equations are no more solved, and the domain outside. As already done in the Whisky code we apply a zeroth-order extrapolation to all the variables at the boundary, i.e. a simple copy of the MHD variables across the

Table 1. Initial conditions for the Riemann problems used to test the code.

Test type	State	ρ	p	v^x	v^y	v^z	B^x	B^y	B^z
Balsara Test 1									
$(\Gamma = 2)$	<i>left</i>	1.000	1.0	0.0	0.0	0.0	0.5	1.0	0.0
	<i>right</i>	0.125	0.1	0.0	0.0	0.0	0.5	-1.0	0.0
Balsara Test 2									
$(\Gamma = 5/3)$	<i>left</i>	1.0	30.0	0.0	0.0	0.0	5.0	6.0	6.0
	<i>right</i>	1.0	1.0	0.0	0.0	0.0	5.0	0.7	0.7
Balsara Test 3									
$(\Gamma = 5/3)$	<i>left</i>	1.0	1000.0	0.0	0.0	0.0	10.0	7.0	7.0
	<i>right</i>	1.0	0.1	0.0	0.0	0.0	10.0	0.7	0.7
Balsara Test 4									
$(\Gamma = 5/3)$	<i>left</i>	1.0	0.1	0.999	0.0	0.0	10.0	7.0	7.0
	<i>right</i>	1.0	0.1	-0.999	0.0	0.0	10.0	-7.0	-7.0
Balsara Test 5									
$(\Gamma = 5/3)$	<i>left</i>	1.08	0.95	0.40	0.3	0.2	2.0	0.3	0.3
	<i>right</i>	1.00	1.0	-0.45	-0.2	0.2	2.0	-0.7	0.5

excision boundary. A different method, based on the use of a linear extrapolation, has instead been implemented in [25] and although it can lead to improved accuracy for smooth flows (and especially when the MHD variables change rapidly near the excision boundary), it also leads to significantly incorrect results when shocks are present (see Sect. 4.2). Great care must therefore be paid at the properties of the flow near the excision boundary and the code presently includes both algorithms.

It is important also to note that other methods, not based on excision techniques, are being developed to improve the stability of numerical codes when black hole are present in the domain. One of these approaches is based on the use of a Kreiss-Oliger dissipation for the field variables inside the apparent horizon [60] and it can be straightforwardly extended also to the MHD case.

3.6. Mesh Refinement

The developments made in *Whisky* for handling non-uniform grids have been extended also to *WhiskyMHD* which can therefore use a “box-in-box” mesh refinement strategy [61]. This allows to reduce the influence of inaccurate boundary conditions at the outer boundaries and for the wave-zone to be included in the computational domain. In practice, we have adopted a Berger-Oliger prescription for the refinement of meshes on different levels [62] and used the numerical infrastructure described in [61], *i.e.*, the *Carpet* mesh refinement driver for *Cactus* (see [63] for details). In addition to this, a simplified form of adaptivity is also in place and which allows for new refined levels to be added at predefined positions during the evolution or for refinement boxes to be moved across the domain to follow, for instance,

regions where higher resolution is needed.

4. Tests

Code-testing represents an important aspect of the development of any newly developed and multidimensional code because it validates that all of the algorithms are implemented correctly and represent a faithful and discretized representation of the continuum equations they are solving. In what follows we report the results for a series of testbeds ranging from the solution of relativistic Riemann problems in flat spacetime, over to the stationary accretion onto a Schwarzschild black hole and up to the evolution of isolated and oscillating magnetized stars. We note that in the tests involving a polytropic EOS the recovery of the primitive variables has been made using the “2D-method” (see Sect. 3.3.1), while a “1D-method” has been used when adopting an ideal-fluid EOS (see Sect. 3.3.2).

4.1. Riemann problems

As customary in the testing of hydrodynamics and magnetohydrodynamics codes we have first validated WhiskyMHD against a set of Riemann problems in a Minkowski spacetime following the series of initial conditions proposed by Balsara [64]. All these tests were run on a grid of unit length with 1600 grid points with the initial discontinuity located at the center of the grid. An ideal equation of state with $\Gamma = 5/3$ was used with the exception of the first test with $\Gamma = 2$ and the initial conditions for all the tests are reported in Table 1.

In all of the tests presented here the numerical solution for the different MHD variables has been compared with the exact one computed with the exact Riemann solver discussed in [49]. This represents an important difference with what done in the past by similar codes as it allows, for the first time, for a quantitative assessment of the code’s ability to evolve correctly all the different waves that can form in relativistic MHD. In figures 2 and 3 the exact solution is represented with a solid line, while the numerical one with different symbols.

In figure 2, in particular, we show the comparison between the numerical and the exact solution at $t = 0.4$ for several MHD variables as computed for the relativistic analogue of the classical Brio-Wu shock tube problem [65, 66]. The initial discontinuity develops a left-going fast rarefaction, a left-going slow compound wave, a contact discontinuity, a right-going slow shock and fast rarefaction. Note that besides presenting the solution with $\alpha = 1$ and $\beta^x = 0$, we have exploited the freedom in choosing these gauges and validated the code also for less trivial values of the lapse and shift [18]. More specifically, shown with different symbols in figure 2 are the numerical solutions with $\alpha = 2$ at time $t = 0.2$ and with $\beta^x = 0.4$ after the latter has been shifted in space by $\beta^x t$. Clearly, all the symbols overlap extremely well, coinciding with the exact solution also in the presence of strong discontinuities. Note that only 160 of the 1600 data points used in the simulation are shown and that the difference between the numerical solution and the exact one at the compound wave is due to the fact that, by construction, our exact solver assumes compound waves never form. We have indeed adopted the same standpoint of Ryu and Jones [67] in the development of their exact Riemann solver in nonrelativistic magnetohydrodynamics. We also remark that it is not yet clear whether compound waves have to be considered acceptable physical solutions of the ideal MHD equations and a debate on this is still ongoing (see, for instance, [68, 69, 70, 71]).

Similarly, shown in figure 3 are the comparisons between the numerical solution for the rest-mass density and the y -component of the magnetic field for the tests number 2 (first row), 3 (second row), 4 (third row) and 5 (fourth row) of Balsara. The first 3 are computed at a time $t = 0.4$ while the test number 5 is computed at $t = 0.55$. Clearly, our code is able

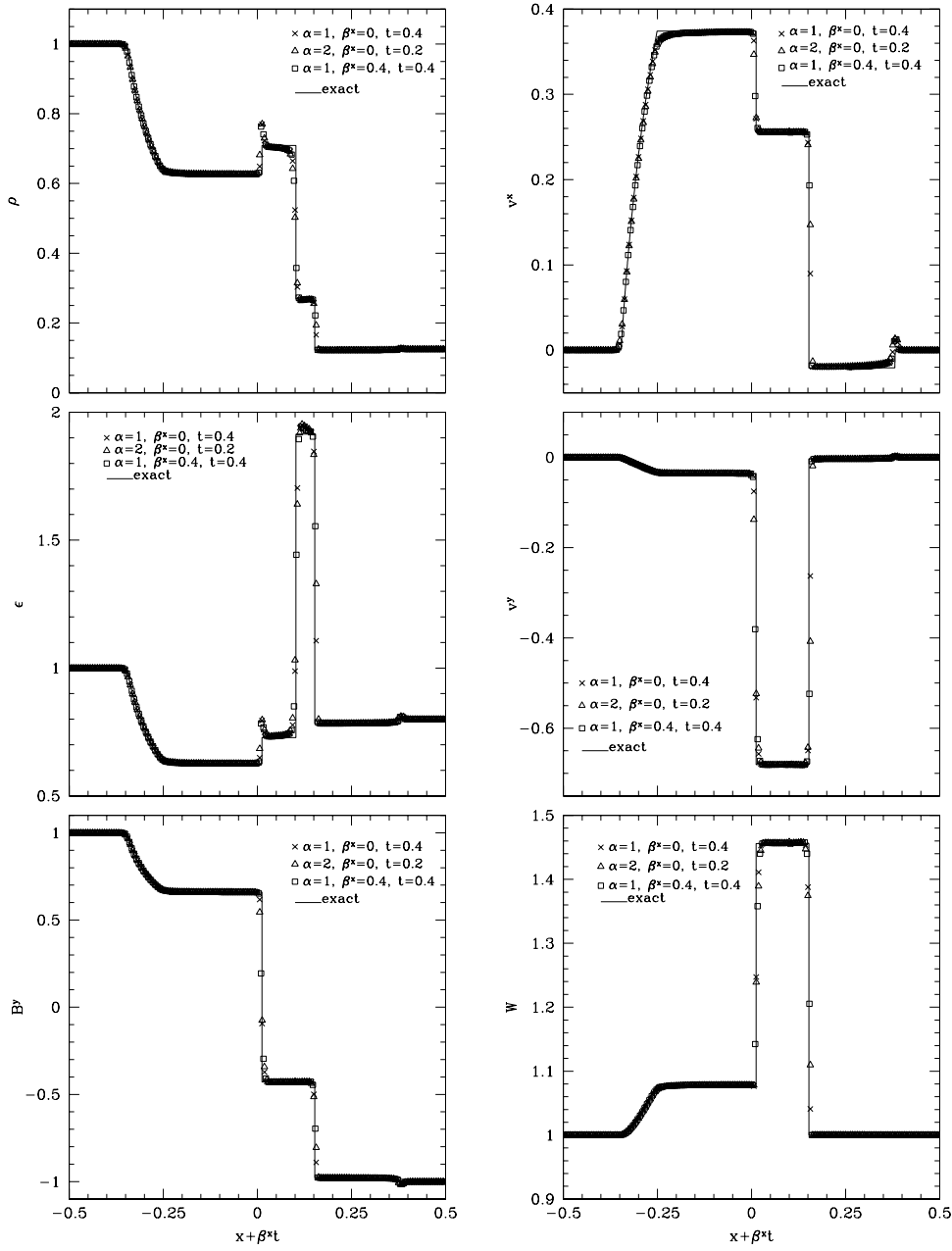


Figure 2. Numerical solution of the test number 1 of Balsara with different values for the lapse α and the shift β^x . The solid line represents the exact solution, the crosses the numerical one at time $t = 0.4$, the open triangle at time $t = 0.2$ but with $\alpha = 2$ and the open squares at $t = 0.4$ but with $\beta^x = 0.4$, in this last case the solution is shifted on the x -axis by the amount $\beta^x t$. Note that only 160 of the 1600 data points used in the numerical solution are shown.

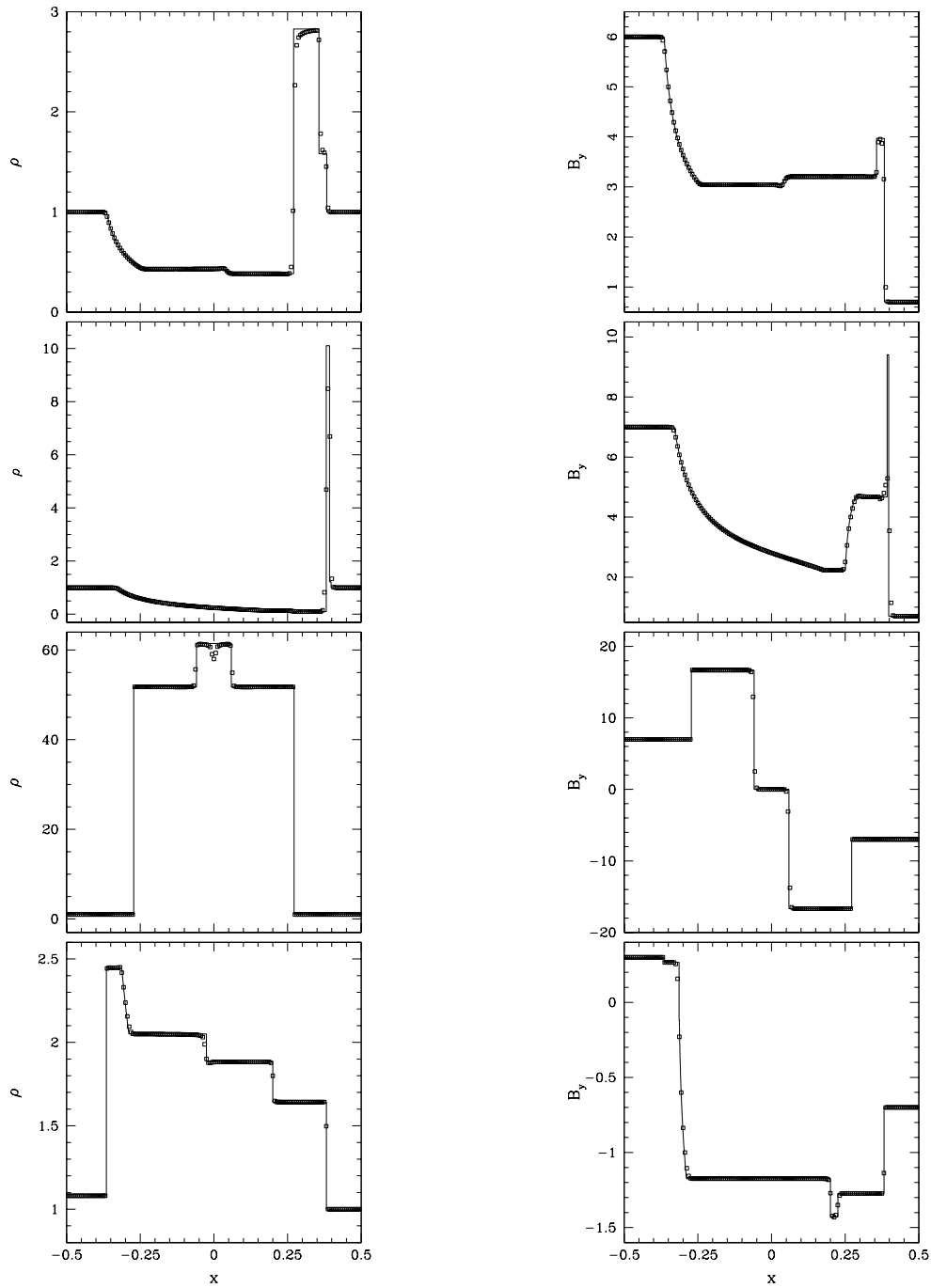


Figure 3. Numerical solution of the tests number 2 (first row), 3 (second row), 4 (third row) and 5 (fourth row) of Balsara. The first 3 are computed at a time $t = 0.4$ while the test number 5 is computed at $t = 0.55$. The solid line represent the exact solution while the open squares the numerical one. Note that only 160 of the 1600 data points used in the numerical solution are shown.

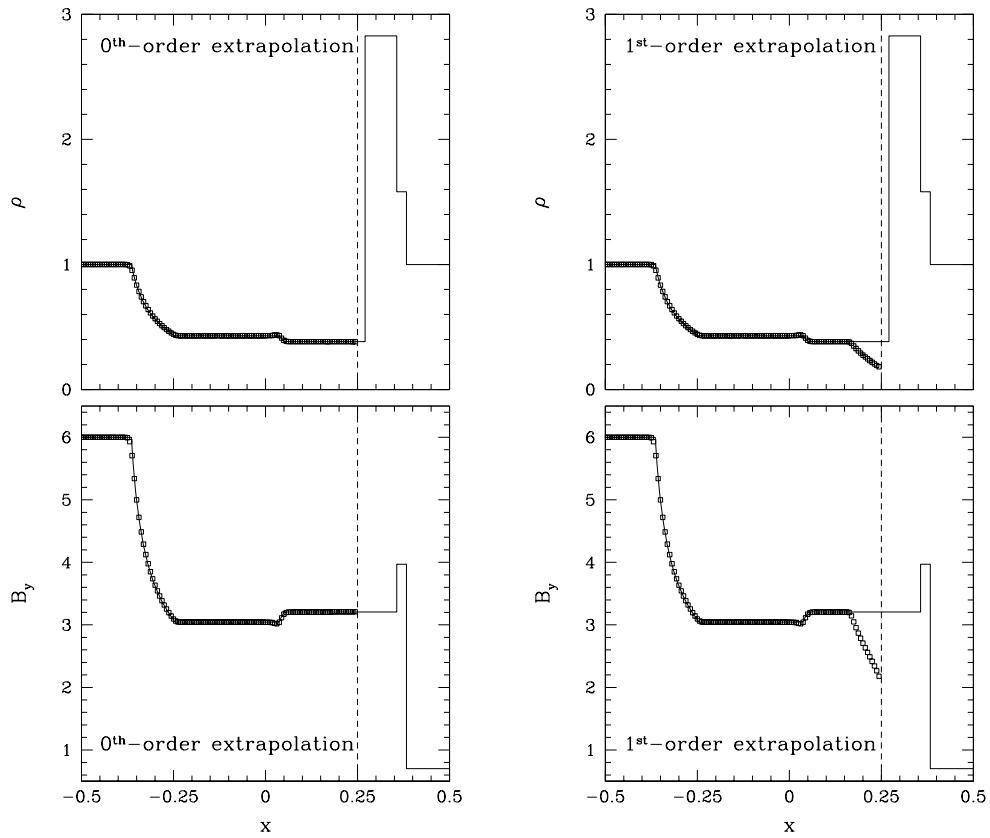


Figure 4. Numerical solution of the test number 2 of Balsara at a time $t = 0.4$ with an excision boundary (dashed vertical line) located at $x = 0.25$; the region at the right of this boundary is not evolved. In the two left panels a zeroth-order extrapolation, *i.e.* a simple copy, was used, while in the two right panels the values of the different variables at the excision boundary were obtained with a linear extrapolation. The solid line represents the exact solution while the open squares the numerical one. The solution is composed of two left-going fast and slow rarefactions, of a contact discontinuity and of two right-going fast and slow shocks. Only 160 of the 1600 data points used in the numerical solution are shown.

to resolve all the different waves present in MHD, showing a very good agreement with the exact solution. Other Riemann problems have been carried out in different directions (either along coordinate axes or along main diagonals) and they all provide the same level of accuracy discussed in figures 2 and 3.

4.2. Excision tests on a flat background

We next show the code's ability to accurately evolve shocks also when an excised region is present in the domain. To this scope, we have used the test number 2 of Balsara excising the region $x^i \in [0.25, 0.5]$ and using the zeroth-order extrapolation scheme. In this case the fast and slow shocks moving to the right go inside the excised region and the solution outside is not affected. This is shown in the left panels of figure 4 which report with small squares the numerical solution for ρ and B_y of the test number 2 of Balsara with an excision

boundary (dashed vertical line) located at $x = 0.25$. The data refers to time $t = 0.4$, when the right-going waves have already gone through the excision boundary as indicated by the exact solution (continuous line).

As a comparison, and to underline its incorrectness in the case of non-smooth flows, we have also considered boundary conditions involving a linear extrapolation of the MHD variables across of the excision boundary as suggested in [25]. This is illustrated in the right panels of figure 4 which are the same as the left ones but for the different boundary condition at the excision boundary. Clearly, in this case the solution outside the excision region is badly affected and a left-going wave is produced which rapidly spoils the solution. Because this happens only when the discontinuity crosses the excision boundary, it is clear that a linear extrapolation is not adequate in this case as it provides an incorrect information on the causal structure of the flow near the boundary. As we will discuss in the following Section, however, a linear extrapolation remains a good, and sometimes preferable, choice in the case of smooth flows.

4.3. Magnetized spherical accretion

This second test proves the ability of the code to evolve accurately stationary accretion solution in a curved but fixed spacetime. In particular, we consider the spherical accretion of a perfect fluid with a radial magnetic field onto a Schwarzschild black hole (this is sometimes referred to as a relativistic Bondi flow). The solution to this problem is already known for the unmagnetized case, but it is simple to show that its form is not affected if a radial magnetic field is added [10]. The initial setup for this test is the same used in [10, 16, 25, 18] and consists of a perfect fluid obeying a polytropic EOS with $\Gamma = 4/3$. The critical radius of the solution is located at $r_c = 8M$ and the rest-mass density at r_c is $\rho_c = 6.25 \times 10^{-2}$. These parameters are sufficient to provide the full description of the accretion onto a solar mass Schwarzschild black hole as described in [72]. We solve the problem on a Cartesian grid going from $x^i = 0$ to $x^i = 11M$.

To avoid problems at the horizon, located at $r = 2M$, the metric is written in terms of ingoing Eddington-Finkelstein coordinates. The excision boundary has the shape of a cubical box of length M so that the domain $[0, M] \times [0, M] \times [0, M]$ is excluded from the evolution. Furthermore, as a boundary condition across the excised cube we have considered both a zero-th and a first-order extrapolation, finding the latter to yield slightly more accurate results (e.g., the overall error is smaller of $\approx 3\%$ for a test case with $b^2/\rho = 25$, where b^2/ρ is the dimensionless magnetic field strength as measured at $r = 2M$). As discussed earlier, this is indeed to be expected for smooth flows as the ones considered here for the relativistic Bondi flow.

We measure the order of accuracy of the code by using the L_1 -norm of the relative error on the rest-mass density

$$\frac{\|\delta\rho\|_1}{\|\rho\|_1} \equiv \frac{\sum_{i,j,k} |\rho_{i,j,k} - \rho_{\text{exact}}(x_i, y_j, z_k)|}{\sum_{i,j,k} \rho_{\text{exact}}(x_i, y_j, z_k)}. \quad (49)$$

We plot this quantity in the left panel of figure 5 as a function of the magnetic-field strength and as computed at time $t = 100M$ for two different resolutions of 100^3 and 150^3 gridpoints, respectively. In addition, the error from the high-resolution simulation is multiplied by 1.5^2 so that the two curves should overlap if the code were second-order convergent. Clearly, the code does not show the expected convergence rate but for relatively weak magnetizations, i.e. $b^2/\rho \lesssim 4$ (we recall that these corresponds nevertheless to rather large magnetic fields of $\approx 10^{19}$ G). This behaviour is indeed similar to what found by Duez *et al.* [25] and has a rather

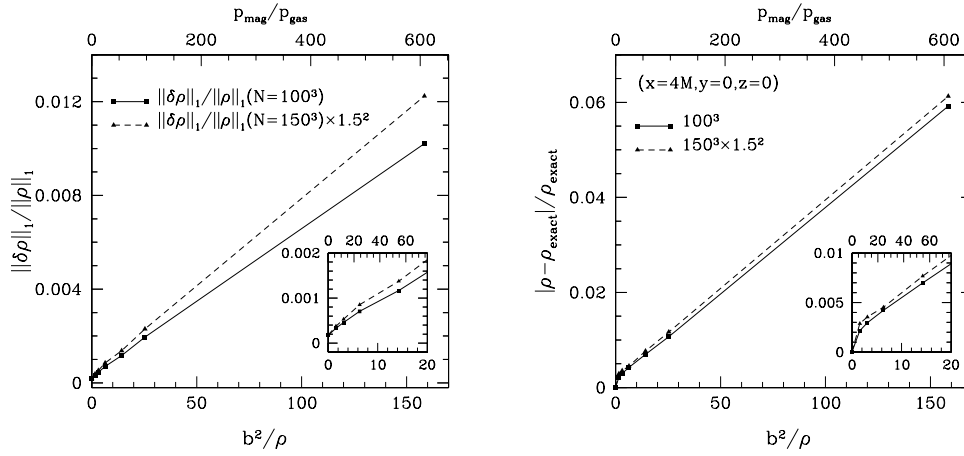


Figure 5. *Left panel:* L_1 norm of the relative error in the rest-mass density for the magnetized spherical accretion test, shown for different values of the magnetic field. Results from 100^3 and 150^3 runs are compared at time $t = 100M$, with the high-resolution curve being multiplied by 1.5^2 so that with a second-order convergence the two lines would overlap. *Right panel:* Relative error computed at point $x = 4M, y = 0, z = 0$; also in this case the high-resolution line is scaled to produce an overlap in the case of second-order convergence. In both cases the insets offer a magnification for small values of the magnetic field.

simple explanation. It is due to the rather large error in gridpoints near the excision boundary, *i.e.*, for $x^i \in [M, 2M]$, which spoil the overall behaviour of the L_1 norm (admittedly not a good measure of the convergence for a solution which is so rapidly varying near the excision boundary). To clarify this, we show in the right panel of figure 5 the same as in the left panel but for the relative error computed at a single gridpoint, *i.e.*, at $x = 4M, y = 0, z = 0$. Clearly the convergence is much closer to second-order in this case (the precise order being ≈ 1.8) and for much larger range in magnetizations.

A closer look at the behaviour of the relative error is offered in figure 6, where it is shown as measured along the x -direction for a magnetization of $b^2/\rho = 25$ (*i.e.*, with $p_{\text{mag}}/p_{\text{gas}} = 97$) and at time $t = 100M$. Also in this case, the high-resolution relative error is multiplied by 1.5^2 so that the two lines overlap if second-order convergent. Clearly this does not happen but also only for a small number of gridpoints near the excision boundary located at $x = M$.

As a final remark, we point out that the simulations of spherical-accretion flows performed here span a range of magnetizations well beyond what considered in the past with codes using Cartesian coordinates (the results reported in [25], for instance, were limited to $b^2/\rho \lesssim 30$). Indeed, no sign of instability has been found and only a moderate loss of accuracy has been measured for magnetic fields as large as $b^2/\rho \lesssim 160$.

4.4. Evolution of a stable magnetized neutron star

As a final test validating the code in a fully dynamical evolution of both the MHD variables and of the spacetime, we now consider the evolution of a stable magnetized neutron star. Although this initial data represents a stationary solution, small oscillations can be triggered by the small but nonzero truncation error. Such oscillations are sometimes considered a nuisance and even suppressed through the introduction of artificial-viscosity terms. On the

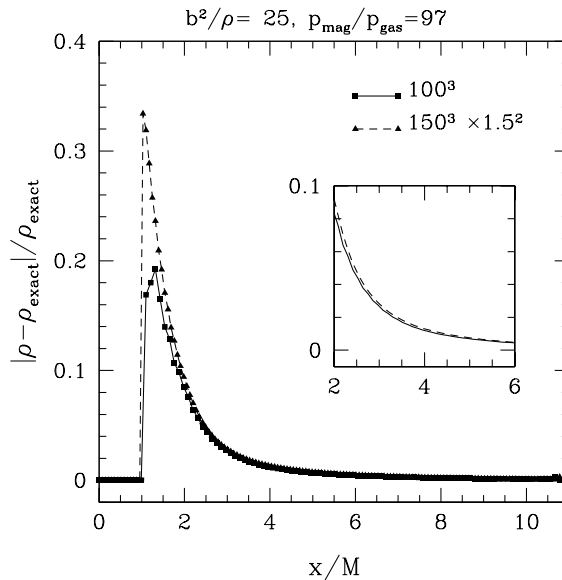


Figure 6. Relative error of the rest-mass density for the magnetized spherical-accretion flow with $b^2/\rho = 25$ (i.e., with $p_{\text{mag}}/p_{\text{gas}} = 97$) along the x -axis at $t = 100M$. The high-resolution error is multiplied by 1.5^2 so that the two lines would overlap with a second-order convergence. Clearly, this does not happen near the excision boundary located at $x = M$.

contrary, since they represent the consistent response of the star to small perturbations, they should be considered as extremely useful. The eigenfunctions and eigenfrequencies of these oscillations, in fact, can serve both as a test of the code, when compared with the expectations coming from perturbation theory (see Appendix B of [37] for a representative example), or to extract information on the properties of the star, when considering regimes which are not yet accessible to perturbative studies (e.g., in the case of nonlinear oscillations or very rapidly rotating stars).

Two options are possible for the construction of the initial data. A first and simpler one was employed extensively in [27, 28, 29] and consists of considering a background purely-hydrodynamical solution in stable dynamical equilibrium and of “adding” a poloidal magnetic field in terms of a purely toroidal vector potential. Besides being essentially arbitrary, the vector potential is chosen to be proportional to the pressure so as to lead to a magnetic field entirely confined within the star. While straightforward, this approach does not construct a magnetized stellar model which is a consistent solution of the Einstein equations and thus inevitably introduces violations of the Hamiltonian and momentum constraints. Such violations, however, are in general negligible for reasonably small magnetic fields.

A second option, and the one employed here, consists of computing the initial conditions as a consistent and accurate solution of the Einstein equations for a stationary, axisymmetric and magnetized star. We have done this by using the spectral code developed by Bocquet *et al.* [73], which solves the full set of Einstein and Maxwell equations to high precision. Assuming an axisymmetric model with a poloidal magnetic field having the dipole moment aligned along the rotation axis, the code is used to build initial configurations of uniformly rotating magnetized neutron stars with different angular velocities and magnetic field strengths.

Magnetic field

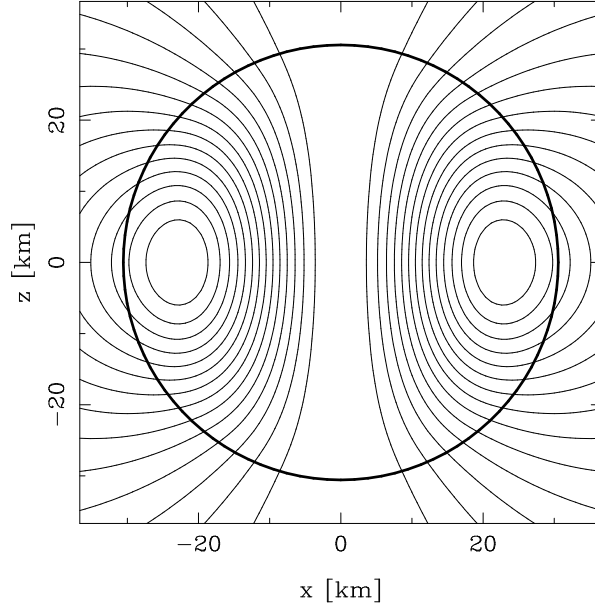


Figure 7. Magnetic field lines of the oscillating magnetized and nonrotating neutron star considered in this paper. The solid thick line represents the star surface.

For simplicity we here consider a nonrotating magnetized neutron star with mass $M = 1.3M_{\odot}$ endowed with a poloidal magnetic field with magnetic dipole along the z -axis and a central magnetic field $B_c = 2.4 \times 10^{14}$ G, corresponding to $\beta = p_{\text{mag}}/p_{\text{gas}} = 10^{-6}$ (this β , which should not be confused with the shift vector β^i , is always monotonically decreasing inside the initial equilibrium model, and is much larger in the atmosphere, where it reaches values $\sim 10^6$). A polytropic equation of state with $\Gamma = 2$ and $K = 372$ was used both for the calculation of the initial model and during the evolution. A representative image of the initial model is presented in figure 7, which shows the magnetic field lines together with the stellar surface (thick solid line). Note that although the star is nonrotating, the presence of a magnetic field replaces the spherical symmetry for an axisymmetrical one.

As a first test, we consider the evolution of the star within the so-called ‘‘Cowling approximation’’, *i.e.* by holding the metric fixed to its initial value and by evolving the MHD variables onto this background spacetime (the evolution is not made only at the outer boundaries, where we use Dirichlet-type boundary conditions). The results of these evolutions are summarized in figure 8, with the left panel showing the evolution of the maximum of the rest-mass density ρ when normalized to its initial value. The three different lines (dotted, dashed and continuous) refer to the three resolutions used of $N = 60^3$, 90^3 , and 120^3 , respectively. The coordinate time on the horizontal axis is expressed in terms of the characteristic ‘‘dynamical timescale’’ $\tau \equiv \sqrt{R^3/M}$, where R is the coordinate radius of the star.

In analogy with what observed in the purely hydrodynamical case [37], the magnetized star is set into oscillation by the small truncation error introduced by the mapping onto a Cartesian coordinate system of the stationary solution found in spherical polar

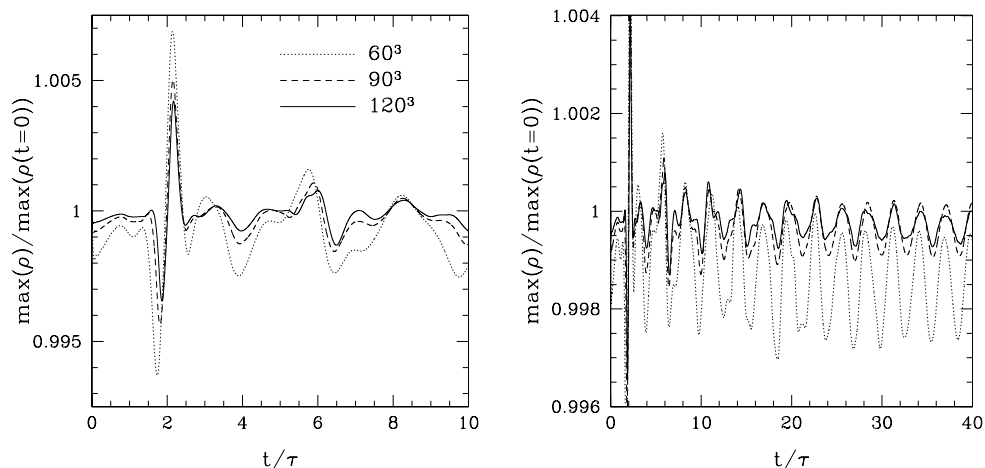


Figure 8. *Left panel:* Maximum of the rest-mass density ρ normalized to its initial value and expressed in terms of the dynamical timescale $\tau \equiv \sqrt{R^3/M}$. The magnetized star is evolved within the Cowling approximation, with different lines referring to different resolutions: $N = 60^3$ (dotted), $N = 90^3$ (dashed line) and $N = 120^3$ (solid line). *Right panel:* The same as in the left one but for a longer timescale.

coordinates [73] (no perturbation coming from the outer boundaries was seen to influence the dynamics of the oscillations). Because of its stochastic nature, the initial perturbation triggers a variety of modes, most of which however decay rather rapidly leaving, after $t \simeq 25\tau$, an essentially harmonic oscillation in the fundamental mode only. This is shown in the right panel of figure 8 which shows the evolution over a longer timescale. Furthermore, in the linear regime considered here, the amplitude of the oscillations is proportional to the magnitude of the truncation error and one expects the former to decrease as the resolution is increased. This is clearly the case for the oscillations shown in figure 8, and the very good overlap among the different timeseries is an indication that the oscillations indeed correspond to eigenmodes and that the code is tracking them correctly at these resolutions.

Next, we consider the evolution of the same initial model discussed above but including also the solution of the Einstein equations so as to make the system fully dynamical (Dirichlet-type boundary conditions are used at the outer boundaries for the MHD variables and radiative ones for the fields). Also in this case, oscillations are triggered by the truncation error, with an amplitude that converges to zero with the increase of the resolution and with an harmonic content that becomes more evident after the initial transient (also in this case, no perturbation coming from the outer boundaries was seen to influence the dynamics of the oscillations). In addition, and in analogy with what observed in the purely hydrodynamical case [74, 75, 37], the oscillations are accompanied by a secular growth which also converges away at the correct rate with increasing grid resolution and that does not influence the long-term evolutions. This is shown in figure 9 which reports the same quantities as in figure 8 but for a fully dynamical evolution. Note also that the secular evolution of the central rest-mass density varies according to the EOS adopted: when using the ideal-fluid EOS, in fact, the secular drift of the central rest-mass density is towards lower densities. However, if the adiabatic condition is enforced, the opposite is true and central rest-mass density evolves towards larger values. Both the evolutions in the Cowling approximation and in dynamical spacetimes, have not shown signs

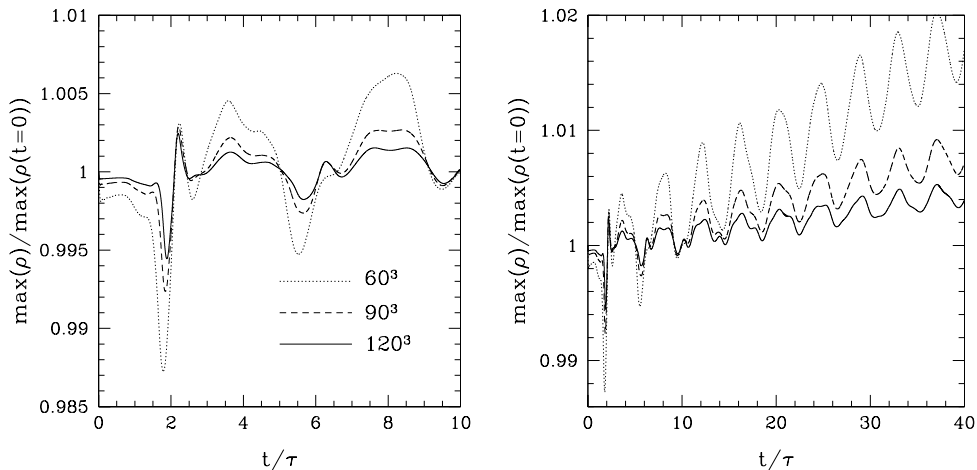


Figure 9. *Left panel:* Maximum of the rest-mass density ρ normalized to its initial value and expressed in terms of the dynamical timescale $\tau \equiv \sqrt{R^3/M}$. The magnetized star is evolved together with the spacetime, with different lines referring to different resolutions: $N = 60^3$ (dotted), $N = 90^3$ (dashed line) and $N = 120^3$ (solid line). *Right panel:* The same as in the left one but for a longer timescale.

of instability at all resolutions considered and up to several tens of dynamical timescales.

As a final remark we underline that the convergence rate is not exactly second-order but slightly smaller, (i.e., 1.7-1.8), because the reconstruction schemes are only first-order accurate at local extrema (i.e. at the centre and at the surface of the star) thus increasing the overall truncation error. Similar estimates were obtained also in the purely hydrodynamical case [37].

5. Conclusions

We have presented a new three-dimensional numerical code in Cartesian coordinates developed to solve the full set of GRMHD equation on a dynamical background. The code is based on high-resolution shock-capturing techniques as implemented on domains with adaptive mesh refinements. This code represents the extension to MHD of the approach already implemented with success in the general-relativistic hydrodynamics code `Whisky` [30].

The code has been validated through an extensive series of testbeds both in special and in general relativity scenarios. In particular, we have first considered a set of Riemann problems in a Minkowski spacetime following a variety of initial conditions. In all of the tests presented, the numerical solution has been compared with the exact one [49], providing, for the first time, a quantitative assessment of the code's ability to evolve correctly and accurately all of the different waves that can form in relativistic MHD. Furthermore, as a demonstration of the proper handling of continuous and discontinuous flows in the presence of an excision region, we have extended the Riemann-problems tests across an excised boundary. In doing so we have revealed the importance of correct boundary conditions and pointed out that those recently proposed in [25] can lead to incorrect solutions for non-smooth flows.

Next, to investigate magnetized fluids in a curved but fixed spacetime, we have

considered the spherical accretion of a perfect fluid with a radial magnetic field onto a Schwarzschild black hole (relativistic Bondi flow). Also in this case, the code has been shown to accurately reproduce the stationary solution and to be convergent at the correct rate for small and large magnetizations. For very large magnetizations, however, the very rapidly varying behaviour of the MHD variables near the excision region prevents from a correct convergence near the horizon, although the code remains second-convergent away from the horizon and is convergent overall. Also for these extremely-high values of the magnetic field, the code has shown to be robust and accurate at regimes where other codes were reported to fail [25].

Finally, we have considered the evolution of magnetized neutron stars in equilibrium and constructed as a consistent solutions of the coupled Einstein-Maxwell equations. Such initial models represent an important difference from those considered by other authors, which were not consistent solution of the Einstein equations initially and whose magnetic field is totally confined within the star [25]. In analogy with what observed in the purely hydrodynamical case [37], these magnetized stars are set into oscillation by the small truncation error. These pulsations, which have been studied both in fixed (Cowling approximation) and in dynamical spacetimes, have been shown to have the correct behaviour under changes of spatial resolution and to correspond to the eigenmodes of relativistic and magnetized stars. Both evolutions in the Cowling approximation and in dynamical spacetimes have not shown signs of instability at all resolutions considered and up to several tens of dynamical timescales.

A number of projects are expected to be carried out with the new code. Firstly, we plan to extend the study on the oscillations of rotating and nonrotating neutron stars with a detailed analysis of the effect of magnetic fields on the frequency of oscillations. It is important to note that only recently some results were obtained in perturbation theory and within the Cowling approximation [76]. Our code will be a complementary tool to the perturbative approaches, using the latter as testbeds and carrying them beyond the regimes of slow rotation and weak magnetizations. Such a study, and the comparison with the frequencies observed in objects such as the soft gamma-repeaters, will provide useful information on the mass and magnetic-field strength of magnetars.

Secondly, we plan to use WhiskyMHD to study the collapse of uniformly and differentially rotating magnetized neutron stars with the aim of extending further the work done in [37, 77, 60] and to highlight the role that this process may have in the phenomenology of short γ -ray bursts. We are especially interested in the calculation of the gravitational-wave signal emitted by these sources and on the influence that magnetic fields may have on it.

Finally, and in view of the total generality with which it has been developed, the code will be used to study the dynamics of binary systems of neutron stars and mixed binaries, with the aim of extending the work carried out in [78] and of considering in a fully general-relativistic context the Newtonian results obtained in [24].

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