

**Tail-induced spin-orbit effect in the gravitational radiation of compact binaries**Luc Blanchet,<sup>1</sup> Alessandra Buonanno,<sup>2</sup> and Guillaume Faye<sup>1</sup><sup>1</sup>*Institut d'Astrophysique de Paris, UMR 7095 CNRS Université Pierre et Marie Curie, 98<sup>bis</sup> boulevard Arago, 75014 Paris, France*<sup>2</sup>*Maryland Center for Fundamental Physics and Joint Space-Science Center, Department of Physics, University of Maryland, College Park, Maryland 20742, USA*

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Gravitational waves contain tail effects which are due to the backscattering of linear waves in the curved space-time geometry around the source. In this paper we improve the knowledge and accuracy of the two-body inspiraling post-Newtonian (PN) dynamics and gravitational-wave signal by computing the spin-orbit terms induced by tail effects. Notably, we derive those terms at 3PN order in the gravitational-wave energy flux, and 2.5PN and 3PN orders in the wave polarizations. This is then used to derive the spin-orbit tail effects in the phasing through 3PN order. Our results can be employed to carry out more accurate comparisons with numerical-relativity simulations and to improve the accuracy of analytical templates aimed at describing the whole process of inspiral, merger, and ringdown.

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**I. INTRODUCTION****A. Motivation**

During the last ten years a network of ground-based laser-interferometer gravitational-wave detectors has been built and has taken data at design sensitivity. It is a worldwide network composed of the Laser Interferometer Gravitational Wave Observatory (LIGO), Virgo, GEO-600, and TAMA, and it has operated in the frequency range 10–10<sup>3</sup> Hz. Coalescing binary systems composed of black holes and/or neutron stars are among the most promising sources for those detectors. By 2016 the gravitational-wave detectors will be upgraded to a sensitivity such that event rates for coalescing binary systems will increase by approximately a factor 1000, thus making likely the first detection of gravitational waves from those systems. In the future, space-based detectors like LISA should detect supermassive black-hole binary systems in the low frequency band 10<sup>-4</sup>–10<sup>-2</sup> Hz.

The search for gravitational waves from coalescing binary systems and the extraction of source parameters are based on the matched-filtering technique, which requires a rather accurate knowledge of the waveform of the incoming signal. In particular, the detection and subsequent data analysis are made by using a bank of templates modeling the gravitational wave emitted by the source. The need for a faithful template bank has driven the development of accurate templates over the last 30 years.

The post-Newtonian (PN) expansion is the most powerful approximation scheme in analytical relativity capable of describing the two-body dynamics and gravitational-wave emission of inspiraling compact binary systems [1]. The PN expansion is an expansion in the ratio of the characteristic orbital velocity of the binary  $v$  to the speed of light  $c$ . However, as the black holes approach each other toward merger, we expect the PN expansion to lose accuracy because the velocity of the holes approaches the speed

of light. At that point, numerical relativity [2–4] plays a crucial role, providing us with the dynamics and gravitational-wave emission of the last cycles of inspiral, followed by the merger and ringdown phases. Furthermore, by properly combining PN predictions and numerical relativity results, it is possible to describe analytically and/or numerically with high accuracy the full gravitational-wave signal [5–8].

Black holes in binary systems can carry spins, and when spins are not aligned with the orbital angular momentum, spins induce precession of the orbital plane (see e.g. Ref. [9]). This adds substantial complexity to the gravitational waveforms, making it indispensable to include spin effects in templates used for the search. Moreover, as found a long time ago [10–18], gravitational waves contain tail effects which are due to the backscattering of linear waves in the curved space-time geometry around the source (and primarily generated by its mass). This causes the gravitational-wave signal to depend on the entire history of the binary system.

In this paper we improve the knowledge and accuracy of the two-body inspiraling dynamics and gravitational-wave signal by computing the spin-orbit (SO) terms induced by tail effects. This is the continuation of our previous work on spins [19,20], where we obtained the next-to-leading 2.5PN SO contributions in the equations of motion and gravitational-wave energy flux. Here, we derive those SO terms at 3PN order in the gravitational-wave energy flux, where they are entirely due to tails. Furthermore, we obtain the SO terms at 2.5PN and 3PN orders in the wave polarizations that are specifically due to tails, leaving aside other SO terms at these orders that come from instantaneous (nontail) linear contributions and which will not be computed here.

We obtain the energy flux in two independent ways, first directly using the radiative multipole moments, and second by differentiating and squaring the gravitational-wave

polarizations. To compute the SO-tail effects in the wave polarizations we solve the two-body dynamics, taking into account spin precession. Assuming quasicircular adiabatic inspiral, we also compute the 3PN SO terms induced by tails in the gravitational-wave phasing. These results can be used to improve the accuracy of inspiraling templates, to carry out a comparison with numerical-relativity predictions, and to improve the accuracy of effective-one-body and phenomenological templates [5–8].

As an important check of our results we obtain the 3PN SO-tail terms in the energy flux in the test-particle limit, and find perfect agreement with earlier PN computations based on black-hole perturbation induced by the motion of a test particle around a massive black hole [21].

This paper is organized as follows. In Sec. II we review the post-Newtonian multipole moment formalism and discuss relevant properties of tails. In Sec. III we describe how spin effects are included in the PN formalism and derive the binary’s evolution equations when black holes carry spins. In Sec. IV we obtain the time evolution of the moving triad and solve the precessing dynamics at the relevant PN order. In Sec. V we compute the 2.5PN and 3PN SO-tail effects in the gravitational waveform and polarizations. Restricting ourselves to quasicircular adiabatic inspiral, we derive in Sec. VI the 3PN SO-tail effects in the energy flux and in the gravitational phasing. Section VII contains our main conclusions. In Appendix A we check that the equations of motion do not contain physical SO effects at 3PN order. In Appendix B we investigate the secular evolution of the spins, and finally in Appendix C we give the explicit gravitational-wave polarizations.

## B. Notation

In this paper we adopt the following notations.  $G$  is the Newton constant and  $c$  is the speed of light. As usual we refer to  $n$ PN as the post-Newtonian terms with formal order  $\mathcal{O}(c^{-2n})$  relative to the Newtonian acceleration in the equations of motion, or to the lowest-order quadrupole-moment formalism for the radiation field. Greek indices are space-time indices, and the Latin ones are space indices. The quantity  $L = i_1 \cdots i_\ell$  denotes a multi-index composed of  $\ell$  multipolar spatial indices  $i_1, \dots, i_\ell$  ranging from 1 to 3. Similarly  $L - 1 = i_1 \cdots i_{\ell-1}$  and  $kL - 2 = ki_1 \cdots i_{\ell-2}$ ;  $N_L = N_{i_1} \cdots N_{i_\ell}$  is the product of  $\ell$  spatial vectors  $N_i$  (similarly for  $x_L = x_{i_1} \cdots x_{i_\ell}$ ). We indicate with  $\partial_L = \partial_{i_1} \cdots \partial_{i_\ell}$  and  $\partial_{kL-2} = \partial_k \partial_{i_1} \cdots \partial_{i_{\ell-2}}$  the product of partial derivatives  $\partial_i = \partial/\partial x^i$ . In the case of summed-up (dummy) multi-indices  $L$ , we do not write the  $\ell$  summations from 1 to 3 over their indices. Furthermore, the symmetric-trace-free (STF) projection is indicated using brackets,  $T_{\langle L \rangle} = \text{STF}[T_L]$ ; thus  $U_L = U_{\langle L \rangle}$  and  $V_L = V_{\langle L \rangle}$  for STF moments. For instance, we write  $x_{\langle i} v_{j \rangle} = \frac{1}{2}(x_i v_j + x_j v_i) - \frac{1}{3} \delta_{ij} \mathbf{x} \cdot \mathbf{v}$ . The transverse-traceless (TT) projection operator is denoted

$\mathcal{P}_{ijkl}^{\text{TT}} = \mathcal{P}_{ik} \mathcal{P}_{jl} - \frac{1}{2} \mathcal{P}_{ij} \mathcal{P}_{kl}$ , where  $\mathcal{P}_{ij} = \delta_{ij} - N_i N_j$  is the projector orthogonal to the unit direction  $\mathbf{N} = \mathbf{X}/R$  of a radiative coordinate system  $X^\mu = (cT, \mathbf{X})$ . The quantity  $\varepsilon_{ijk}$  is the Levi-Civita antisymmetric symbol such that  $\varepsilon_{123} = 1$ . Finally, we denote time derivatives with a superscript  $(n)$ , and we indicate the symmetrization operation with round parentheses.

## II. WAVE GENERATION FORMALISM

The gravitational waveform  $h_{ij}^{\text{TT}}$ , generated by an isolated source described by a stress-energy tensor  $T^{\mu\nu}$  with compact support, and propagating in the asymptotic regions of the source, is the TT projection of the metric deviation at the leading-order  $1/R$  in the distance to the source. It is parametrized by STF mass-type moments  $U_L$  and current-type ones  $V_L$ , which constitute the observables of the waveform at infinity from the source and are called the radiative moments [22]. The general expression of the TT waveform, in a suitable radiative coordinate system  $X^\mu = (cT, \mathbf{X})$ , reads, when neglecting terms of the order of  $1/R^2$  or higher,

$$h_{ij}^{\text{TT}} = \frac{4G}{c^2 R} \mathcal{P}_{ijkl}^{\text{TT}} \sum_{\ell=2}^{+\infty} \frac{N_{L-2}}{c^\ell \ell!} \times \left[ U_{kL-2} - \frac{2\ell}{c(\ell+1)} N_m \varepsilon_{mnk} V_{l n L-2} \right]. \quad (2.1)$$

Here the radiative moments  $U_L$  and  $V_L$  are functions of the retarded time  $T_R \equiv T - R/c$  in the radiative coordinate system (we denote  $R = |\mathbf{X}|$ ). The integer  $\ell$  refers to the multipolar order, and  $\mathbf{N} = \mathbf{X}/R = (N_i)$  is the unit vector pointing from the source to the far away detector. The TT projection operator  $\mathcal{P}_{ijkl}^{\text{TT}}$  and other notations are defined in Sec. IB. With  $\mathbf{P} = (P_i)$  and  $\mathbf{Q} = (Q_i)$  denoting two unit polarization vectors, orthogonal and transverse to the direction of propagation  $\mathbf{N}$ , the two “plus” and “cross” polarization states of the waveform read as

$$h_+ = \frac{P_i P_j - Q_i Q_j}{2} h_{ij}^{\text{TT}}, \quad (2.2a)$$

$$h_\times = \frac{P_i Q_j + P_j Q_i}{2} h_{ij}^{\text{TT}}. \quad (2.2b)$$

Our convention for the choice of the polarization vectors  $\mathbf{P}$  and  $\mathbf{Q}$  in the case of binary systems will be specified in Fig. 2. Plugging Eq. (2.1) into the standard expression for the gravitational-wave energy flux, we get [22]

$$\mathcal{F} = \sum_{\ell=2}^{+\infty} \frac{G}{c^{2\ell+1}} \left[ \frac{(\ell+1)(\ell+2)}{(\ell-1)\ell!(2\ell+1)!!} U_L^{(1)} U_L^{(1)} + \frac{4\ell(\ell+2)}{c^2(\ell-1)(\ell+1)!(2\ell+1)!!} V_L^{(1)} V_L^{(1)} \right]. \quad (2.3)$$

### A. Expression of the radiative moments

In the multipolar-post-Minkowskian formalism [13–15], the radiative moments are expressed in terms of two other sets of moments, referred to as the “canonical” moments  $M_L$ ,  $S_L$ , and which are relevant to the description of the source’s near zone. The relation between the radiative moments  $U_L$ ,  $V_L$  and the canonical ones  $M_L$ ,  $S_L$  encodes all the nonlinearities in the wave propagation between the source and the detector [15]. Those relations may be reexpanded in a PN way and are then seen to contain, at the leading 1.5PN order, the contribution of the so-called gravitational-wave tails, due to backscattering of linear waves onto the space-time curvature associated with the total mass of the source itself. The explicit expressions at 1.5PN order are [15,23]

$$U_L(T_R) = M_L^{(\ell)} + \frac{2GM}{c^3} \int_{-\infty}^{T_R} dt M_L^{(\ell+2)}(t) \left[ \ln\left(\frac{T_R - t}{2\tau_0}\right) + \kappa_\ell \right] + \mathcal{O}\left(\frac{1}{c^3}\right)_{\text{nontail}}, \quad (2.4a)$$

$$V_L(T_R) = S_L^{(\ell)} + \frac{2GM}{c^3} \int_{-\infty}^{T_R} dt S_L^{(\ell+2)}(t) \left[ \ln\left(\frac{T_R - t}{2\tau_0}\right) + \pi_\ell \right] + \mathcal{O}\left(\frac{1}{c^3}\right)_{\text{nontail}}, \quad (2.4b)$$

where  $M$  is the Arnowitt-Deser-Misner mass associated with the source. It also coincides with the mass monopole moment. The remainders  $\mathcal{O}(c^{-3})_{\text{nontail}}$  in Eqs. (2.4) denote some correction terms which are at least of order 1.5PN and are instantaneous or contain the nonlinear memory effect which will not be considered in the present computation. Here  $\kappa_\ell$  and  $\pi_\ell$  denote some numerical constants given by [23]

$$\kappa_\ell = \frac{2\ell^2 + 5\ell + 4}{\ell(\ell + 1)(\ell + 2)} + \sum_{k=1}^{\ell-2} \frac{1}{k}, \quad (2.5a)$$

$$\pi_\ell = \frac{\ell - 1}{\ell(\ell + 1)} + \sum_{k=1}^{\ell-1} \frac{1}{k}. \quad (2.5b)$$

The constant  $\tau_0$  in Eqs. (2.4) is a freely specifiable time scale entering the relation between the radiative time  $T_R$  and the corresponding retarded time in harmonic coordinates.

The canonical moments  $M_L$  and  $S_L$  are themselves linked to six sets of multipole moments characterizing the source, collectively called the “source” moments and denoted  $I_L$ ,  $J_L$ ,  $W_L$ ,  $X_L$ ,  $Y_L$ ,  $Z_L$ . The point is that those source moments are known as explicit integrals extending over the pseudo-energy-momentum tensor of the matter fields and the gravitational field of the source [23,24]. In the following we shall essentially need  $I_L$  and  $J_L$  which represent the main mass and current moments of the source. The other moments  $W_L$ ,  $X_L$ ,  $Y_L$ , and  $Z_L$  play a little role because they are associated with a possible gauge transformation performed at linear order. It turns out that

the difference between the canonical moments  $M_L$ ,  $S_L$  and the source moments  $I_L$ ,  $J_L$  arises only at the 2.5PN order:

$$M_L = I_L + \mathcal{O}\left(\frac{1}{c^5}\right), \quad (2.6a)$$

$$S_L = J_L + \mathcal{O}\left(\frac{1}{c^5}\right). \quad (2.6b)$$

For the present application it will be sufficient to consider the source mass moments  $I_L$  at 1PN order and the current ones  $J_L$  at Newtonian order (see the discussion in Sec. III B). These are given by [25]

$$I_L = \int d^3x \left[ \hat{x}_L \sigma + \frac{1}{2c^2(2\ell + 3)} \hat{x}_L |\mathbf{x}|^2 \sigma^{(2)} - \frac{4(2\ell + 1)}{c^2(\ell + 1)(2\ell + 3)} \hat{x}_{iL} \sigma_i^{(1)} \right] + \mathcal{O}\left(\frac{1}{c^4}\right), \quad (2.7a)$$

$$J_L = \varepsilon_{ij(i\ell} \int d^3x \hat{x}_{L-1)i} \sigma_j + \mathcal{O}\left(\frac{1}{c^2}\right). \quad (2.7b)$$

The other moments we shall need are the mass monopole moment  $M$  and the monopole of the moment  $W_L$ , which are given by

$$M = \int d^3x \sigma + \mathcal{O}\left(\frac{1}{c^2}\right), \quad (2.8a)$$

$$W = \frac{1}{3} \int d^3x x^i \sigma_i + \mathcal{O}\left(\frac{1}{c^2}\right). \quad (2.8b)$$

The mass, current, and tensor densities  $\sigma$ ,  $\sigma_i$ ,  $\sigma_{ij}$  in Eqs. (2.7) and (2.8), are defined as (where  $T^{ii} \equiv \delta_{ij} T^{ij}$ )

$$\sigma = \frac{T^{00} + T^{ii}}{c^2}, \quad \sigma_i = \frac{T^{0i}}{c}, \quad \sigma_{ij} = T^{ij}. \quad (2.9)$$

We recall that, e.g.,  $\sigma^{(n)}$  in Eqs. (2.7) means taking  $n$ -time derivatives.

The spin parts of the source moments in Eqs. (2.7) will come from the model we adopt for the stress-energy tensor  $T^{\mu\nu}$  appropriate for spinning compact binaries (see details in Sec. III). Importantly, we notice that for the accuracy required by our calculation of the spin effects due to tails, all integrands in Eqs. (2.7) have compact support. This is in contrast with the spin effects at 2.5PN order which necessitate noncompact supported higher-order terms in the source moments [20]. (We also find that the second term in  $I_L$  can be ignored in the present application to spins.)

### B. Computing the tail integrals

The tail integrals in Eqs. (2.4) extend over the entire past of the evolving source and it is *a priori* a nontrivial task to compute them. Here we recall, based on Refs. [15,16], that the tails are actually very weakly sensitive (in a post-Newtonian sense) to the past history of the source, and can essentially be computed using the *current* dynamics, i.e. at current time  $T_R$ , of the source.

We have to compute, e.g., the integral appearing in the radiative mass multipole moment (2.4a), in which we can replace, following (2.6a), the canonical moment  $M_L$  by the source moment  $I_L$ . Thus,

$$\mathcal{U}_L(T_R) = \int_{-\infty}^{T_R} dt I_L^{(\ell+2)}(t) \ln\left(\frac{T_R - t}{2\tau'_0}\right), \quad (2.10)$$

where we pose  $\tau'_0 = \tau_0 e^{-\kappa\epsilon}$ .

Let us introduce a constant time interval  $\mathcal{T}$  to split the integral (2.10) into some contribution coming from the “recent past” and extending from the current time  $T_R$  to  $T_R - \mathcal{T}$ , and the remaining contribution called the “remote past,” from  $T_R - \mathcal{T}$  to  $-\infty$  in the past. The recent past can be thought of as corresponding to the most recent orbital period of a compact binary system, while the remote past will include the details (eventually unknown) of the formation and early past evolution of the compact binary. However, we shall prove that our result is independent of the chosen time scale  $\mathcal{T}$ .

To control the convergence of the tail integral (2.10) in the past we make a physical assumption regarding the behavior of the multipole moment  $I_L(t)$  when  $t \rightarrow -\infty$ . We assume that at very early times the source was formed from a bunch of freely falling particles initially moving on some hyperbolic-like orbits, and forming at a later time a gravitationally bound system by emission of gravitational radiation. The gravitational motion of initially free particles is given by  $x^i(t) = V^i t + W^i \ln(-t) + X^i + o(1)$ , where  $V^i$  and  $X^i$  denote constant vectors, and  $W^i = GmV^i/V^3$  (see Ref. [26] for a proof; to simplify we consider the relative motion of two particles with total mass  $m$ ). Here the Landau remainder  $o$  symbol satisfies  $\partial^n o(1)/\partial t^n = o(1/t^n)$ . From that physical assumption we find that the multipole moment behaves, when  $t \rightarrow -\infty$ , like

$$I_L(t) = A_L t^\ell + B_L t^{\ell-1} \ln(-t) + C_L t^{\ell-1} + o(t^{\ell-1}), \quad (2.11)$$

where  $A_L$ ,  $B_L$ , and  $C_L$  are constant tensors. The time derivatives of the moment appearing in Eq. (2.10) are therefore dominantly like

$$I_L^{(\ell+2)}(t) = D_L t^{-3} + o(t^{-3}), \quad (2.12)$$

which ensures that the integral (2.10) is convergent.

Next we integrate the remote-past integral (from  $T_R - \mathcal{T}$  to  $-\infty$ ) by parts and make use of our assumption (2.11) and (2.12) to arrive at (posing  $t = T_R - \mathcal{T}x$ )

$$\begin{aligned} \mathcal{U}_L(T_R) &= I_L^{(\ell+1)}(T_R) \ln\left(\frac{\mathcal{T}}{2\tau'_0}\right) + \mathcal{T} \int_0^1 dx \ln x I_L^{(\ell+2)}(T_R - \mathcal{T}x) \\ &+ \int_1^{+\infty} \frac{dx}{x} I_L^{(\ell+1)}(T_R - \mathcal{T}x). \end{aligned} \quad (2.13)$$

At this stage it is convenient to perform a Fourier decomposition of the multipole moment, i.e.

$$I_L(t) = \int_{-\infty}^{+\infty} \frac{d\Omega}{2\pi} \tilde{I}_L(\Omega) e^{-i\Omega t}. \quad (2.14)$$

[The Fourier coefficients satisfy  $\tilde{I}_L^*(\Omega) = \tilde{I}_L(-\Omega)$  since the moment is real.] Inserting (2.14) into (2.13) we obtain a closed-form result in the Fourier domain thanks to the mathematical formula [27]

$$\lambda \int_0^1 dx \ln x e^{i\lambda x} + i \int_1^{+\infty} \frac{dx}{x} e^{i\lambda x} = -\frac{\pi}{2} s(\lambda) - i(\ln|\lambda| + \gamma_E), \quad (2.15)$$

where  $\lambda = \Omega\mathcal{T}$ , with  $s(\lambda)$  and  $|\lambda|$  denoting the sign and the absolute value, and  $\gamma_E$  being the Euler constant. Finally the result reads

$$\begin{aligned} \mathcal{U}_L(T_R) &= i \int_{-\infty}^{+\infty} \frac{d\Omega}{2\pi} (-i\Omega)^{\ell+1} \tilde{I}_L(\Omega) e^{-i\Omega T_R} \\ &\times \left[ \frac{\pi}{2} s(\Omega) + i(\ln(2|\Omega|\tau'_0) + \gamma_E) \right]. \end{aligned} \quad (2.16)$$

We observe that the arbitrary time scale  $\mathcal{T}$  has been canceled from this result.

Later we shall apply this result to the computation of the waveform and energy flux of a spinning compact binary system. *A priori*, since the tail integral (2.10) depends on all the past history of the binary (with the current binary’s dynamics being the result of its long evolution by gravitational radiation emission), we expect that the binary’s continuous spectrum of frequencies  $\Omega$  should contain all orbital frequencies at any epoch in the past, say  $\omega(t)$  with  $t \leq T_R$ , besides the current orbital frequency  $\omega(T_R)$ . However, it has been shown in the Appendix of Ref. [16] that one can actually compute the tail integral by considering only the *current* frequency  $\omega(T_R)$ . Indeed the error made by this procedure is small in a post-Newtonian sense, being of the order of  $\mathcal{O}(\xi \ln \xi)$ , where  $\xi = \dot{\omega}/\omega^2$  denotes the adiabatic parameter associated with the gravitational radiation emission, and evaluated at the current time  $T_R$ . In a PN expansion we have  $\xi(T_R) = \mathcal{O}(1/c^5)$  so the error made by replacing the past dynamics by the current one is of the order of  $\mathcal{O}(\ln c/c^5)$  and can be neglected. The proof given in Ref. [16] is based on a simple model of binary evolution in the past, where an always circular orbit is decaying by radiation following the lowest-order quadrupole formula, and spins are neglected. In this paper we shall assume that this result remains valid for spinning binaries.

### III. APPLICATIONS TO SPINNING BINARIES

#### A. Spin vectors for pointlike objects

Following our previous work [19,20] we base our calculations on the model of point particles with spins [28–40]. The stress-energy tensor  $T^{\mu\nu}$  of a system of spinning particles is the sum of a monopolar piece, made of Dirac delta functions, plus the dipolar or spin piece, made of gradients of delta functions:

$$T^{\mu\nu} = c^2 \sum_A \int_{-\infty}^{+\infty} d\tau_A \left\{ p_A^{(\mu} u_A^{\nu)} \frac{\delta^{(4)}(x - y_A)}{\sqrt{-g_A}} - \frac{1}{c} \nabla_\rho \left[ \bar{S}_A^{\rho(\mu} u_A^{\nu)} \frac{\delta^{(4)}(x - y_A)}{\sqrt{-g_A}} \right] \right\}, \quad (3.1)$$

where  $\delta^{(4)}$  is the four-dimensional Dirac function,  $x^\mu$  is the field point,  $y_A^\mu$  is the worldline of particle  $A$ ,  $u_A^\mu = dy_A^\mu / (cd\tau_A)$  is the four-velocity, such that  $g_{\mu\nu}^A u_A^\mu u_A^\nu = -1$ , where  $g_{\mu\nu}^A \equiv g_{\mu\nu}(y_A)$  denotes the metric at the particle's location,  $p_A^\mu$  is the linear momentum of the particle, and  $\bar{S}_A^{\mu\nu}$  denotes its antisymmetric spin angular momentum. Our notation and conventions are the same as in Refs. [19,20] which provide more details, except that here we shall denote using an overbar (i.e.  $\bar{S}_A^{\mu\nu}$ ) the original spin variable used in [19,20]. Note that with our convention the spin variable has the dimension of an angular momentum times  $c$ .

In order to fix unphysical degrees of freedom associated with an arbitrariness in the definition of  $\bar{S}^{\mu\nu}$  in the case of point particles (and associated with the freedom in the choice for the location of the center-of-mass worldline of extended bodies), we adopt the covariant supplementary spin condition, also called the Tulczyjew condition [33,34]:

$$\bar{S}_A^{\mu\nu} p_\nu^A = 0, \quad (3.2)$$

which allows the natural definition of the spin four-vector  $\bar{S}_\mu^A$  in such a way that

$$\bar{S}_A^{\mu\nu} = -\frac{1}{\sqrt{-g_A}} \varepsilon^{\mu\nu\rho\sigma} \frac{p_\rho^A}{m_{AC}} \bar{S}_\sigma^A, \quad (3.3)$$

where  $\varepsilon^{\mu\nu\rho\sigma}$  is the four-dimensional antisymmetric Levi-Civita symbol (such that  $\varepsilon^{0123} = 1$ ). For the spin vector  $\bar{S}_\mu^A$  itself, we choose a four-vector that is purely spatial in the particle's instantaneous rest frame, which means that in any frame

$$\bar{S}_\mu^A u_A^\mu = 0. \quad (3.4)$$

This choice is also adopted in Refs. [41–44]. As a consequence of the condition (3.2), we can check that the mass defined by  $m_A^2 c^2 = -p_A^\mu p_\mu^A$  is constant along the trajectories, i.e.  $dm_A/d\tau_A = 0$ .

Important simplifications occur in the case of SO interactions, which are linear in the spins. Neglecting quadratic (spin-spin) interactions, the linear momentum is simply linked to the four-velocity as  $p_A^\mu = m_A c u_A^\mu$ , so the supplementary spin condition (3.2) reduces to  $\bar{S}_A^{\mu\nu} u_\nu^A = 0$ , and the equation of evolution of the spins is given by

$$D \bar{S}_\mu^A / d\tau_A = 0, \quad (3.5)$$

which means that the spin is parallelly transported along the particle's trajectory.

Following [19,20] we adopt in a first stage as the vector spin variable the *contravariant* components of the vector  $\bar{S}_A^i$ , which are obtained by raising the index on  $\bar{S}_A^i$  by means of the spatial metric  $\gamma_A^{ij}$ , denoting the inverse of the covariant spatial metric  $\gamma_{ij}^A \equiv g_{ij}^A$  evaluated at point  $A$  (i.e. such that  $\gamma_A^{ik} \gamma_{kj}^A = \delta^i_j$ ). Hence our initial spin variable is

$$\bar{S}_A^i = \gamma_A^{ij} \bar{S}_j^A. \quad (3.6)$$

This definition of the spin vector  $\bar{S}_A = (\bar{S}_A^i)$  agrees with the choice already made in Refs. [43,44].

At the leading SO approximation, the contravariant spin variables  $\bar{S}_A$  defined by Eq. (3.6) coincide with the spin variables with constant magnitude broadly used in the literature (see, e.g., Ref. [42]). At the next-to-leading order, the variables  $\bar{S}_A$  differ from constant-magnitude spins, and their relationship has been worked out at 2PN order in Eq. (7.4) of Ref. [20]. In the present paper we shall denote the constant-magnitude spins by  $S_A$  (although they were denoted  $S_A^c$  in Refs. [19,20]). We know that it is actually better when presenting final results to switch to the constant-magnitude spins  $S_A$  since they have a simpler precession equation (and turn out to be secularly conserved, i.e., over a radiation-reaction time scale; see Ref. [45] and Appendix B below).

For two bodies ( $A = 1, 2$ ) the relationship between the constant-magnitude spins and the original spin variables up to 1PN order is

$$S_1 = \bar{S}_1 + \frac{1}{c^2} \left[ -\frac{1}{2} (v_1 \bar{S}_1) v_1 + \frac{Gm_2}{r_{12}} \bar{S}_1 \right] + \mathcal{O}\left(\frac{1}{c^4}\right), \quad (3.7)$$

together with the relation for the other particle obtained by exchanging all particle labels. We denote by  $v_A = dy_A/dt$  the coordinate velocity of the particle  $A$  (with mass  $m_A$ ) and by  $r_{12} = |y_1 - y_2|$  the relative distance. See Ref. [20] for more accurate formulas extending (3.7) to 2PN order. In the present paper we shall consistently work only with the constant-magnitude spins  $S_A$ .

In the case of binary systems it is convenient to pose

$$S = S_1 + S_2, \quad (3.8a)$$

$$\Sigma = \frac{S_2}{X_2} - \frac{S_1}{X_1}, \quad (3.8b)$$

where  $X_1 = m_1/m$  and  $X_2 = m_2/m$  (with  $m = m_1 + m_2$ ). In addition we find it useful to occasionally use the dimensionless spin variables

$$s = \frac{S}{Gm^2}, \quad \sigma = \frac{\Sigma}{Gm^2}. \quad (3.9)$$

### B. Multipole moments with spin-orbit effects

The matter-source densities (2.9) depend on the components of the stress-energy tensor. At the leading PN order, the spin contributions therein (indicated by the subscript  $S$ ) reduce to

$$\sigma_S = -\frac{2}{c^3} \varepsilon_{ijk} v_1^i S_1^j \partial_k \delta_1 + 1 \leftrightarrow 2 + \mathcal{O}\left(\frac{1}{c^5}\right), \quad (3.10a)$$

$$\sigma_{S_i} = -\frac{1}{2c} \varepsilon_{ijk} S_1^j \partial_k \delta_1 + 1 \leftrightarrow 2 + \mathcal{O}\left(\frac{1}{c^3}\right), \quad (3.10b)$$

$$\sigma_{S_{ij}} = -\frac{1}{c} \varepsilon_{kl(i} v_1^j) S_1^k \partial_l \delta_1 + 1 \leftrightarrow 2 + \mathcal{O}\left(\frac{1}{c^3}\right), \quad (3.10c)$$

where  $\delta_1(\mathbf{x}, t) = \delta[\mathbf{x} - \mathbf{y}_1(t)]$  means the three-dimensional Dirac delta function evaluated on particle 1, and  $1 \leftrightarrow 2$  means the same quantity but corresponding to particle 2.

In Ref. [20] the SO terms have been computed in the source mass quadrupole moment  $I_{ij}$  up to next-to-leading 2.5PN order and the source current quadrupole moment  $J_{ij}$  up to next-to-leading 1.5PN order. All the other source moments were computed at the leading SO order. Those results are sufficient for our purpose. Actually, to compute the specific contributions of tails we need only the moments at leading SO order, given for general  $\ell$  by

$$I_{S^L} = \frac{2\ell}{c^3(\ell+1)} [\ell v_1^i S_1^j \varepsilon_{ij\langle i_\ell} v_1^{L-1} y_1^{L-1} \rangle - (\ell-1) y_1^i S_1^j \varepsilon_{ij\langle i_\ell} v_1^{L-1} y_1^{L-2} \rangle] + 1 \leftrightarrow 2 + \mathcal{O}\left(\frac{1}{c^5}\right), \quad (3.11a)$$

$$J_{S^L} = \frac{\ell+1}{2c} y_1^{L-1} S_1^{i_\ell} + 1 \leftrightarrow 2 + \mathcal{O}\left(\frac{1}{c^3}\right). \quad (3.11b)$$

Because the leading SO terms scale as  $\mathcal{O}(1/c^3)$  in the mass source moments, and as  $\mathcal{O}(1/c)$  in the current source moments, the number of nonlinear terms needed in the radiative moments [Eqs. (5.1) below] is small. We refer to Sec. V of [20] for higher-order expressions of SO contributions of the source quadrupole moments.

### C. Equations of motion with spin-orbit effects

Here we investigate the case where the binary's orbit is nearly circular, i.e., whose radius is constant apart from small perturbations induced by the spins (as usual, we neglect the gravitational radiation damping at 2.5PN order). We denote by  $\mathbf{x} = \mathbf{y}_1 - \mathbf{y}_2$  the relative position of the particles (and  $\mathbf{v} = d\mathbf{x}/dt$ ). Following Ref. [42] we introduce an orthonormal moving triad  $\{\mathbf{n}, \boldsymbol{\lambda}, \boldsymbol{\ell}\}$  defined by  $\mathbf{n} = \mathbf{x}/r$  as before,  $\boldsymbol{\ell} = \mathbf{L}_N/|\mathbf{L}_N|$ , where  $\mathbf{L}_N \equiv m\nu\mathbf{x} \times \mathbf{v}$  with  $\nu = X_1 X_2$  denotes the Newtonian orbital angular momentum and  $\nu$  the symmetric mass ratio, and  $\boldsymbol{\lambda} = \boldsymbol{\ell} \times \mathbf{n}$ . Those vectors are represented on Fig. 1, which shows the geometry of the system. The orbital frequency  $\omega$  is defined for general, not necessarily circular orbits, by

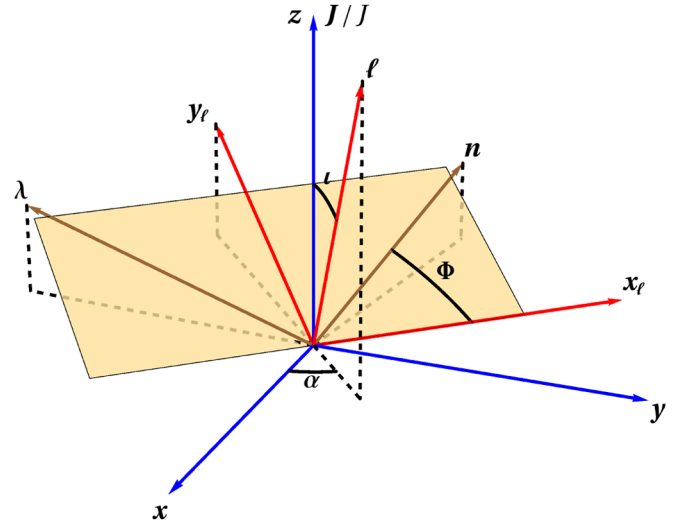


FIG. 1 (color online). We show (i) the source frame defined by the orthonormal basis  $(x, y, z)$ , (ii) the instantaneous orbital plane which is described by the orthonormal basis  $(x_\ell, y_\ell, \ell)$ , (iii) the moving triad  $(\mathbf{n}, \boldsymbol{\lambda}, \boldsymbol{\ell})$ , and (iv) the direction of the total angular momentum  $\mathbf{J}$  (agreeing, by definition, with the  $z$  direction). Dashed lines show projections into the  $x - y$  plane.

$\mathbf{v} = \dot{r}\mathbf{n} + r\omega\boldsymbol{\lambda}$ , where  $\dot{r}$  represents the derivative of  $r$  with respect to the coordinate time  $t$ . It is also equal to the scalar product of  $\mathbf{n}$  and  $\mathbf{v}$  which we denote as  $(\mathbf{n}\mathbf{v}) = \dot{r}$ . The components of the acceleration  $\mathbf{a} = d\mathbf{v}/dt$  along the basis  $\{\mathbf{n}, \boldsymbol{\lambda}, \boldsymbol{\ell}\}$  are then given by

$$\mathbf{n} \cdot \mathbf{a} = \ddot{r} - r\omega^2, \quad (3.12a)$$

$$\boldsymbol{\lambda} \cdot \mathbf{a} = r\dot{\omega} + 2\dot{r}\omega, \quad (3.12b)$$

$$\boldsymbol{\ell} \cdot \mathbf{a} = -r\omega \left( \boldsymbol{\lambda} \cdot \frac{d\boldsymbol{\ell}}{dt} \right). \quad (3.12c)$$

We project out the spins on this orthonormal basis, defining  $\mathbf{S} = S_n\mathbf{n} + S_\lambda\boldsymbol{\lambda} + S_\ell\boldsymbol{\ell}$ , and similarly for  $\boldsymbol{\Sigma}$ . Next we impose the restriction to quasicircular precessing orbits which is defined by the conditions  $\ddot{r} = 0 = \dot{r}$  so that  $v^2 = r^2\omega^2$  (neglecting radiation-reaction damping terms). In this way we find [19] that the equations of the relative motion in the frame of the center of mass are

$$\frac{d\mathbf{v}}{dt} = -\omega r[\omega\mathbf{n} + \omega_{\text{prec}}\boldsymbol{\ell}] + \mathcal{O}\left(\frac{1}{c^6}\right). \quad (3.13)$$

There is no component of the acceleration along  $\boldsymbol{\lambda}$ . Comparing with Eqs. (3.12) in the case of circular orbits, we see that  $\omega$  is indeed the orbital frequency, while what we call the ‘‘precessional frequency’’  $\omega_{\text{prec}} = \boldsymbol{\lambda} \cdot d\boldsymbol{\ell}/dt$  is proportional to the variation of  $\boldsymbol{\ell}$  in the direction of the velocity  $\mathbf{v} = r\omega\boldsymbol{\lambda}$ . We know that  $\omega^2$  is given by

$$\omega^2 = \frac{Gm}{r^3} \{1 + \gamma(-3 + \nu) + \gamma^{3/2}(-5s_\ell - 3\delta\sigma_\ell)\} + \mathcal{O}\left(\frac{1}{c^4}\right), \quad (3.14)$$

where we denote  $\delta \equiv X_1 - X_2$  and  $s_\ell \equiv (s\ell) = \mathbf{s} \cdot \boldsymbol{\ell}$ , where the spin variables are defined by Eq. (3.9). The PN parameter is  $\gamma \equiv Gm/(rc^2)$  and we have included only the 1PN nonspin term and the leading SO correction at 1.5PN order. On the other hand, we get [19]

$$\omega_{\text{prec}} = -\omega\gamma^{3/2}(7s_n + 3\delta\sigma_n) + \mathcal{O}\left(\frac{1}{c^4}\right), \quad (3.15)$$

where  $s_n \equiv (sn) = \mathbf{s} \cdot \mathbf{n}$ . At the leading 1.5PN SO order the orbital frequency (3.14), as well as  $\omega_{\text{prec}}$ , remain unchanged if we were to substitute some other variables to the spins  $\mathbf{S}$ ,  $\boldsymbol{\Sigma}$ . However, when working at a higher PN approximation, it is more convenient to use the spin variables  $\mathbf{S}$ ,  $\boldsymbol{\Sigma}$ , built from the constant-magnitude spins. The main advantage of the spins  $\mathbf{S}_A$  is that they satisfy the usual-looking precession equations

$$\frac{d\mathbf{S}_A}{dt} = \boldsymbol{\Omega}_A \times \mathbf{S}_A, \quad (3.16)$$

showing that the spins precess around the direction of  $\boldsymbol{\Omega}_A$ , and at the rate  $\Omega_A = |\boldsymbol{\Omega}_A|$ . Equation (3.16) could, in principle, be extended to any PN order (at the linear SO level). The precession's angular-frequency vectors  $\boldsymbol{\Omega}_A$  have been computed up to the 2PN order for circular orbits in Ref. [20]. Here, we shall only need the 1PN leading order:

$$\boldsymbol{\Omega}_1 = \omega\gamma \left[ \frac{3}{4} + \frac{\nu}{2} - \frac{3}{4}\delta \right] \boldsymbol{\ell} + \mathcal{O}\left(\frac{1}{c^4}\right). \quad (3.17)$$

To obtain  $\boldsymbol{\Omega}_2$  we simply have to change  $\delta$  into  $-\delta$ . Both precession frequencies are constant in magnitude and independent of the spins in the 1.5PN dynamics.

The equations of motion (3.13) and the precession equations (3.16) together leave invariant the total angular momentum,

$$\mathbf{J} = \mathbf{L} + \frac{1}{c}\mathbf{S}, \quad \frac{d\mathbf{J}}{dt} = 0, \quad (3.18)$$

where  $\mathbf{L}$  denotes the orbital angular momentum. For future reference we give the components of  $\mathbf{L}$  along the triad basis at 1PN order for nonspin effects and at the leading 1.5PN order for spin ones [19,20]:

$$L_\ell = \frac{Gm^2\nu}{c} x^{-1/2} \left[ 1 + \left( \frac{3}{2} + \frac{\nu}{6} \right) x \right], \quad (3.19a)$$

$$L_n = \frac{\nu x}{c} \left[ \frac{1}{2} S_n + \frac{1}{2} \delta \Sigma_n \right], \quad (3.19b)$$

$$L_\lambda = \frac{\nu x}{c} [-3S_\lambda - \delta \Sigma_\lambda]. \quad (3.19c)$$

Note that the components  $L_n$  and  $L_\lambda$  are due to spin effects arising at order  $\mathcal{O}(c^{-3})$ . See Eq. (7.10) of Ref. [20].

#### IV. EVOLUTION OF THE TRIAD $\{\mathbf{n}, \boldsymbol{\lambda}, \boldsymbol{\ell}\}$

Using Eq. (3.13) the time derivatives of the three moving triad vectors  $\{\mathbf{n}, \boldsymbol{\lambda}, \boldsymbol{\ell}\}$  can be expressed with respect to that triad basis as

$$\frac{d\mathbf{n}}{dt} = \boldsymbol{\omega} \boldsymbol{\lambda}, \quad (4.1a)$$

$$\frac{d\boldsymbol{\lambda}}{dt} = -\boldsymbol{\omega} \mathbf{n} - \boldsymbol{\omega}_{\text{prec}} \boldsymbol{\ell}, \quad (4.1b)$$

$$\frac{d\boldsymbol{\ell}}{dt} = \boldsymbol{\omega}_{\text{prec}} \boldsymbol{\lambda}. \quad (4.1c)$$

Equivalently, introducing the orbital precession vector  $\boldsymbol{\omega} = \omega \boldsymbol{\ell}$  and spin precession vector  $\boldsymbol{\omega}_{\text{prec}} = -\boldsymbol{\omega}_{\text{prec}} \mathbf{n}$ , these equations can be written as

$$\frac{d\mathbf{n}}{dt} = \boldsymbol{\omega} \times \mathbf{n}, \quad (4.2a)$$

$$\frac{d\boldsymbol{\lambda}}{dt} = (\boldsymbol{\omega} + \boldsymbol{\omega}_{\text{prec}}) \times \boldsymbol{\lambda}, \quad (4.2b)$$

$$\frac{d\boldsymbol{\ell}}{dt} = \boldsymbol{\omega}_{\text{prec}} \times \boldsymbol{\ell}. \quad (4.2c)$$

We recall that the spin precession frequency is given by Eq. (3.15) or, equivalently,

$$\omega_{\text{prec}} = -\omega x^{3/2} (7s_n + 3\delta\sigma_n) + \mathcal{O}\left(\frac{1}{c^4}\right), \quad (4.3)$$

where we pose  $x \equiv (Gm\omega/c^3)^{2/3}$  which defines a gauge-invariant PN parameter agreeing with  $\gamma$  at leading order.

We shall now solve the evolution equations for the moving triad  $\{\mathbf{n}, \boldsymbol{\lambda}, \boldsymbol{\ell}\}$  at the 1.5PN order in the conservative dynamics (i.e., neglecting radiation reaction). We recall that the spin variables we use in this calculation are those with constant magnitude, denoted here  $\mathbf{S}_A$ .

It is convenient to introduce a fixed (inertial) orthonormal basis  $\{\mathbf{x}, \mathbf{y}, \mathbf{z}\}$ ; see Figs. 1 and 2. For a given value of the total angular momentum  $\mathbf{J}$  (a constant vector), and of the direction  $\mathbf{N} = \mathbf{X}/R$  of the detector as seen from the source, a canonical choice of the basis vectors is as follows: (i)  $\mathbf{z}$  is defined to be the normalized value of  $\mathbf{J}$ , namely,  $\mathbf{J}/J$ ; (ii)  $\mathbf{y}$  is orthogonal to the plane spanned by  $\mathbf{N}$  and  $\mathbf{z}$  and points to the direction that corresponds to the positive orientation of the acute angle  $\langle \mathbf{z}, \mathbf{N} \rangle$ , i.e.  $\mathbf{y} = \mathbf{z} \times \mathbf{N} / |\mathbf{z} \times \mathbf{N}|$ ; (iii)  $\mathbf{x}$  completes the triad. We see that  $\mathbf{x}$ ,  $\mathbf{z}$ , and  $\mathbf{N}$  are coplanar by construction. Then, we introduce the standard spherical coordinates with the inclination angle measured from the zenith direction  $\mathbf{z}$  and the azimuthal angle measured from  $\mathbf{x}$ . The spherical coordinates of  $\mathbf{N}$  and  $\boldsymbol{\ell}$  are denoted as  $(\theta, \varphi)$  and  $(\alpha, \iota)$ , respectively, and since  $\mathbf{N}$  lies in the same plane as  $\mathbf{x}$  and  $\mathbf{z}$ , we have  $\varphi = 0$  (see Fig. 2). Since  $\iota$  is the angle between the total and orbital angular momenta, we have

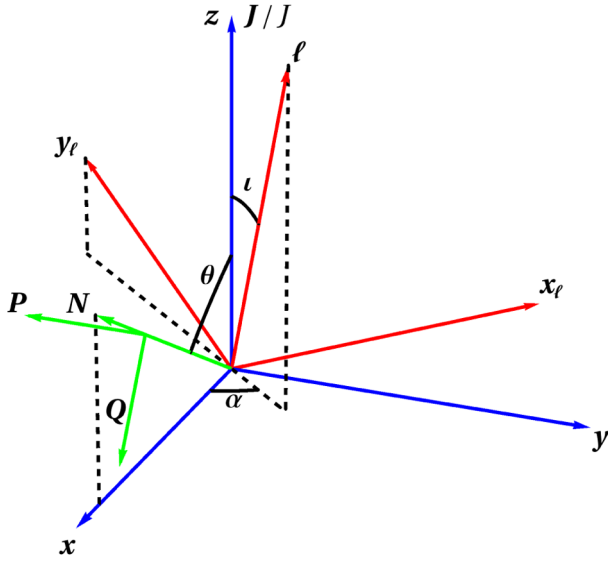


FIG. 2 (color online). Similar to Fig. 1 but with the direction of the source  $N$  indicated together with a choice of convention for the two polarization vectors  $P$  and  $Q$ .

$$\sin \iota = \frac{|\mathbf{J} \times \boldsymbol{\ell}|}{J}. \quad (4.4)$$

The angles  $(\alpha, \iota)$  are referred to as the precession angles.

We now derive the time evolution of our triad vectors from that of the precession angles  $(\alpha, \iota)$ , and of an appropriate phase  $\Phi$  that specifies the position of  $\mathbf{n}$  with respect to some reference direction. Following Ref. [46], we introduce the unit vectors

$$\mathbf{x}_\ell = \frac{\mathbf{J} \times \boldsymbol{\ell}}{|\mathbf{J} \times \boldsymbol{\ell}|}, \quad \mathbf{y}_\ell = \boldsymbol{\ell} \times \mathbf{x}_\ell, \quad (4.5)$$

such that  $\{\mathbf{x}_\ell, \mathbf{y}_\ell, \boldsymbol{\ell}\}$  is an orthonormal basis. The phase angle  $\Phi$  is defined by (see Fig. 1)

$$\Phi = \langle \mathbf{x}_\ell, \mathbf{n} \rangle = \langle \mathbf{y}_\ell, \boldsymbol{\lambda} \rangle. \quad (4.6)$$

The rotation takes place in the instantaneous orbital plane spanned by  $\mathbf{n}$  and  $\boldsymbol{\lambda}$ , and we have

$$\mathbf{n} = \cos \Phi \mathbf{x}_\ell + \sin \Phi \mathbf{y}_\ell, \quad (4.7a)$$

$$\boldsymbol{\lambda} = -\sin \Phi \mathbf{x}_\ell + \cos \Phi \mathbf{y}_\ell, \quad (4.7b)$$

from which we deduce

$$e^{-i\Phi} = \mathbf{x}_\ell \cdot (\mathbf{n} + i\boldsymbol{\lambda}) = \frac{J_\lambda - iJ_n}{\sqrt{J_n^2 + J_\lambda^2}}. \quad (4.8)$$

Combining (4.8) with (4.4) we also get

$$\sin \iota e^{-i\Phi} = \frac{J_\lambda - iJ_n}{J}. \quad (4.9)$$

By identifying the right-hand sides of Eqs. (4.1) or (4.2) with the time derivatives of the identities (4.7), we obtain the following system of equations for the variations of  $\alpha$ ,  $\iota$ , and  $\Phi$ , equivalent to the system (4.1),

$$\frac{d\alpha}{dt} = -\omega_{\text{prec}} \frac{\sin \Phi}{\sin \iota}, \quad (4.10a)$$

$$\frac{d\iota}{dt} = -\omega_{\text{prec}} \cos \Phi, \quad (4.10b)$$

$$\frac{d\Phi}{dt} = \omega + \omega_{\text{prec}} \frac{\sin \Phi}{\tan \iota}. \quad (4.10c)$$

On the other hand, using the total angular momentum (3.18) together with the components of the orbital angular momentum given by Eqs. (3.19)—notably the fact that  $L_n$  and  $L_\lambda$  are due to SO terms dominantly of order  $\mathcal{O}(c^{-3})$ —we deduce that  $\sin \iota$  is a small quantity of order  $\mathcal{O}(1/c)$ . From this fact, we conclude by direct integration of the sum of Eqs. (4.10a) and (4.10c) that

$$\Phi + \alpha = \phi + \mathcal{O}\left(\frac{1}{c^4}\right), \quad (4.11)$$

in which we have defined the “carrier” phase as

$$\phi = \int \omega dt = \omega(t - t_0) + \phi_0, \quad (4.12)$$

with  $\phi_0$  the value of the carrier phase at some arbitrary initial time  $t_0$ . We recall that the orbital frequency (3.14) is constant in a first approximation for circular motion.

The combination  $\Phi + \alpha$  being known by Eq. (4.11), we can further express the precession angles  $\iota$  and  $\alpha$  in a first approximation in terms of the components  $S_n$  and  $S_\lambda$  of the total spin  $\mathbf{S} = \mathbf{S}_1 + \mathbf{S}_2$ . From (4.4) we find (discarding nonlinear spin contributions)

$$\sin \iota = \frac{\sqrt{S_n^2 + S_\lambda^2}}{cL_N} + \mathcal{O}\left(\frac{1}{c^3}\right), \quad (4.13)$$

where we recall that  $L_N = mvr^2\omega$  denotes the Newtonian orbital angular momentum. On the other hand, using also Eq. (4.9) and the relation (4.11) we obtain at leading order

$$e^{i\alpha} = \frac{S_\lambda - iS_n}{\sqrt{S_n^2 + S_\lambda^2}} e^{i\phi} + \mathcal{O}\left(\frac{1}{c^2}\right). \quad (4.14)$$

[See also the more precise Eqs. (4.21) and (4.22).]

It remains now to obtain the explicit time variation of the components of the individual spins  $S_n^A$ ,  $S_\lambda^A$ , and  $S_\ell^A$ . Using (4.13) and (4.14) [and also (4.11)] we shall then be able to obtain the explicit time variation of the precession angles and phase. Combining (3.16) and (4.1) we obtain the precession equations for the three unknowns  $S_n^A$ ,  $S_\lambda^A$ , and  $S_\ell^A$  in the form of the following first-order system (valid at any PN approximation):



$$\frac{dS_n^A}{dt} = (\omega - \Omega_A)S_\lambda^A, \quad (4.15a)$$

$$\frac{dS_\lambda^A}{dt} = -(\omega - \Omega_A)S_n^A - \omega_{\text{prec}}S_\ell^A, \quad (4.15b)$$

$$\frac{dS_\ell^A}{dt} = \omega_{\text{prec}}S_\lambda^A, \quad (4.15c)$$

where  $\Omega_A$  is the norm of the precession vector of the spin  $A$  as given by (3.17), and the precession frequency  $\omega_{\text{prec}}$  is explicitly given by (4.3). Actually, the terms involving  $\omega_{\text{prec}}$  in the right-hand sides of (4.15) can be neglected because they are quadratic in the spins. Thus, staying at the linear SO level, we find that Eqs. (4.15) can be decoupled and integrated as

$$S_n^A = S_\perp^A \cos\psi_A, \quad (4.16a)$$

$$S_\lambda^A = -S_\perp^A \sin\psi_A, \quad (4.16b)$$

$$S_\ell^A = S_\parallel^A. \quad (4.16c)$$

Here  $S_\perp^A$  and  $S_\parallel^A$  denote two constants for each spin  $A$ , and agree with the projections (which are constant at the linear SO level) of the spins perpendicular and parallel to the direction of the orbital angular momentum  $\ell$ . The phase of each of the spins is given by

$$\psi_A = (\omega - \Omega_A)(t - t_0) + \psi_A^0, \quad (4.17)$$

where  $\psi_A^0$  is the constant initial phase at time  $t_0$ .

With those results we obtain an explicit solution for the precession angles by substituting Eqs. (4.16) into the results (4.13) and (4.14). We find that  $\iota(t)$  is given at the 0.5PN level by

$$\begin{aligned} \sin\iota &= \frac{x^{1/2}}{\nu} \sqrt{(s_\perp^1)^2 + (s_\perp^2)^2 + 2s_\perp^1 s_\perp^2 \cos(\psi_1 - \psi_2)} \\ &+ \mathcal{O}\left(\frac{1}{c^3}\right), \end{aligned} \quad (4.18)$$

where we recall that  $s_\perp^A = S_\perp^A/(Gm^2)$ . Knowing  $\iota(t)$  we deduce  $\alpha(t)$  from

$$\sin\iota e^{i\alpha} = -i \frac{x^{1/2}}{\nu} e^{i\phi} (s_\perp^1 e^{-i\psi_1} + s_\perp^2 e^{-i\psi_2}) + \mathcal{O}\left(\frac{1}{c^3}\right). \quad (4.19)$$

The difference of spin phases  $\psi_{12} \equiv \psi_1 - \psi_2$  readily follows from Eqs. (3.17) and (4.17) at 1PN order as

$$\psi_{12} = \psi_{12}^0 + \frac{3}{2} \omega x \delta(t - t_0) + \mathcal{O}\left(\frac{1}{c^4}\right). \quad (4.20)$$

Moreover, Eq. (4.9) can be written more explicitly at the 1.5PN level as

$$\sin\iota e^{-i\Phi} = -i \frac{J_+}{L_\ell} + \mathcal{O}\left(\frac{1}{c^4}\right), \quad (4.21)$$

where  $J_+ \equiv J_n + iJ_\lambda$  is given at the 1.5PN order by

$$\begin{aligned} J_+ &= \frac{S_\perp^1}{c} \left\{ e^{-i\psi_1} \left[ 1 + x \left( -\frac{1}{8} - \frac{3}{4} \nu + \frac{\delta}{8} \right) \right] \right. \\ &\quad \left. + e^{i\psi_1} x \left( \frac{3}{8} + \frac{\nu}{4} - \frac{3\delta}{8} \right) \right\} + 1 \leftrightarrow 2 + \mathcal{O}\left(\frac{1}{c^4}\right), \end{aligned} \quad (4.22)$$

and the 1PN orbital angular momentum  $L_\ell$  is known from Eq. (3.19a).

As a check of the previous solution we observe that if we take the time derivative of Eq. (4.4), then evaluate the total angular momentum  $\mathbf{J}$  given by (3.18) together with the components of the orbital angular momentum  $\mathbf{L}$  provided in (3.19), and use the solution (4.15) for the evolution of the spin components, we obtain

$$\frac{d\iota}{dt} = -\omega_{\text{prec}} \frac{S_\lambda}{\sqrt{S_n^2 + S_\lambda^2}} + \mathcal{O}\left(\frac{1}{c^4}\right), \quad (4.23)$$

which is consistent with (4.10b) once (4.8) is employed.

Finally, we express the triad vectors  $\mathbf{n}(t)$ ,  $\mathbf{\lambda}(t)$ , and  $\ell(t)$  in terms of the precession angles and the carrier phase, and in terms of the initial triad and angles at the initial instant  $t_0$ , modulo terms of order  $\mathcal{O}(c^{-4})$ . To do this we notice that the triad  $\{\mathbf{n}, \mathbf{\lambda}, \ell\}$  at time  $t$  is obtained from the inertial triad  $\{\mathbf{x}, \mathbf{y}, \mathbf{z}\}$  by the rotation associated with the three Euler angles  $\alpha$ ,  $\iota$ , and  $\Phi$ . Similarly, the initial triad  $\{\mathbf{n}_0, \mathbf{\lambda}_0, \ell_0\}$  at time  $t_0$  is obtained by the rotation associated with  $\alpha_0$ ,  $\iota_0$ , and  $\Phi_0$ . So, combining those two rotations we readily obtain  $\{\mathbf{n}, \mathbf{\lambda}, \ell\}$  in terms of  $\{\mathbf{n}_0, \mathbf{\lambda}_0, \ell_0\}$ . Using Eq. (4.11) to eliminate the phase  $\Phi$  in favor of the carrier phase  $\phi$ —this introduces small remainder terms  $\mathcal{O}(c^{-4})$ —and neglecting all terms quadratic in the spins, we get

$$\begin{aligned} \mathbf{n} &= \cos(\phi - \phi_0) \mathbf{n}_0 + \sin(\phi - \phi_0) \mathbf{\lambda}_0 + (\sin\iota \sin(\phi - \alpha) \\ &\quad - \sin\iota_0 \sin(\phi - \alpha_0)) \ell_0 + \mathcal{O}\left(\frac{1}{c^4}\right), \end{aligned} \quad (4.24a)$$

$$\begin{aligned} \mathbf{\lambda} &= -\sin(\phi - \phi_0) \mathbf{n}_0 + \cos(\phi - \phi_0) \mathbf{\lambda}_0 + (\sin\iota \cos(\phi - \alpha) \\ &\quad - \sin\iota_0 \cos(\phi - \alpha_0)) \ell_0 + \mathcal{O}\left(\frac{1}{c^4}\right), \end{aligned} \quad (4.24b)$$

$$\begin{aligned} \ell &= \ell_0 + (-\sin\iota \sin(\phi_0 - \alpha) + \sin\iota_0 \sin(\phi_0 - \alpha_0)) \mathbf{n}_0 \\ &\quad + (-\sin\iota \cos(\phi_0 - \alpha) + \sin\iota_0 \cos(\phi_0 - \alpha_0)) \mathbf{\lambda}_0 + \mathcal{O}\left(\frac{1}{c^4}\right). \end{aligned} \quad (4.24c)$$

Since we have found in Eqs. (4.18) and (4.19) an explicit solution for the precession angles  $\iota(t)$  and  $\alpha(t)$ , the time dependence of  $\mathbf{n}$ ,  $\mathbf{\lambda}$ , and  $\ell$  is completely known. Note that in practical computations it is often more convenient to work not with  $\mathbf{n}$  and  $\mathbf{\lambda}$  but with the complex null vector  $\mathbf{m} = (\mathbf{n} + i\mathbf{\lambda})/\sqrt{2}$  and its conjugate.

## V. COMPUTATION OF THE WAVEFORM

Here we shall compute the SO terms coming from all nonlinear (i.e. of formal order  $G^2$ ) contributions associated

with tails consistent with the 2.5PN and 3PN orders in the waveform. We shall need only to focus on the tails entering the mass and current quadrupoles  $U_{ij}$  and  $V_{ij}$  (having  $\ell = 2$ ) and on the current octupole  $V_{ijk}$  ( $\ell = 3$ ). They indeed contain, when specialized to spinning compact binary systems, the SO contributions we are interested in. The reason is that the leading SO terms start at the 0.5PN order  $\mathcal{O}(1/c)$  in the current moments, but only at the 1.5PN order  $\mathcal{O}(1/c^3)$  in the mass moments [see Eqs. (3.11) above].

In addition to the tail integrals shown in Eq. (2.4), we shall also compute some terms of order  $G^2$  at 2.5PN or 1.5PN order, but which are nonhereditary, i.e., merely depend on the instantaneous retarded time  $T_R$ . Those corrections are given in full form by Eqs. (5.4)–(5.6) of Ref. [47], but here we shall need only, for the same reason as before, two terms involving the source current dipole moment or angular momentum  $J_i$ ; the other terms will not contribute to the SO effect at 3PN order. Furthermore, we have to include some additional corrections of similar nature originating from the differences between the canonical and the source moments in Eqs. (2.6). Those are given in the general case by Eqs. (5.9)–(5.10) of Ref. [47], but we can check that only the contribution in the mass quadrupole  $U_{ij}$  *a priori* matters here.

The relevant contributions in the radiative moments, including only the terms needed for the applications below (see Ref. [47] for more complete expressions) read

$$\begin{aligned} \delta U_{ij} = & I_{ij}^{(2)} + \frac{2Gm}{c^3} \int_{-\infty}^{T_R} dt \left[ \ln\left(\frac{T_R - t}{2\tau_0}\right) + \frac{11}{12} \right] I_{ij}^{(4)}(t) \\ & + \frac{G}{c^5} \left( \frac{1}{3} \varepsilon_{ab\langle i} I_{j\rangle a}^{(4)} J_b + 4[W^{(2)} I_{ij} - W^{(1)} I_{ij}^{(1)}]^{(2)} \right), \end{aligned} \quad (5.1a)$$

$$\delta V_{ij} = J_{ij}^{(2)} + \frac{2Gm}{c^3} \int_{-\infty}^{T_R} dt \left[ \ln\left(\frac{T_R - t}{2\tau_0}\right) + \frac{7}{6} \right] J_{ij}^{(4)}(t), \quad (5.1b)$$

$$\begin{aligned} \delta V_{ijk} = & J_{ijk}^{(3)} + \frac{2Gm}{c^3} \int_{-\infty}^{T_R} dt \left[ \ln\left(\frac{T_R - t}{2\tau_0}\right) + \frac{5}{3} \right] J_{ijk}^{(5)}(t) \\ & - \frac{2G}{c^3} J_{(i} I_{jk)}^{(4)}, \end{aligned} \quad (5.1c)$$

where we have replaced  $M$  with  $m = m_1 + m_2$  which is valid at the dominant order. Moreover, we shall find that the terms in the right-hand side of Eq. (5.1a) which depend on the moment  $W$  vanish at the considered order. Indeed, by inserting the value of  $\sigma_i$  from (3.10b) into the potential  $W$  defined by (2.8b) and integrating by parts, we find that the result is zero.

The corresponding gravitational waveform, for which all three moments  $U_{ij}$ ,  $V_{ij}$ , and  $V_{ijk}$  are important, is then given by

$$\begin{aligned} \delta h_{ij}^{\text{TT}} = & \frac{2G}{c^4 R} \mathcal{P}_{ijkl}^{\text{TT}} \left[ \delta U_{kl} - \frac{4}{3c} N_a \varepsilon_{ab\langle k} \delta V_{l\rangle b} \right. \\ & \left. - \frac{1}{2c^2} N_{am} \varepsilon_{ab\langle k} \delta V_{l\rangle bm} \right]. \end{aligned} \quad (5.2)$$

The main task consists in evaluating the tail integrals (2.10) whose Fourier transforms have already been obtained in Eq. (2.16). The result (2.16) heavily relied on a physical assumption concerning the system in the remote past, namely, that it was formed by freely falling incoming particles; see Eq. (2.11).

We reviewed in Sec. II B that one can insert in the result (2.16) the binary's *current* frequency spectrum, i.e. at time  $T_R$ , modulo small error terms of the order of the adiabatic parameter of the inspiral, or, more precisely, of negligible order  $\mathcal{O}(\ln c/c^5)$ . In the case of spinning compact binaries this means that we have to include in the spectrum the current orbital frequency  $\omega \equiv \omega(T_R)$ , and also the precession frequencies  $\Omega_1 \equiv \Omega_1(T_R)$  and  $\Omega_2 \equiv \Omega_2(T_R)$  of the two spins. This follows from the explicit solution of the triad  $\{\mathbf{n}, \boldsymbol{\lambda}, \boldsymbol{\ell}\}$  and of the precession equations (see Sec. III). [Notice that the precession angles  $\alpha$  and  $\iota$  always appear through the product  $\sin \iota e^{i\alpha}$  given by equations such as (4.19).] Hence we can take for the Fourier components of the multipole moments

$$\tilde{I}_L(\Omega) = 2\pi \sum_{n,n_1,n_2} A_L^{n,n_1,n_2} \delta(\Omega - \omega_{n,n_1,n_2}), \quad (5.3)$$

where the frequency modes are some  $\omega_{n,n_1,n_2} = n\omega + n_1\Omega_1 + n_2\Omega_2$ . The result (2.16) then becomes

$$\begin{aligned} \mathcal{U}_L(T_R) = & \sum_{n,n_1,n_2} i A_L^{n,n_1,n_2} (-i\omega_{n,n_1,n_2})^{\ell+1} e^{-i\omega_{n,n_1,n_2} T_R} \\ & \times \left[ \frac{\pi}{2} s(\omega_{n,n_1,n_2}) + i(\ln(2|\omega_{n,n_1,n_2}| \tau'_0) + \gamma_E) \right]. \end{aligned} \quad (5.4)$$

We recall from Eq. (3.17) that the precession frequencies  $\Omega_1$  and  $\Omega_2$  are small quantities of order 1PN. This means, in particular, that because of the explicit factor  $\omega_{n,n_1,n_2}^{\ell+1}$  in Eq. (5.4) (which arises from taking the time derivatives of the multipole moment and integrating), the modes for which  $n = 0$  in tail integrals are very small, at least of order 4.5PN, and can be neglected.

The SO terms in the radiative tails originate primarily from the spins present in the sources of the integrals, i.e. appropriate derivatives of multipole moments as shown in Eqs. (5.1). The SO contributions in the source moments have been given in Eqs. (3.11). However, there are other crucial SO terms which originate from the nonspin parts of the moments. They come from time differentiations of the triad vectors using the evolution equations (4.1), which produce spin terms contained in the precession frequency  $\omega_{\text{prec}}$ , and from time differentiations of spins themselves *via* the precession equations (3.16). SO contributions may

also be generated by the tail integration itself due to the precession of the triad basis  $\{\mathbf{n}, \boldsymbol{\lambda}, \boldsymbol{\ell}\}$  according to the formula (4.24). On the other hand, other SO terms come from the reduction to circular orbits when we eliminate the orbital separation  $r$  in favor of a function of the orbital frequency  $\omega$  obtained from inverting the relation (3.14). Note that since the two latter corrections are of 1.5PN relative order, they cannot come from anywhere but the tail integral of the Newtonian quadrupole moment.

During the practical computation, we make explicit the time dependence of the derivatives of multipole moments, computed in the center-of-mass frame as functions of

the relative position, the relative velocity, and both spins. For circular orbits,  $\mathbf{x}$  and  $\mathbf{v}$  depend only on  $r$  and  $\omega$ , which are approximately constant on dynamical time scales, and on the unit vectors  $\mathbf{n}$  and  $\boldsymbol{\lambda}$ . Thus, the whole time dependence arises through that of  $\mathbf{n}$  and  $\boldsymbol{\lambda}$  (and  $\boldsymbol{\ell} = \mathbf{n} \times \boldsymbol{\lambda}$ ), and is provided by our explicit solution (4.24), together with the precessing angles  $\alpha(t)$  and  $\iota(t)$  given by Eqs. (4.18) and (4.19), or (4.21) and (4.22) with more precision.

The complete results for the spin-dependent parts of the radiative moments, in which we use the shorthand notation for spins (3.9) and where the basis vectors  $\{\mathbf{n}, \boldsymbol{\lambda}, \boldsymbol{\ell}\}$  are evaluated at the current time  $T_R$ , are then

$$\begin{aligned} \delta U_{ij} = & 2m\nu x^4 c^2 \left[ \frac{1}{3} (71s_n + 35\delta\sigma_n) \left( \pi n^{(i}\ell^{j)} - 2\lambda^{(i}\ell^{j)} \left( \ln(4\omega\tau_0) + \gamma_E - \frac{11}{12} \right) \right) - \frac{1}{3} (29s_\lambda + 17\delta\sigma_\lambda) \right. \\ & \times \left( \pi\lambda^{(i}\ell^{j)} + 2n^{(i}\ell^{j)} \left( \ln(4\omega\tau_0) + \gamma_E - \frac{11}{12} \right) \right) + 4 \left( s_\ell + \frac{\delta}{3}\sigma_\ell \right) \left( -4n^{(i}\lambda^{j)} \left( \ln(4\omega\tau_0) + \gamma_E - \frac{11}{12} \right) + \pi(n^{ij} - \lambda^{ij}) \right) \\ & \left. - 2 \left( n^{(i}\ell^{j)} \left( \frac{19}{3}s_\lambda + 3\delta\sigma_\lambda \right) + \lambda^{(i}\ell^{j)} \left( \frac{19}{3}s_n + 3\delta\sigma_n \right) \right) - \frac{8}{3}s_\ell n^{(i}\lambda^{j)} \right], \end{aligned} \quad (5.5a)$$

$$\delta V_{ij} = -3m\nu x^{7/2} c^3 \left[ \lambda^{(i}\sigma^{j)} \left( \ln(2\omega\tau_0) + \gamma_E - \frac{7}{6} \right) - \frac{\pi}{2} n^{(i}\sigma^{j)} \right], \quad (5.5b)$$

$$\delta V_{ijk} = -16m\nu x^4 c^4 \left[ s^{(k}(n^{ij)} - \lambda^{ij}) + 2(s^{(k} + \delta\sigma^{(k)} \left( (n^{ij)} - \lambda^{ij} \right) \left( \ln(4\omega\tau_0) + \gamma_E - \frac{5}{3} \right) + \pi n^i \lambda^j) \right]. \quad (5.5c)$$

Insertion of the above quantities into Eq. (5.2) yields the nonlinearly induced SO contributions at 2.5PN and 3PN orders in the waveform as

$$\begin{aligned} \delta \hat{h}_{ij}^{\text{TT}} = & \left\{ x^{5/2} \left[ \lambda^k \left( \ln(2\tau_0\omega) + \gamma_E - \frac{7}{6} \right) - \frac{\pi}{2} n^k \right] \left[ (N \times \boldsymbol{\sigma})^i \delta_{jk} + N^a \boldsymbol{\varepsilon}_{aki} \sigma^j \right] + x^3 \left[ \frac{4}{3} \left( s^k(n^{cd} - \lambda^{cd}) + 2(s^k + \delta\sigma^k) \right) \right. \right. \\ & \times \left( (n^{cd} - \lambda^{cd}) \left( \ln(4\tau_0\omega) + \gamma_E - \frac{5}{3} \right) + \pi n^{(c}\lambda^{d)} \right) \left( N^{ak} \boldsymbol{\varepsilon}_{aci} \delta_{jd} + N^{ca} \boldsymbol{\varepsilon}_{aki} \delta_{jd} + N^{ad} \boldsymbol{\varepsilon}_{aci} \delta_{jk} - \frac{2}{5} N^{aj} \boldsymbol{\varepsilon}_{iac} \delta_{kd} \right) \\ & + 4 \left( s_\ell + \frac{\delta\sigma_\ell}{3} \right) \left( -4n^i \lambda^j \left( \ln(4\tau_0\omega) + \gamma_E - \frac{11}{12} \right) + \pi(n^{ij} - \lambda^{ij}) \right) - \frac{8}{3} s_\ell n^i \lambda^j - 2 \left( n^i \ell^j \left( \frac{19}{3}s_\lambda + 3\delta\sigma_\lambda \right) \right. \\ & \left. + \lambda^i \ell^j \left( \frac{19}{3}s_n + 3\delta\sigma_n \right) \right) + \frac{1}{3} (71s_n + 35\delta\sigma_n) \left( \pi n^i \ell^j - 2\lambda^i \ell^j \left( \ln(4\omega\tau_0) + \gamma_E - \frac{11}{12} \right) \right) \\ & \left. - \frac{1}{3} (29s_\lambda + 17\delta\sigma_\lambda) \left( \pi\lambda^i \ell^j + 2n^i \ell^j \left( \ln(4\omega\tau_0) + \gamma_E - \frac{11}{12} \right) \right) \right] \right\}^{\text{TT}}, \end{aligned} \quad (5.6)$$

for which we have conveniently introduced the rescaled waveform  $\delta \hat{h}_{ij}^{\text{TT}}$  defined by

$$\delta h_{ij}^{\text{TT}} = \frac{4Gm\nu x}{Rc^2} \delta \hat{h}_{ij}^{\text{TT}}. \quad (5.7)$$

The two gravitational-wave polarizations  $h_+$  and  $h_\times$  are given in Appendix C. We have checked that the test-particle limit  $\nu \rightarrow 0$  of the  $-2$  spin-weighted spherical modes resulting from the above waveform agrees with the results of Ref. [21] (given explicitly in Ref. [48]) based on black-hole perturbation theory.

## VI. ENERGY FLUX AND ORBITAL PHASING

The case of the gravitational energy flux is simpler than for the waveform, notably because we need only the

contributions from the mass and current quadrupole moments, i.e.

$$\delta \mathcal{F} = \frac{G}{c^5} \left[ \frac{2}{5} U_{ij}^{(1)} \delta U_{ij}^{(1)} + \frac{32}{45c^2} V_{ij}^{(1)} \delta V_{ij}^{(1)} \right]. \quad (6.1)$$

The 3PN SO effects in the energy flux have been computed in two different ways. In the first way, we compute the time derivative of the radiative moments  $U_{ij}$  and  $V_{ij}$  whose SO-tail contributions are given in Eqs. (5.5), and then square these radiative moments to get the flux (6.1). The second way is completely equivalent, but entirely done by hand. It consists of writing all the separate pieces composing the energy flux (6.1), made of the coupling between some instantaneous moment (evaluated at the current instant  $T_R$ ) times a hereditary tail integral. The SO terms

have to be included in either the instantaneous moment in front of the integral, or in the tail integral itself. This then gives several “direct” SO contributions coming from tails at relative 1.5PN order (for the mass quadrupole tail) or 0.5PN order (for the current quadrupole tail) which are then added together. In addition, there is the crucial contribution due to the reduction to circular orbits of the standard (nonspin) tail integral at 1.5PN order, for which the relation between the orbital separation  $r$  and the orbital frequency  $\omega$  [as given by the inverse of Eq. (3.14)] provides a supplementary SO term at relative 1.5PN order, which thus contributes *in fine* at the same 3PN level as the direct SO-tail terms.

Finally, we obtain the following net result for the SO-tail contribution at 3PN order in the total energy flux:

$$\delta \mathcal{F} = \frac{32}{5} \frac{c^5}{G} x^8 \nu^2 \left[ -16\pi s_\ell - \frac{31\pi}{6} \delta\sigma_\ell \right], \quad (6.2)$$

$$\begin{aligned} \mathcal{F} = & \frac{32}{5} \frac{c^5}{G} x^5 \nu^2 \left\{ 1 + x \left( -\frac{1247}{336} - \frac{35}{12} \nu \right) + x^{3/2} \left( 4\pi - 4s_\ell - \frac{5}{4} \delta\sigma_\ell \right) + x^2 \left( -\frac{44711}{9072} + \frac{9271}{504} \nu + \frac{65}{18} \nu^2 \right) \right. \\ & + x^{5/2} \left( -\frac{8191}{672} \pi - \frac{9}{2} s_\ell - \frac{13}{16} \delta\sigma_\ell + \nu \left[ -\frac{583}{24} \pi + \frac{272}{9} s_\ell + \frac{43}{4} \delta\sigma_\ell \right] \right) + x^3 \left( \frac{6643739519}{69854400} + \frac{16}{3} \pi^2 \right. \\ & \left. \left. - \frac{1712}{105} \gamma_E - \frac{856}{105} \ln(16x) - 16\pi s_\ell - \frac{31\pi}{6} \delta\sigma_\ell + \nu \left[ -\frac{134543}{7776} + \frac{41}{48} \pi^2 \right] - \frac{94403}{3024} \nu^2 - \frac{775}{324} \nu^3 \right) \right\}. \quad (6.3) \end{aligned}$$

We are consistently using the constant-magnitude spins  $S_A$  that are related to the original variables  $\bar{S}_A$  of Refs. [19,20] by Eq. (3.7); see also Eqs. (7.4) of Ref. [20]. The nonspin terms are given, e.g., in Ref. [1]. We find perfect agreement in the perturbative limit  $\nu \rightarrow 0$  with black-hole perturbation calculations reported in Ref. [21]. On the other hand, the total conservative energy  $E$  of the binary is not affected by the SO terms at the 3PN order (we check this point in Appendix A); hence we have

$$\begin{aligned} E = & -\frac{1}{2} m \nu c^2 x \left\{ 1 + x \left( -\frac{3}{4} - \frac{\nu}{12} \right) + x^{3/2} \left( \frac{14}{3} s_\ell + 2\delta\sigma_\ell \right) \right. \\ & + x^2 \left( -\frac{27}{8} + \frac{19}{8} \nu - \frac{\nu^2}{24} \right) + x^{5/2} \left( 11s_\ell + 3\delta\sigma_\ell \right. \\ & \left. + \nu \left[ -\frac{61}{9} s_\ell - \frac{10}{3} \delta\sigma_\ell \right] \right) + x^3 \left( -\frac{675}{64} + \left[ \frac{34445}{576} \right. \right. \\ & \left. \left. - \frac{205}{96} \pi^2 \right] \nu - \frac{155}{96} \nu^2 - \frac{35}{5184} \nu^3 \right) \right\}. \quad (6.4) \end{aligned}$$

Following Ref. [20] we shall next use the standard energy balance argument to deduce the evolution of the orbital frequency even in the presence of spins. To this end we have to check that the constant-magnitude spins are secularly constant (i.e., constant over a long radiation-reaction time scale) up to the right level, 1.5PN order in the present case. In Ref. [20] we have referred to the work [45] for proof that this is correct up to relative 1PN order, i.e.

where we recall that  $s_\ell = s \cdot \ell$  and  $\sigma_\ell = \sigma \cdot \ell$ , with the spin variables  $s$  and  $\sigma$  being defined by Eqs. (3.7), (3.8), and (3.9). Let us remark that in the energy flux the 3PN SO term is entirely constituted by the SO tails we have obtained in (6.2). So the complete 3PN SO term in the flux is provided by Eq. (6.2). Contrary to the waveform computed in Sec. V, there are no other SO terms coming from linear source moments at that order.

Because the energy flux and the resulting orbital phasing are so important for gravitational-wave observations, we shall now give the complete formula for the total flux, including all nonspin terms and all linear SO terms up to 3PN order (but neglecting nonlinear SS interactions). However, we shall not write the known nonspin 3.5PN terms in the flux (due to nonspin tails [17]) because some yet uncalculated SO effects should conjointly appear at that order. The 3PN energy flux, complete except for SS interactions, then reads

considering radiation-reaction effects up to 3.5PN order. In Appendix B below we extend the argument of Ref. [45] to the relative 1.5PN order, which essentially means adding the tail-induced part of the radiation reaction at 4PN order. This check being done, we can thus neglect  $\langle ds_\ell/dt \rangle$  and  $\langle d\sigma_\ell/dt \rangle$  in average over a radiation-reaction time scale.

An alternative way to see this is to directly compute the variation of the projection of the spins along the Newtonian orbital angular momentum, i.e.  $S_\ell^A = S_A \cdot \ell$ , using the precession equations (3.16) appropriate for constant-magnitude spins. We readily find that  $dS_\ell^A/dt = S_A \cdot [d\ell/dt + \ell \times \Omega_A]$ , which shows that  $dS_\ell^A/dt$  is at least quadratic in the spins for circular orbits. This readily follows from the facts that  $\ell$  remains constant in the absence of spins, and that, as we have seen in Eq. (3.17),  $\Omega_A$  for circular orbits points in the direction of  $\ell$  modulo spin corrections. Thus we have  $dS_\ell^A/dt = 0$  at the linear SO level (neglecting quadratic SS couplings). The argument is, in principle, valid up to any PN order, but is restricted to circular orbits.

The conclusion is that the constant-magnitude spin terms can be considered as constant when computing the averaged evolution  $\langle dE/dt \rangle$  of the energy given by Eq. (6.4). Equating then  $\langle dE/dt \rangle$  to  $-\mathcal{F}$ , where  $\mathcal{F}$  is given by (6.3), we obtain the secular variation of the frequency  $\langle \dot{\omega} \rangle$ —denoted  $\dot{\omega}$  for simplicity—as (neglecting SS contributions)

$$\begin{aligned}
\frac{\dot{\omega}}{\omega^2} = & \frac{96}{5} \nu x^{5/2} \left\{ 1 + x \left( -\frac{743}{336} - \frac{11}{4} \nu \right) + x^{3/2} \left( 4\pi - \frac{47}{3} s_\ell - \frac{25}{4} \delta\sigma_\ell \right) + x^2 \left( \frac{34\,103}{18\,144} + \frac{13\,661}{2016} \nu + \frac{59}{18} \nu^2 \right) \right. \\
& + x^{5/2} \left( -\frac{4159}{672} \pi - \frac{5861}{144} s_\ell - \frac{809}{84} \delta\sigma_\ell + \nu \left[ -\frac{189}{8} \pi + \frac{1001}{12} s_\ell + \frac{281}{8} \delta\sigma_\ell \right] \right) + x^3 \left( \frac{16\,447\,322\,263}{139\,708\,800} + \frac{16}{3} \pi^2 \right. \\
& \left. \left. - \frac{1712}{105} \gamma_E - \frac{856}{105} \ln(16x) - \frac{188\pi}{3} s_\ell - \frac{151\pi}{6} \delta\sigma_\ell + \nu \left[ -\frac{56\,198\,689}{217\,728} + \frac{451}{48} \pi^2 \right] + \frac{541}{896} \nu^2 - \frac{5605}{2592} \nu^3 \right) \right\}. \quad (6.5)
\end{aligned}$$

By integrating this out using standard PN rules for multiplying, dividing, and integrating PN expressions, we obtain the secular evolution of the carrier phase [defined by  $\phi = \int \omega dt$ ; see Eq. (4.12)] as

$$\begin{aligned}
\phi = & \phi_0 - \frac{1}{32\nu} \left[ x^{-5/2} + x^{-3/2} \left( \frac{3715}{1008} + \frac{55}{12} \nu \right) + x^{-1} \left( -10\pi + \frac{235}{6} s_\ell + \frac{125}{8} \delta\sigma_\ell \right) + x^{-1/2} \left( \frac{15\,293\,365}{1\,016\,064} + \frac{27\,145}{1008} \nu \right. \right. \\
& \left. \left. + \frac{3085}{144} \nu^2 \right) + \ln x \left( \frac{38\,645}{1344} \pi - \frac{554\,345}{2016} s_\ell - \frac{41\,745}{448} \delta\sigma_\ell + \nu \left[ -\frac{65}{16} \pi - \frac{55}{8} s_\ell + \frac{15}{8} \delta\sigma_\ell \right] \right) \right. \\
& \left. + x^{1/2} \left( \frac{12\,348\,611\,926\,451}{18\,776\,862\,720} - \frac{160}{3} \pi^2 - \frac{1712}{21} \gamma_E - \frac{856}{21} \ln(16x) + \frac{940\pi}{3} s_\ell + \frac{745\pi}{6} \delta\sigma_\ell \right. \right. \\
& \left. \left. + \nu \left[ -\frac{15\,737\,765\,635}{12\,192\,768} + \frac{2255}{48} \pi^2 \right] + \frac{76\,055}{6912} \nu^2 - \frac{127\,825}{5184} \nu^3 \right) \right]. \quad (6.6)
\end{aligned}$$

We recall that to the carrier phase we also have to add the precessional correction, arising from the changing orientation of the orbital plane. We have proved in Eq. (4.11) that at the 1PN order the total phase  $\Phi$  is given by  $\Phi = \phi - \alpha + \mathcal{O}(c^{-4})$ . Thus the precessional correction is given by  $-\alpha$  and is explicitly provided by the solution (4.18) and (4.19). Alternatively, the precessional correction can be computed numerically [9].

## VII. CONCLUSION

So far, the search for gravitational waves with LIGO and Virgo detectors has focused on nonspinning compact binaries [49–53], although in Ref. [54] single-spin templates were employed, for the first time, to search for inspiraling spinning compact objects. It is timely and necessary to develop more accurate templates which include spin effects. Extrapolating results from the nonspinning case, we expect that, for maximally spinning objects, reasonably accurate templates would need to be computed at least through 3.5PN order.

During the last years, motivated by the search for gravitational waves, SO effects have been computed in the two-body equations of motion through 3.5PN order [19,41,43,44,55–57] and energy flux through 2.5PN order [20,42–45]. Moreover, SS effects have been calculated through 3PN order in the conservative dynamics [41,42,58–68] and multipole moments [69].

In this paper, building on our previous work [19,20], we have improved the accuracy of the energy flux and gravitational waveform by computing SO terms induced by tail effects [10–18]. Those effects are due to the backscattering of linear waves in the curved space-time geometry around the source. Using the multipolar PN formalism developed in Refs. [13–15,23,24], we have identified and computed

the radiative multipole moments responsible of tail terms involving SO couplings. More specifically, we have computed those SO-tail contributions to the energy flux at 3PN order and to the gravitational waveforms at 2.5PN and 3PN order. Those SO tails constitute the complete coefficient at 3PN order in the energy flux. In particular, we find that the energy flux is in complete agreement with the result of black-hole perturbations in the test-particle limit [21]. Our computation is restricted to quasicircular inspiraling orbits, and uses the two-body precessional dynamics at 1.5PN order.

The computation of SO-tail terms in the waveform was summarized in Sec. V, and some building blocks and foundation for calculating tail effects in the PN formalism were reviewed in Sec. II. For the first time, we have computed tail terms when precession effects in the two-body dynamics are also included. The relevant results for the waveform are given in Eq. (5.6) and in Appendix C. The SO-tail effects in the energy flux and phasing at 3PN order are given in Sec. VI; see, in particular, Eqs. (6.2), (6.3), and (6.6).

Considering the vigorous synergy which is currently taking place between analytical and numerical relativity for building faithful templates [5–8], we expect that the results developed in this paper will help the construction of more accurate analytical templates describing the entire process of inspiral, merger, and ringdown of black holes in the presence of spins.

In the near future we plan to complete the knowledge of SO effects in the gravitational waveform at 3PN order, by computing the nontail (i.e., instantaneous) SO couplings at 2PN and 3PN orders, and the corresponding  $-2$  spin-weighted spherical harmonics (or gravitational modes). This will constitute a step further with respect to

Ref. [46] which computed SO effects in the gravitational modes through 1.5PN order.

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### APPENDIX A: 3PN SPIN TERMS IN THE EQUATIONS OF MOTION

In this appendix, we check that there are no SO terms at the 3PN order in the total conservative invariant energy of the binary given by Eq. (6.4). Indeed, we find that the 3PN SO terms in the binary's equations of motion (say, in harmonic coordinates) can be gauged away. The result is to be expected because we know that the first SO modification of the radiation-reaction damping force arises at the 4PN order rather than 3PN [45].

We compute the near-zone PN metric by solving the Einstein field equations in harmonic coordinates for the stress-energy tensor (3.1). We find by direct PN iteration of the metric, parametrized by means of retarded potentials  $V, V_i, \dots$  (see Ref. [19] for more details), that the contribution of SO terms at 3PN order in this gauge is given by

$$\delta g_{00} = \frac{2Gm_1}{3c^5} (r_1 \dot{a}_1) + \frac{4}{3c^8} \varepsilon_{ijk} S_1^j \ddot{a}_1^i r_1^k + \frac{1}{c^8} \text{cst}(t) + 1 \leftrightarrow 2, \quad (\text{A1a})$$

$$\delta g_{0i} = \frac{4Gm_1}{c^4} a_1^i - \frac{10G}{3c^7} \varepsilon_{ijk} S_1^j \dot{a}_1^k + 1 \leftrightarrow 2, \quad (\text{A1b})$$

$$\delta g_{ij} = 0, \quad (\text{A1c})$$

where we indicate with  $\text{cst}(t)$  some irrelevant  $\mathcal{O}(c^{-8})$  constant term in space, where we keep the SO parts of the acceleration unreplaced, and where  $1 \leftrightarrow 2$  refers to the same expression but for particle 2.

The metric (A1) yields a 3PN contribution in the equations of motion of spinning particles which can be calculated from the Papapetrou [28,29] equations of motion (see e.g. Sec. III in [19]). The result for the acceleration of particle 1 is

$$\delta a_1^i = \frac{G^2}{c^6 r^4} \varepsilon_{ijk} (m_1 S_2^k - m_2 S_1^k) \left[ (15(nv)^2 - 3v^2) n^j + 2 \frac{Gm}{r} n^j - 6(nv)v^j \right]. \quad (\text{A2})$$

We observe that  $\delta a_1^i$  is symmetric under the exchange of particles 1 and 2. A closer inspection reveals that it is in fact given by the second total time derivative of a certain vector, namely,

$$\delta a_1^i = \delta a_2^i = \frac{d^2 \delta X^i}{dt^2}, \quad (\text{A3})$$

and we find that  $\delta X = \nu \gamma^3 (\mathbf{x} \times \boldsymbol{\sigma})$  in the notation of Sec. III C. This is precisely the effect of the gauge transformation associated with a shift of coordinates  $x^{i'} = x^i + \delta X^i$ . We thus conclude that the 3PN SO terms in the equations of motion are pure gauge and cannot affect the binary's invariant energy (6.4).

### APPENDIX B: 4PN SPIN SECULAR EVOLUTION

Here we show that the constant-magnitude spins are secularly constant, i.e. constant over a long radiation-reaction time scale, up to the 4PN order corresponding to the 1.5PN relative order. In Ref. [45] this has already been proved up to 1PN relative order; here we extend the argument to 1.5PN order. [In the main text after Eq. (6.4) we present an alternative argument valid at any PN order but restricted to circular orbits.]

Following Ref. [45] we describe our source by a set of well-separated extended bodies  $A$ , supposed to be Newtonian in a first stage. We define the spin  $S_A^i$  of each of the bodies in the usual Newtonian way as an integral extending over the volume of body  $A$ ,

$$S_A^i = \varepsilon_{ijk} \int_A d^3 \mathbf{x} \rho_*(x^j - x_A^j) v^k, \quad (\text{B1})$$

where  $\rho_*$  denotes the Newtonian (baryonic) mass density, and  $x_A^j$  is the Newtonian center-of-mass position of the body  $A$ . At Newtonian level the spin (B1) agrees with the definition employed in the present paper. The ‘‘baryonic’’ spin (B1) is only used for the purpose of this appendix. The equation of evolution of the baryonic spin reads as

$$\frac{dS_A^i}{dt} = \varepsilon_{ijk} \int_A d^3 \mathbf{x} \rho_*(x^j - x_A^j) a^k. \quad (\text{B2})$$

The spin precession equation follows from inserting into (B2) an explicit solution for the acceleration in terms of positions and velocities. The resulting equation is then simplified using some virial relations appropriate to the case where the compact body is ‘‘stationary’’; see Ref. [45]. The secular evolution of the spin is then obtained by considering the radiation-reaction piece of the acceleration in Eq. (B2).

At the dominant 2.5PN level the radiation-reaction acceleration inside an isolated body in harmonic coordinates is given by (see, e.g., Ref. [70])

$$a_{2.5\text{PN}}^i = \frac{G}{c^5} \left[ \frac{3}{5} x^j I_{ij}^{(5)} + 2 \frac{d}{dt} (v^j I_{ij}^{(3)}) + I_{jk}^{(3)} \partial_i U_{jk} \right], \quad (\text{B3})$$

where  $I_{ij}$  is the source's STF quadrupole moment (at Newtonian order), and  $U_{ij}$  is the Newtonian potential tensor defined by

$$U_{ij}(\mathbf{x}, t) = G \int d^3 \mathbf{x}' \rho_*(\mathbf{x}', t) \frac{(x^i - x'^i)(x^j - x'^j)}{|\mathbf{x} - \mathbf{x}'|^3}. \quad (\text{B4})$$

(We have  $U_{ii} = U$ , the usual Newtonian scalar potential.) It can be shown [45] that the only contribution at 2.5PN order to the spin precession equation comes from the velocity-dependent part of Eq. (B3), i.e.

$$a_{2.5\text{PN}}^i = \frac{2G}{c^5} v^j I_{ij}^{(4)} + \dots \quad (\text{B5})$$

The other pieces in the 2.5PN acceleration vanish when the size of the body tends to zero (compact-body limit) and may be ignored. Using a virial relation [45] we readily obtain

$$\left( \frac{dS_A^i}{dt} \right)_{2.5\text{PN}} = - \frac{G}{c^5} I_{ij}^{(4)}(t) S_A^j. \quad (\text{B6})$$

Because the spin is constant in the lowest approximation, the latter result is a total time derivative:

$$\left( \frac{dS_A^i}{dt} \right)_{2.5\text{PN}} = \frac{d}{dt} \left[ - \frac{G}{c^5} I_{ij}^{(3)}(t) S_A^j \right], \quad (\text{B7})$$

which can be moved to the left-hand side and absorbed into a negligible redefinition of the spin variable at 2.5PN order. When specialized to two compact bodies the result (B7) becomes

$$\left( \frac{dS_1^i}{dt} \right)_{2.5\text{PN}} = \frac{d}{dt} \left\{ \frac{G^2 m_1 m_2}{c^5 r^2} \left[ -6(nv)(nS_1)n^i + 4(nS_1)v^i + 4(vS_1)n^i - \frac{2}{3}(nv)S_1^i \right] \right\}. \quad (\text{B8})$$

[Note that the latter expression depends on the specific definition of the spin one is using, i.e. in the present case the Newtonian spin defined for extended bodies by (B1); for the spin variable used in [19,20] the expression is expected to be different, but still in the form of a total time derivative.]

The results (B7) or (B8) show that there is no secular evolution for the spin at the dominant 2.5PN order (see also [71] and references therein for related discussions). Note that this conclusion actually applies to any spin variable at dominant order. However, it has been shown in Ref. [45] that in the case of the constant-magnitude spin there is

also no secular evolution of the spins at the next-to-leading 3.5PN order in the radiation reaction. At next-to-leading order this result [45] applies specifically to the constant-magnitude spins and uses the radiation-reaction acceleration up to the 3.5PN order.

We now extend the argument by including the higher-order 4PN correction term (i.e. 1.5PN radiation-reaction order) which is known to be due to gravitational-wave tails [14,72]. That extension is actually straightforward since it basically needs only the Newtonian result (B7). The reason is that including the effect of tails in the radiation reaction simply amounts to replacing the source Newtonian quadrupole moment  $I_{ij}$  by the tail-corrected expression [72]

$$I_{ij}^{\text{tail}}(t) = I_{ij}(t) + \frac{4GM}{c^3} \int_{-\infty}^t dt' I_{ij}^{(2)}(t') \ln\left(\frac{t-t'}{2\tau_0}\right), \quad (\text{B9})$$

where  $\tau_0$  denotes some arbitrary time scale, for instance, the one which appears in Eqs. (2.4). Note the factor  $4GM/c^3$  in front of the tail integral which is twice the factor  $2GM/c^3$  in front of the tail integrals in (2.4). This factor ensures the consistency between the work done by the radiation-reaction force in the local source and the total energy flux radiated at infinity from the source [15].

Thus the radiation-reaction force including the 4PN tails takes (in harmonic coordinates) the same form as in Eq. (B3) but with  $I_{ij}$  replaced by  $I_{ij}^{\text{tail}}$ . This shows that the previous Newtonian argument still holds for the 2.5PN + 4PN radiation-reaction force and that the effect on the precession equation is still in the form of some irrelevant total time derivative:

$$\left( \frac{dS_A^i}{dt} \right)_{2.5\text{PN}+4\text{PN}} = \frac{d}{dt} \left[ - \frac{G}{c^5} I_{ij}^{(3)}(t) S_A^j \right]. \quad (\text{B10})$$

Hence our conclusion is that the constant-magnitude spins are secularly constant up to 4PN order corresponding to 1.5PN radiation-reaction order.

### APPENDIX C: GRAVITATIONAL-WAVE POLARIZATIONS

We derive in this appendix the two gravitational-wave polarizations. They are computed from the projection formulas (2.2), using the expression (5.6) [together with Eq. (5.7)] for  $\delta h_{ij}^{\text{TT}}$ . We adopt the convention shown in Fig. 2 for the polarization vectors. To shorten the result, we denote the projections of the polarization basis  $\{\mathbf{N}, \mathbf{P}, \mathbf{Q}\}$  onto the moving triad  $\{\mathbf{n}, \boldsymbol{\lambda}, \boldsymbol{\ell}\}$  by e.g.  $P_n, P_\lambda, P_\ell$ . With this notation, we have

$$\begin{aligned}
\delta h_+ = & \frac{Gm\nu}{c^2 R} \left\{ x^{7/2} \left[ \pi(2(P_n Q_\ell + P_\ell Q_n) \sigma_\ell - 2(P_\ell Q_\ell - P_n Q_n + P_\lambda Q_\lambda) \sigma_n + 2(P_\lambda Q_n + P_n Q_\lambda) \sigma_\lambda) \right. \right. \\
& + 4(-P_\lambda Q_\ell + P_\ell Q_\lambda) \sigma_\ell - (P_\lambda Q_n + P_n Q_\lambda) \sigma_n + (P_\ell Q_\ell + P_n Q_n - P_\lambda Q_\lambda) \sigma_\lambda \left. \left( \ln(2\tau_0 \omega) - \frac{7}{6} + \gamma_E \right) \right] \\
& + x^4 \left[ -\frac{16}{3} (P_n P_\lambda + N_n P_n Q_\ell - N_\lambda P_\lambda Q_\ell + N_n P_\ell Q_n + N_\ell P_n Q_n - N_\lambda P_\ell Q_\lambda - N_\ell P_\lambda Q_\lambda - Q_n Q_\lambda) s_\ell \right. \\
& - \frac{4}{3} (19P_\ell P_\lambda + 12N_n P_n Q_n - 4N_\lambda P_\lambda Q_n - 4N_\lambda P_n Q_\lambda - 4N_n P_\lambda Q_\lambda - 19Q_\ell Q_\lambda) s_n \\
& - \frac{4}{3} (19P_\ell P_n + 4N_\lambda P_n Q_n + 4N_n P_\lambda Q_n - 19Q_\ell Q_n + 4N_n P_n Q_\lambda - 12N_\lambda P_\lambda Q_\lambda) s_\lambda + \delta(-12(P_\ell P_\lambda - Q_\ell Q_\lambda) \sigma_n \\
& - 12(P_\ell P_n - Q_\ell Q_n) \sigma_\lambda) + \pi \left( 8(P_n^2 - P_\lambda^2 - Q_n^2 + Q_\lambda^2) s_\ell + \frac{142}{3} (P_\ell P_n - Q_\ell Q_n) s_n - \frac{58}{3} (P_\ell P_\lambda - Q_\ell Q_\lambda) s_\lambda \right. \\
& - \frac{16}{3} (N_\lambda P_n Q_\ell + N_n P_\lambda Q_\ell + N_\lambda P_\ell Q_n + N_\ell P_\lambda Q_n + N_n P_\ell Q_\lambda + N_\ell P_n Q_\lambda) (s_\ell + \delta \sigma_\ell) - \frac{32}{3} (N_\lambda P_n Q_n + N_n P_\lambda Q_n \\
& + N_n P_n Q_\lambda) (s_n + \delta \sigma_n) + \delta \left( \frac{8}{3} (P_n^2 - P_\lambda^2 - Q_n^2 + Q_\lambda^2) \sigma_\ell + \frac{70}{3} (P_\ell P_n - Q_\ell Q_n) \sigma_n - \frac{34}{3} (P_\ell P_\lambda - Q_\ell Q_\lambda) \sigma_\lambda \right) \\
& - \frac{32}{3} (N_\lambda P_\lambda Q_n + N_\lambda P_n Q_\lambda + N_n P_\lambda Q_\lambda) (s_\lambda + \delta \sigma_\lambda) \left. \right) + \left( -\frac{32}{3} (N_n P_n Q_\ell - N_\lambda P_\lambda Q_\ell + N_n P_\ell Q_n \right. \\
& + N_\ell P_n Q_n - N_\lambda P_\ell Q_\lambda - N_\ell P_\lambda Q_\lambda) (s_\ell + \delta \sigma_\ell) - \frac{32}{3} (3N_n P_n Q_n - N_\lambda P_\lambda Q_n - N_\lambda P_n Q_\lambda - N_n P_\lambda Q_\lambda) (s_n + \delta \sigma_n) \\
& - \frac{32}{3} (N_\lambda P_n Q_n + N_n P_\lambda Q_n + N_n P_n Q_\lambda - 3N_\lambda P_\lambda Q_\lambda) (s_\lambda + \delta \sigma_\lambda) \left. \right) \left( \ln(4\tau_0 \omega) - \frac{5}{3} + \gamma_E \right) \\
& + \left( -32(P_n P_\lambda - Q_n Q_\lambda) s_\ell - \frac{284}{3} (P_\ell P_\lambda - Q_\ell Q_\lambda) s_n - \frac{116}{3} (P_\ell P_n - Q_\ell Q_n) s_\lambda \right. \\
& \left. + \delta \left( -\frac{32}{3} (P_n P_\lambda - Q_n Q_\lambda) \sigma_\ell - \frac{140}{3} (P_\ell P_\lambda - Q_\ell Q_\lambda) \sigma_n - \frac{68}{3} (P_\ell P_n - Q_\ell Q_n) \sigma_\lambda \right) \right) \left( \ln(4\tau_0 \omega) - \frac{11}{12} + \gamma_E \right) \left. \right\}, \tag{C1a}
\end{aligned}$$

$$\begin{aligned}
\delta h_\times = & \frac{Gm\nu}{c^2 R} \left\{ x^{7/2} \left[ \pi(-2(P_\ell P_n - Q_\ell Q_n) \sigma_\ell + (P_\ell^2 - P_n^2 + P_\lambda^2 - Q_\ell^2 + Q_n^2 - Q_\lambda^2) \sigma_n - 2(P_n P_\lambda - Q_n Q_\lambda) \sigma_\lambda) \right. \right. \\
& + (4(P_\ell P_\lambda - Q_\ell Q_\lambda) \sigma_\ell + 4(P_n P_\lambda - Q_n Q_\lambda) \sigma_n - 2(P_\ell^2 + P_n^2 - P_\lambda^2 - Q_\ell^2 - Q_n^2 + Q_\lambda^2) \sigma_\lambda) \left. \left( \ln(2\tau_0 \omega) - \frac{7}{6} + \gamma_E \right) \right] \\
& + x^4 \left[ \frac{8}{3} (2N_n P_\ell P_n + N_\ell P_n^2 - 2N_\lambda P_\ell P_\lambda - N_\ell P_\lambda^2 - 2P_\lambda Q_n - 2N_n Q_\ell Q_n - N_\ell Q_n^2 - 2P_n Q_\lambda \right. \\
& + 2N_\lambda Q_\ell Q_\lambda + N_\ell Q_\lambda^2) s_\ell + \frac{4}{3} (6N_n P_n^2 - 4N_\lambda P_n P_\lambda - 2N_n P_\lambda^2 - 19P_\lambda Q_\ell - 6N_n Q_n^2 - 19P_\ell Q_\lambda + 4N_\lambda Q_n Q_\lambda \\
& + 2N_n Q_\lambda^2) s_n + \frac{4}{3} (2N_\lambda P_n^2 + 4N_n P_n P_\lambda - 6N_\lambda P_\lambda^2 - 19P_n Q_\ell - 19P_\ell Q_n - 2N_\lambda Q_n^2 - 4N_n Q_n Q_\lambda + 6N_\lambda Q_\lambda^2) s_\lambda \\
& + \delta(-12(P_\lambda Q_\ell + P_\ell Q_\lambda) \sigma_n - 12(P_n Q_\ell + P_\ell Q_n) \sigma_\lambda) + \pi \left( 16(P_n Q_n - P_\lambda Q_\lambda) s_\ell + \frac{142}{3} (P_n Q_\ell + P_\ell Q_n) s_n \right. \\
& - \frac{58}{3} (P_\lambda Q_\ell + P_\ell Q_\lambda) s_\lambda \\
& + \frac{16}{3} (N_\lambda P_\ell P_n + N_n P_\ell P_\lambda + N_\ell P_n P_\lambda - N_\lambda Q_\ell Q_n - N_n Q_\ell Q_\lambda - N_\ell Q_n Q_\lambda) (s_\ell + \delta \sigma_\ell) + \frac{16}{3} (N_\lambda P_n^2 + 2N_n P_n P_\lambda \\
& - N_\lambda Q_n^2 - 2N_n Q_n Q_\lambda) (s_n + \delta \sigma_n) + \delta \left( \frac{16}{3} (P_n Q_n - P_\lambda Q_\lambda) \sigma_\ell + \frac{70}{3} (P_n Q_\ell + P_\ell Q_n) \sigma_n - \frac{34}{3} (P_\lambda Q_\ell + P_\ell Q_\lambda) \sigma_\lambda \right) \\
& \left. + \frac{16}{3} (2N_\lambda P_n P_\lambda + N_n P_\lambda^2 - 2N_\lambda Q_n Q_\lambda - N_n Q_\lambda^2) (s_\lambda + \delta \sigma_\lambda) \right) \left. \right\}
\end{aligned}$$



$$\begin{aligned}
& + \left( \frac{16}{3} (2N_n P_\ell P_n + N_\ell P_n^2 - 2N_\lambda P_\ell P_\lambda - N_\ell P_\lambda^2 - 2N_n Q_\ell Q_n - N_\ell Q_n^2 + 2N_\lambda Q_\ell Q_\lambda + N_\ell Q_\lambda^2) (s_\ell + \delta\sigma_\ell) \right. \\
& + \frac{16}{3} (3N_n P_n^2 - 2N_\lambda P_n P_\lambda - N_n P_\lambda^2 - 3N_n Q_n^2 + 2N_\lambda Q_n Q_\lambda + N_n Q_\lambda^2) (s_n + \delta\sigma_n) \\
& + \frac{16}{3} (N_\lambda P_n^2 + 2N_n P_n P_\lambda - 3N_\lambda P_\lambda^2 - N_\lambda Q_n^2 - 2N_n Q_n Q_\lambda + 3N_\lambda Q_\lambda^2) (s_\lambda + \delta\sigma_\lambda) \left. \right) \left( \ln(4\tau_0\omega) - \frac{5}{3} + \gamma_E \right) \\
& + \left( -32(P_\lambda Q_n + P_n Q_\lambda) s_\ell - \frac{284}{3} (P_\lambda Q_\ell + P_\ell Q_\lambda) s_n - \frac{116}{3} (P_n Q_\ell + P_\ell Q_n) s_\lambda \right. \\
& \left. + \delta \left( -\frac{32}{3} (P_\lambda Q_n + P_n Q_\lambda) \sigma_\ell - \frac{140}{3} (P_\lambda Q_\ell + P_\ell Q_\lambda) \sigma_n - \frac{68}{3} (P_n Q_\ell + P_\ell Q_n) \sigma_\lambda \right) \right) \left( \ln(4\tau_0\omega) - \frac{11}{12} + \gamma_E \right) \left. \right\}.
\end{aligned} \tag{C1b}$$

The authors can provide on demand a file containing the results in MATHEMATICA input format.

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