

Charge-induced force-noise on free-falling test masses: results from LISA Pathfinder

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We report on electrostatic measurements made on board the European Space Agency mission LISA Pathfinder. Detailed measurements of the charge-induced electrostatic forces exerted on free-falling test masses (TMs) inside the capacitive gravitational reference sensor are the first made in a relevant environment for a space-based gravitational wave detector. Employing a combination of charge control and electric-field compensation, we show that the level of charge-induced acceleration noise on a single TM can be maintained at a level close to $1.0 \text{ fm s}^{-2} \text{ Hz}^{-1/2}$ across the 0.1-100 mHz frequency band that is crucial to an observatory such as LISA. Using dedicated measurements that detect these effects in the differential acceleration between the two test masses, we resolve the stochastic nature of the TM charge build up due to interplanetary cosmic rays and the TM charge-to-force coupling through stray electric fields in the sensor. All our measurements are in good agreement with predictions based on a relatively simple electrostatic model of the LISA Pathfinder instrument.

Sensitive gravitational experiments employ quasi-free-falling isolated test masses as a reference system for the measurement of the local curvature of space-time. Electrostatic free charge and stray potentials introduce unwanted disturbances that can limit measurement precision. The effect is relevant for gravitational wave observatories both in space [1, 2] and on-ground [3], tests of the equivalence principle [4] and measurements of relativistic effects on precessing gyroscopes [5].

The LISA Pathfinder spacecraft [6], a technology-demonstration experiment for a space-based gravitational wave observatory, LISA [7], was launched on December 3 2015. The aim of the mission was to demonstrate the ability to fly free-falling test masses in a single spacecraft with a differential acceleration noise below $30 \text{ fm s}^{-2} \text{ Hz}^{-1/2}$. The sensitivity of the instrument has far-exceeded its design specification, achieving a level close to the LISA goal from 0.1-100 mHz and around $5 \text{ fm s}^{-2} \text{ Hz}^{-1/2}$ in the mHz band [8]. In this paper we describe the measurements and techniques used to minimise charge-related electrostatic forces and evaluate their contribution to the differential acceleration noise of the test masses.

The LISA Pathfinder test masses, identical

to those for LISA, are 46-mm cubes of mass 1.928 kg made from a gold-platinum alloy. They sit within a 6 degree-of-freedom capacitive position sensor and actuator, the gravitational reference sensor (GRS) [9, 10]. The masses are separated from the walls of the sensor by gaps of between 2.9 and 4 mm and have no grounding wire. All test mass and sensor surfaces are gold-coated. The large gaps mitigate the impact of surface forces and the absence of a grounding wire eliminates thermal noise associated with mechanical damping that dominates the low-frequency performance of accelerometers in existing geodesy and fundamental physics missions.

The achieved level of sensitivity to the differential acceleration of the test masses level is made possible by an additional high-precision readout along the x -axis provided by a laser interferometer [11–13] with a readout noise of $35 \text{ fm Hz}^{-1/2}$ measured above 60 mHz [8].

The GRS consists of a system of 12 electrodes for TM position sensing and actuation, and a further 6 for capacitive biasing of the test mass at 100 kHz. Actuation is achieved with audio-frequency sinusoidal voltages. DC or slowly-varying ($f \sim \text{mHz}$) voltage signals can be applied to measure TM charge and balance stray

electrostatic fields, as will be discussed shortly. Voltages on all electrodes originate from the GRS front-end electronics [14].

High-energy cosmic rays and solar energetic particles, mostly protons, penetrate the spacecraft and instrument shielding depositing charge on the test mass, either by stopping directly or by secondary emission [15–18]. Limiting charge accumulation on the electrically isolated TMs is needed to control electrostatic forces, the subject of this paper. In LISA Pathfinder, non-contact discharge is achieved by illuminating the sensor and test-mass surfaces with UV light and transferring charge by photoemission [19] in a similar way to that already demonstrated on GP-B [20]. A detailed account of the performance of the LISA Pathfinder UV discharge system will be provided in a subsequent article.

As well as providing desired actuation forces, the GRS is a source of unwanted electrostatic disturbances on the test mass [10, 21]. With the TM centered, the dominant source of electrostatic force noise is the interaction between the TM charge, q , and stray electric fields represented by an effective potential difference between opposite sides of the TM, Δ_x [22]. The resulting force along the x -axis, following the notation of [2], is

$$F_x(q) = -\frac{q}{C_T} \left| \frac{\partial C_x}{\partial x} \right| \Delta_x, \quad (1)$$

where $\frac{\partial C_x}{\partial x}$ is the derivative of a single sensing electrode capacitance with respect to TM displacement along x and C_T is the total capacitance of the test mass with respect to the GRS [23].

In LISA Pathfinder, the principle science observable is the differential force per unit mass acting on the two TMs, $\Delta g \equiv \frac{F_{2x}}{m_2} - \frac{F_{1x}}{m_1}$. The measurement is thus sensitive to the in-band fluctuations of both Δ_x and q for the two TM. Force noise is produced by a non-zero charge q , coupling with fluctuations in the average potential difference Δ_x and, likewise, stochastic charge fluctuations mixing with any non-zero

potential difference. We measure these effects with a number of dedicated techniques.

The test-mass charge, q can be detected in its effect on the TM potential, $\delta V_{TM} = \frac{\delta q}{C_T}$, measured by applying sinusoidally varying voltages with amplitude V_{MOD} and frequency, f_{MOD} on the x -axis electrodes, a technique well-demonstrated in ground-based investigations [24, 25]. The resulting force on the TM is $F_x(f_{MOD}) = -4 \left| \frac{\partial C_x}{\partial x} \right| V_{MOD} V_{TM}$. A continuous measurement provides an extended time series of $q(t)$ from which the low-frequency behavior of the charge build up can be studied.

To measure the relevant stray potential-difference Δ_x , we introduce a potential $\pm V_{COMP}$ to each x electrode. Following the method described in [2] it is possible to estimate $\frac{\partial F_x}{\partial q}$ as a function of V_{COMP} measuring the change in Δg as the charge of one test mass is increased in steps by photoemission under UV illumination. By choosing a value for V_{COMP} that provides an equal and opposite potential difference to Δ_x , we can cancel $\frac{dF_x}{dq}$ to first order.

The LISA Pathfinder sensitivity is sufficient that the effects of in-band fluctuations of q and Δ_x , described by their power spectral densities (PSDs), S_q and S_{Δ_x} are measurable directly in Δg by exaggerating Δ_x or q respectively.

Simulations and ground-based laboratory measurements provide indications of the expected behavior of the test mass charge and stray potentials. High-energy physics simulations [16, 17] predict that a net positive charging rate of 40-70 elementary charges per second (es^{-1}) from galactic cosmic rays (GCR) at the minimum of the 11-year solar activity cycle (20-40 es^{-1} at maximum when GCR flux is suppressed) with a charge-noise equivalent to that produced by a rate of single charges, λ_{eff} , of 200-400 s^{-1} . The amplitude spectral density (ASD) of the test-mass charge is expected to have the form $S_q^{1/2} = \frac{e\sqrt{2\lambda_{\text{eff}}}}{2\pi f}$ with an amplitude of 0.6-0.7 fC $\text{Hz}^{-1/2}$ at 1 mHz.

Δ_x originates both from surface patch potentials within the sensor and the GRS electronics.

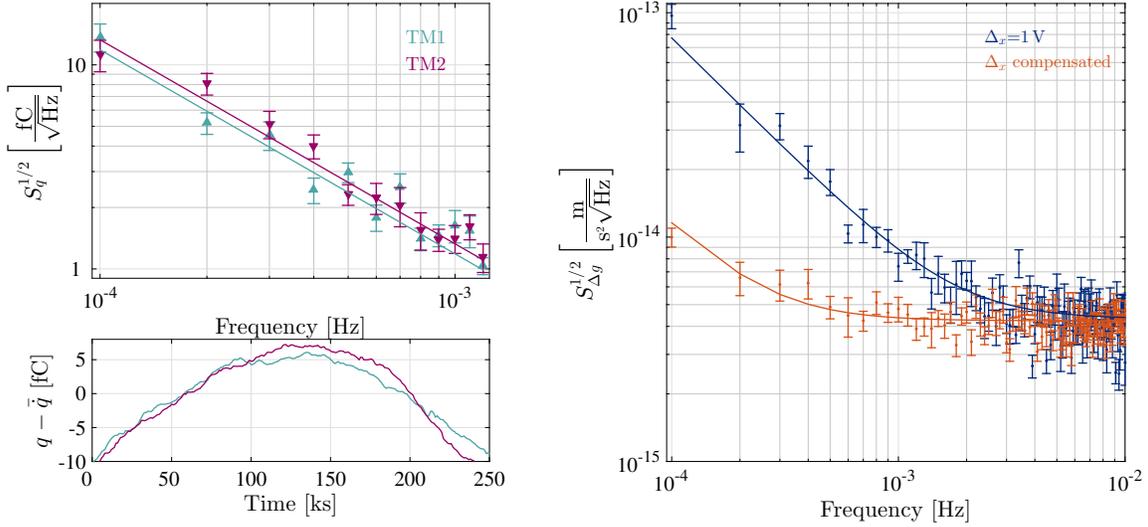


FIG. 1. Measurements of TM charge fluctuations. Upper-left: the ASD of the charge on TM 1 (Δ) and TM 2 (∇) in the LISA band with $1/f$ fits. Lower-left: the charge time series after removal of the linear trend due to the average charge rate over the course of the 3-day measurement. Right: consecutive measurements of the ASD of Δg with exaggerated Δ_x (dark-blue) and with Δ_x compensated to $\lesssim 3$ mV (red). Continuous curves show the result of a combined fit to the background noise and Δ_x -dependent $1/f$ excess.

Measurements with representative systems in laboratory tests have found static levels of up to 100 mV [2, 26–28]. Tests before launch showed $S_{\Delta_x}^{1/2}$ coming from voltage fluctuations from the spacecraft electronics to be $30 \mu\text{V Hz}^{-1/2}$ at 1 mHz [29]. Torsion pendulum measurements using a representative TM and GRS and similar electronics have placed $2\text{-}\sigma$ upper-limits on the total fluctuations, including patch potentials of $80 \mu\text{V Hz}^{-1/2}$ at 1 mHz and $290 \mu\text{V Hz}^{-1/2}$ at 0.1 mHz [2].

A ~ 3 -day measurement of q was made injecting $V_{\text{MOD}}=3$ V at $f_{\text{MOD}}=6$ and 9 mHz on TM 1 and TM 2 respectively. The charge was calculated by heterodyne demodulation of Δg , with the applied V_{MOD} as the phase reference.

The average charging rates were $+22.9 \text{ es}^{-1}$ and $+24.5 \text{ es}^{-1}$ on TM 1 and 2 respectively. Over 10000-s periods, the charge rate is observed to vary by $\pm 2 \text{ es}^{-1}$, caused by a combination of low frequency noise and drift. The fluctuations around the mean charging rates are

shown in the lower-left panel of Figure 1; two $> 5\sigma$ glitches in the TM2 charge fluctuations have been removed. S_q was calculated with the Welch method, averaging 11 detrended, 40000-s Blackman-Harris (BH) spectral windows with 50%-overlap. A f^{-2} fit was applied to the PSD, down-sampled by a factor 4 to remove data correlated by spectral windowing. The resulting ASD are shown in the upper left panel of Figure 1 and a summary of the results is given in Table I.

The f^{-2} dependence of the charge PSD and observed absence of correlation in the two charge time-series are consistent with the model of independent Poissonian processes for the two TMs, at least down to 0.1 mHz. A common drift in the charge rate at very-low frequency (visible in the time series as a quadratic dependence after removal of the linear trends due to the average charging rates) correlates well with measurements in the on-board particle monitor and is therefore likely caused by changes in the

TABLE I. Test mass charging properties

	TM1	TM2	
\dot{q}	+22.9	+24.5	e s^{-1}
λ_{eff}	1060 ± 90	1360 ± 130	s^{-1}
$\lambda_{\text{eff}(1+2)}^{\text{a}}$	2200 ± 260		s^{-1}

^a Determined from fit to Δg with $\Delta_x=1\text{V}$

incident particle flux. The measured charge-noise levels have roughly 5 times the expected noise power, with effective charge rates between 1000 and 1400 s^{-1} . Possible causes for an excess are a larger-than-expected number of high-multiplicity charging events produced by very-high energy ($\sim\text{TeV}$) cosmic rays, or a large population of low-energy ($\sim\text{eV}$) secondaries emitted from TM and GRS surfaces. These two energy regimes are the source of most uncertainty in the charging predictions [16].

In this measurement, made some 3-4 years before solar minimum, we find test-mass charging rates within the expected range but measurably different on the two TMs. The difference in the charge rates may originate in the different Volt-scale AC electrostatic fields used for force actuation in the two GRS. If confirmed, this would favor secondary electrons as the source of excess noise. Further measurements characterizing the charge-rate behavior in detail will be the subject of future work.

The spectral density of the charge noise can also be determined from a measurement of Δg with an exaggerated potential difference Δ_x . The right panel of Figure 1 shows two measurements of the ASD of Δg calculated with the same method described for S_q . The first lasting ~ 2.5 -days with $\Delta_{x_1}=\Delta_{x_2}=1\text{ V}$ and calculated by averaging 9 overlapping, 40000-s BH windows. The second ~ 1 -day later with both Δ_x compensated to $\lesssim 3\text{ mV}$ (as described below) using 15 windows covering nearly 4-days. We perform a combined fit to the spectra in the frequency range $0.1 \leq f \leq 20\text{ mHz}$ assuming a stationary background and an excess due to random charging proportional to Δ_x . We find the

excess noise in Δg in the presence of the applied electric field is compatible with a total effective charge rate $\lambda_{\text{eff}_1} + \lambda_{\text{eff}_2} = 2220 \pm 260\text{ s}^{-1}$ in good agreement with the dedicated measurement of the charge fluctuations on each test mass shown in Table I. The charge noise observed in these two measurements separated by 60-days is stationary to better than 10% and shows no measurable departure from a pure Poissonian behavior.

In order to calculate Δ_x and the required compensation voltages, $\frac{dF}{dq}$ was determined from Δg using four charge steps of $\sim 0.6\text{ pC}$. The charge was measured throughout with $V_{\text{MOD}}=50\text{ mV}$ and $f_{\text{MOD}}=5\text{ mHz}$. Figure 2 shows q and Δg as a function of time through one of these measurements. The charge steps were repeated with V_{COMP} of $-20, 0, +20\text{ mV}$ and the dependence of $\frac{dF}{dq}$ on V_{COMP} confirms our electrostatic model to better than 2%. Two measurements on each GRS were made 45-days apart, the second with Δ_x on TM1 compensated within 3 mV. At this level, the contribution to $S_{\Delta g}^{1/2}$ from random charging is $0.2\text{ fm s}^{-2}\text{ Hz}^{-1/2}$ at 0.1 mHz. A further three measurements were made 7 months later. The first on TM1 and a final measurement on each TM after reducing the temperature of the sensor from 22 to 11°C.

The calculated values for Δ_x , corrected for applied compensation, are given in Table II and plotted in Figure 2 against the system pumping time. We note that the rotational stray-field imbalance, Δ_ϕ and Δ_η , that couple TM charge into torque in analogous fashion to Eqn 1, have been measured in the same experiments to be roughly -32 and $+36\text{ mV}$ for TM1 and $+119$ and $+84\text{ mV}$ for TM2. When considered with the uncompensated values measured for Δ_x , roughly -20 mV and 0 mV for TM1 and TM2, the stray fields in the GRS would seem to be similar in magnitude to those observed in various measurements on GRS prototype hardware on ground [2, 26, 27]. Small but significant changes in Δ_x are observed for the two TMs, consistent with drifts of slightly less

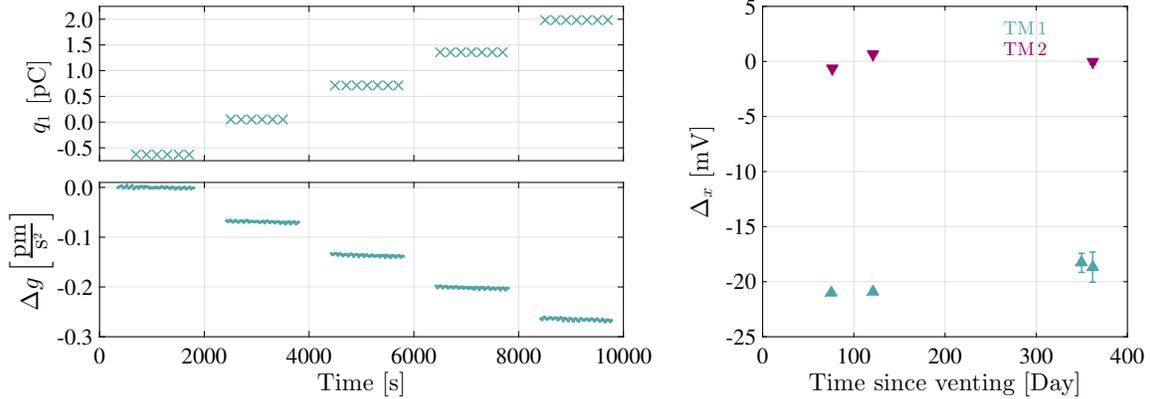


FIG. 2. Estimation of Δ_x . Left: time-series of charge steps in q_1 (\times) and Δg (\bullet) for measurement on TM1 on day 110 of 2016 with no applied compensation. Data during UV illumination periods lasting ~ 100 s have been removed. Right: Seven measurements of uncompensated Δ_x on TM1 (Δ) and TM2 (∇), plotted against the time elapsed since opening the vacuum chambers containing the GRS to space.

TABLE II. Estimates of uncompensated Δ_x

Date	Δ_{x1}	Δ_{x2}	
2016-110	-21.02 ± 0.07	-0.67 ± 0.07	mV
2016-155	-20.93 ± 0.04	$+0.66 \pm 0.03$	mV
2017-018	-18.3 ± 0.9	—	mV
2017-030	-18.7 ± 1.4	-0.1 ± 0.2	mV

than mV/month, roughly an order of magnitude below typical drift values observed, for a limited number of samples, on ground [2, 28]. This suggests that only very infrequent repetition of the measurement and compensation scheme will be necessary in LISA to keep acceleration noise from TM charge fluctuations below a tolerable level.

The spectral density of stray voltage fluctuations was measured by increasing the test-mass charge and using a similar method to that used to observe the charge noise effect on $S_{\Delta g}^{1/2}$. Figure 3 shows two measurements of the ASD of Δg . The first over nearly 2-days (average of 6, overlapping, 40000-s windows) with normal levels of TM potential: $\langle V_{\text{TM1}} \rangle = -16$ mV and $\langle V_{\text{TM2}} \rangle = -24$ mV. followed within a day by a measurement of just over 2-days (8, 40000-s

windows) at $\langle V_{\text{TM1}} \rangle = -1066$ mV and $\langle V_{\text{TM2}} \rangle = -1058$ mV. Fitting a polynomial to the excess in the PSD of Δg proportional to V_{TM}^2 of the form $S_{\Delta g} = \left[A^2 \left(\frac{1 \text{ mHz}}{f} \right)^2 + B^2 \left(\frac{1 \text{ mHz}}{f} \right) \right] \left(\frac{V_{\text{TM}}}{\text{TV}} \right)^2$ we find $A = 3.5 \pm 0.7 \text{ fm s}^{-2} \text{ Hz}^{-1/2}$ and $B = 6.4 \pm 0.5 \text{ fm s}^{-2} \text{ Hz}^{-1/2}$ ($\chi^2 = 1.2$). This converts to an ASD of the fluctuations in Δ_x of $34 \pm 2 \mu\text{V Hz}^{-1/2}$ at 1 mHz and $190 \pm 30 \mu\text{V Hz}^{-1/2}$ at 0.1 mHz. The measured fluctuations in Δ_x are thus clearly resolved and are consistent with the upper limits placed for a GRS prototype on ground [2]. They are also consistent with ground measurements of the low-frequency actuation-circuitry voltage noise, which is indistinguishable from stray surface potential fluctuations in our acceleration noise measurement.

Throughout the majority of the operational phase of the mission, V_{TM} has been controlled within ± 80 mV of zero by discharging under UV illumination during weekly or fortnightly interruptions to science measurements for orbital correction. The RMS V_{TM} in a two-week measurement period such as that described in [8] is typically < 40 mV and the contribution to $S_{\Delta g}^{1/2}$ at 0.1 mHz is $1.6 \pm 0.2 \text{ fm s}^{-2} \text{ Hz}^{-1/2}$.

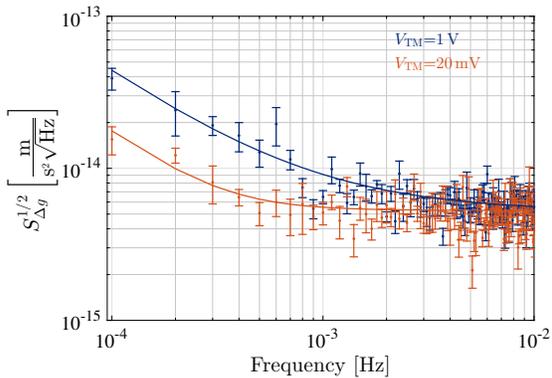


FIG. 3. Spectral density of consecutive measurements of Δg with $V_{TM}=1$ V and 20 mV. A combined fit to the background and V_{TM} -dependent component is shown with continuous curves.

We have presented the most sensitive measurements of charge-related electrostatic forces on free-falling test masses relevant for sensitive gravitational experiments in space. Technology and mitigation methods developed for minimizing these forces on test masses in capacitive position sensors have been demonstrated and are directly transferable to LISA. Their contribution to the acceleration noise of individual test masses has been reduced to a level of roughly $1 \text{ fm s}^{-2} \text{ Hz}^{-1/2}$ at 0.1 mHz, compatible with the budgeted allocation for electrostatic forces contributing to the total acceleration noise in LISA [7].

The optimal method of test-mass charge control in a future space-based GW observatory will minimize the disruption to observations and the total contribution to acceleration noise. A continuous UV discharge scheme trades additional random-charging noise for a reduced coupling between V_{TM} and fluctuating stray potentials. Since the latter dominates the charge-related noise in the periodic discharge scenario described in this paper by a factor of nearly 50 in power, it is likely some improvement can be obtained with a continuous scheme. This will be studied in a future experiment.

The methods and technology described here

can enable a new generation of instrumentation for gravity-gradiometry and fundamental physics with improved performance in the milli-Hz regime.

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