

# Intermediate mass-ratio black hole binaries: Applicability of small mass-ratio perturbation theory

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The inspiral phasing of binary black holes at intermediate mass-ratios ( $m_2/m_1 \sim 10^{-3}$ ) is important for gravitational wave observations, but not accessible to standard modeling techniques: The accuracy of the small mass-ratio (SMR) expansion is unknown at intermediate mass-ratios, whereas numerical relativity simulations cannot reach this regime. This article assesses the accuracy of the SMR expansion by extracting the first three terms of the SMR expansion from numerical relativity data for non-spinning, quasi-circular binaries. We recover the leading term predicted by SMR theory and obtain a robust prediction of the next-to-leading term. The influence of higher order terms is bounded to be small, indicating that the SMR series truncated at next-to-leading order is quite accurate at intermediate mass-ratios and even at nearly comparable mass binaries. We estimate the range of applicability for SMR and post-Newtonian series for non-spinning, quasi-circular inspirals.

Inspiraling and merging black hole (BH) binaries are the most numerous source of gravitational waves (GW) observed by the LIGO and Virgo detectors [1, 2] and are one of the key science targets for third generation ground-based GW detectors [3], as well as the space-based LISA observatory [4]. The mass-ratio  $q \equiv m_2/m_1 \leq 1$  is one of the key parameters in the dynamics of these systems. The LIGO and Virgo observations [5–7] mostly report  $q$  close to unity, with GW190412 [8] and GW190814 [9] the first systems with clearly unequal masses ( $q \sim 0.28$  and  $q \sim 0.11$ ).

In the future, observations of binaries with lower  $q$  are expected: Continued observations with the current detectors [10] may reveal binaries with smaller  $q$ . Third generation ground based detectors with improved low frequency sensitivity will be able to detect the capture of stellar mass BHs by intermediate mass BHs with mass-ratios down to  $q \sim 10^{-3}$  [11]. LISA will observe the mergers of massive BHs of millions of solar masses. While the majority of these are expected to have  $q \gtrsim 0.1$ , there could be a significant tail of events down to  $q \sim 0.01$  [12, 13]. LISA will also be sensitive to mergers of intermediate mass BHs with massive BHs ( $q \sim 10^{-3}$ ) and extreme mass-ratio inspirals ( $q \sim 10^{-5}$ ) as sensitive probes of black hole physics [4].

The modeling of inspiraling binaries at all mass-ratios is therefore of paramount importance for detection and analysis of GW sources. The three primary modeling approaches are post-Newtonian slow-velocity perturbation theory [14], numerical relativity (NR), i.e. direct numerical integration of the full non-linear Einstein equations [15], and small mass-ratio (SMR) perturbation theory [16]. Effective one body methods [17] provide a means to combine and resum information from all three approaches and also from newer developments like post-Minkowski expansions [18].

This article examines whether the SMR and NR approaches combined can accurately model binaries with any mass-ratio or whether there is a ‘gap’ at intermediate mass-ratios where neither SMR nor NR is sufficiently accurate. The SMR approximation expands the dynamics of a coalescing binary in powers of  $q$  or the symmetric mass-ratio  $\nu \equiv m_1 m_2 / (m_1 + m_2)^2 = q + \mathcal{O}(q^2)$ . At leading order, the secondary object follows a geodesic in the background space-time generated by the primary. The impact of the secondary’s mass on the dynamics can be included as an effective force term, the gravitational self-force (GSF). Calculation of the GSF has progressed rapidly over the past two decades (see [19] for a review), but the full next-to-leading order contribution to the orbital phasing has not yet been obtained. While the main motivation for SMR lies in extreme-mass-ratio inspirals, there is increasing evidence [20–26] that the SMR may be applicable even at comparable masses.

Numerical relativity directly solves the full non-linear Einstein equations [15]. The vast majority of simulations performed to date are at comparable masses, with only very few simulations at  $q \lesssim 0.1$  (see, e.g. [27], but note [28, 29] for simulations at  $q = 1/18$  to  $q = 1/128$ ). The limited coverage in  $q$  has two causes. First, the number of orbits the binary spends in the strong field region grows  $\propto \nu^{-1}$ . Second, because of the Courant-limit on the time-step of the numerical simulations, the number of time-steps per orbit increases  $\propto q^{-1}$ . Combined, these effects cause an increase in computational cost at least quadratically in mass-ratio. The need for higher numerical resolution to resolve the ever smaller secondary (as  $q \rightarrow 0$ ), and to preserve phase-accuracy over the increasingly longer inspiral will increase computational cost further.

Given the expectation of binaries at all mass-ratios, the question arises how to model intermediate mass-ratio binaries at small separation: post-Newtonian theory is not

accurate close to merger owing to the high velocities; numerical relativity simulations are limited to large mass-ratios,  $q \gtrsim 0.1$ ; and the SMR approximation is presently only available at leading order in  $q$ , and thus may be inaccurate at intermediate mass-ratios. This letter investigates the existence of a mass-ratio gap where none of the modeling approaches is applicable. We analyse NR simulations at mass-ratios  $0.1 \leq q \leq 1$  computed with the `SpEC`-code [27, 30] and extract the first three terms in the SMR expansion of the orbital phasing. Analysing these terms, we conclude that SMR results at next-to-leading order can likely bridge the mass-ratio gap up to mass-ratios  $q$  large enough to be covered by numerical relativity.

*Methodology.*— We use geometric units such that  $c = G = 1$  and examine the orbital phase extracted from the gravitational radiation at future null infinity,

$$\phi \equiv \frac{1}{2} \arg h_{22}. \quad (1)$$

Here  $h_{22}$  is the spin-weight  $s = -2$  spherical harmonic  $(\ell, m) = (2, 2)$  mode of the complex GW strain. The current work focuses on non-precessing binaries where Eq. (1) is sufficient.

Introducing the orbital frequency,

$$\Omega \equiv \frac{d\phi}{dt}, \quad (2)$$

we consider the orbital phase as a function of the orbital frequency,  $\phi(\Omega)$ . In the SMR approximation,  $\phi$  can be calculated by a two timescale expansion [31] leading to a power series in the mass-ratio, known as the post-adiabatic (PA) expansion,

$$\phi(\Omega) = \sum_{n=0}^{\infty} \nu^{n-1} \phi_{n\text{PA}}(M\Omega). \quad (3)$$

Here,  $\phi_{n\text{PA}}$  are functions of  $M\Omega$ , where  $M \equiv m_1 + m_2$  is the total mass of the binary. Alternatively, one can consider  $\phi_{n\text{PA}}$  as functions of  $m_1\Omega$  and/or expand in  $q$  as the small parameter (cf. Fig. 3 below).

The leading order term [31, 32]  $\phi_{0\text{PA}}$  (called “adiabatic” or “0-post-adiabatic”) is independent of the choice of expansion parameter or mass-normalization. It can be computed by energy balance,

$$\frac{d\phi_{0\text{PA}}}{d\Omega} = \nu \Omega \frac{dE}{d\Omega} \left( \frac{dE}{dt} \right)^{-1}, \quad (4)$$

where  $E(\Omega)$  is the specific energy of the circular geodesic with orbital frequency  $\Omega$ , and  $dE/dt$  its energy loss to GWs. We compute  $dE/dt$  with the `BLACK HOLE PERTURBATION TOOLKIT` [33], utilizing the arbitrary precision Teukolsky code developed in [34–40], and denote the result as  $\phi_{0\text{PA}}^{\text{SMR}}$  below.

The 1PA term in the expansion requires knowledge of the full first-order GSF for nearly circular orbits, and the dissipative part of second-order GSF for quasi-circular inspirals [31, 32]. Calculation of the first order GSF for non-spinning binaries is now routine [41–45]. The calculation of second order GSF for quasi-circular orbits, however, remains an open challenge in GSF theory, although steady progress has been made [46–53].

We use numerical relativity simulations from the `SpEC`-code, which utilizes the quasi-local angular momentum formalism to monitor the black hole spins [54–57], iterative eccentricity reduction to achieve orbital eccentricities  $e \lesssim 10^{-4}$  [58, 59], and solves the Einstein evolution equations in the generalized harmonic formulation [60–63] with constraint damping and minimally reflective outer boundary conditions [63–65] (see [27] for more details). Because of the use of spectral methods and a dual-frame approach [66] `SpEC` achieves very high accuracies even for long inspiral simulations that cover a comparatively large range in orbital frequencies. Gravitational radiation is extracted using the Regge-Wheeler-Zerilli formalism, extrapolated to future null infinity [27, 67], and corrected for center of mass drifts [68].

This study utilizes 55 NR simulations of non-spinning quasi-circular inspirals from the public SXS catalog [69, 70] with mass-ratios  $q \in [0.1, 1]$ . The initial orbital frequency is in the range  $M\Omega \sim 0.015 \dots 0.02$ . Simulations with smaller  $q$  tend to start at the higher frequencies, to achieve a computationally manageable overall duration of the simulations. All simulations are available at multiple numerical resolutions for convergence tests.

The orbital phase  $\phi^{\text{NR}}(M\Omega)$  is determined by locally fitting a low order polynomial in  $t$  to  $\phi^{\text{NR}}(t)$ . The width of the fitting window is variable such that at low frequencies it encompasses several radial oscillations of any residual eccentricity in the simulations, while at larger frequencies it is small enough to avoid systematic bias due to the rapidly changing frequency. The constant of integration when integrating Eq. (2) is chosen such that  $\phi = 0$  at  $M\Omega = 0.046$ . At a given value of  $M\Omega$ , the post-adiabatic coefficients  $\phi_{n\text{PA}}(M\Omega)$  are determined by fitting a polynomial in  $\nu$  to the data-points  $(\nu_A, \phi_A^{\text{NR}}(M\Omega))$ , where  $A = 1, \dots, 55$  labels the NR simulations, and  $\nu_A$  is the symmetric mass-ratio of each simulation. This fit is repeated for many values of  $M\Omega$ . Error estimates are obtained by repeating this procedure with (i) medium-resolution NR simulations; (ii) using the Weyl-scalar  $\Psi_4$  instead of the GW strain in Eq. (1); (iii) vary the order with which the GW strain is extrapolated to future null infinity; and (iv) changing the number of terms in the fit of form Eq. (3) between three and four. The range of these calculations is reported as error bar in our results. At each frequency, only those NR simulations are used that have a starting frequency below; for  $M\Omega \lesssim 0.02$ ,

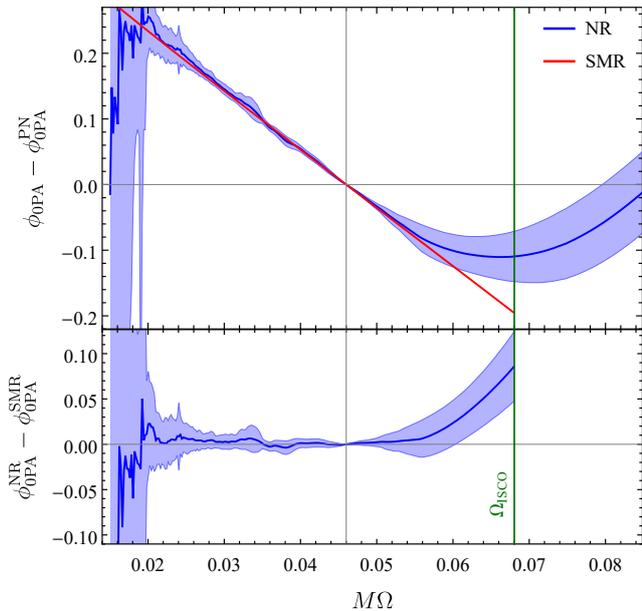


Figure 1. **Top:** Leading order in mass-ratio contribution to the orbital phasing of the quasi-circular inspiral of non-spinning black holes. Shown are the result derived here from NR simulations (‘NR’), as well as the small-mass-ratio perturbation theory (‘SMR’). For both curves, the 3.5PN result [14] was subtracted for clarity of plotting. **Bottom:** Difference between SMR and the NR result. The shaded areas indicate the estimated uncertainty of the numerical calculation and  $\Omega_{ISCO}$  indicates the last stable orbit for  $\nu = 0$ .

the reduced number of available NR simulations causes larger error bars.

*Results.*— The leading order term  $\phi_{0PA}(M\Omega)$  can be extracted with good accuracy from the NR simulations, as shown in Fig. 1. To reduce the dynamic range on the y-axis, this figure shows the difference to the post-Newtonian  $\phi_{0PA}^{PN}$  result at order  $(v/c)^7$ , taken from [14]. The blue curve represents the result of our analysis of NR simulations (with error bar), whereas the red line is the leading order SMR result computed by Eq. (4). The agreement between the two is quite remarkable, and is a first indication that the PA expansion of the phase in the mass-ratio is well-behaved for comparable mass-ratios.

At higher frequencies,  $M\Omega \gtrsim 0.055$ , we find an apparently systematic deviation between NR and the SMR result. This deviation may arise from a breakdown of the PA expansion near the last stable orbit as the binary transitions from inspiral to plunge. Studies of this transition regime [71, 72] lead to order  $\nu^{-1/5}$  corrections to Eq. (3). Including such a term in our fit does indeed eliminate the systematic deviation at  $M\Omega \gtrsim 0.055$ . However, the additional term is nearly degenerate with the 0PA and 1PA terms at low frequencies making it impossible to get robust numerical results for  $\phi_{1PA}$  and higher. Therefore, we proceed in our analysis without such tran-

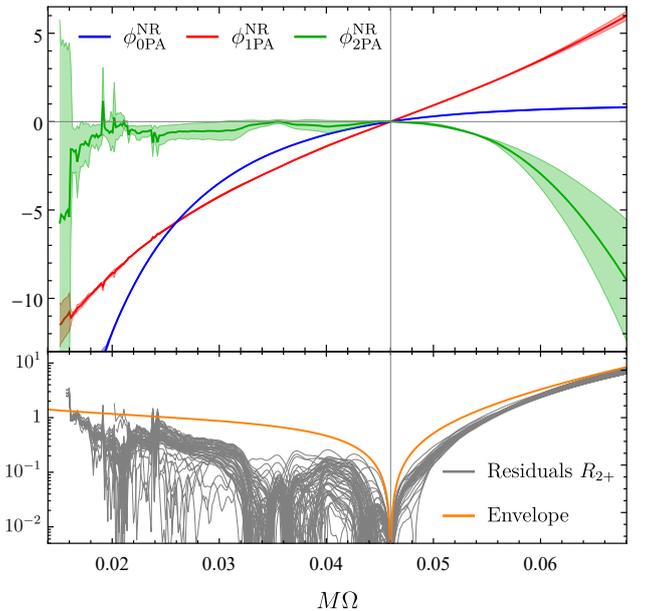


Figure 2. **Top:** The three leading terms in the mass-ratio expansion of the orbital phase, as computed here. **Bottom:** Residuals  $R_{2+}$  for all 55 NR simulations indicating the combined contributions of 2PA and higher, as well as an envelope bounding these residuals in a  $\nu$ -independent manner.

sition terms.

Given how well the numerically extracted 0PA term agrees with its SMR prediction, we henceforth set it to the SMR value when fitting for the higher order PA terms. Figure 2 shows the 1PA and 2PA term obtained from the NR simulations, together with the 0PA term already discussed in Fig. 1. The coefficients  $\phi_{nPA}$  are of comparable magnitude in the frequency range covered by our analysis, suggesting that the PA series is convergent at equal masses. Moreover, for frequencies  $M\Omega \lesssim 0.05$ , the 2PA coefficient is almost consistent with zero, i.e. the 0PA and 1PA terms already capture essentially all variation due to mass-ratio in the numerical data at these frequencies. In fact, “goodness-of-fit” indicators, such as the adjusted  $R^2$  value, show only marginal improvements when adding terms to the fit beyond the 1PA coefficient.

The lower panel of Fig. 2 provides a different view on the importance of terms beyond  $\phi_{1PA}$ : For each of the 55 NR simulations, this panel plots

$$R_{2+} \equiv \frac{1}{\nu} \left( \phi^{NR} - \frac{1}{\nu} \phi_{0PA}^{NR} - \phi_{1PA}^{NR} \right), \quad (5)$$

i.e. the contribution of all terms  $n \geq 2$  in Eq. (3), with overall  $\nu$  scaling compensated. All  $R_{2+}$  can be bounded independent of mass-ratio by an envelope function, consisting of the known 3.5PN terms of  $\phi_{2PA}$  and a higher order polynomial in  $M\Omega$  fitted by eye.

So far, we have expanded in symmetric mass-ratio  $\nu$ ,

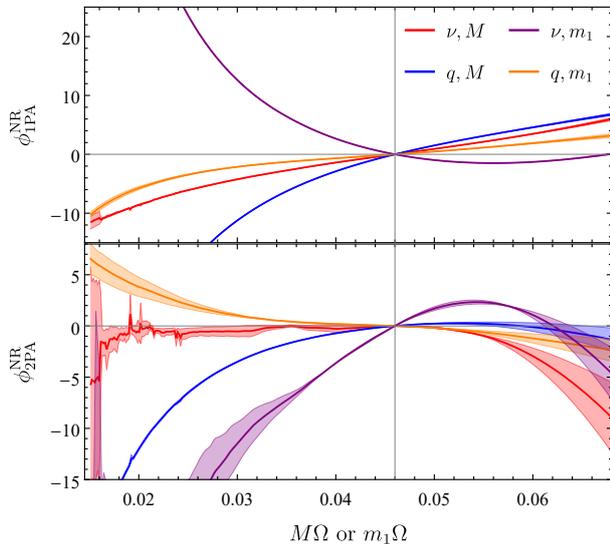


Figure 3. Impact of the choice of expansion parameters on the 1PA and 2PA contributions to the orbital phasing. The different curves differ in whether in Eq. (3) is expanded in powers of  $\nu$  or in  $q$ , and whether the  $\phi$ -functions are written in terms of  $M\Omega$  or  $m_1\Omega$ . The combination  $\nu, M$  yields an exceptionally small 2PA term at low frequencies.

while scaling orbital frequencies by total mass  $M$ , cf. Eq. (3). One can also use the mass-ratio  $q = m_2/m_1$  as the small parameter, and/or scale orbital frequency by the large body’s mass  $m_1$ . This yields four variations, all of which agree at the leading 0PA-order. Figure 3 presents the results for the 1PA and 2PA contributions. In all four cases, the extracted 1PA and 2PA coefficients remain of similar magnitude, implying that the expansion is not dominated by higher order terms. However, the 2PA term is remarkably small only when expanding using the symmetric mass ratio  $\nu$  and total mass  $M$ . The choice  $\nu, M$  is indeed preferred as it is invariant under exchange of the two bodies  $1 \leftrightarrow 2$  [20, 73].

*Discussion.*– The phasing of inspiraling BH binaries is of utmost importance for GW astronomy to find signals, determine their parameters and to perform tests of general relativity. Binaries at intermediate mass-ratios  $q \sim 10^{-3}$  are in a regime not accessible to NR, while potentially out of reach for SMR perturbation theory. This situation is compounded by the difficulty of calculations of the SMR expansion, for which today only the leading order (called the zeroth post-adiabatic order) is fully known. Here, we extract the first three terms of the SMR-expansion from NR simulations at comparable masses,  $q \geq 0.1$ , and use these results to perform the first comparison between NR and SMR expanded results for a gauge invariant quantity that includes both dissipative and conservative effects, namely the accumulated orbital phase as a function of orbital frequency  $\phi(M\Omega)$ . We have successfully extracted the post-adiabatic expansion

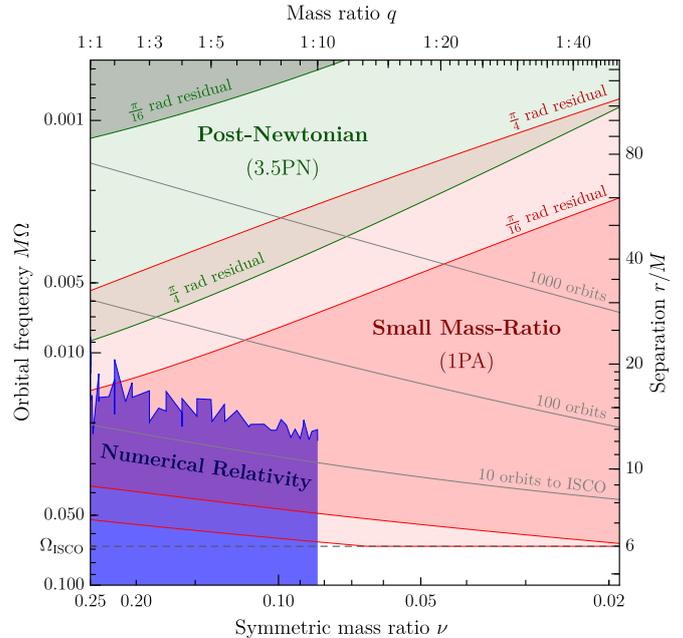


Figure 4. Region of applicability of different approximation techniques for non-spinning quasi-circular binary black hole inspiral. The shaded regions indicate ranges within which the cumulative orbital phase-error is less than  $\pi/4$  and  $\pi/16$  radians, respectively.

sion of this quantity as a power series in the mass-ratio from non-spinning quasi-circular NR simulations.

The leading adiabatic (0PA) term agrees with the result from SMR calculations. In addition we obtain a robust determination of the 1PA term, serving as a concrete prediction for the ongoing SMR calculation of this term, which requires the dissipative part of the second order gravitational self-force. We also estimate the 2PA term  $\phi_{2PA}$  from the NR data. Its amplitude is comparable to  $\phi_{0PA}$  and  $\phi_{1PA}$  for the frequency-range considered here, indicating that the PA-expansion remains well-behaved. In particular, when the PA series is expanded in powers of the symmetric mass-ratio while keeping the total mass fixed, the 2PA and higher order terms are consistent with zero within the numerical accuracy for  $0.015 \lesssim M\Omega \lesssim 0.05$ . For higher frequencies (approaching the last stable circular orbit), we find indications of a transition-regime to plunge where the series in integer powers of  $\nu$  is no longer applicable.

Our analysis allows us to delineate the regions of applicability of SMR, NR and PN in a quantitative way, as shown in Fig. 4: Assuming  $\phi_{1PA}$  will become available through GSF calculations, the envelope to the  $R_{2+}$  in Fig. 2 gives a bound on the secular contributions of higher PA terms. The red shaded areas in Fig. 4 show the largest  $M\Omega$  interval that can be covered such that the total accumulated phase-error due to  $\geq 2PA$  terms is below a certain value. The region of applicability of

SMR increases toward smaller mass-ratios, but is still non-negligible even at comparable masses. The post-Newtonian errors are estimated by fits against  $\phi^{\text{NR}}(M\Omega)$ , cf. top panel of Fig. 1. The green shaded areas indicate regions where the cumulative 3.5-PN phase-error for the entire inspiral up to the given frequency is below a certain value. Finally, the blue shaded area indicates the region covered by the NR simulations used here. These simulations have phase-accuracy better than the  $\pi/16$  contour line, indicating that the usability of NR is not limited by accuracy but rather by the length of the simulations. The three modelling approaches deliver complementary information, covering different regions of the parameter space. The region of validity of each method depends on the desired accuracy, and it also depends on the use of the waveforms: For GW astronomy, only the accuracy within the frequency band of the relevant GW detectors is important, and this will depend on the total mass of the binary. Moreover, the needed accuracy will depend on the signal-to-noise ratio at which it is observed.

We note that the adiabatic  $\phi_{\text{OPA}}$  term is never accurate enough in the metric of Fig. 4, because  $\phi_{\text{1PA}}$  contributes tens of radians in the frequency range considered, independent of the mass-ratio. This underlines the importance of calculating the 1PA term (and therefore the second order gravitational self-force) for modelling binaries of any mass-ratio. Furthermore, the application of the 1PA approximation for low frequencies is limited by a  $(M\Omega)^{-1/3}$  divergence of the 2PA term. This motivates the development of models that incorporate both SMR and PN results, e.g. using effective-one-body theory [74–77].

The results in this paper come with two important caveats. First, our results are limited to non-spinning quasi-circular black hole binaries. Adding spin or eccentricity makes the waveform considerably more complex and could make the convergence of the PA series significantly worse. Future studies are needed to explore the full parameter space. Even for non-spinning quasi-circular case, NR simulations at smaller mass-ratio are needed to investigate the transition to plunge, as well as longer simulations, to extend our analysis to smaller frequencies.

Second, the current analysis applies only to the inspiral, since the PA expansion is known to breakdown at the last stable orbit. Our results motivate the development of 1PA accurate models that also include plunge, merger, and ringdown, as has previously been done at OPA order [78].

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