

Quantum Populations near Black-Hole Singularities

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Schwarzschild black-hole interiors border on space-like singularities representing classical information leaks. We show that local quantum physics is decoupled from these leaks due to dynamically generated boundaries, called Zeno borders. Beyond Zeno borders black-hole interiors become asymptotically silent, and quantum fields evolve freely towards the geodesic singularity with vanishing probability measure for populating the geodesic boundary. Thus Zeno borders represent a probabilistic completion of Schwarzschild black holes within the semiclassical framework.

I. INTRODUCTION

Schwarzschild black-hole interiors contain geodesic borders separating them from space-like singularities. Any information migrating across a geodesic border towards the singularity is irretrievably lost by causality. If Schwarzschild singularities absorb information, the corresponding evolution qualifies as paradoxical since it violates sacrosanct rules of information processing [1]. It is generally expected that spacetime fluctuations deform the Schwarzschild geometry near its geodesic border to yield a consistent quantum evolution. While the details of this dynamical regularization mechanism are unknown, they are important for global aspects of quantum information processing by black holes, such as the black-hole information paradox [2–4].

In this article, we show that Schwarzschild singularities border on asymptotically silent spacetime regions, that is regions inhibiting spatial quantum correlations irrespective of the initial field configuration. More importantly, they accommodate so-called Zeno borders, which mark a stack of hypersurfaces terminated by a geodesic border with the following property: The probability measure for populating quantum information within the stack decreases monotonously towards the singularity and vanishes at the geodesic border. As a consequence, quantum events cannot probe the geodesic border and quantum information cannot migrate across the geodesic border, irrespective of any quantum completion of gravity. In other words, Zeno borders render black-hole interiors leak-proof and represent, to the best of our knowledge, the first explicit quantum completion of a geodesically incomplete spacetime within the semiclassical framework [5]. Quantum field theoretic completeness conceptualises the intrinsic consistency of quantum field theory in curved space-times in the sense that the theory itself is protected from a breakdown when reaching the geodesic border. Zeno regions incorporate this dynamical protection mechanism which guarantees the predictability of

the theory itself and likewise passes the regularity on to physical observables.

To substantiate this probabilistic completion we proceed as follows: after stating the geometrical set-up, we will present an intuitive argument based on scaling relations in the micropædia which introduces the idea of the Zeno region from the perspective of quantum measurements. The exact argument is then given in the macropædia followed by an interpretation based on observables showing that the probabilistic completion closes the geodesic border with respect to quantum field theory.

II. GEOMETRICAL PRELIMINARIES

Schwarzschild black holes are the warped geometries $\mathcal{B} := P_{<} \times_t S^2$, with $P_{<}$ denoting the region $t < t_g := 2M$ in the (t, r) -half plane $(0, r_g) \times \mathbb{R}^+$, where the projection $t : \mathcal{B} \rightarrow (0, r_g)$ is the Schwarzschild time and the projection $r : \mathcal{B} \rightarrow \mathbb{R}^+$ is the Schwarzschild radius function, and S^2 denotes the unit two-sphere. Note that in these conventions the Schwarzschild coordinate vector field ∂_t is time-like, and ∂_r is space-like on \mathcal{B} . Compared to the Schwarzschild exterior spacetime, Schwarzschild time and radius function have interchanged their meaning. Taking this into account, the quadratic form associated with Schwarzschild black holes is $ds^2 = -s^{-1}(t)dt^2 + s(t)dr^2 + t^2w$, where $s(t) := |1 - r_g/t|$ is the Schwarzschild function, and w denotes the line element associated with the Euclidean metric on S^2 , equipped with the usual spherical coordinates (ϑ, φ) . \mathcal{B} corresponds to a spatially homogeneous and anisotropic cosmological solution of general relativity. It can be foliated into Cauchy hypersurfaces $(0, r_g) \times \Sigma$ along the temporal direction, where Σ is the folio of spatial hypersurfaces Σ_t labeled by Schwarzschild time. The metric field associated with the above quadratic form will be denoted by g , and its pull-back to Σ by g_Σ .

In this geometry, the geodesic motion of a point particle that is initially equatorial relative to Schwarzschild spherical coordinates is bound to remain equatorial, $\vartheta = \pi/2$. The so-called energy equation $E^2 = (dt/ds)^2 + V_{\text{eff}}$ holds, where $E := s(t)dr/ds$ and $L := t^2d\varphi/ds$ are constants, which have an intuitive interpretation in the

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exterior as asymptotic energy per unit mass and angular momentum per unit mass, respectively. In fact, in the exterior, the definition of L formally coincides with Kepler's second law. The effective potential is given by $V_{\text{eff}} := -(1 + L^2/t^2)s(t)$. Close to the endpoint at $t = 0$, the effective potential is bounded from above, $V_{\text{eff}} = -L^2 r_g/t^3$ plus less singular contributions. Thus the classical motion generated by V_{eff} is incomplete at $t = 0$. Potential incompleteness implies that \mathcal{B} is geodesic incomplete [6], which in turn qualifies Σ_0 as a geodesic information sink. Geodesic incompleteness, however, does not imply quantum incompleteness (and vice versa).

III. MICROPÆDIA

Let us first give an intuitive argument based on scaling relations showing that quantum information in \mathcal{B} cannot migrate across Σ_{0+} , before providing precise statements. In \mathcal{B} , scalar quantum fields Φ^*, Φ charged under $U(1)$ evolve according to the Lagrange density $\mathcal{L} = \mathcal{L}_0 + \mathcal{L}_{\text{int}}$, with the first term denoting the free theory $\mathcal{L}_0 = \Phi^* \square \Phi$. The intuitive argument will be given in the absence of interactions, $\mathcal{L}_{\text{int}} \equiv 0$. Close to the space-like singularity Σ_0 bordering on \mathcal{B} , $ds^2 \cong -(t/r_g)dt^2 + (r_g/t)dr^2 + t^2 w$, where \cong means equality up to sub-leading contributions in each term as Σ_0 is approached [7]. In this asymptotic regime, $\square \cong (-r_g/t)(\partial_t^2 + (1/t)\partial_t) + (t/r_g)\partial_r^2 + (1/t^2)\partial_{\mathcal{A}}^2$. Here $\partial_{\mathcal{A}}^2$ denotes the usual angular part of the Laplace operator in \mathbb{R}^3 in Schwarzschild spherical coordinates.

The corresponding Green function G is sourced by $\delta(t - t')\delta(\sigma - \sigma')/\sqrt{-\det(g)}$, with σ and σ' denoting Schwarzschild spatial coordinates of events localized on Σ_t and $\Sigma_{t'}$, respectively, and satisfies

$$\begin{aligned} D(t)G(t, \sigma; t', \sigma') &= \delta(t - t')\delta(\sigma - \sigma'), \\ D(t) &:= -r_g \partial_t(t \partial_t) + (t/r_g)t^2 \partial_r^2 + \partial_{\mathcal{A}}^2. \end{aligned} \quad (1)$$

In order to estimate the asymptotic relevance of each term in the differential operator, consider $D(\varepsilon\tau)$ in the limit $\varepsilon \rightarrow 0^+$. The effective potential for free fields scales asymptotically as $1/t^2$ and develops a repulsive barrier. It is well-known from quantum mechanics that potentials of this type cannot be penetrated via tunneling processes.

Following the geometrical description of space-like singularities by Belinskii, Khalatnikov, and Lifshitz [8], temporal variations dominate over spatial variations in the region bordering on Σ_{0+} . Therefore, $D(\varepsilon\tau) \cong (1/\varepsilon)\partial_\tau(\tau\partial_\tau)$. The time-dependent part of the source distribution scales like $\delta(\varepsilon\tau - t')$. This effectively allows to split the Green function $G = T(t, t')P(\sigma, \sigma')$ in the vicinity of Σ_{0+} , with the asymptotic dynamics given by $\partial_\tau(\tau\partial_\tau T) \cong 0$. Here, all identifiers labeling the eigenvalue problem of the Laplace operator have been suppressed for ease of notation. We find the asymptotic solution $T(t, t') \cong C_0(t') + C_1(t') \ln(t/r_g)$, with $C_{0,1} < \infty$ depending on the initial values given on $\Sigma_{t'}$.

In order to appreciate the rather mild divergence of the asymptotic solution T , we introduce a model for

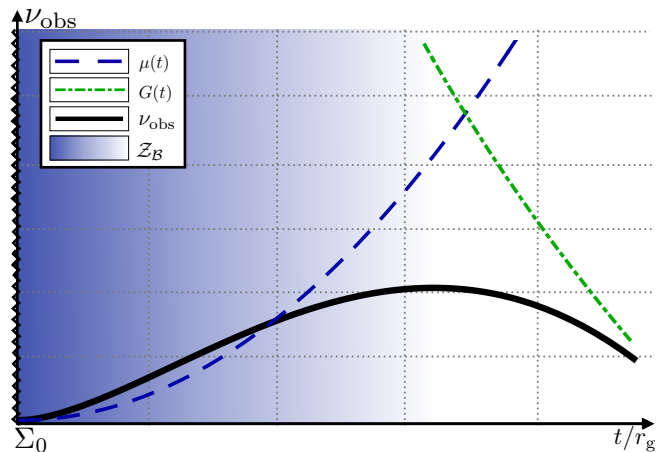


FIG. 1. Vertex density ν_{obs} for measurement processes inside Schwarzschild black holes. The communication between emitter and absorber, described by the asymptotic Green function G , diverges towards the geodesic border Σ_0 , but the integration measure μ goes sufficiently fast to zero to render ν_{obs} finite at all times. Furthermore, within the Zeno region (shaded), see Fig. 2 and the *macropædia* for details, ν_{obs} is monotonously decreasing and vanishes towards Σ_0 .

a physical detection with an emitter F_{em} localized on $\Sigma_{t'}$, $t' \in (0, t_*)$ in the asymptotic domain, where t_* denotes a fiducial time in this regime, and an absorber F_{ab} on $\Sigma_{\varepsilon\tau}$. A general measurement is given by a vertex density and involves the communication channel between source and the detector. For instance, consider $F_{\text{ab}} = \delta(t - \varepsilon\tau) f_{\text{ab}}(\sigma)$, with f_{ab} encoding the spatial extension of the detector on $\Sigma_{\varepsilon\tau}$. This blueprint effectively replaces part of Σ_{0+} with a detector volume that can resolve arbitrary frequencies. Note that the asymptotic regime is controlled by the parameter ε while τ represents a constant instant of time. The classical measurement process is described by the on-shell vertex density $\nu_{\text{obs}} = \sqrt{-\det(g)} F_{\text{ab}}^* \Xi_{\text{os}} + \text{c.c.}$ with $\Xi_{\text{os}} = G * F_{\text{em}}$ denoting a convolution of F_{em} with a bi-local kernel given by G with respect to x' [9]. In the region bordering on Σ_{0+} , as specified by the support properties of emitter and absorber, $\nu_{\text{obs}} \cong t^2 \ln(t) \delta(t - \varepsilon\tau) \sin(\vartheta) f_{\text{ab}} F_{\text{em}}$, where F_{em} contains the exclusive information on the emission process and depends only on source parameters. In particular, F_{em} is finite in accordance with the homogeneous interior \mathcal{B} . For fixed t' and $\varepsilon \rightarrow 0$, the measurement of the emitter's influence on the detector gives a vanishing response, see figure 1, $\nu_{\text{obs}} \cong 0$, in the distributional sense. This implies that no information carried by local bookkeeping devices reach Σ_{0+} .

It is possible to be more specific about the emitter and absorber. As an example, the energy momentum tensor for the complex scalar field scales like $\mathcal{T} = d\Phi \otimes d\Phi^* \propto 1/(\varepsilon\tau)^2$ on $\Sigma_{\varepsilon\tau}$ and develops a singularity towards Σ_{0+} . Our detector model is given by $F_{\text{ab}} = \delta(t - \varepsilon\tau) f_{\text{ab}}(\sigma)$,

with f_{ab} encoding the spatial extension of the detector on $\Sigma_{\varepsilon\tau}$ or a more physical absorber of an ideal gas $\tilde{F}_{\text{ab}} = M(\varepsilon\tau)U \otimes U$, where $U \cong \sqrt{\varepsilon\tau/r_g}dt$ is the normal, and $M(t)$ denotes the spatial volume integral over an energy density. From a phenomenological point of view, F_{ab} is required to have nontrivial support towards Σ_{0+} . Then, $\tilde{\nu}_{\text{obs}} \cong 0$, as well, which only confirms that the asymptotic description of the tree-level measurement process is independent of the tensor providing the principal communication channel.

Before closing the micropædia, let us briefly discuss the asymptotic diagnostics of Noether charges. Let $\Phi : \mathcal{B} \rightarrow \mathbb{C}$ carry a $U(1)$ charge. The four-dimensional $U(1)$ current density is $j = i(\Phi^* P \Phi - c.c.)$, where P denotes the four-momentum operator. Projecting the current density onto U from before, we find the following scaling relation for the charge density ρ localized on $\Sigma_{\varepsilon\tau}$: $\rho(\varepsilon\tau) \cong \rho(t_*)(t_*/\varepsilon\tau)^{3/2}$, which formally diverges as Σ_{0+} is approached. Physical measurements of the charge $Q(\varepsilon\tau)$, however, yield a finite result. In fact $Q(\varepsilon\tau) = Q(t_*)$. Black holes cannot be discharged through the geodesic singularity bordering on Σ_{0+} . Any active information sink would necessarily lead to charge depletion. Note that this discussion of asymptotic charge conservation is fully based on local physics inside black holes, and no reference to the usual global characterization in the exterior is made.

IV. MACROPÆDIA

A more rigorous argument is based on the Schrödinger representation of local quantum physics which we summarise for convenience below. Let $(a, b) \times \Sigma$ be a Cauchy foliation of a globally hyperbolic spacetime (\mathcal{M}, g) with a geodesic border Σ_a in the folio Σ of hypersurfaces. For $t \in (a, b)$, we denote by $\mathcal{C}(\Sigma_t)$ the space of instantaneous field configurations $\phi : \Sigma_t \rightarrow \mathbb{C}$. Consider the \mathbb{C} -vector space of measurable wave functionals $\Psi_t : \mathcal{C}(\Sigma_t) \rightarrow \mathbb{C}$ whose modulus is square integrable with respect to the formal functional measure $\mathcal{D}\phi$. Let $L^2(\mathcal{C}(\Sigma_t), \mathcal{D}\phi)$ be the quotient of this space by the subspace of wave functionals vanishing $\mathcal{D}\phi$ -almost everywhere in $\mathcal{C}(\Sigma_t)$. We introduce the usual norm on $L^2(\mathcal{C}(\Sigma_t), \mathcal{D}\phi)$. If \mathcal{U} is a formally measurable subset of $\mathcal{C}(\Sigma_t)$ and $\mathcal{X}_{\mathcal{U}}$ the corresponding indicator functional, then $\|\mathcal{X}_{\mathcal{U}}\Psi_t\|^2$ is the probability for the field configuration on Σ_t to be given by some $\phi \in \mathcal{U}$. This interpretation requires normalizable wave functionals. The configuration field operator $\Phi[f]$ is just the multiplication operator with $\phi[f]$, where f denotes some smooth test function of compact support on Σ_t such that $\langle \Psi_t | \Phi[f] | \Psi_t \rangle = \|\sqrt{\phi[f]}\Psi_t\|^2$ is bounded. The conjugated momentum field operator $\Pi[f]$ is the functional derivative $-i(\det(g_{\Sigma_t}))^{-1/2}\delta/\delta\phi$ in the direction of f . Heisenberg's fundamental uncertainty relation is a consequence of $[\Phi[f_1], \Pi[f_2]] = i(f_1, f_2)$, where (\cdot, \cdot) denotes the scalar product of smearing functions on Σ_t .

Consider the evolution semigroup given by the con-

tinuous map $(a, b) \rightarrow L^2(\mathcal{C}(\Sigma_t), \mathcal{D}\phi)$, defined by $t \rightarrow \mathcal{E}(t, t_0)|\Psi_{t_0}\rangle$ for an initial time $t_0 \in (a, b)$. A probabilistic interpretation is only possible for $\|\mathcal{E}(t, t_0)\| \leq 1$, that is for contractive evolution semigroups. While this includes unitarity as a special case for Minkowski spacetime, its generator density \mathcal{H} is only required to be accretive, $\text{Re}(\langle \Psi_t | \mathcal{H} | \Psi_t \rangle) \geq 0$, instead of being self-adjoint. On generic spacetimes, the pair (contractive, accretive) supersedes (unitary, self-adjoint) as a requirement on evolution semigroups in order to allow for a probabilistic interpretation. This has important ramifications for physical spacetimes as support for quantum events:

Let $(L^2(\mathcal{C}(\Sigma_{t_0}), \mathcal{H}))$ be the initial Schrödinger representation of a quantum field theory in a globally hyperbolic spacetime (M, g) with the Cauchy foliation $(a, b) \times \Sigma$ containing a geodesic boundary Σ_a in the folio Σ of hypersurfaces. A Zeno region $\mathcal{Z}_{\mathcal{M}}$ in \mathcal{M} is a stack of Cauchy hypersurfaces in Σ with the following property: The evolution semigroup generated by \mathcal{H} is strictly contractive in $\mathcal{Z}_{\mathcal{M}}$ and $\|\mathcal{E}(c, t_0)\| \rightarrow 0$ as $c \rightarrow a$.

Evidently, $\mathcal{Z}_{\mathcal{M}}$ is bounded by some Cauchy hypersurface Σ_z , called the Zeno border, and the geodesic border Σ_a . Let us justify the introduction of Zeno regions: If a spacelike singularity borders on a Zeno region, then the singularity cannot be probed by quantum fields and the quantum evolution in its immediate vicinity allows for a probabilistic interpretation. This distinguishes the physical spacetime as the principal support of quantum events from the geometrical spacetime model. Therefore Zeno regions imply a probabilistic or quantum completion of a geodesic incomplete spacetime within the semiclassical framework.

Let us now turn back to Schwarzschild black-hole interiors \mathcal{B} , Cauchy-foliated as $(0, r_g) \times \Sigma$ with respect to the Schwarzschild time function. Consider the space $\mathcal{C}(\Sigma_{t_0})$ of instantaneous field configuration $\phi, \phi^* : \Sigma_{t_0} \rightarrow \mathbb{C}$ and its Schrödinger representation $(L^2(\mathcal{C}(\Sigma_{t_0}), \mathcal{H}))$ in \mathcal{B} with Hamilton density $\mathcal{H} = \mathcal{G}(\Pi^*, \Pi) + \mathcal{V}(\Phi^*, \Phi)$, where $\mathcal{G}(\Pi^*, \Pi) := \sqrt{-g_{tt}} \Pi^* \Pi / \det(g_{\Sigma_t})$, and $\Pi := -i\delta/\delta\Phi$ denotes the momentum field conjugated to Φ [10]. The effective potential density $\mathcal{V} := \sqrt{-g_{tt}}(|\text{grad}\Phi|^2 + m^2|\Phi|^2)$ is a pure multiplication operator, with m being the mass of the scalar field.

For the ground-state wave functional we make the following ansatz

$$\Psi_t[\phi^*, \phi] = \mathcal{N}_t \exp(-[V^*]\mathcal{K}_t[V]), \quad (2)$$

which will be justified a posteriori. Here, $V = (\phi, \phi^*)^T$, and

$$[V^*]\mathcal{K}_t[V] := \frac{1}{2} \int_{\Sigma_t} d\mu_{x,y} V^*(x) K_t(x, y) V(y) \quad (3)$$

is a quadratic functional with a bi-local representation given by the matrix K . In the absence of interactions, K is diagonal and only its trace enters in (3). Close to the Schwarzschild singularity Σ_0 the quantum evolution trivialises [11], that is $\mathcal{H} \approx \mathcal{G}(\Pi^*, \Pi)$. As a consequence, the trace of the bi-local kernel matrix K becomes a contact term on $\Sigma_t \in \Sigma$ near the geodesic border,

$\text{tr}(K_{\varepsilon\tau})(x, y) \cong k(\varepsilon\tau)\delta^{(3)}(x, y)$ with [12]

$$\begin{aligned} \text{Im}(k(\varepsilon\tau)) &\cong -\frac{2}{(\varepsilon\tau)^3} \frac{1}{|\ln(\varepsilon\tau)|}, \\ \text{Re}(k(\varepsilon\tau)) &\cong |\text{Im}(C)| \frac{|\text{Im}(k(\varepsilon\tau))|}{|\ln(\varepsilon\tau)|}, \end{aligned} \quad (4)$$

where \cong refers to the limit $\varepsilon \rightarrow 0$. Our result is reminiscent of the analysis of Kasner spacetimes by Belinskii, Khalatnikov and Lifshitz: In the vicinity of space-like singularities, but still in the domain of general relativity, the spatial variations of local quantities are insignificant compared to temporal gradients. This is also similar to asymptotic silence [13] which describes how the particle horizon shrinks in homogeneous but anisotropic spacetimes bordering on a spacelike singularity. All spatial correlations vanish eventually.

The asymptotic kernel (4) implies that the immediate vicinity of the geodesic border is a Zeno region \mathcal{Z}_B since

$$\lim_{\varepsilon \rightarrow 0} \Psi_{\varepsilon\tau}[\phi^*, \phi] = \lim_{\varepsilon \rightarrow 0} |\ln(\varepsilon\tau)|^{-\Lambda v(\Sigma_{\varepsilon\tau})} = 0, \quad (5)$$

where Λ denotes a short-distance cut-off and $v(\Sigma_{\varepsilon\tau})$ a volume regularization. Note that any hypersurface contained in \mathcal{Z}_B has an infinite volume, (5) is robust against the detailed regularization prescription. The Zeno region \mathcal{Z}_B secludes the Schwarzschild black-hole's interior not granting any sojourn to quantum probes on the singular hypersurface bordering on \mathcal{B} . Thence the physical space-time is closed for quantum probes despite being geodesically incomplete in the mathematical notion. In other words, it is the portion of \mathcal{B} that can be probed via the support of quantum events, see Fig. 2.

Intuitively \mathcal{Z}_B marks the stack of hypersurfaces which enforces the probabilistic completion such that the geodesic border is closed in the quantum theoretical sector. Quantum field theory, hence, introduces a new scale τ_Z in the black hole interior. To classify τ_Z further, we compare it to the scales coming from classical gravity, i.e. r_g , and from quantum gravity, i.e. the Planck time t_{P1} .

Following the definition of \mathcal{Z}_B , the exact location is given by an interplay between the background geometry and the probing quanta. The relevant quantity to estimate τ_Z is the position of the maximum in the probability density $w_{\varepsilon\tau} = |\Psi_{\varepsilon\tau}[\phi^*, \phi]|^2$ which is closest to Σ_0 . Since the asymptotic solution is confined to a small region, we need to extend it by including less singular terms of order $\mathcal{O}(1/\varepsilon)$ such that $\Psi_{\varepsilon\tau}$ covers the evolution of test fields in a larger portion of \mathcal{B} and a possible maximum becomes visible. In this range, we find oscillatory solutions to (1) given by Bessel functions. This allows to estimate the position of τ_Z at which the maximum occurs which also determines the beginning of the Zeno region. To give a proof of concept, we perform a mode-sum decomposition in $l \in \mathbb{N}_0$, where we split off the time-dependent contribution of the kernel (note, for the full kernel the sum over all l has to be performed)

$$k_\ell(\varepsilon\tau) = \frac{-i}{\varepsilon^3 \tau^2} \partial_\tau \ln \left(c_1 J_0 \left(\sqrt{\ell^2 \varepsilon \tau} \right) + c_2 Y_0 \left(\sqrt{\ell^2 \varepsilon \tau} \right) \right). \quad (6)$$

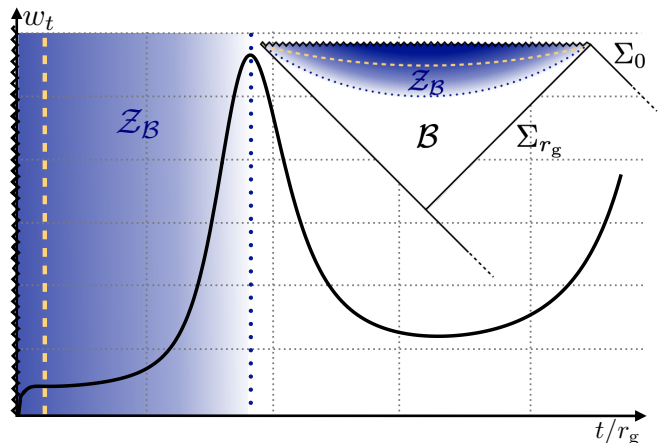


FIG. 2. Probability density $w_t^\ell = |\Psi_t^\ell[\phi^*, \phi]|^2$ for the distribution of a field configuration ϕ for a fixed ℓ in the folium Σ of \mathcal{B} . Wavy lines mark the geodesic border Σ_0 , dashed lines represent the onset of asymptotic silence, and the shaded region between Σ_0 and the dotted line is the Zeno region \mathcal{Z}_B . The upper right corner displays the corresponding Penrose diagram.

Here, $\ell^2 = 4l(l+1)$, $c_i \in \mathbb{C}$ constants of integration, and J_0 and Y_0 are the Bessel function of the first and second kind. The location of Σ_Z^ℓ can be expressed as a three-parameter family of hypersurfaces determined by ℓ^2 as well as the initial conditions represented by c_1 and c_2 of the quantum probes. By performing a series expansion for small arguments $\ell^2 \varepsilon \tau \ll 1$, we can approximate the position of the maximum in w_t^ℓ for ℓ^2 -values that allow for a suitable covering in the τ -direction within this approximation. Thus, we find for a given ℓ

$$\tau_Z^\ell \approx \frac{r_g}{|\mathcal{C}(c_1, c_2)|^2 \ell^2}, \quad (7)$$

where $\mathcal{C}(c_1, c_2)$ is a constant depending on the initial conditions. Apparently, increasing the angular momentum ℓ results in shifting the Zeno border closer to the geodesic border while for the case $\ell = 0$ we find the logarithmic (non-oscillating) solution in (5) where, the Zeno border lies at infinity, i.e. it covers all of the interior. Although the exact position of \mathcal{Z}_B depends on the specifics of the solution, we can compare τ_Z^ℓ to the high-energy scale provided by the Planck energy. To this aim, we consider Christensen's general form of a renormalised stress-energy tensor [14] with Synge's world function given in [15] and look at its trace $\langle T(x) \rangle_{\text{ren}} = \frac{1}{8\pi^2} \sqrt{-\det(g)} \left(\frac{1}{360} K + C \right) + \mathcal{O}(1/m^2) \approx r_g^2 / (224\pi^2 \tau^4)$ where $K = 12r_g^2/\tau^6$, the Kretschmann scalar and $C = \frac{1}{432} r_g^2/\tau^6$ other contractions of the Riemann tensor. The trace is usually the most divergent contribution since it includes the inverse metric. By comparing it with the Planck scale t_{P1} , we are able to determine τ_* where presumably quantum gravity effects become important.

Therefore, we demand $\langle T(\tau_*) \rangle_{\text{ren}} = t_{\text{Pl}}^{-2}$. Hence, we locate the high-energy scale at $\tau_* \approx \frac{1}{7} \sqrt{r_g t_{\text{Pl}}}$. For stellar-mass Schwarzschild black holes, the Zeno border will then be located $\tau_* < \tau_Z^\ell \leq r_g$ for all initial conditions coming from sub-Planckian energy densities.

Next we consider observables in \mathcal{Z}_B and introduce an auxiliary source functional $\mathcal{J} : \mathcal{C}(\Sigma_t) \rightarrow \mathbb{C}$, describing the absorption and emission of an instantaneous field configuration by the associated local current density $J : \Sigma_t \rightarrow \mathbb{C}$. Let $\Psi_t^J[\phi] := \langle \phi | \exp(\mathcal{J}_t)[\Phi] | \Psi_t \rangle$ which allows to replace compositions of the configuration operator Φ by the corresponding succession of functional derivatives with respect to the current. In the presence of the auxiliary source,

$$\langle \Psi_t^J | |\Phi[f]|^2 | \Psi_t^J \rangle = [f] \delta_J^2 \exp \left\{ \frac{1}{4} \frac{1}{\sqrt{\det(g_\Sigma)}} [J] [\text{Re}(\mathcal{K}_t)]^{-1} [J] \right\} \|\Psi_t^J[\phi]\|^2 \quad (8)$$

with $[f] \delta_J^2$ denoting the second functional derivative with respect to J , smeared with an appropriate field configuration f . In the absence of an auxiliary source, the ground-state expectation (8) is real and semi-positive definite. It is monotonously decreasing in \mathcal{Z}_B towards Σ_0 and vanishes on Σ_0 . Similarly,

$$\langle \Psi_t^J | |\Pi|^2 [f] | \Psi_t^J \rangle = [f] \mathcal{K}_t [f] \|\Psi_t^J[\phi]\|^2 - \det(g_\Sigma) k^2(t) \langle \Psi_t^J | |\Phi[f]|^2 | \Psi_t^J \rangle_{\Sigma_t}. \quad (9)$$

Both, the real and imaginary part of (9) are monotonously decreasing in \mathcal{Z}_B towards Σ_0 and vanish on Σ_0 when the auxiliary current is switched off.

Within the Zeno region, the ground-state expectation values (8) and (9) are controlled by the contractive representation of the evolution semigroup. As a result $\langle \Psi_t^J | \mathcal{H}[f] | \Psi_t^J \rangle$ does receive decreasing probabilistic support within \mathcal{Z}_B and has vanishing probabilistic support on Σ_0 . The Zeno region completes \mathcal{B} since information cannot leak through the geodesic information

sink Σ_0 . These results remain true if (self-) interactions are included [11] which can be understood at the qualitative level from the approximate operator identity $\mathcal{H} \cong \mathcal{G}(\Pi^*, \Pi)$.

V. DISCUSSION

The mathematical spacetime model of Schwarzschild black-hole interiors is geodesic incomplete. This result is often used to argue for the breakdown of general relativity as predicted by general relativity. In this letter we have shown that interiors of Schwarzschild black holes contain so-called Zeno regions, i.e. stacks of hypersurfaces in the vicinity of the geodesic spacetime singularity, on which observables enjoy a vanishing probabilistic support towards the geodesic spacetime border. In this sense Zeno regions represent a probabilistic completion of Schwarzschild black holes within the usual semiclassical framework. Moreover, geodesic incompleteness of Schwarzschild black holes qualifies the mathematical spacetime model, but certainly not the physical spacetime with a Zeno region in the immediate vicinity of the geodesic singularity. This statement is justified since Zeno regions allow for a consistent evolution of quantum fields but are beyond the scope of classical point particle mechanics. It is interesting to speculate whether Zeno regions are terminated by Hawking's hidden surface to which the principle of ignorance applies [16], but this we leave for further investigations.

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