

Coarse-graining of inhomogeneous dust flow in General Relativity via isometric embeddings

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Abstract.

We present a new approach to coarse-graining of variables describing dust flow in GR. It is based on assigning quasi-local shear, twist and expansion to 2-dimensional surfaces with the help of isometric embeddings into the 3-dimensional Euclidean space and deriving the time evolution equations for them. In contrast to the popular Buchert's scheme it allows to coarse-grain tensorial quantities in a coordinate-independent way. The framework can be used to estimate backreaction in inhomogeneous cosmological models.

Keywords: relativistic cosmology, cosmological inhomogeneities, backreaction problem

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1. INTRODUCTION

In this paper we discuss an alternative to the well-known Buchert's averaging scheme [1] for inhomogeneous cosmological models with dust. The new procedure allows for coarse-graining of both scalar (expansion) and tensorial (shear and vorticity) part of the fluid velocity gradient in a coordinate-independent way. Given a 3+1 decomposition of the spacetime we assign the coarse-grained expansion, shear and vorticity to three-dimensional comoving domains of the fluid. The resulting non-local quantities depend on the 3+1 splitting, but not on the coordinate system introduced on the spatial slices. If the fluid is irrotational, there exists a preferred, orthogonal splitting and the formalism is entirely covariant; otherwise one needs to fix the constant time slices by an additional condition.

The procedure is based on a reformulation of the definition of the volume average of the velocity gradient over a comoving domain in Newtonian theory. It turns out that it is possible to reexpress that volume average as a functional of the time derivative of the metric tensor induced on the domain's two-dimensional boundary. It has been noted before [2, 3] that such volume-averaged velocity gradients satisfy an evolution equation which is very similar to the evolution equation of the local velocity gradient, the only difference being a surface integral of the inhomogeneities of the flow over the domain's boundary. This volume average can be rewritten as a *surface* integral of the fluid velocity due to the divergence theorem. The key observation is that the surface values of the velocity field can be almost uniquely reconstructed from the comoving time derivative of the metric induced on the domain's boundary as long as the that boundary has the topology of S^2 and its metric has a positive curvature. The only part of the velocity field which cannot be reconstructed is the total linear and angular velocity, responsible for translations and rotations of the surface as a whole. However, since they do not contribute to the average shear and expansion, these quantities can be evaluated just from the reconstructed velocity field. Thus it is possible to reexpress in Newtonian theory the average shear and expansion as functionals of the derivative of the induced metric on the boundary.

The reformulated definition carries over easily to the relativistic case. If we fix a comoving, three-dimensional domain in a 3+1 decomposed spacetime, we can find the isometric embeddings of the domain's two-dimensional boundaries into the three-dimensional Euclidean space and evaluate the coarse-grained expansion and shear. If the background metric is not flat, the resulting quantities are not equivalent to any volume integrals any more, but nevertheless they can serve as a measure the deformation of a finite fluid element. As we show in the paper, the new definition satisfies simple reasonability conditions: it yields the right answer in the limit of the domain shrinking to a point and in FLRW metrics.

We also present a way to define the coarse-grained vorticity from the projection of particle four-velocities to the spatial constant-time slice if the dust is rotating. The definition is independent of the coordinates introduced on the spatial slice, but in contrast to the irrotational case we do not have a simple, unique definition of a preferred 3+1

splitting and the coarse-grained quantities do depend on it.

Finally we present the evolution equation for the coarse-grained velocity gradient and show that its structure is very similar to the evolution equation for its local counterpart. Two additional terms in the form of surface integrals may serve as a measure of backreaction. i.e. the impact of the inhomogeneities of the flow on the large scale motion of the dust within the domain.

The proposed procedure is more mathematically involved than Buchert's, as it requires finding the isometric embeddings of S^2 surfaces, which is equivalent to solving a system of coupled, non-linear PDE's. In most cases this is impossible to do analytically, but numerical methods have already been developed for that problem [4, 5]. Let us also note that in contrast to Buchert's approach we do not have evolution equations for the coarse-grained curvature and matter density. The dust was chosen as the matter model in this paper just for its simplicity; the formalism introduced here can be applied to any other fluid in General Relativity with only minor modifications. For a more exhaustive treatment of the topic, including rigorous derivations of the results, see [6].

2. COARSE-GRAINING VIA ISOMETRIC EMBEDDINGS

Consider gravitating, Newtonian, pressureless dust in a flat space. It is described by the velocity field v^i , density ρ and the Newtonian potential ϕ . The gradient of the velocity field $Q^i_j = v^i_{,j}$ describes how an infinitesimal fluid or dust element is deformed and how it rotates during its motion. It can be decomposed into the scalar expansion, symmetric traceless shear and antisymmetric vorticity

$$Q_{ij} = \frac{1}{3} \theta \delta_{ij} + \sigma_{(ij)} + \omega_{[ij]}.$$

Q_{ij} satisfies the evolution equation along the flow

$$\frac{D}{dt} Q^i_j = -Q^i_k Q^k_j - \phi^i_{,j}, \quad (1)$$

where $\frac{D}{dt}$ is the convective derivative (see [6, 2, 3]). This equation has a coarse-grained counterpart which one obtains by taking the volume average of Q^i_j over a comoving domain G_t

$$\langle Q^i_j \rangle = \frac{1}{V} \int_{G_t} Q^i_j d^3x = \frac{1}{V} \int_{G_t} v^i_{,j} d^3x. \quad (2)$$

Note that $\langle Q^i_j \rangle$ is effectively a surface integral over the boundary of ∂G_t

$$\langle Q^i_j \rangle = \frac{1}{V} \int_{\partial G_t} v^i n_j d\sigma. \quad (3)$$

Its evolution equation has a similar structure to (1),

$$\frac{\partial}{\partial t} \langle Q^i_j \rangle = -\langle Q^i_k \rangle \langle Q^k_j \rangle - \langle \phi^i_{,j} \rangle + B^i_j \quad (4)$$

with $\langle \phi^i_{,j} \rangle$ being the volume average of the Hessian of the potential over G_t . The additional term

$$B^i_j = \frac{1}{V} \int_{\partial G_t} (\delta v^k \delta v^i_{,j} n_k - \delta v^k \delta v^i_{,k} n_j) d\sigma, \quad (5)$$

measures how the inhomogeneities of the fluid velocity field $\delta v^i = v^i - \langle Q^i_j \rangle x^j$ affect its large scale motion in the domain G_t [2, 3, 6].

In the relativistic case, on a non-flat background, we do not have a well-defined, coordinate-independent notion of volume averages of tensor objects, so we cannot use (2) to obtain a relativistic counterpart of (4). It turns out, however, that definition (2) can be reformulated in a geometric, coordinate independent manner which has a natural extension to non-flat backgrounds. We will first deal with the symmetric part of $\langle Q_{ij} \rangle$, i.e. shear and expansion. We will show

that under certain technical assumptions they can be reexpressed as functionals of the metric induced on the boundary ∂G_t and its convective (comoving with the dust) time derivative.

The main tool we are going to employ is the isometric embedding theorem for surfaces of S^2 topology, conjectured by Weyl and proved by Lewy, Alexandrov, Pogorelov, Nirenberg and Cohn–Vossen (see [7], [8]):

Theorem 2.1 (Isometric embedding theorem for S^2) *Given a compact, orientable surface S homeomorphic to S^2 , with positive metric q whose scalar curvature $R > 0$. Then*

- *there exists an isometric embedding*

$$f : S \mapsto \mathbf{E}^3$$

into the 3–dimensional Euclidean space;

- *the embedding is unique up to rigid rotations, translations and reflexions.*

If we fix the orientation, the theorem states that for a given S^2 surface with a positive metric of positive curvature can be recognized as a submanifold in \mathbf{E}^3 and this can be done only in one way, up to moving the surface around and rotating it as a whole. Surfaces satisfying the hypothesis of theorem 2.1 will be called admissible.

Consider now the boundary ∂G_t of the coarse–graining domain, parametrized by two coordinates θ^A labeling individual particles. It is described, together with its time evolution, by equations

$$x^i = \zeta^i(t, \theta^A).$$

The induced metric has the form of

$$q_{AB}(t, \theta^A) = \zeta^i_{,A} \zeta^j_{,B} \delta_{ij} \quad (6)$$

and its time derivative in coordinates θ^A

$$\dot{q}_{AB} = 2v^i_{(A} \zeta^j_{,B)} \delta_{ij}.$$

Assume now that ∂G_t is admissible at least for some time and that we know its position in space at one instance of one time $t = t_0$. We will show that in that case $\langle Q_{(ij)} \rangle$ can be read out solely from the induced metric (6), even if we do not have each individual particle’s velocity and position.

Given $q_{AB}(t)$ we can find the isometric embeddings of ∂G_t , consistent with its orientation, into \mathbf{E}^3 at each time t . We assume that at t_0 the embedding is known, but at other times we need to solve the following system of PDE’s

$$q_{AB} = \chi^i_{,A} \chi^j_{,B} \delta_{ij} \quad (7)$$

for three functions $\chi^i(t, \theta^A)$. We construct this way a one–parameter family of embeddings

$$x^i = \chi^i(t, \theta^A)$$

for which we assume the additional condition

$$\chi^i(t_0, \theta^A) = \zeta^i(t_0, \theta^A).$$

The first statement of theorem 2.1 assures that functions $\chi^i(t, \theta^A)$ can always be found, although they are non–unique. The true particle positions of ∂G_t , given by $\zeta^i(t, \theta^A)$, constitute trivially a family of isometric embeddings, so from the the uniqueness part of 2.1 we know that functions χ^i and ζ^i are related via

$$\chi^i(t, \theta^A) = R^i_{,j}(t) \zeta^j(t, \theta^A) + W^i(t),$$

where $R^i_{,j}$ is a rotation. If we take a time derivative of this equation at $t = t_0$, we obtain

$$\left. \frac{\partial \chi^i}{\partial t} \right|_{t=t_0} = u^i(\theta^A) = v^i(t_0, \theta^A) + \Omega^i_{,j} \zeta^j(t_0, \theta^A) + E^i.$$

with constants $\Omega_{ij} = -\Omega_{ji}$ and E^i . Thus the velocity field u^i , reconstructed from an arbitrary family isometric embeddings, differs from the true one only by the restrictions to ∂G_t of a rotational ($\Omega^i_j x^j$) and a translational (E^i) vector field. This difference does not influence the value of integral (3) symmetrized in the two indices

$$\int_{\partial G_t} u_{(i} n_{j)} d\sigma = \int_{\partial G_t} v_{(i} n_{j)} d\sigma = \langle Q_{(ij)} \rangle, \quad (8)$$

so we can use the reconstructed u instead of v to evaluate the shear and expansion.

In fact we do not have to find the embeddings explicitly to obtain the reconstructed velocity field. By differentiating (7) with respect to t we obtain an equation relating u^i directly to the convective time derivative of q_{AB} at $t = t_0$

$$2u^i_{(A} \zeta^j_{B)} \delta_{ij} = \dot{q}_{AB}.$$

The linear operator $\mathcal{P}_\zeta[Y]_{AB} = 2Y^i_{(A} \zeta^j_{B)} \delta_{ij}$ appearing on the left hand side yields the variation of the surfaces' induced metric if the embedding itself is being dragged by a vector field Y^i . By the virtue of theorem 2.1 for an admissible surface it has a six-dimensional kernel consisting of rigid motions and a "non-unique inverse" \mathcal{P}_ζ^{-1} . We can now rewrite the definition of the symmetric part coarse-grained velocity gradient explicitly as a functional of \dot{q}_{AB}

$$\langle Q_{(ij)} \rangle = \frac{1}{V} \int_{\partial G_t} \mathcal{P}_\zeta^{-1}[\dot{q}_{AB}]_{(i} n_{j)} d\sigma. \quad (9)$$

3. VELOCITY GRADIENT IN GENERAL RELATIVITY

Formula (9) is clearly more complicated than (2), but its main advantage lies in the fact that it can be generalized to fluids on a non-flat background in a simple and natural way. Consider a (possibly curved) spacetime, filled with dust described by density ρ and four-velocity field u . The integral curves of u are geodesics, i.e. $\nabla_u u^\mu = 0$. The four-velocity gradient $Z_{\mu\nu} = \nabla_\nu u_\mu$ describes the deformations of an infinitesimal fluid element in its motion. It is orthogonal to u in both indices and therefore effectively a three-dimensional object $Z_{\mu\nu}$ satisfies the relativistic counterpart of (1)

$$\nabla_u Z^\mu{}_\nu = -Z^\mu{}_\rho Z^\rho{}_\nu - R^\mu{}_{\alpha\nu\beta} u^\alpha u^\beta. \quad (10)$$

Unless stated otherwise, in this paper we will assume the flow to be irrotational, i.e. $Z_{\mu\nu} = Z_{(\mu\nu)}$. We will now show how to derive the coarse-grained counterpart of (10) in which $\langle Z_{ij} \rangle$ corresponds to a finite fluid volume rather than an infinitesimal element.

In comoving and orthogonal coordinates (t, y^i) one can perform the standard ADM decomposition of the metric into the spatial part h_{ij} , with vanishing N^i and $N = 1$ due to the coordinate system choice. Consider a three-dimensional tube of particle worldlines ∂C enclosing a cylinder C (see fig. 1). The tube is described by equations $y^i = \xi^i(\theta^A)$, ξ^i being three time-independent functions and θ^A comoving coordinates on the tube labeling individual particles. The intersections of the cylinder with constant time slices Σ_t , denoted in this paper by ∂C_t , constitute two-dimensional surfaces of spherical topology. Assuming that the surfaces are admissible we assign the coarse-grained velocity gradient to their interior in the following way: we take the induced metric

$$q_{AB}(t, \theta^A) = \xi^i_{,A} \xi^j_{,B} h_{ij}(t) \quad (11)$$

on ∂C_t and find an arbitrary family of isometric embeddings $f_t : \partial C_t \rightarrow \partial D_t \subset \mathbf{E}^3$. They are described by equations

$$x^a = \chi^a(t, \theta^A)$$

with three functions χ^{a1} satisfying a system of non-linear PDE's

$$q_{AB}(t, \theta^A) = \chi^a_{,A} \chi^b_{,B} \delta_{ab}.$$

¹ Indices a, b, c, \dots run from 1 to 3, just like i, j, k, \dots . We will reserve the former for geometric objects in \mathbf{E}^3 and the latter for objects on a constant time slice of the spacetime.

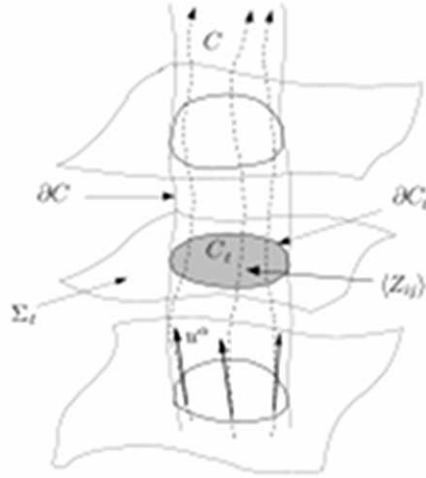


FIGURE 1. Coarse-graining of $Z_{\mu\nu}$ over the slices of a tube of particles

These are subject to gauge transformations of the form of rigid motions

$$\chi^a(t, \theta^A) \rightarrow R^a_b(t) \chi^b(t, \theta^A) + W^a(t). \quad (12)$$

We can now define a fictitious rather than reconstructed velocity field v on the surface image in \mathbf{E}^3 by taking the time derivative of each particle's positions given by χ^a . The vector field can also be recovered directly from $\dot{q}_{AB} = \xi^i_{(A} \xi^j_{B)} \dot{h}_{ij}(t)$ by the inverse of \mathcal{P}_χ . We can then use v to evaluate the coarse-grained $Z_{(ij)}$ as tensors in \mathbf{E}^3

$$\langle \tilde{Z}_{(ab)} \rangle = \frac{1}{V_0} \int_{\partial D_t} v_{(a} n_{b)} d\sigma \quad (13)$$

(V_0 denotes here the volume enclosed by ∂D_t in \mathbf{E}^3). Because of (8) these integrals are independent of the vector field v chosen, but they do depend on the orientation of the embedded surface and under gauge transformations (12) they transform according to

$$\langle \tilde{Z}_{ab} \rangle \rightarrow R^c_a R^d_b \langle \tilde{Z}_{cd} \rangle.$$

In order to get rid of that gauge dependence we need to map $\langle \tilde{Z}_{ab} \rangle$ back to the spacetime. Each isometric embedding f_t defines canonical isometries g_p between the tangent spaces $T_p \Sigma_t$ to the constant time slices at points $p \in \partial C_t$ and \mathbf{E}^3 . Any vector $X \in T_p \Sigma_t$ can be decomposed into the tangent part to ∂C_t and the projection to the outward pointing normal n^i

$$X^i = X_T^i + X_n n^i.$$

We then take

$$g_p(X)^a = (f_t^* X_T)^a + X_n m^a,$$

where m^a is the outward pointing normal to the image of the surface at $f(p)$ (see fig. 2). g_p preserves the metric tensor and can be used to map geometric objects from $T_p \Sigma_t$ to \mathbf{E}^3 or back. In particular, one can pull $\langle \tilde{Z}_{ab} \rangle$ back to each point p in ∂C_t by g_p^*

$$\langle Z_{ij} \rangle_p = g_p^* (\langle \tilde{Z} \rangle)_{ij}.$$

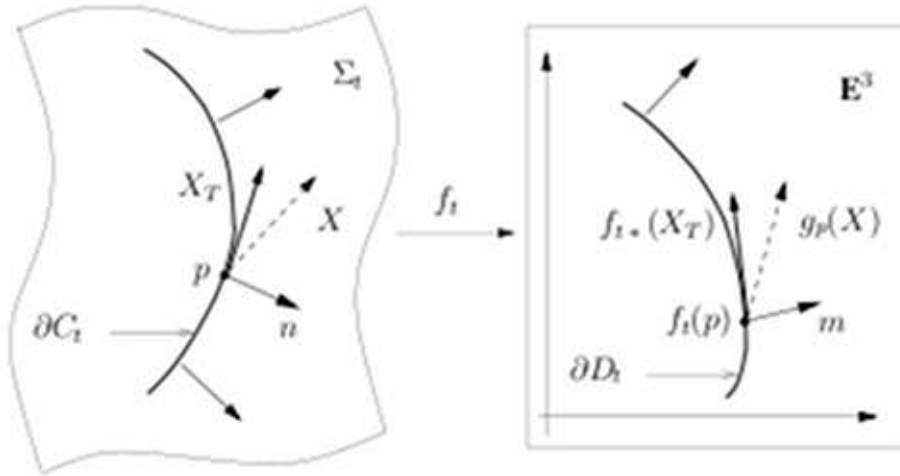


FIGURE 2. Canonical isometry $g_p : T_p \Sigma_t \rightarrow \mathbb{E}^3$

It is straightforward to verify that the tensor field $\langle Z_{ij} \rangle$ constructed on ∂C_t this way is insensitive to transformations (12) and thus completely independent of the choice of isometric embeddings f_t .

If the antisymmetric part of $Z_{\mu\nu}$, i.e. the vorticity, does not vanish, it can be coarse-grained as well, although in a different way. Note that the 3+1 splitting cannot be orthogonal in that case and the four-velocity field u^μ has a non-vanishing projection to the constant time slice, denoted by w^i . We define now the antisymmetric part of $\langle \tilde{Z}_{ab} \rangle$ as

$$\langle \tilde{Z}_{[ab]} \rangle = \frac{1}{V_0} \int_{\partial D_t} g(w)_{[a} n_{b]} d\sigma. \quad (14)$$

Note however that, in contrast to the rotating case, in a spacetime filled with rotational dust it is difficult to single out a preferred 3+1 splitting. Unfortunately, just like in the Buchert's formalism, the values of $\langle Z_{ij} \rangle$ depend on the splitting one chooses in general. Throughout the rest of the paper we will only deal with the non-rotating dust and in that case $\langle Z_{[ij]} \rangle = 0$.

3.1. The coarse-graining procedure in the narrow tube limit and in FLRW metrics

Outside the Newtonian setting we lose the equivalence of $\langle Z_{ij} \rangle$ to the volume average of the local velocity gradient. Nevertheless the newly defined quantities turn out to have the correct limiting behavior, namely if we shrink the coarse-graining domain to a single point, we recover the local Z_{ij} at that point. More precisely, consider a one-parameter family of tubes described by equation

$$y^i = \lambda \xi^i(\theta^A). \quad (15)$$

with a positive parameter λ . As $\lambda \rightarrow 0$, the tube shrinks down to the single worldline $y^i = 0$. One can prove that for small λ the values of $\langle Z_{ij} \rangle$, including the antisymmetric part if the vorticity doesn't vanish, can be expanded at every point p of the tube in powers of λ as

$$\langle Z_{ij} \rangle = Z_{ij}|_{y^i=0} + O(\lambda). \quad (16)$$

Intuitively one can understand this fact in the following way: expressions (13) and (2) are perfectly equivalent if the background metric is flat and if the velocity gradient is constant over the coarse-graining domain. In general none of these conditions is satisfied on Σ_t , but the more the domain under consideration shrinks, the less it "feels" the background curvature and the inhomogeneities in Z_{ij} . As it shrinks down to a point, both curvature and inhomogeneities

drop out and $\langle Z_{ij} \rangle$ is equal to the volume average of the local Z_{ij} over a small domain around the point $y^i = 0$. It is now straightforward to prove that such volume average approaches then the value of Z_{ij} at $y^i = 0$. For a rigorous argument see [6].

In order to test the coarse-graining procedure we may also apply it to known homogeneous cosmological solutions. For a dust-filled Friedman–Lemaître–Robertson–Walker metric in the natural 3+1 splitting the it yields the expected answer for any admissible surface:

$$\langle \tilde{Z}_{ab} \rangle = H(t) \delta_{ab}, \quad (17)$$

i.e. the coarse-grained velocity gradient consist only of the scalar part given by the Hubble parameter. This result is easy to see if we notice that the metric q_{AB} induced on the surface, like any other object in a FLRW solution comoving with the dust, undergoes a homogeneous rescaling in time described by $\dot{q}_{AB} = 2H(t) q_{AB}$. The same metric variation is induced on the image of the embedded surface by the homothetic vector field $v^a = H(t) x^a$ in \mathbf{E}^3 . It is now straightforward to verify that the restriction of this vector field to ∂D_t gives (17) when plugged into (13). This result holds for open, closed and flat FLRW metrics alike.

4. TIME EVOLUTION EQUATION FOR THE COARSE-GRAINED QUANTITIES

In order to derive the time evolution equations for the coarse-grained velocity gradient, one must first introduce a preferred time derivative of a tensor field in \mathbf{E}^3 , which will serve as a generalization of the covariant derivative along the geodesic in (10). We define it via

$$\begin{aligned} \mathcal{D}T^{ab\dots}_{cd\dots} &= \frac{\partial}{\partial t} T^{ab\dots}_{cd\dots} + W^a_z T^{zb\dots}_{cd\dots} + W^b_z T^{az\dots}_{cd\dots} + \dots + \\ &\quad - W^z_c T^{ab\dots}_{zd\dots} - W^z_d T^{ab\dots}_{cz\dots} + \dots, \end{aligned}$$

where

$$W_{ab} = -\frac{1}{V_0} \int_{\partial D_t} v_{[a} n_{b]} d\sigma.$$

\mathcal{D} preserves the Euclidean metric in \mathbf{E}^3 and has the right narrow tube limit, i.e. it reduces to ∇_u along the limiting geodesic [6].

We now fix a tube of particle worldlines C_t , given by functions $\xi^i(\theta^A)$, whose sections ∂C_t are admissible surfaces. We find the isometric embeddings and calculate $\langle \tilde{Z}_{ab} \rangle$. This allows for decomposing the *local* Z_{ij} , pushed forward to $\partial D_t \subset \mathbf{E}^3$ via g_p , and the velocity field v^a into the the coarse-grained, large scale part and local inhomogeneities

$$g_{p*}(Z)_{ab} = \langle \tilde{Z}_{ab} \rangle + \delta Z_{ab} \quad (18)$$

$$v^a = \langle \tilde{Z}^a_b \rangle x^b + \delta v^a, \quad (19)$$

both defined on ∂D_t only.

For the sake of brevity we introduce the following notation for surface integrals of the type of (3) and (13):

$$\mathcal{N}[X_a]_b = \frac{1}{V_0} \int_{\partial D_t} X_a n_b d\sigma.$$

We will also denote by $\mathcal{R}[\cdot]$ the unique inverse of \mathcal{P}_χ satisfying

$$\mathcal{N}\mathcal{R}[r_{AB}]_{[cd]} = 0$$

for any r_{AB} , plus an irrelevant condition fixing the constant part E^a , for example fixing the surface's centroid at the origin. Now the definition of the coarse-grained Z_{ij} can be written down in a slightly more compact manner

$$\langle \tilde{Z}_{(ab)} \rangle = \mathcal{N}\mathcal{R}[\dot{q}_{AB}]_{ab} = \mathcal{N}\mathcal{R}\left[2Z_{(ij)} \xi^i_{,A} \xi^j_{,B}\right]_{ab} \quad (20)$$

in which we have used the time derivative of (11) to substitute $2Z_{(ij)} \xi^i_{,A} \xi^j_{,B}$ for \dot{q}_{AB} . Note that the combination $\mathcal{NR} \left[2\xi^i_{,A} \xi^j_{,B} \right]$ plays in our formalism the role of the coarse-graining operator for symmetric tensors of rank 2 on Σ_t .

The evolution equation for $\langle \tilde{Z}_{ab} \rangle$ takes now the form of

$$\mathcal{D} \langle \tilde{Z}_{ab} \rangle = -\langle \tilde{Z}_{ac} \rangle \langle \tilde{Z}^c_b \rangle - \langle R_{a0b0} \rangle + B_{ab} + \tilde{B}_{ab}, \quad (21)$$

where

$$\langle R_{a0b0} \rangle = \mathcal{NR} \left[2R_{i0j0} \xi^i_{,A} \xi^j_{,B} \right]_{(ab)}$$

is obviously the coarse-grained contraction of the Riemann tensor with $u^\mu u^\nu$, considered as a symmetric tensor on Σ_t [6]. We can see that two new terms have appeared in comparison with (10). The first one is the symmetrized version of the familiar Newtonian backreaction term (5):

$$B_{ab} = \mathcal{N} \left[\delta v^c_{,c} \delta v_{(a} \right]_{b)} - \mathcal{N} \left[\delta v^c_{,(a} \delta v_{b)} \right]_c. \quad (22)$$

The second one is entirely relativistic and has a more complicated structure

$$\begin{aligned} \tilde{B}_{ab} = & \mathcal{NR} \left[4 \langle \tilde{Z}_{cd} \rangle \left(\delta Z^d_e \chi^e_{(,A} - \delta v^d_{(,A} \right) \chi^c_{,B)} \right]_{ab} + \\ & + \mathcal{NR} \left[2 \left(\delta Z_{ce} \delta Z^e_d \chi^c_{,A} \chi^d_{,B} - \delta v^c_{,A} \delta v^d_{,B} \delta_{cd} \right) \right]_{ab}. \end{aligned}$$

Just like in the Newtonian case, both backreaction terms are surface integrals divided by volume. In contrast to (22) however, \tilde{B}_{ab} involves linear terms in perturbations. Both B_{ab} and \tilde{B}_{ab} together measure the influence of inhomogeneities on the time evolution of the coarse-grained expansion and shear.

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